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Spatially resolved Kennicutt–Schmidt relation at $z \approx 7$ and its connection with the interstellar medium properties

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ABSTRACT

We exploit moderately resolved [O III], [C II] and dust continuum ALMA observations to derive the gas density (n), the gas-phase metallicity (Z), and the deviation from the Kennicutt–Schmidt (KS) relation (κ_s) on \approx sub – kpc scales in the interstellar medium (ISM) of five bright Lyman Break Galaxies at the Epoch of Reionization ($z \approx 7$). To do so, we use GLAM, a state-of-art, physically motivated Bayesian model that links the [C II] and [O III] surface brightness ($\Sigma_{\text{[CII]}}$, $\Sigma_{\text{[OIII]}}$) and the SFR surface density (Σ_{SFR}) to n , κ_s , and Z . All five sources are characterized by a central starbursting region, where the Σ_{gas} versus Σ_{SFR} align $\approx 10 \times$ above the KS relation ($\kappa_s \approx 10$). This translates into gas depletion times in the range $t_{\text{dep}} \approx 80 - 250$ Myr. The inner starbursting centres are characterized by higher gas density ($\log(n/\text{cm}^{-3}) \approx 2.5-3.0$) and higher metallicity ($\log(Z/Z_{\odot}) \approx -0.5$) than the galaxy outskirts. We derive marginally negative radial metallicity gradients ($\nabla \log Z \approx -0.03 \pm 0.07$ dex/kpc), and a dust temperature ($T_d \approx 32 - 38$ K) that anticorrelates with the gas depletion time.

Key words: galaxies: evolution – galaxies: high-redshift – galaxies: ISM – dark ages, reionization, first stars.

1 INTRODUCTION

The Epoch of Reionization (EoR) represents a critical phase of the Universe evolution, and its study is one of the frontiers in modern astrophysics (e.g. Robertson 2022). During the EoR, the first galaxies started to rapidly form stars, which in turn began producing photons able to ionize the surrounding gas – first the interstellar medium (ISM), eventually the intergalactic medium (Dayal & Ferrara 2018, for a review). For this reason, shedding light on how the gas is converted into stars (e.g. Tacconi, Genzel & Sternberg 2020), and how this process is influenced by the ISM properties holds the key to understanding the evolution of cosmic reionization.

At low and intermediate redshifts, the so-called Kennicutt–Schmidt (KS) relation¹ $\Sigma_{\text{SFR}} \approx 10^{-12} \kappa_s \Sigma_{\text{gas}}^{1.4}$, linking the star formation rate (SFR) and the gas surface densities (Σ_{SFR} , Σ_{gas} , respectively) is well established (Schmidt 1959; Kennicutt 1998; Heiderman et al. 2010; de los Reyes & Kennicutt 2019). The ‘burstiness’ parameter, κ_s , was first introduced in Ferrara et al. (2019) to quantify the deviation from the KS relation that might occur in the high- z Universe. Galaxies with $\kappa_s > 1$ show a larger SFR per unit area with respect to those located on the KS relation, i.e. they

tend to be starburst. At high redshifts, values in the range $\kappa_s = 10-100$ have been measured for massive/rare sub-millimetre galaxies for which spatially resolved data of cold gas tracers, namely low- J CO lines, are available (e.g. Hodge et al. 2015; Chen et al. 2017). Spatially resolved low- J CO detections in galaxies representative of the bulk population in the EoR are instead time demanding even with state-of-art radio/sub-mm facilities, unless taking advantage of strong gravitational lensing (Nagy et al. 2023, at $z \approx 1$). This is due to the efficient CO photodissociation at low metallicity and dust abundance (Bolatto, Wolfire & Leroy 2013; Wolfire, Vallini & Chevance 2022), and because of the increasing temperature of the CMB (da Cunha et al. 2013; Vallini et al. 2015) against which the lines are observed. Only a few mid- J CO detections have been reported so far (e.g. Pavesi et al. 2019; Ono et al. 2022) but none of them are spatially resolved even by the Atacama Large Millimetre/submillimetre Array (ALMA), thus hampering the measure of the size of the emitting area and ultimately the derivation of Σ_{gas} .

In recent years an alternative, indirect, method has been proposed to infer the location of EoR sources with respect to the KS relation. This is done by linking their κ_s to the relative surface brightness ratios of bright neutral (e.g. [C II] 158 μm) versus ionized (e.g. [O III] 88 μm , CIII] λ 1907, 1909 doublet) gas tracers (Ferrara et al. 2019; Vallini et al. 2020, 2021; Markov et al. 2022) that can be spatially resolved by ALMA (Akins et al. 2022; Herrera-Camus et al. 2022; Molyneux et al. 2022; Witstok et al. 2022; Posses et al. 2023) and JWST (e.g. Hsiao et al. 2023). A starburst source has a larger ionization

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¹The SFR (gas) surface density is expressed in units of $M_{\odot} \text{yr}^{-1} \text{kpc}^{-2}$ ($M_{\odot} \text{kpc}^{-2}$).

parameter, U , producing a correspondingly larger ionized gas column density, as compared to a galaxy with the same Σ_{gas} but lying on the KS relation (Ferrara et al. 2019). These conditions boost (quench) ionized (neutral) gas tracers and, together with the gas density (n) and metallicity (Z), concur in determining the surface brightness ratios (Kohandel et al. 2023).

By leveraging this method, Vallini et al. (2021) analysed the nine EoR Lyman Break Galaxies (LBGs) that had joint (albeit only barely resolved) [C II]-[O III] detections at the time (Laporte et al. 2017; Tamura et al. 2019; Bakx et al. 2020; Carniani et al. 2020; Harikane et al. 2020), obtaining $\kappa_s = 10\text{--}100$. These high burstiness parameters, in agreement with expectations from cosmological zoom-in simulations (Pallottini et al. 2019, 2022), suggest ISM conditions favouring an efficient conversion of gas into stars (short depletion times), with starburst episodes producing bright [O III] emission from H II regions (e.g. Cormier et al. 2019; Harikane et al. 2020). Also, the gas metallicity and density were found to be relatively high ($Z = 0.2\text{--}0.5 Z_{\odot}$, and $n = 10^{2-3} \text{ cm}^{-3}$, respectively) in agreement with independent analysis carried out on the same objects (Jones et al. 2020; Yang & Lidz 2020).

A recent study of three $z \approx 7$ galaxies from REBELS (Bouwens et al. 2022) supports the tight relation between galaxy burstiness and [O III]/[C II] ratios. In this case, low [O III]/[C II] have been explained with the weak ionizing field resulting from the non-starbursting nature of the sources (Algera et al. 2023). The lack of recent bursts is also likely the cause (e.g. Sommovigo et al. 2020) of their cold dust temperatures. The sources analysed by Algera et al. (2023) seem, however, to be an outlier with respect to the average conditions of EoR galaxies with below-average [O III] $\lambda\lambda 4959, 5007 + \text{H}\beta$ equivalent widths compared to the known high- z population.

From the theoretical side, an increasing number of simulations and models developed to interpret [C II] (e.g. Vallini et al. 2015; Lagache, Cousin & Chatzikos 2018; Pallottini et al. 2022), [O III] (e.g. Moriwaki et al. 2018; Arata et al. 2020; Katz et al. 2022) and dust continuum emission (e.g. Behrens et al. 2018; Di Cesare et al. 2023), find high turbulence (e.g. Kohandel et al. 2020), strong radiation fields (e.g. Katz et al. 2022), high densities, and warm dust temperatures (e.g. Sommovigo et al. 2021), to be common on sub-kpc scales in the ISM of star-forming galaxies in the EoR.

The goal of this work is to push further the study of the link between the KS relation, and ISM/dust properties in the first galaxies by leveraging the spatially resolved [C II], [O III], and dust continuum data recently presented by Witstok et al. (2022) in a sample of five bright LBGs at $z \approx 7$. Our aim is to investigate the sub-kpc relation between the burstiness parameter, gas density, and metallicity and study their connection with global values that can be inferred from unresolved data.

The paper is structured as follows: in Section 2, we summarize the sample and data used in this analysis; in Section 3, we illustrate the model. The results are outlined in Section 4 while we discuss the implications and present our conclusions in the final Section 5.

2 SAMPLE AND DATA

Details regarding the sample and data reduction can be found in Witstok et al. (2022); however, we summarize the key points here. We considered all available ALMA data sets of [C II] $\lambda 158\mu\text{m}$ (2015.1.01111.S, 2017.1.00604.S, 2019.1.01611.S, PI: Smit, 2018.1.00085.S, PI: Schouws, 2018.1.01359.S, PI: Aravena, 2015.1.00540.S, 2018.1.00933.S, PI: Bowler) and [O III] $\lambda 88\mu\text{m}$ (2018.1.00429.S, 2019.1.01524.S, PI: Smit) for the sample of

LBGs at $z \sim 7$: COS-3018555981 (COS-3018, hereafter), COS-2987030247 (COS-2987, hereafter), UVISTA-Z-001, UVISTA-Z-007, and UVISTA-Z-019. Data were calibrated and reduced with the automated pipeline of the Common Astronomy Software Application (CASA; McMullin et al. 2007). In cases where the continuum is robustly detected (i.e. next to the [C II] emission in COS-3018, UVISTA-Z-001, and UVISTA-Z-019 and next to the [O III] emission in UVISTA-Z-001), we first performed continuum subtraction using the UVCNTSUB task in CASA. After this step, we created images with the TCLEAN task both under natural and several Briggs weightings. We tuned the weighting and/or taper scheme to match the beam sizes as closely as possible, using natural weighting (and a small taper, if required) for the line observed with highest spatial resolution and Briggs weighting for the other. The robust parameter has been tuned to the highest resolution achievable while maintaining a reasonable signal-to-noise ratio (SNR). The resulting matched beam sizes ($\theta \approx 0.4\text{--}0.5$ arcsec) for [C II] and [O III] are listed in table 2 of Witstok et al. (2022). Finally, we regridded images of the [O III] and [C II], obtained by integrating along the frequency axis over the full-width at half maximum of the line, to a common coordinate mesh with the REPROJECT package in ASTROPY.

