



LJMU Research Online

Vallely, PJ, Prieto, JL, Stanek, KZ, Kochanek, CS, Sukhbold, T, Bersier, D, Brown, JS, Chen, P, Dong, S, Falco, E, Berlind, P, Calkins, M, Koff, RA, Kiyota, S, Brimacombe, J, Shappee, BJ, Holoien, TW-S, Thompson, TA and Stritzinger, MD

The highly luminous Type Ibn supernova ASASSN-14ms

<http://researchonline.ljmu.ac.uk/id/eprint/8458/>

Article

Citation (please note it is advisable to refer to the publisher's version if you intend to cite from this work)

Vallely, PJ, Prieto, JL, Stanek, KZ, Kochanek, CS, Sukhbold, T, Bersier, D, Brown, JS, Chen, P, Dong, S, Falco, E, Berlind, P, Calkins, M, Koff, RA, Kiyota, S, Brimacombe, J, Shappee, BJ, Holoien, TW-S, Thompson, TA and Stritzinger, MD (2017) The highly luminous Type Ibn supernova ASASSN-

LJMU has developed **LJMU Research Online** for users to access the research output of the University more effectively. Copyright © and Moral Rights for the papers on this site are retained by the individual authors and/or other copyright owners. Users may download and/or print one copy of any article(s) in LJMU Research Online to facilitate their private study or for non-commercial research. You may not engage in further distribution of the material or use it for any profit-making activities or any commercial gain.

The version presented here may differ from the published version or from the version of the record. Please see the repository URL above for details on accessing the published version and note that access may require a subscription.

For more information please contact researchonline@ljmu.ac.uk

<http://researchonline.ljmu.ac.uk/>

The highly luminous Type Ibn supernova ASASSN-14ms

P. J. Vallely,^{1★} J. L. Prieto,^{2,3} K. Z. Stanek,^{1,4} C. S. Kochanek,^{1,4} T. Sukhbold,^{1,4}
D. Bersier,⁵ J. S. Brown,¹ P. Chen,^{6,7} S. Dong,⁷ E. Falco,⁸ P. Berlind,⁸ M. Calkins,⁸
R. A. Koff,⁹ S. Kiyota,¹⁰ J. Brimacombe,¹¹ B. J. Shappee,¹² T. W.-S. Holoien,¹³
T. A. Thompson^{1,4} and M. D. Stritzinger¹⁴

¹Department of Astronomy, The Ohio State University, 140 West 18th Avenue, Columbus, OH 43210, USA

²Núcleo de Astronomía de la Facultad de Ingeniería, Universidad Diego Portales, Av. Ejército 441, Santiago, Chile

³Millennium Institute of Astrophysics, Santiago, Chile

⁴Center for Cosmology and AstroParticle Physics, The Ohio State University, 191 W. Woodruff Ave., Columbus, OH 43210, USA

⁵Astrophysics Research Institute, Liverpool John Moores University, 146 Brownlow Hill, Liverpool L3 5RF, UK

⁶Department of Astronomy, Peking University, Yi He Yuan Road 5, Hai Dian District, Beijing 100871, China

⁷Kavli Institute for Astronomy and Astrophysics, Peking University, Yi He Yuan Road 5, Hai Dian District, Beijing 100871, China

⁸Harvard-Smithsonian Center for Astrophysics, 60 Garden St., Cambridge, MA 02138, USA

⁹Antelope Hills Observatory 980 Antelope DR W Bennett, CO 80102, USA

¹⁰Variable Stars Observers League in Japan (VSOLJ), 7-1 Kitahatsutomi, Kamagaya 273-0126, Japan

¹¹Coral Towers Observatory, Cairns, Queensland 4870, Australia

¹²Institute for Astronomy, University of Hawaii, 2680 Woodlawn Drive, Honolulu, HI 96822, USA

¹³Carnegie Observatories, 813 Santa Barbara Street, Pasadena, CA 91101, USA

¹⁴Department of Physics and Astronomy, Aarhus University, Ny Munkegade 120, DK-8000 Aarhus C, Denmark