Both UV and IR luminosity maps were regridded to the same pixel grid as the [O III] and [C II] lines. We used the same imaging parameters for the dust continuum at rest-frame wavelength $\lambda_{\text{emit}} \sim 160\mu\text{m}$ (since band-6 measurements had the most significant detections) to achieve a beam nearly identical to that of the [C II] line. The IR luminosity ($L_{\text{IR}}, 8\text{--}1000\mu\text{m}$) was calculated using the global best-fitting spectral energy distribution (SED; see section 4 of Witstok et al. 2022 and table 3 for the best-fitting dust temperatures, $T_d \approx 29\text{--}60$ K) as a template for rescaling the $\sim 160\mu\text{m}$ flux in each pixel. The UV continuum was convolved with an effective Richardson (1972) and Lucy (1974) beam to match the point-spread function of the dust-continuum emission.

3 MODEL

The derivation and study of the ISM properties presented in this paper is based on GLAM² (Vallini et al. 2020, 2021, hereafter V20, V21). GLAM is a tool to perform Bayesian inference that is based on a physically motivated model for the analytical treatment of the radiative transfer of ionizing ($h\nu > 13.6$ eV, Extreme Ultraviolet, EUV) and non-ionizing (6 eV $< h\nu < 13.6$ eV, Far Ultraviolet, FUV) photons in the ISM of galaxies (Ferrara et al. 2019, hereafter F19).

F19 enables the computation of the surface brightness of lines excited either in the ionized and/or in the photodissociation region (PDR; Hollenbach & Tielens 1999; Wolfire, Vallini & Chevance 2022) of a gas slab illuminated by ultraviolet (UV) radiation from newly formed stars. The surface brightness of the lines is determined by the average gas density (n) of the H II/PDR environment – characterized by electron density n_e and neutral gas density n_{H} , respectively³ – the dust-to-gas ratio, ($\mathcal{D} \propto Z$, where Z is the gas metallicity), and ionization parameter, U . The latter, can be expressed in terms of observed quantities by deriving its relation ($U \propto \Sigma_{\text{SFR}}/\Sigma_{\text{gas}}$, see eqs. 38 and 40 in F19) with the SFR surface density (Σ_{SFR}) and the gas surface density (Σ_{gas}), which in turn are

²GLAM: Galaxy Line Analyzer with MCMC is publicly available at <https://vallini.github.io/MCMC-galaxyline-analyzer/>

³Both n_e and n_{H} can be expressed as a function of n . In the ionized layer, $n_e = x_e n \approx n$ assuming an ionized fraction $x_e \approx 1$, while in the PDR the neutral gas density $n_{\text{H}} = (1 - x_e)n \approx n$ given that $x_e \approx 0$.

connected through the star formation law. This leaves us with the κ_s parameter, describing the burstiness of the galaxy. GLAM adopts a Markov Chain Monte Carlo (MCMC) algorithm (implemented with `emcee`, Foreman-Mackey et al. 2013) to search for the posterior probability of the model parameters (n , κ_s , Z) that reproduce the observed [C II] surface brightness ($\Sigma_{\text{[C II]}}$), [O III] surface brightness ($\Sigma_{\text{[O III]}}$), and the SFR surface density (Σ_{SFR}). GLAM accounts for the observed errors ($\delta_{\text{[C II]}}$, $\delta_{\text{[O III]}}$, and δ_{SFR}) and can accept also different lines (e.g. CIII] λ 1909 instead of [O III]) as input (V20, Markov et al. 2022).

It is worth noting that the F19 model assumes a fixed O/C ratio (Asplund et al. 2009), with carbon and oxygen abundances linearly scaling with metallicity and a constant gas temperature in the ionized layer ($T = 10000$ K) and PDR ($T = 100$ K). The impact of the latter assumption has been tested against numerical radiative transfer calculations performed with CLOUDY (Ferland et al. 2017) over a wide range of ionization parameters and metallicities. Overall, the agreement is very good (see fig. 3 of F19), with CLOUDY confirming both the amplitude and linear slope of the increasing [C II] flux with gas metallicity, along with the saturation of the [C II] flux for increasing U . In spite of the inevitable simplifications of an analytical model such as that of F19, the differences with CLOUDY are relatively small (e.g. the [C II] flux is overestimated by F19 at most by ≈ 2 at any Z). Moreover, as outlined in V21, the temperature in the ionized layer does not have a strong impact on the predicted [O III] flux. The 88 μm line, and the other transition in the doublet ([O III] at 52 μm), have similar excitation energy ($T_{\text{ex},88} \approx 160$ K and $T_{\text{ex},52} \approx 260$ K) but different critical densities, hence for $T > 1000$ K their ratio is only affected by the gas density (Palay et al. 2012).

Up to now, the exploitation of GLAM in high- z galaxies has been limited to the derivation of the galaxy-averaged (n , κ_s , Z) values (V20, V21, Markov et al. 2022) because of the relatively low ($>$ kpc scales) spatial resolution of the [C II], [O III], and CIII] data (Smit et al. 2018; Carniani et al. 2020; Markov et al. 2022, respectively) used as input in the code. The recent work by Witstok et al. (2022), which gathers moderate resolution (\approx kpc) [C II] and [O III] observations in a sample of $z \approx 7$ galaxies, regridded to a common coordinate mesh of sub-kpc pixels, allows us for the first time the use of GLAM on a pixel-by-pixel basis, and the characterization at sub-kpc scales of the ISM properties of EoR sources.

We perform two types of analysis with GLAM. First, we fit a 2D Gaussian profile to the [C II], [O III], UV, and IR continuum maps, to compute the global size of the emission ($r_{\text{[C II]}}$, $r_{\text{[O III]}}$, r_{UV}) and, with that, infer the mean $\Sigma_{\text{[C II]}} = L_{\text{[C II]}}/\pi r_{\text{[C II]}}^2$, $\Sigma_{\text{[O III]}} = L_{\text{[O III]}}/\pi r_{\text{[O III]}}^2$, $\Sigma_{\text{SFR}} = (\text{SFR}_{\text{UV}} + \text{SFR}_{\text{IR}})/\pi r_{\text{UV}}^2$. The choice of a 2D Gaussian profile, instead of e.g. an exponential one, is for consistency with the V21 analysis $z \approx 6 - 9$ sources with barely resolved observations. Second, we feed to the model the $\Sigma_{\text{[C II]}}^i$, $\Sigma_{\text{[O III]}}^i$, Σ_{SFR}^i of each i -pixel of the grid (pixel size $\approx 0.3 - 0.8$ kpc, depending on the source), for which all the three quantities are above the 3σ level, to obtain spatially resolved derivation of the ISM parameters.

4 RESULTS

In this section, we present our results and their implications, starting with an overview of the spatially resolved versus global gas density, metallicity, and burstiness values computed with GLAM (Section 4.1). We then focus on key quantities from which we can infer insights on the ISM enrichment and baryon cycle (Section 4.2), the conversion of the gas into stars (Section 4.3), and the dust properties (Section 4.4) in the EoR.

4.1 Spatially resolved versus global ISM properties

In Fig. 1, we present the n , κ_s , and Z maps for the five galaxies in the Witstok et al. (2022) sample, produced with GLAM by simultaneously fitting the $\Sigma_{\text{[C II]}}^i$, $\Sigma_{\text{[O III]}}^i$, and the Σ_{SFR}^i in each pixel. We note that the central regions are characterized by higher gas density, burstiness parameter, and metallicity, suggesting an inside-out star formation scenario. The sources are likely experiencing a burst in star formation in connection with the central [O III] bright regions. The connection between recent bursts of star formation and high [O III]/[C II] ratios has been discussed on global galactic scales by several authors (Katz et al. 2017; Arata et al. 2020; Pallottini et al. 2022; Sugahara et al. 2022; Kohandel et al. 2023), but this is the first time that we obtain a quantitative measure of spatially resolved trends in burstiness within EoR galaxies.

As outlined in Section 3, for sources detected in the dust continuum (COS-3018, UVISTA-Z-001, and UVISTA-Z-019) we also considered the obscured SFR (SFR_{IR}) when deriving the total Σ_{SFR} . The SFR_{IR} is computed from the L_{IR} (see Section 2) using the conversion from Kennicutt & Evans (2012) in those pixels where the continuum is detected at $\geq 3\sigma$. For COS-3018 and UVISTA-Z-001 adding SFR_{IR} does not alter the smooth decreasing radial trends of (n , κ_s , Z) from the galaxy centre towards the periphery, but in UVISTA-Z-019 the SFR_{IR} produces a sharp gradient in the κ_s and n values towards the centre of the source. In the IR-detected central region both κ_s and n have higher values with respect to the neighbouring regions that are only UV detected. This finding can be even more prominent should the galaxy centres be characterized by warmer T_{d} than the value derived from the global SED fitting procedure (see Section 2).

In Fig. 2, we analyse the probability distribution function (PDF) of n , κ_s , and Z derived on pixel-by-pixel basis within the five galaxies. The relative error on the (n , κ_s , Z) parameters, see Appendix A for the corresponding maps, are in the range $\Delta n/n \sim 0.3 - 0.5$, $\Delta \kappa_s/\kappa_s \sim 0.15 - 0.20$, $\Delta Z/Z \sim 0.2 - 0.5$, respectively, depending on the source. The density distribution in all the galaxies is fairly narrow (≈ 0.4 dex between the minimum and maximum value) and the peak of the distribution ranges from $\log(n/\text{cm}^{-3}) = 2.5$ (UVISTA-Z-007), to $\log(n/\text{cm}^{-3}) = 2.9$ (UVISTA-Z-019), albeit higher resolution data might reveal larger variability in the density within the ISM of the sources. For comparison, we also report the *global value* and uncertainties for the same parameters obtained using the mean $\Sigma_{\text{[C II]}}$, $\Sigma_{\text{[O III]}}$, and Σ_{SFR} of each source. The global n (see Table 1) for each source is very close to the peak of the corresponding PDF over the pixels. This implies that using GLAM for deriving the gas density of a galaxy using the average [C II], [O III], and SFR surface density would return values that are representative of the actual ISM conditions within the source.

Our derived gas densities are slightly higher than the electron density, $\log(n_e/\text{cm}^{-3}) \approx 2.2$, inferred by Fujimoto et al. (2022) using [O III] λ 88 μm and [O III] λ 5007 JWST data in a $z \approx 8.4$ source. This is expected because the electron densities derived with standard methods based on optical/UV line ratios (e.g. Kewley et al. 2019) are sensitive to the conditions in the H II regions only, which, despite being connected with the surrounding environment, have an overall lower density than the neutral/molecular gas in the PDRs (fig. 1 of Vallini et al. 2021). It is also interesting to compare our results with those obtained by Davies et al. (2021) on the redshift evolution of the (electron) density. Davies et al. (2021) find an increasing trend with redshift from $\log(n_e/\text{cm}^{-3}) \approx 1.5$ at $z \approx 0$ to $\log(n_e/\text{cm}^{-3}) \approx 2.4$ at $z \approx 2.3$. Such a positive correlation is likely connected with the evolution in the normalization of the star formation main sequence,

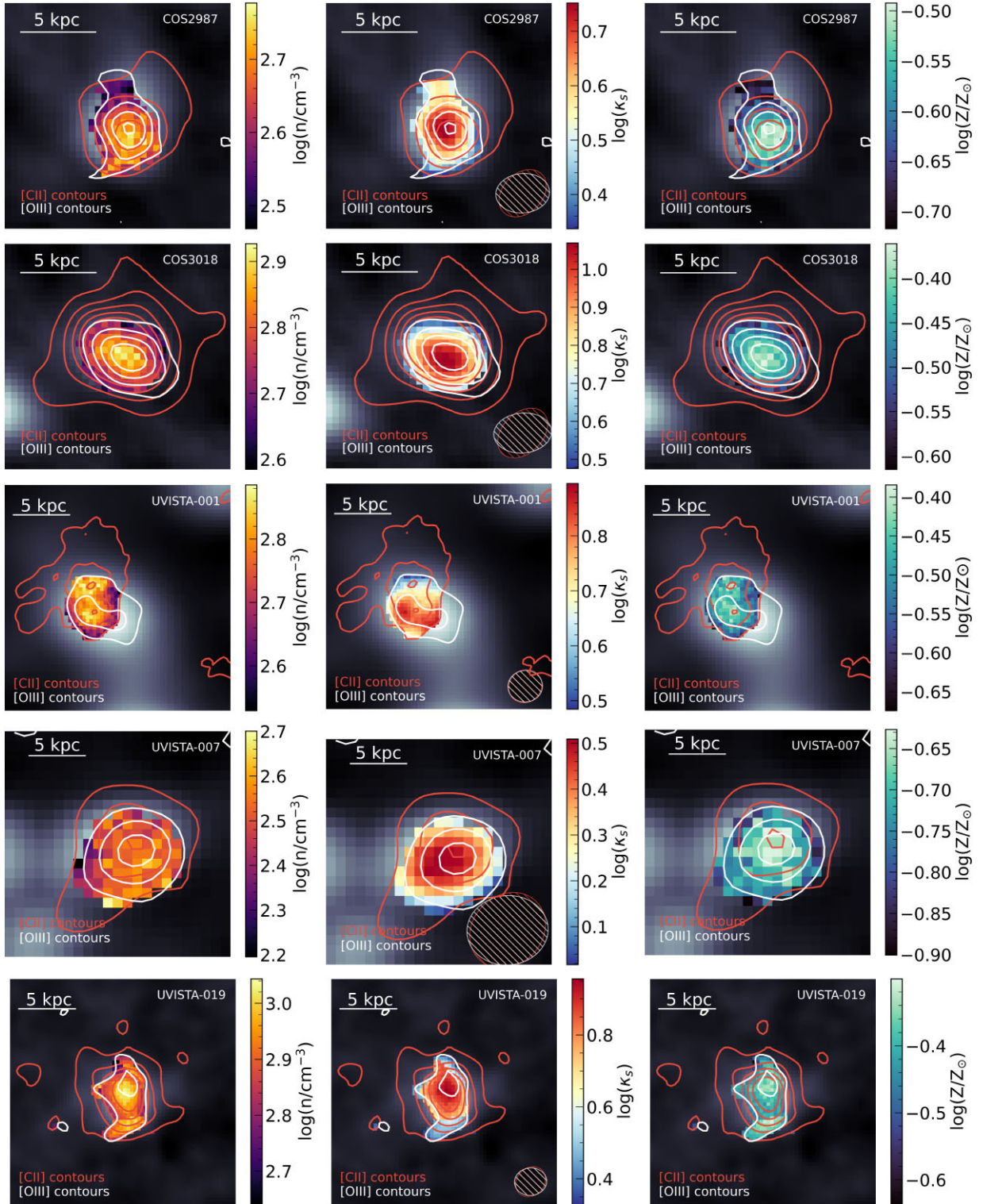


Figure 1. The derived gas density (n , left column), deviation from the KS relation (κ_s , central column), and gas metallicity (Z , right column) for the five galaxies (COS-2987, COS-3018, UVISTA-Z-001, UVISTA-Z-007, and UVISTA-Z-019, from top to bottom) analysed in this work with GLAM. The [C II] and [O III] contours (red and white, respectively) are overplotted onto the HST rest-frame UV images (in background). All the contours start at 3σ . The matched [C II] and [O III] beam sizes used for the present analysis are indicated in the bottom-right corner of the κ_s maps (central column).

and with the H II regions being embedded in parent giant molecular clouds (GMCs) characterized by higher densities at high- z (see also Sommovigo et al. 2020). Note that the total gas density derived with GLAM is explicitly linked to that of the GMCs as, by construction,

the density is parametrized in term of the H II region-PDR complexes tracing [O III] and [C II], respectively.

Our results at $z \approx 7$ are also in agreement with the density increase ($n_e \geq 300 \text{ cm}^{-3}$) with redshift recently found by Isobe et al. (2023)

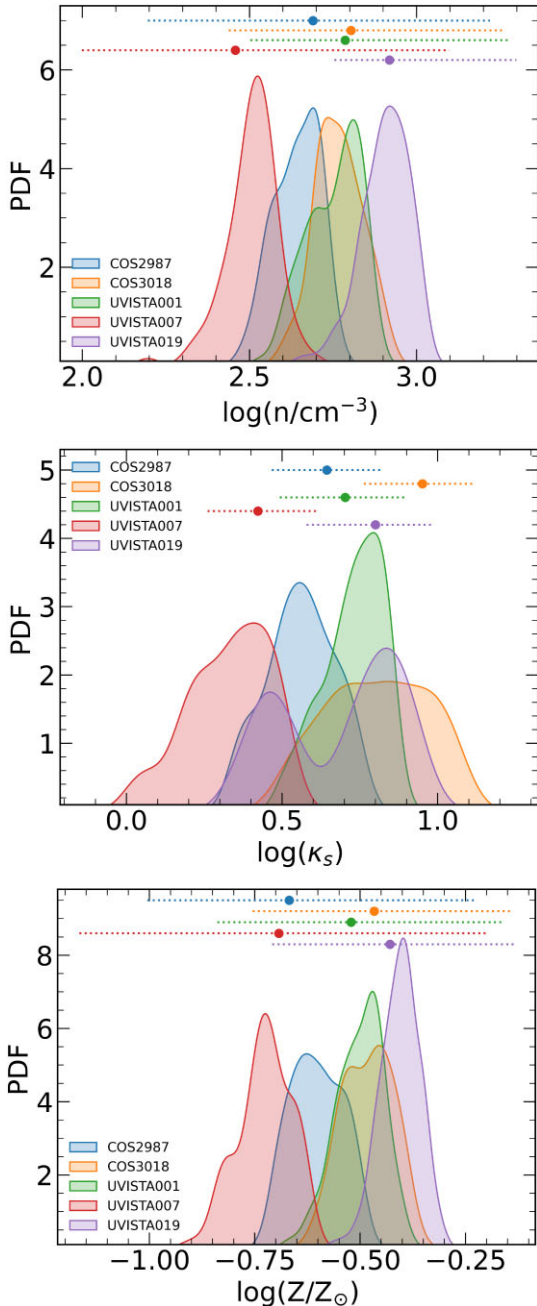


Figure 2. PDF of the gas density (n , upper panel), deviation from the KS relation (κ_s , centre), and metallicity (Z , lower panel) over the pixel maps in the five galaxies. The coloured dots (dotted lines) represent the global value (uncertainty) for n , κ_s , Z that derived with GLAM for each galaxy (same colour code of the PDFs) when using the global Σ_{CII} , Σ_{OIII} , and Σ_{SFR} .