Accepted 2017 December 20. Received 2017 December 18; in original form 2017 October 30

ABSTRACT

We present photometric and spectroscopic follow-up observations of the highly luminous Type Ibn supernova ASASSN-14ms, which was discovered on UT 2014-12-26.61 at $m_V \sim 16.5$. With a peak absolute V -band magnitude brighter than -20.5 , a peak bolometric luminosity of 1.7×10^{44} erg s⁻¹, and a total radiated energy of 2.1×10^{50} erg, ASASSN-14ms is one of the most luminous Type Ibn supernovae yet discovered. In simple models, the most likely power source for this event is a combination of the radioactive decay of ⁵⁶Ni and ⁵⁶Co at late times and the interaction of supernova ejecta with the progenitor's circumstellar medium at early times, although we cannot rule out the possibility of a magnetar-powered light curve. The presence of a dense circumstellar medium is indicated by the intermediate-width He I features in the spectra. The faint ($m_g \sim 21.6$) host galaxy SDSS J130408.52+521846.4 has an oxygen abundance below $12 + \log(O/H) \lesssim 8.3$, a stellar mass of $M_* \sim 2.6 \times 10^8 M_\odot$, and a star formation rate of $SFR \sim 0.02 M_\odot \text{ yr}^{-1}$.

Key words: techniques: photometric – techniques: spectroscopic – supernovae: general – supernovae: individual: ASASSN-14ms.

1 INTRODUCTION

A particularly interesting group of supernovae (SNe) are the Type IIn and Type Ibn events. These SNe are characterized by the presence of narrow- or intermediate-width hydrogen (IIn) and helium (Ibn) emission features in their spectra (Filippenko 1997; Pastorello et al. 2008a). Their unusual spectra are probably due to interaction between the SN ejecta and a relatively dense circumstellar medium (CSM) surrounding the progenitor star. As such, they provide

valuable insight into the progenitor star by probing its mass-loss rate and composition immediately prior to explosion (Schlegel 1990; Hosseinzadeh et al. 2017).

Although there is some evidence that the historical M 31 supernova SN 1885A was an SN Ibn (Pastorello et al. 2008a), the first clear detection of narrow helium emission lines in the spectra of a hydrogen-deficient (Type Ib/c) SN was reported for SN 1999cq by Matheson et al. (2000). In the ensuing years, a number of additional events showing similar He I emission features were discovered (SN 2000er and SN 2002ao, Pastorello et al. 2008a; SN 2006jc, Foley et al. 2007), and the small class of objects were termed Type Ibn SN by Pastorello et al. (2008a). With our addition of

* E-mail: vallely.7@osu.edu

ASASSN-14ms, there are now ~ 30 such events known (Hossein-zadeh et al. 2017; Shivvers et al. 2016; Prieto et al. 2014; Smartt et al. 2016; Gorbikov et al. 2014). There are also two examples of transitional cases that show strong narrow-/intermediate-width emission features of both hydrogen and helium (SN 2005la, Pastorello et al. 2008b; SN 2011hw, Smith et al. 2012). These transitional case events seem to be rare even compared to the small numbers of SNe Ibn.

The archetypal SNe Ibn is SN 2006jc, the most well-studied member of this class of transients (Anupama et al. 2009; Hossein-zadeh et al. 2017; Pastorello et al. 2008a, 2007). Nakano et al. (2006) noted that SN 2006jc was spatially coincident with a weaker optical transient which had been observed two years earlier. One potential explanation is the explosion of a Wolf–Rayet star that had either recently undergone a luminous blue variable (LBV) like outburst or had a massive LBV binary companion (Pastorello et al. 2007). Such scenarios would be consistent with the idea that the progenitors of Type Ib/c SNe are massive stars ($M_{\text{ZAMS}} \gtrsim 8 M_{\odot}$).