exploiting $[\text{OII}]\lambda\lambda 3726, 3729$ fluxes in $z \approx 4.0 - 9$ sources. Isobe et al. (2023) identify an increase of the electron density with redshift that can be approximated as $n_e \propto (1+z)^p$, with $p \sim 1 - 2$. The exponent is explained by a combination of the compact morphology towards high- z , and the reduction of the electron density due to high electron temperatures of high- z metal-poor nebulae. Our method favours the $p \approx 2$ solution that implies $n_e \approx 500 - 1000 \text{ cm}^{-3}$ at $z \approx 7$.

All five galaxies lie above the KS relation, i.e. they are characterized by $\kappa_s > 1$. Overall the κ_s PDFs span a range between $\log(\kappa_s) \approx$

0.3 (UVISTA-Z-007), up to $\log(\kappa_s) \approx 1.3$ (UVISTA-Z-019). Note that for UVISTA-Z-019 we recover a clear bi-modality in the PDF of the burstiness parameter. This is because the dust-continuum detected region is more bursty than the outer part that is instead undetected with ALMA in continuum at $\approx 160 \mu\text{m}$. As for the gas density, the κ_s derived from global values is close to the peak of the corresponding PDF over the pixels, albeit the PDFs show a larger scatter, while the error on the global κ_s is rather small (≈ 0.2 dex). Moreover, the global κ_s tends to be skewed towards the higher end of the PDFs (see the case of UVISTA-Z-019).

Finally, we report the PDFs of the gas-phase metallicity obtained with GLAM in every pixel. As previously discussed for the density, and burstiness parameter, also the global Z of each galaxy is close to the peak of the corresponding PDFs. The uncertainties in the global values are much larger than the typical width of the PDFs and comparable to the pixel relative error. All the galaxies have sub-solar gas metallicities ranging from $\log Z/Z_\odot = -0.75$ of UVISTA-Z-007 up to the $\log(Z/Z_\odot) = -0.25$ in some of the pixels of UVISTA-Z-019. As a caveat, we stress that the GLAM model assumes a fixed O/C ratio; the O/C in the early Universe is expected to be higher than that measured at $z = 0$ (e.g. Maiolino & Mannucci 2019) thus the Z inferred from the (high) $\Sigma_{\text{OIII}}/\Sigma_{\text{CII}}$ ratios might be underestimated by GLAM.

4.2 Metallicity gradients

Metallicity gradients are sensitive probes of the complex network of processes that regulates gas inflows/outflows, feedback, and mixing within the ISM of galaxies across cosmic time (e.g. Sánchez Almeida et al. 2014). Theoretical models and numerical simulations (e.g. Gibson et al. 2013; Ma et al. 2017; Hemler et al. 2021; Sharda et al. 2021; Tissera et al. 2022) addressed the physical mechanisms shaping the metallicity gradient within galaxies and its evolution with redshift. In particular, negative gradients represent one of the strongest pieces of evidence for the inside-out galaxy formation scenario in which the nucleus forms first and more metals enrich the galaxy centre as compared to the disc. A flattening in the gas-phase galaxy metallicity can instead be the signature of star formation and strong stellar feedback (e.g. Gibson et al. 2013; Ma et al. 2017), merger events (e.g. Rupke, Kewley & Barnes 2010) and pristine gas inflows towards the central regions (e.g. Ceverino et al. 2016). The combined effects of these physical processes modulate the evolution of the metallicity gradients as function of redshift.

Thanks to the pixel-by-pixel derivation of the ISM properties within the five galaxies in our sample, it is possible to study how the gas metallicity, but also the galaxy burstiness and gas density, vary within each source as a function of the galactocentric radius. To do so we use the `RadialProfile` class within the `photutils` package to compute the azimuthally averaged value, and the corresponding uncertainty, of all the three GLAM parameters over circular annuli of ≈ 0.8 kpc (2 pixels) width. The exception is UVISTA-Z-007 for which the radial spacing corresponds to 1 pixel, due to the lower spatial resolution of the data. We assume the centre ($r = 0$) to be located at peak of the $[\text{O III}]$ emission. The error on the radial profiles is computed by providing to the `RadialProfile` routine the map of the $1\sigma_i$ errors of each parameter in each i -th pixel as computed by GLAM (see Appendix A).

The result of this procedure is outlined in Fig. 3 where we show the gas metallicity profiles for the five sources analysed in this work. We refer the interested reader to Appendix B for details regarding the radial profiles of the gas density and burstiness. We note that Z gradients as a function of galactocentric radius are marginally

Table 1. Derived physical properties for the galaxies analysed in this work. We list the galaxy names (column 1), redshift (column 2), global values (see text for details) for the gas density (column 3), burstiness (column 4), and gas metallicity (column 5). In column 6, we report the radial gradient $\nabla(\log Z)$. Finally, in column 7, we report the median value of the depletion time over the spatially resolved pixels, and its $\pm\sigma$ deviation.

name	redshift	$\log(n/\text{cm}^{-3})$	$\log(\kappa_s)$	$\log(Z/Z_\odot)$	$\nabla(\log Z)$	t_{dep} (Myr)
COS-2987	6.807	$2.69^{+0.53}_{-0.49}$	$0.64^{+0.18}_{-0.18}$	$-0.67^{+0.44}_{-0.34}$	-0.05 ± 0.09	190 ± 56
COS-3018	6.854	$2.80^{+0.46}_{-0.37}$	$0.95^{+0.17}_{-0.19}$	$-0.47^{+0.33}_{-0.29}$	-0.06 ± 0.08	100 ± 40
UVISTA-Z-001	7.060	$2.79^{+0.49}_{-0.28}$	$0.70^{+0.19}_{-0.21}$	$-0.52^{+0.36}_{-0.32}$	-0.03 ± 0.06	96 ± 21
UVISTA-Z-007	6.749	$2.46^{+0.64}_{-0.46}$	$0.42^{+0.20}_{-0.16}$	$-0.69^{+0.49}_{-0.47}$	-0.04 ± 0.10	280 ± 130
UVISTA-Z-019	6.754	$2.92^{+0.38}_{-0.16}$	$0.80^{+0.19}_{-0.22}$	$-0.43^{+0.29}_{-0.28}$	0.00 ± 0.03	84 ± 14

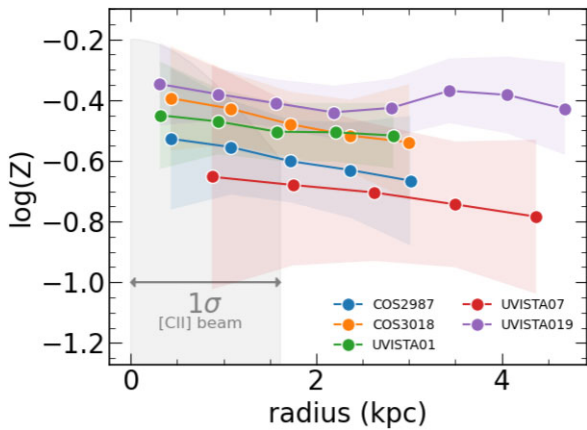


Figure 3. Radial profiles of the gas metallicity Z for the five sources analysed in this work. The grey shaded region denotes the 1σ width of the median [C II] beam. The shaded coloured regions represent the 1σ error, see the text for details on the calculation.

negative, but consistent with being flat within the errors. The median in the sample $\nabla(\log Z) \approx -0.03 \pm 0.07$ dex/kpc and the values for each source can be found in Table 1. Our analysis suggests that the metal enrichment in the central regions might be connected with recent burst (high κ_s) of star formation and that, at the same time, the vigorous starburst (and possibly flickering SFR, e.g. Pallottini & Ferrara 2023) implies copious energy injection via stellar feedback, thus flattening the gradient.

In Fig. 4, we report the gradient estimates for the five sources as a function of their redshift, comparing them with observations and theoretical models in the literature. Negative metallicity gradients in the radial direction have been confirmed in most galaxies at $z \approx 0$ (e.g. Stanghellini & Haywood 2010; Belfiore et al. 2017; Stanghellini & Haywood 2018) whereas there is evidence of an evolution of the metallicity gradients towards a flattening at high- z (Wuyts et al. 2016; Carton et al. 2018; Curti et al. 2020) albeit with large scatter (see e.g. the negative slopes reported by Jones et al. 2013).

In the coming years, high- z samples with spatially resolved metallicity measurements will rapidly expand thanks to JWST that will allow observing nebular lines (e.g. [O III] λ 5007, H β , [O II] λ λ 3726, 3729) that are routinely used as metallicity estimators. Wang et al. (2022) reported the first JWST determination of the metallicity gradient in a $z \approx 3.2$ galaxy, finding a strongly positive gradient likely due to the interaction with a nearby object. With our method, we do not find evidence of positive gradients in our sources at $z \approx 7$ and this is likely connected to the fact that the star formation burst, as traced by [O III], is localized at the centre of the sources. Nevertheless, as a caveat, we warn the reader that our derivation of the metallicity using GLAM represents an indirect methodology

and thus the uncertainty in the comparison with gradients in the O/H abundance obtained with other methods is large. In particular, systematics in the Z determination can arise from the fact that the likely enhancement O/C ratio at low metallicity (e.g. Maiolino & Mannucci 2019) is not accounted for in GLAM (see Section 3). This assumption implies that the Z inferred with GLAM from the $\Sigma_{\text{[OIII]}}$ and $\Sigma_{\text{[CII]}}$ fitting, might be biased towards higher values to compensate for a lower than expected oxygen abundance at sub-solar metallicity.

4.3 Kennicutt–Schmidt relation and gas depletion time

In Fig. 5 we show the $\Sigma_{\text{SFR}}-\Sigma_{\text{gas}}$ (KS) relation for the five galaxies analysed in this work. We consider both the 2-D probability density distribution (PDF) of the values within the pixels, and the location of the five galaxies in the KS plane when deriving the Σ_{gas} from the global κ_s and Σ_{SFR} .

In line with previous works (V21), we confirm that the inferred location in the KS plane of a source, obtained considering the global $\Sigma_{\text{[CII]}}$ and $\Sigma_{\text{[OIII]}}$ values, traces the most starbursting (high κ_s) patches within the ISM of the objects. The global values fall indeed in the upper- κ_s end of the 2-D distribution of the pixels within each galaxy. The global and spatially resolved $\Sigma_{\text{gas}}-\Sigma_{\text{SFR}}$ values place our galaxies above the $z = 0$ relation for spiral galaxies (de los Reyes & Kennicutt 2019) in the region populated by starburst sources in the local Universe (Kennicutt & De Los Reyes 2021).

Interestingly, their location is in good agreement with the position in the KS plane of simulated star-forming galaxies at $z \approx 7$ extracted from the SERRA zoom-in cosmological simulation (Pallottini et al. 2022) that cover a range of stellar masses ($10^8 M_\odot \lesssim M_* \lesssim 5 \times 10^{10} M_\odot$) and SFR ($\text{SFR} \approx 1 - 100 M_\odot \text{ yr}^{-1}$) that encompass that of the five LBGs analysed in this work.

Our analysis suggests that luminous LBGs in the EoR are characterized by an efficient conversion of the gas into stars. To put this conclusion into a broader context we perform a comparison with three luminous $z \approx 6$ LBGs (J0235-0532, J1211-0118 and J0217-0208) first targeted by Harikane et al. (2020) with ALMA in [C II] and [O III] characterized by UV luminosities ($\approx 3 \times 10^{11} L_\odot$) and [O III]/[C II] ratios ($\approx 3 - 8$) similar to those of our galaxy sample (see Witstok et al. 2022). V21 studied J0235-0532, J1211-0118 and J0217-0208 with GLAM, albeit using only the barely resolved [C II] and [O III] data available at the time, deriving their burstiness parameter, gas-phase metallicity and density. Thanks to the detection of the CO(6–5) line in J0235-0532, and the upper limits in the other two LBGs (Ono et al. 2022), it is possible to compare their location in the KS plane from GLAM with that obtained using the CO as fiducial gas proxy, which, however, involves many uncertain parameters. In fact, (i) the conversion of the CO(6–5) flux into the

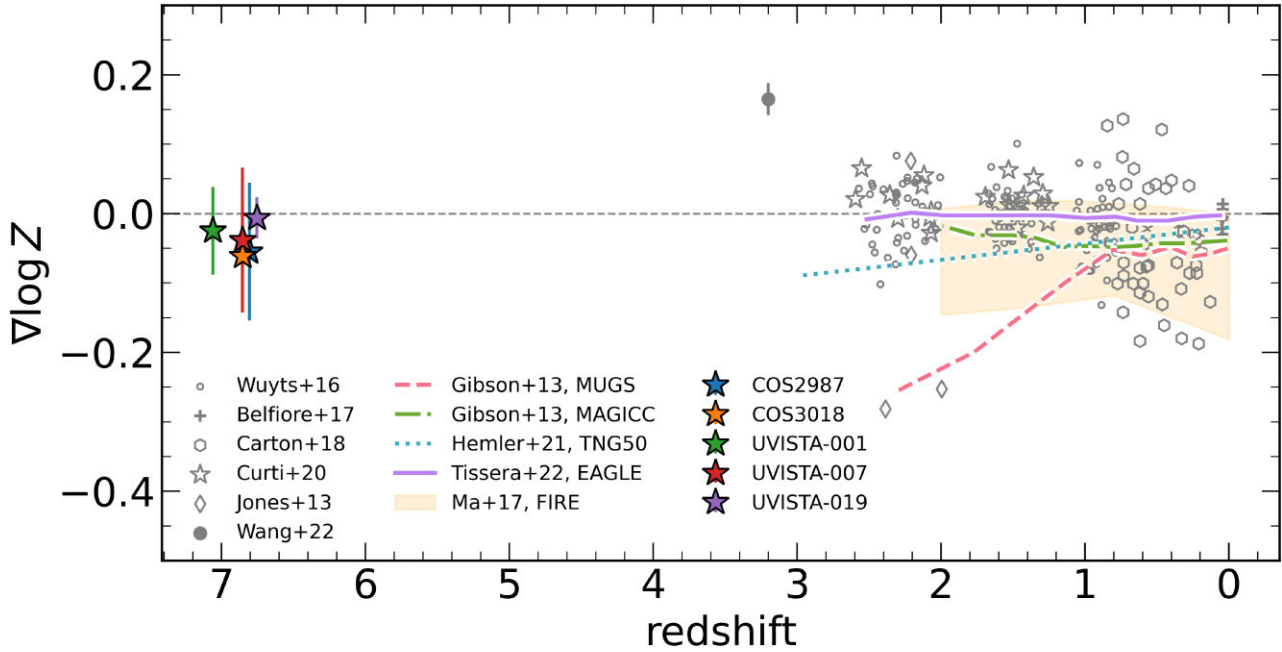


Figure 4. The redshift evolution of the metallicity gradient. Data, plotted without errors for *clarity*, are taken from Curti et al. (2020, stars), Carton et al. (2018, hexagons), Belfiore et al. (2017, crosses), Wuyts et al. (2016, small circles), Jones et al. (2013, diamonds), and Wang et al. (2022, filled circle). Simulated gradients taken from the literature are shown with lines, Gibson et al. (2013, pink dashed line: MUGS, normal feedback, green dot-dashed line: MAGICC enhanced feedback), Hemler et al. (2021, IllustrisTNG, cyan dotted line), Tissera et al. (2022, EAGLE, purple solid line), and Ma et al. (2017, FIRE, orange shaded area). The gradients inferred for our five sources are marginally negative but consistent with a slope equal to 0.

$\text{CO}(1-0)^4$ depends on the CO Spectral Line Energy Distribution (COSLED) excitation, which is observationally unconstrained in LBGs in the EoR (Pavesi et al. 2019), (ii) the CO-to- H_2 conversion factor is also highly uncertain in high- z sources as it depends on metallicity (Bolatto, Wolfire & Leroy 2013) and (iii) the actual size of the molecular gas distribution cannot be derived from unresolved CO observations. As discussed by Ono et al. (2022), assuming $\alpha_{\text{CO},\text{MW}} = 4.3 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ and the CO(6–5)/CO(1–0) ratio of $z \approx 1$ star-forming galaxies (Daddi et al. 2010), all the three sources lie below the KS relation, at odds with their high $[\text{O III}]/[\text{C II}]$ ratios that are usually powered by ongoing bursts of star formation (Arata et al. 2020; Vallini et al. 2021; Kohandel et al. 2023). If, instead, one assumes the $\alpha_{\text{CO}} = 1.4 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ and CO(6–5)/CO(1–0) ratio derived by Vallini et al. (2018) for Althaea—a simulated LBG extracted from the precursor of the SERRA zoom-in simulation (Pallottini et al. 2017, 2019, 2022)—the sources are compatible with the location of starburst galaxies in agreement within the errors (see Fig. 5) with the location inferred with GLAM in V21. Given that the CO(6–5) traces dense/warm ($n_{\text{crit}} \approx 3 \times 10^5 \text{cm}^{-3}$, $T_{\text{ex}} \approx 116 \text{K}$) molecular gas (Wolfire, Vallini & Chevance 2022, for a recent review) the line is expected to be luminous in bursty galaxies that are experiencing on-going star formation within dense GMCs (Vallini et al. 2018). UVISTA-Z-019 would be therefore an ideal target for CO follow-up being the most dense and bursty among the five LBGs in our sample with both CO(6–5) and CO(7–6) falling into ALMA band 3. Moreover, its continuum detection translates into $\text{SFR}_{\text{IR}}/\text{SFR}_{\text{TOT}} = 0.7$, namely 70 per cent of the star formation is dust obscured and dust shielding is one of the key necessary conditions mitigating the CO dissociation in star