Among SNe Ibn, Hossein-zadeh et al. (2017) further distinguish between the ‘P Cygni’ and ‘emission’ subclasses based on differences in the early spectra. Events in the P Cygni subclass are characterized by the presence of narrow P Cygni profiles in their spectra, while those in the emission subclass are dominated by broader emission lines. Virtually, all known SNe Ibn fall into one of these two subclasses, with approximately half of the total population falling into each classification. It is currently unclear whether these distinctions trace two different CSM configurations or a continuum of CSM properties.

Here, we describe the discovery and follow-up observations of the highly luminous Type Ibn SN ASASSN-14ms and present simple physical models for its unusual light curve. In Section 2, we describe the discovery of this transient by All-Sky Automated Survey for Supernovae (ASAS-SN), report the results of our photometric and spectroscopic follow-up observations, and also briefly discuss the spectroscopic characteristics of its faint host galaxy, SDSS J130408.52+521846.4 at $z = 0.0540$. In Section 3, we compare the photometric and spectroscopic observations of ASASSN-14ms to those of other SNe Ibn. In Section 4, we briefly describe a few simple physical light-curve models and then use these models to fit the observed light curve of ASASSN-14ms. In Section 5, we discuss our conclusions.

2 OBSERVATIONS AND ANALYSIS

2.1 Discovery

ASASSN-14ms was discovered by the ASAS-SN (Shappee et al. 2014; Holoiu et al. 2017) in images obtained on JD 2457018.11 (UT 2014-12-26.61) at J2000 RA $13^{\text{h}}04^{\text{m}}08^{\text{s}}.69$ and Dec. $+52^{\circ}18'46''.5$. These images were obtained using the quadruple 14 cm ASAS-SN telescope Brutus, which is deployed at the Haleakala station of the Las Cumbres Observatory Global Telescope Network (LCOGT; Brown et al. 2013). Its discovery was confirmed two days later by amateur astronomer S. Kiyota using a Planewave CDK 0.61 m telescope located at the Sierra Remote Observatory (Kiyota et al. 2014). Photometry obtained on JD 2457018.11 gave an apparent V -band magnitude of 16.5 mag, and the transient was not detected in images obtained on JD 2457007.10 (UT 2014-12-15.60) and earlier down to a limiting magnitude of $V \sim 17.5$ mag. We find no evidence in archival ASAS-SN photometry for a pre-SN outburst analogous to that observed for SN 2006jc. However, with

ASASSN-14ms peaking at $V \sim 16.5$ and a survey depth of $V \sim 17.5$, such an outburst would have been too faint to be detected.

ASASSN-14ms is located $\sim 1'1$ South and $\sim 0'8$ East of the centre of SDSS J130408.52+521846.4. Prior to this work, no spectra of this galaxy were publicly available. A follow-up spectrum obtained using the Multi-Object Double Spectrographs (MODS; Pogge et al. 2010) on the Large Binocular Telescope indicates that the host galaxy has a redshift of $z = 0.0540 \pm 0.0001$. Adopting $H_0 = 69.6$ and $\Omega_M = 0.286$ and assuming a flat universe yields a luminosity distance of 242 Mpc (Wright 2006). The host galaxy spectrum is discussed in greater detail in Section 2.4.

2.2 Photometry

Photometric observations were obtained over the course of 110 d following the discovery of ASASSN-14ms. The SN remained bright enough to be detected in ASAS-SN images for nearly a month after discovery. ASAS-SN images are processed in an automated pipeline using the ISIS image subtraction package (Alard & Lupton 1998; Alard 2000). Using the IRAF APPHOT package, we performed aperture photometry on the subtracted images and then calibrated the results using the AAVSO Photometric All-Sky Survey (APASS; Henden et al. 2015).