forming regions (Wolfire, Hollenbach & McKee 2010). Note that the UVISTA-Z-019 IR luminosity (and the obscured fraction of the star formation) $L_{\text{IR}} = 3.1 \times 10^{11} L_{\odot}$, ($\text{SFR}_{\text{IR}}/\text{SFR}_{\text{TOT}} = 0.7$) is similar to that of J0235–0532 ($L_{\text{IR}} = 5.8 \times 10^{11} L_{\odot}$, $\text{SFR}_{\text{IR}}/\text{SFR}_{\text{TOT}} = 0.6$, respectively), the CO(6-5)-detected LBG from the Ono et al. (2022) sample.

High burstiness parameters translate into short gas depletion times (e.g. Tacconi et al. 2013; Tacconi, Genzel & Sternberg 2020), defined on spatially resolved scales as $t_{\text{dep}} = \Sigma_{\text{gas}}/\Sigma_{\text{SFR}} \propto \kappa_s^{-1} \Sigma_{\text{gas}}^{-0.4}$. The evolution of the depletion time with redshift (see Fig. 6) is a fundamental quantity shaping galaxy evolution, as it quantifies the typical timescales for the conversion of the gas into stars, thus ultimately the galaxy and stellar build up from the Dark Ages to the present day and the efficiency with which galaxies build up their stellar mass. In particular, Tacconi, Genzel & Sternberg (2020) found that the integrated depletion timescale, namely that derived as $t_{\text{dep}} = M_{\text{gas}}/\text{SFR}$, depends mainly on the redshift and offset from the main sequence ($t_{\text{dep}} \propto (1+z)^{-1} \times \Delta_{\text{MS}}^{-0.5}$). The trend for main sequence galaxies has been overall confirmed in the $z \approx 4.5 - 5.8$ redshift range by the ALPINE (Le Fèvre et al. 2020) survey (Dessauges-Zavadsky et al. 2020) by inferring the gas mass from the $[\text{C II}]$ luminosity (see Zanella et al. 2018).

From the spatially resolved κ_s within our sources we obtain median t_{dep} ranging from ≈ 280 Myr of UVISTA-Z-007 to ≈ 80 Myr of UVISTA-Z-019, in agreement with the redshift evolution proposed by Tacconi, Genzel & Sternberg (2020). In particular, these values encompass the tight range between the extrapolation of t_{dep} out to $z \approx 8$ for main sequence galaxies and that for starburst (deviation from the main sequence, $\Delta_{\text{MS}} = 10$) sources.

Finally, we also compare the spatially resolved t_{dep} obtained from GLAM against that derived on global scales by using the Zanella et al. (2018) conversion factor to infer the gas mass

⁴The CO(1–0) luminosity can be then converted into the molecular mass via the CO-to- H_2 conversion factor, $M_{\text{mol}} = \alpha_{\text{CO}} L'_{\text{CO}(1-0)}$

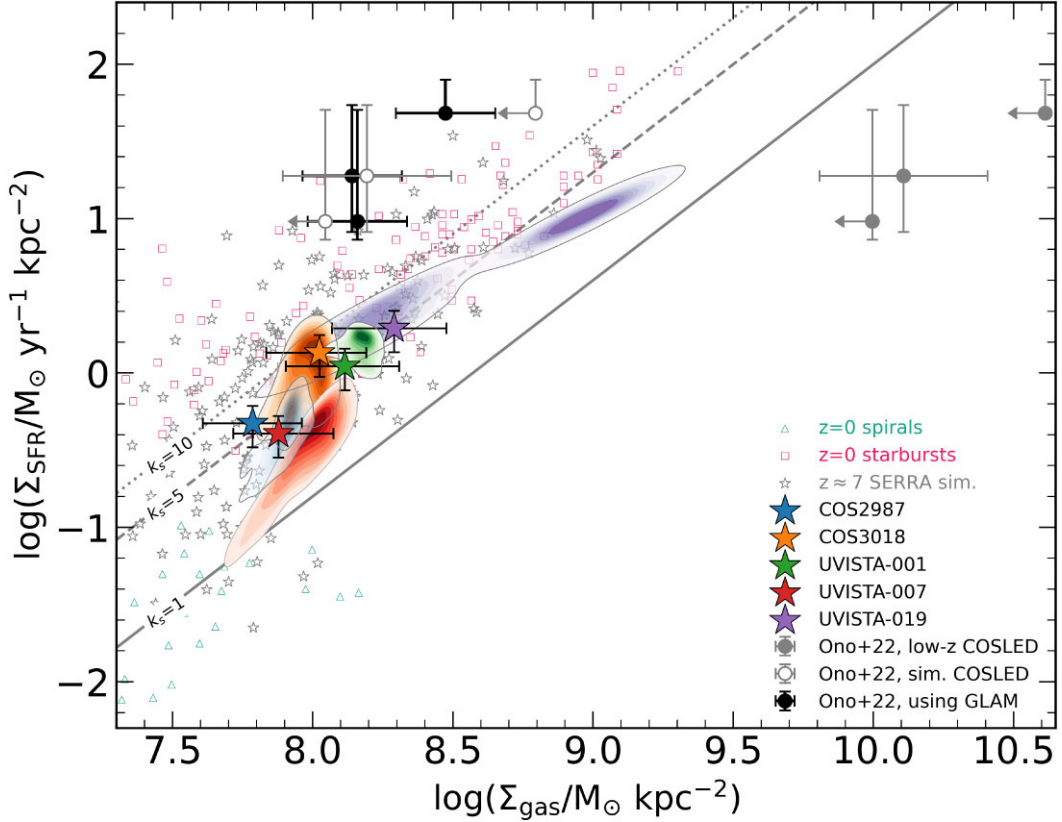


Figure 5. Relation between SFR and total (atomic plus molecular) gas surface densities $\Sigma_{\text{SFR}} = \kappa_s 10^{-12} \Sigma_{\text{gas}}^{1.4}$ (Heiderman et al. 2010). The solid (dashed, dotted) line represents the relation for $\kappa_s = 1$ ($\kappa_s = 5$, $\kappa_s = 10$, respectively). Local spiral galaxies (de los Reyes & Kennicutt 2019) and starburst galaxies (Kennicutt & De Los Reyes 2021) are indicated with green triangles and magenta squares, respectively. For the local spiral (starburst) galaxies the molecular component of the total gas surface density is derived assuming the Milky Way CO-to-H₂ conversion factor $\alpha_{\text{CO,MW}} = 4.3 \text{ M}_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$ (ULIRG conversion factor $\alpha_{\text{CO,ULIRG}} = 0.86 \text{ M}_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$). Simulated $z \approx 7$ galaxies from SERRA (Pallottini et al. 2022) are indicated with empty grey stars. The 2D density distribution of Σ_{SFR} versus Σ_{gas} of the pixels within each of the five $z \approx 7$ galaxies analysed in this work are represented with shaded coloured regions. The coloured stars indicate the location of the sources when considering their global values. We complement the plot with three LBGs at $z \approx 6$ that have been detected in [C II], [O III] (Harikane et al. 2020), and recently followed up in CO(6–5) by Ono et al. (2022). Their location (grey filled versus grey empty circles) depends on the assumed CO excitation (low- z COSLED versus high- z simulated COSLED) and CO-to-H₂ conversion factor ($\alpha_{\text{CO,MW}}$ versus $\alpha_{\text{CO}} = 1.4 \text{ M}_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$ derived by Vallini et al. (2018) in the pilot SERRA simulation (Pallottini et al. 2017)). The location in the KS plane of the three LBGs derived from [C II] and [O III] data using GLAM (V21) is indicated with black circles.

$M_{\text{gas}} = \alpha_{[\text{CII}]} L_{[\text{CII}]}$. This also enables a fair comparison with the depletion times in ALPINE (Dessauges-Zavadsky et al. 2020), which appear still marginally higher than those derived in our five galaxies.

For COS-2987 the global value of the t_{dep} is in excellent agreement the median of the spatially resolved one. For the other four sources the two methods return t_{dep} that agree within the errors (see Fig. 6). More precisely, in COS-3018, UVISTA-Z-001, UVISTA-Z-019 the global value is higher than the median of the spatially resolved one. Overall we interpret this trend as the probe that GLAM, by using both the information on [C II] and [O III], is more sensitive to the starbursting regions within galaxies, whose depletion time is expected to be shorter than the average value. The exception is UVISTA-Z-007 that is the galaxy with the poorer spatial resolution of the [C II] and [O III] data for which the median of the spatially resolved value is lower than the global t_{dep} . We point out that thanks to the derivation of the t_{dep} with GLAM the five galaxies analysed here allow to constrain the extrapolation of the redshift evolution of the depletion time in the EoR, and the method is definitively promising as an alternative in galaxies for which the CO detection might be challenging.