We obtained g -band images using the 1m Las Cumbres Observatory telescope deployed at McDonald Observatory, and V -band images from a 0.25 m telescope at the Antelope Hills Observatory operated by R.A. Koff. Most of the post-maximum photometry was obtained using the 2 m Liverpool Telescope (LT; Steele et al. 2004) at the Observatorio del Roque de los Muchachos, which provided $griz$ observations spanning nearly 80 d. All images were reduced after bias/dark-frame and flat-field corrections. We extract the photometry using the IRAF APPHOT package, and the photometric calibrations are performed using APASS for V band and the first data release of Panoramic Survey Telescope and Rapid Response System (Flewelling et al. 2016) for Sloan Digital Sky Survey (SDSS) g , r , i and z bands. In order to correct for the host galaxy contribution, we performed image subtraction on the LT observations using images of the host obtained long after the SN faded (2017 April). We note that for the last three epochs of photometric observations the SN is several magnitudes dimmer than the host, making the photometry somewhat suspect.

We also obtained images in the U , B , V , $UVW1$, $UVW2$ and $UVM2$ bands with *Swift*’s Ultraviolet Optical Telescope (UVOT; Roming et al. 2005). The *Swift*/UVOT photometry is extracted using a $5''.0$ aperture and a sky annulus with an inner radius of $15''.0$ and an outer radius of $30''.0$ with the UVOTSOURCE task in the HEASOFT package. The *Swift*/UVOT photometry is calibrated in the Vega magnitude system based on Breeveld et al. (2011). We note that typical offsets between the *Swift* V -band and Johnson V -band photometries are small (~ 0.05 mag; Godoy-Rivera et al. 2017), and we add a conservative estimate of the uncertainty introduced by the offset of 0.075 mag in quadrature to the measured photometric uncertainty. The ground-based photometry obtained for ASASSN-14ms is reported in Table 1, and the *Swift* photometry is reported in Table 2. The complete set of photometry is presented in Fig. 1.

With a peak absolute V -band magnitude of -20.5 ± 0.1 mag, ASASSN-14ms is consistent with being the most luminous SNe Ibn yet discovered. The only other events to attain comparably high peak absolute V -band magnitudes are SN 2015U (PSN J07285387+3349106) and iPTF15ul, at approximately -20.3 mag (Pastorello et al. 2015d; Tsvetkov, Volkov & Pavlyuk 2015) and -20.8 mag (Hossein-zadeh et al. 2017), respectively. Within

Table 1. Ground-based photometric observations.

JD	Telescope	V	g	r	i	z
2457018.1	ASAS-SN: 14cm	16.86 ± 0.14	–	–	–	–
2457019.1	ASAS-SN: 14cm	16.64 ± 0.11	–	–	–	–
2457020.906	LCOGT: 1m	–	16.471 ± 0.059	–	–	–
2457020.911	LCOGT: 1m	–	16.460 ± 0.059	–	–	–
2457031.1	ASAS-SN: 14cm	16.92 ± 0.32	–	–	–	–
2457032.1	ASAS-SN: 14cm	16.91 ± 0.25	–	–	–	–
2457035.1	ASAS-SN: 14cm	17.00 ± 0.18	–	–	–	–
2457036.1	ASAS-SN: 14cm	17.16 ± 0.22	–	–	–	–
2457038.1	ASAS-SN: 14cm	16.80 ± 0.15	–	–	–	–
2457039.6	LT: 2m	–	17.15 ± 0.063	17.17 ± 0.066	17.39 ± 0.053	17.50 ± 0.065
2457040.0	AHO: 25cm	17.181 ± 0.065	–	–	–	–
2457041.0	ASAS-SN: 14cm	17.20 ± 0.19	–	–	–	–
2457042.6	LT: 2m	–	17.42 ± 0.083	17.35 ± 0.16	17.64 ± 0.13	17.64 ± 0.085
2457044.0	ASAS-SN: 14cm	16.84 ± 0.17	–	–	–	–
2457045.6	LT: 2m	–	17.73 ± 0.098	17.58 ± 0.16	17.79 ± 0.13	17.72 ± 0.11
⋮	⋮	⋮	⋮	⋮	⋮	⋮

This table is available in its entirety in machine-readable format in the online version of the paper.