4.4 Linking dust temperature and gas depletion time

In the last few years, several works dealing with Far Infrared (FIR) stacked SED fitting across cosmic time ($z \approx 0-10$), have inferred the presence of a tight correlation between T_{d} and redshift (Schreiber et al. 2018; Bouwens et al. 2020; Viero et al. 2022, see also Liang et al. 2019). Sommovigo et al. (2022a, b) proposed a theoretical explanation for the $T_{\text{d}} - z$ relation based on the evolution of the total gas depletion time. They show that, in a simplified single-phase ISM model, $T_{\text{d}} \propto t_{\text{dep}}^{-1/6}$ (see equation 10 of Sommovigo et al. 2022a also reported in Fig. 7), resulting in a mild increase of T_{d} with redshift ($T_{\text{d}} \propto (1+z)^{0.4}$) due to the shorter t_{dep} at early epochs (Fig. 6) produced by the more vigorous cosmic accretion. At any fixed epoch, the scatter in T_{d} is produced by variations in metallicity and optical depth τ_{eff} (see Fig. 7), with lower Z (higher τ_{eff}) resulting in warmer dust.

To investigate this scenario, in Fig. 7, we report the dust temperature for the sources in our sample as a function of the depletion times derived in the previous section. Unfortunately, for most of our sources (and in general for galaxies towards the EoR, e.g. Béthermin et al. 2020; Gruppioni et al. 2020; Inami et al. 2022),

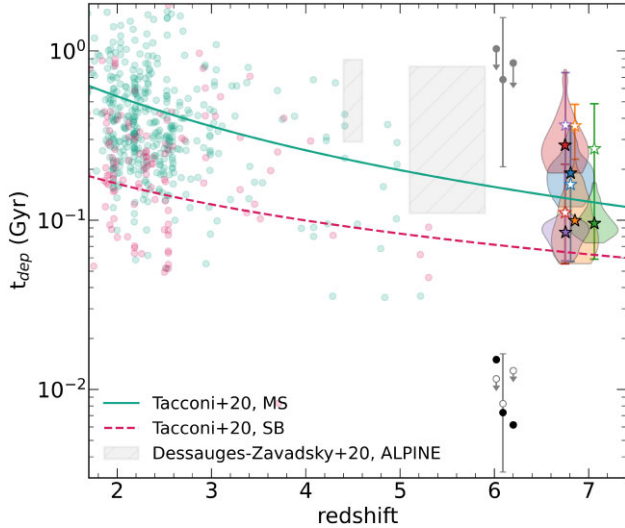


Figure 6. Redshift evolution of the depletion time. The coloured violin plots represent the distribution of t_{dep} within our galaxies as inferred from the spatially resolved κ_s . The median value (also reported in Table 1) is highlighted with a filled coloured star. For comparison, empty stars with colour bars represent the integrated depletion time ($t_{\text{dep}} = M_{\text{gas}}/\text{SFR}$) derived using the Zanella et al. (2018) relation for inferring the gas mass from the [C II] luminosity, and considering the SFR_{tot} by Witstok et al. (2022). We adopt the same colour-code of the previous figures.

The extrapolation out to $z = 7.5$ of the best fit relations from Tacconi, Genzel & Sternberg (2020) for main sequence galaxies (green points) and starburst galaxies (magenta points) are indicated with solid and dashed lines, respectively. The t_{dep} from ALPINE in two redshift bins (Dessauges-Zavadsky et al. 2020) is indicated with grey shaded areas. The t_{dep} inferred by Ono et al. (2022) for the Harikane et al. (2020) LBGs is indicated with filled and empty grey circles (same colour-code of Fig. 5) while that inferred with GLAM with filled black circles.

only a single/two ALMA dust continuum detections are generally available. Thus, the T_d derived from SED fitting is highly uncertain ($\Delta T_d/T_d \gtrsim 60$ per cent, see Witstok et al. 2022) hampering an unambiguous analysis.

To overcome this problem, and study in detail the $T_d - t_{\text{dep}}$ relation in all the sample, we use an alternative method (Sommovigo et al. 2021) based on the combination of the ALMA FIR continuum data point with the [C II] luminosity information. The latter is used as a proxy of the total gas (and dust) mass, so that the single continuum measurement can be exploited to constrain the dust temperature. In this case, this allows us to constrain T_d to a greater precision ($\Delta T_d/T_d \gtrsim 20$ per cent). We obtain $T_d \approx 35$ K for the 3 dust-continuum detected galaxies, and $T_d \lesssim 30$ K for the 2 continuum-undetected ones. These values are consistent within the uncertainties with those derived from the SED fitting by Witstok et al. (2022), where available.

As shown in Fig. 7, we find that galaxies with shorter depletion times host warmer dust. Our results are in agreement with the physically motivated prediction for the optically thin ($\tau_{\text{eff}} = 0.05$) case (Sommovigo et al. 2022a, b) when accounting for the metallicity variation within the sample.

If future ALMA dust continuum observations at shorter wavelengths will confirm the warmer T_d suggested by the median values for UVISTA-Z-001 and UVISTA-Z-019 by Witstok et al. (2022), this might indicate that these galaxies are characterized by $\tau_{\text{eff}} = 1$ and thus a spatially segregated scenario between dust and UV emission. This can be further confirmed by high-spatial resolution

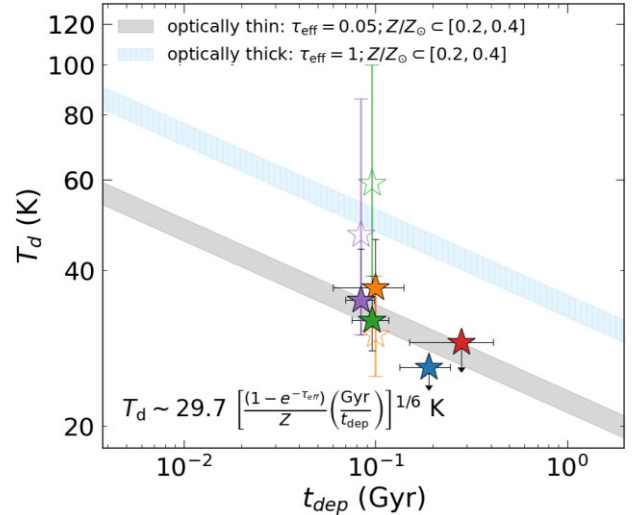


Figure 7. The dust temperature (T_d) as a function of the depletion time (t_{dep}), the colour-code is the same as in previous figures. Filled stars represent T_d derived with the Sommovigo et al. (2021) method, while empty stars those obtained by Witstok et al. (2022) by fitting the SED (probed by 2 photometric points) for the five LBGs analysed in this paper. The theoretical relation between T_d and t_{dep} is shown in the lower-left corner. The grey (blue) shaded region highlights the solutions for $\tau_{\text{eff}} = 0.05$ ($\tau_{\text{eff}} = 1$) and varying the metallicity in the range spanned by our estimates with GLAM for our sources.

ALMA observations (tracing the dust-obscured star formation) in conjunction with JWST (tracing the un-obscured one).

5 SUMMARY

In this paper, we have used GLAM (Vallini et al. 2020, 2021) to derive the ISM properties (gas density, deviation from the KS relation, and gas metallicity) in five UV luminous LBGs at $z \approx 7$ for which moderately resolved [C II] and [O III] observations are available. We have compared the pixel-by-pixel values for the ISM parameters with the global ones derived with the same methodology using instead average $\Sigma_{[\text{C II}]}$ and $\Sigma_{[\text{O III}]}$ surface brightness values. We confirm the conclusion by V21, namely that global values are biased towards the most luminous ISM regions. The main results from our spatially resolved analysis are the following:

(i) The distribution of the gas density in the five LBGs is narrow and peaks in the range $\log(n/\text{cm}^{-3}) = 2.5 - 3.0$, depending on the source. The gas densities obtained are higher than typical values in local galaxies, hence suggesting an overall increase in the mean gas density in the ISM at the early epochs.

(ii) We derived radial profiles for the metallicity, density, and burstiness. In particular, the metallicity shows a mildly negative radial gradient that, within the uncertainties, is compatible with being flat.

(iii) All five galaxies lie above the KS relation by a factor of $\approx 3 - 10$, in perfect agreement with expectations from cosmological zoom-in simulations (Pallottini et al. 2022) at the same redshift. The κ_s value is higher in the centre. In some cases, we obtain a bimodal distribution in regions where dust continuum emission is detected, suggesting the presence of dense, dust-obscured, highly star-forming regions.

(iv) We predict that bursty galaxies with dense gas (such as UVISTA-Z-019) would be an ideal target for ALMA follow-ups in CO(6-5) as mid- J CO lines trace warm/dense molecular gas.

(v) The gas depletion times, derived from the KS relation, are in the range $t_{\text{dep}} \approx 80\text{--}250$ Myr. The t_{dep} of the five sources fall between that predicted by the extrapolation out to $z \approx 7$ of for MS and SB galaxies of the Tacconi, Genzel & Sternberg (2020) relation.