Table 2. *Swift* UVOT photometric observations.

JD	V	B	U	UVW1	UVM2	UVW2
2457023.0	16.328 ± 0.078	16.273 ± 0.048	14.943 ± 0.042	14.561 ± 0.045	14.418 ± 0.052	14.606 ± 0.042
2457023.1	16.380 ± 0.079	16.240 ± 0.047	14.906 ± 0.042	14.586 ± 0.045	14.457 ± 0.043	14.606 ± 0.042
2457025.1	16.396 ± 0.055	16.289 ± 0.039	14.956 ± 0.037	14.751 ± 0.042	14.708 ± 0.045	14.935 ± 0.040

This table is available in machine-readable format in the online version of the paper.

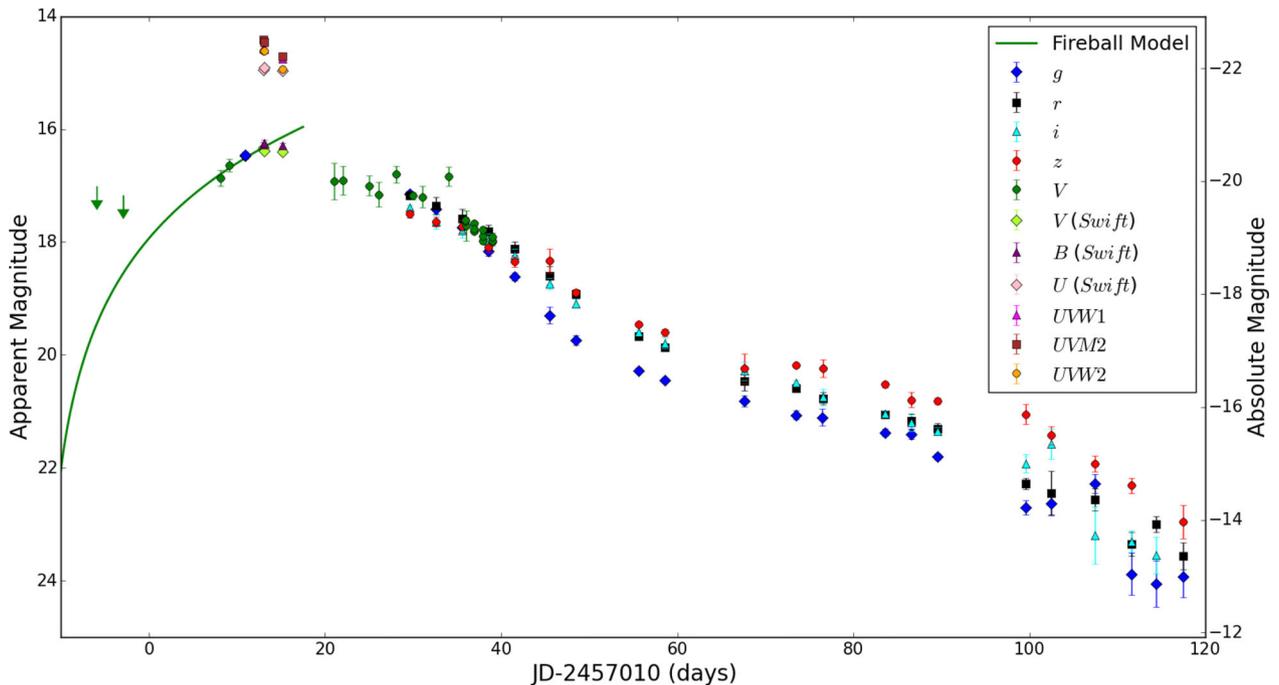


Figure 1. The complete photometry of ASASSN-14ms with the marker shapes and colours indicating the filter. Error bars are shown for all points, but can be smaller than the symbol used to represent the data. The two arrows show 3σ V-band non-detection upper limits from ASAS-SN, and the green curve shows the rising-phase expanding fireball fit discussed briefly in Section 2.2.

published uncertainties, the peak V-band magnitudes of all three events are consistent with one another, and significant uncertainties associated with estimating the host galaxy extinction for SN 2015U and iPTF15ul make it difficult to confidently determine which event is truly the most luminous at peak.