(vi) The dust temperature of the five sources correlates with t_{dep} , as predicted by theoretical models (Sommovigo et al. 2021) for an optically thin medium. We confirm that the redshift evolution of the dust temperature might be the imprint of a more efficient conversion of the gas into stars.

The work presented in this paper highlights the huge potential of the synergy between physically motivated line emission models and spatially resolved observations of [C II] and [O III] for constraining a wealth of ISM properties within galaxies in the EoR. Simply using a handful of observables – traced down to kpc scales – allows to determine the KS relation, metallicity profiles, and impact of gas accretion on the dust continuum properties. Further follow-ups at higher spatial resolution in a larger samples of sources already detected in [C II] will allow putting the results presented here on a statistically more robust basis.

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DATA AVAILABILITY

GLAM can be accessed on Github at https://vallini.github.io/MCMC_galaxylineanalyzer/. HST data underlying this article are available in the MAST archive at 10.17909/6gya-3b10 (GO 13793), 10.17909/T9-JHSFM392 (GO 16506) and from <https://archive.stsci.edu/prepds/3d-hst/> (the 3D-HST Treasury Program). The ALMA data are available in the ALMA science archive at <https://almascience.eso.org/asax/> under the following project codes: 2015.1.01111.S, 2018.1.01359.S, 2018.1.00429.S, 2018.1.01551.S, 2017.1.00604.S, 2015.1.00540.S, 2018.1.00085.S, 2018.1.00933.S, 2019.1.01611.S, 2019.1.01524.S, 2018.1.00085.S, and 2019.1.01611.S.

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APPENDIX A: ERROR MAPS

In Fig. A1, we report the relative errors on the (n, κ_s, Z) parameters derived with GLAM on pixel-by-pixel basis. The error on gas density ranges between 30 per cent and 50 per cent depending on the source. The burstiness parameter is the ISM property that is better constrained, with errors ranging between 15 per cent and 20 per cent. Finally, the metallicity errors range between 25 per cent and 50 per cent. We note that while the errors for the gas density and metallicity are lower in the central regions for all the galaxies, the trend is the opposite for κ_s in UVISTA-Z-007 and UVISTA-Z-019, where the parameter is less constrained in the centre. The uncertainties on (n, κ_s, Z) are influenced by the SNR of the inputs ($\Sigma_{\text{[C II]}}$, $\Sigma_{\text{[O III]}}$ and Σ_{SFR}), which is higher in the centre where the emission is brighter. However, this is not the only aspect influencing the uncertainty on the parameters derived from the MCMC. In fact, the analytical functions describing the [C II] and [O III] fluxes are both characterized by a plateau in the flux at large Σ_{SFR} (see F19), which makes the model more degenerate. For this reason there are regimes in which, the central part of the sources being characterized by large Σ_{SFR} , the GLAM parameters end up being less precisely constrained albeit the SNR of the input data is higher. Improving the precision on the SFR tracers, with better rest-frame UV and IR data, can help in alleviating this issue.

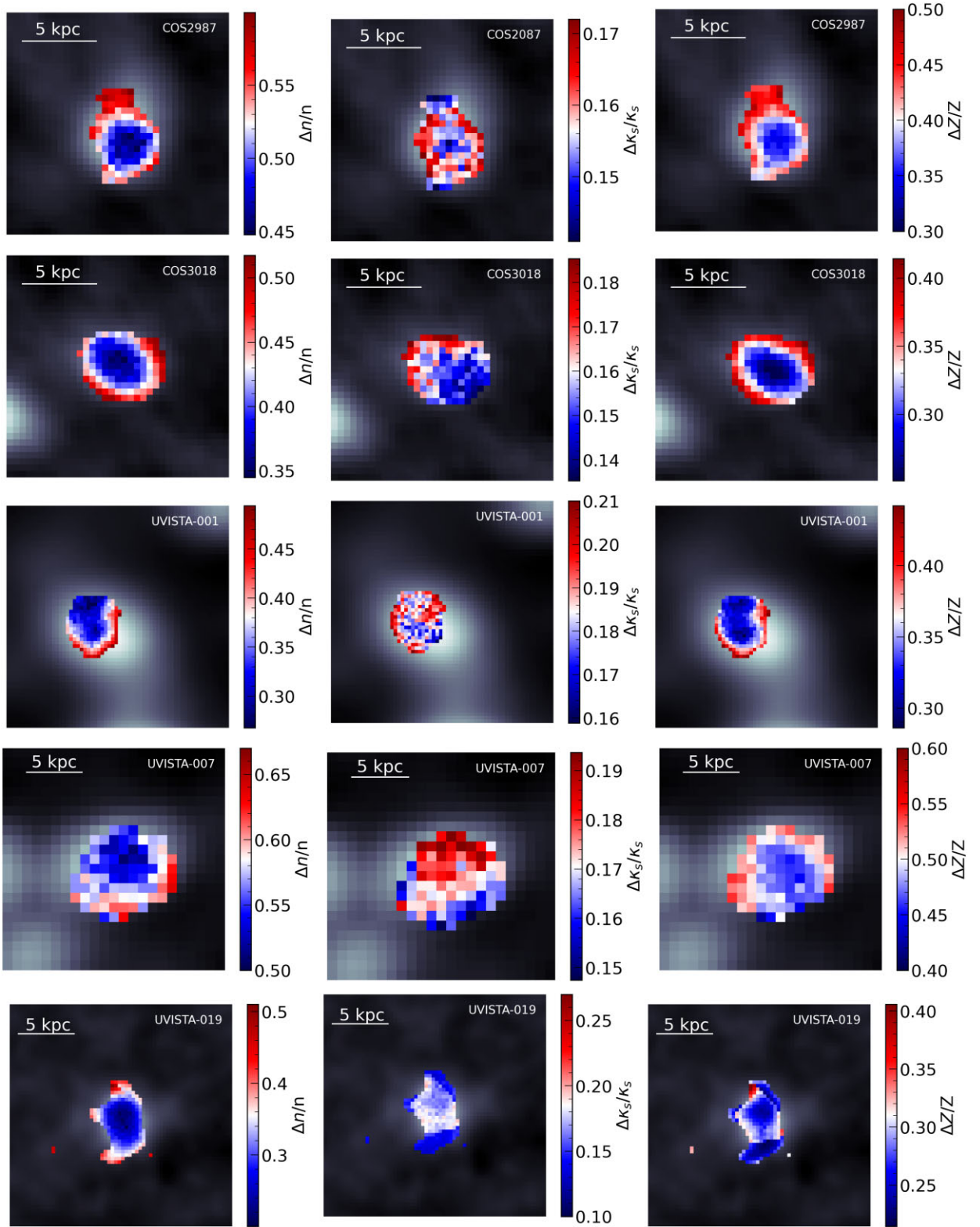


Figure A1. Maps of the relative errors on the gas density (left column), burstiness parameter (central column), and metallicity (right column) as obtained from GLAM on pixel-by-pixel basis.

APPENDIX B: DENSITY AND BURSTINESS GRADIENTS

We computed the gradient for the gas density and the burstiness parameter in the same way discussed for the metallicity ones. The results for our five sources are shown in Fig. B1. We note that while the density profile is consistent with being flat (the median gradient among the five galaxies in our sample is $\nabla(\log n) \approx -0.01 \pm 0.02 \text{ dex kpc}^{-1}$), the burstiness parameter shows a steeper decrease (median $\nabla(\log \kappa_s) \approx -0.12 \pm 0.03 \text{ dex kpc}^{-1}$).

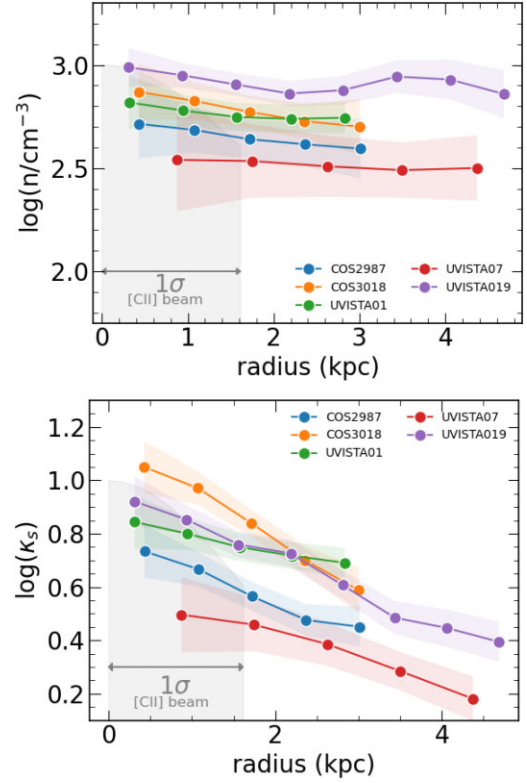


Figure B1. Radial profiles of the gas density (n , upper panel) and deviation from the KS relation (κ_s , lower panel). The grey shaded region denotes the 1σ width of the median [C II] beam. The shaded coloured regions represent the 1σ error, see Section 4 for details on the calculation.

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