ASASSN-14ms was detected relatively close to maximum light. As a result, the rise of the light curve is poorly characterized. Nevertheless, we can fit a simple expanding fireball model to this limited data to estimate the explosion date. The key assumption of the expanding fireball model is that the luminosity of the SN should scale

as $f = \alpha \times (t - t_0)^2$ (Riess et al. 1999; Nugent et al. 2011). We fit this functional form to the V -band fluxes and obtained an estimated explosion date of JD 2456998.0 (UT 2014-12-06.5). However, since this fit uses only a handful of points near the end of the rising phase it is unclear how accurate this value is.

To estimate the date of maximum light, we calculated an uncertainty-weighted fit to the combined set of near-peak V -band observations using a third-degree polynomial. This interpolation yields JD 2457025.7 \pm 0.5 (UT 2015-01-03) for the epoch of maximum light and a maximum apparent V -band magnitude of 16.4. We will reference our spectroscopic observations to this date.

We estimate bolometric luminosities using a Markov Chain Monte Carlo routine to fit a blackbody curve to the observed spectral energy distribution (SED). We first fit the dates with sufficient filter coverage (i.e. dates with *Swift* and LT observations) to produce well-constrained temperature estimates. We then built a temperature prior by fitting a fifth-degree polynomial to these constrained temperature estimates. For photometry obtained within the range of the *Swift* and LT observations, we simply take this polynomial interpolation to be the temperature prior. For the handful of pre-maximum V -band observations without multifilter coverage, we use a constant temperature prior equal to the 17 300 K value determined for the earliest date *Swift* observations in order to avoid extrapolation. The temperature prior is applied when fitting dates for which we have observations in less than four filters.

All photometry was corrected for a Galactic foreground extinction of $E(B - V) = 0.01$ mag (Schlafly & Finkbeiner 2011) before being fit. The UV colours of the transient are relatively blue and we see no evidence of Na ID absorption in the spectra, so we make no additional correction for extinction in the host galaxy. The resulting bolometric light curve is shown with 1σ error estimates in Fig. 2. Uncertainties are shown for all points, although for many of the *Swift* and LT dates, the symbols used to plot the data are larger than the uncertainties. The bolometric light curve of ASASSN-14ms is compared to those of other SNe Ibn in Section 3.

We can integrate the bolometric luminosity over time in the SN rest frame to estimate that the total radiated energy is $(2.1 \pm 0.3) \times 10^{50}$ erg. This is a lower limit since it does not include emission outside the time span of our observations. In Section 4, we discuss refinements to this estimate based on our semi-analytic physical models.

2.3 Spectroscopy

We obtained six optical spectra of ASASSN-14ms, all after maximum light, along with an optical spectrum of the host galaxy SDSS J130408.52+521846.4. The first three spectra were obtained at phases of +7, +11, and +12 d relative to the V -band peak using the Fast Spectrograph on the (1.5m) Tillinghast Telescope (FAST; Fabricant et al. 1998; Tokarz & Roll 1997) located at the Fred Lawrence Whipple Observatory (FLWO). Two spectra were obtained using the MODS mounted on the 8.4 m Large Binocular Telescope (Pogge et al. 2010) at phases of +18 and +44 d. The sixth spectrum was obtained at a phase of +39 d using The Ohio State Multi-Object Spectrograph (OSMOS) mounted on the 2.4 m Hiltner Telescope at the MDM Observatory (Martini et al. 2011). Finally, a spectrum of the host galaxy was obtained using MODS long after the SN had faded. This spectrum is discussed further in Section 2.4. These spectroscopic observations are summarized in Table 3.

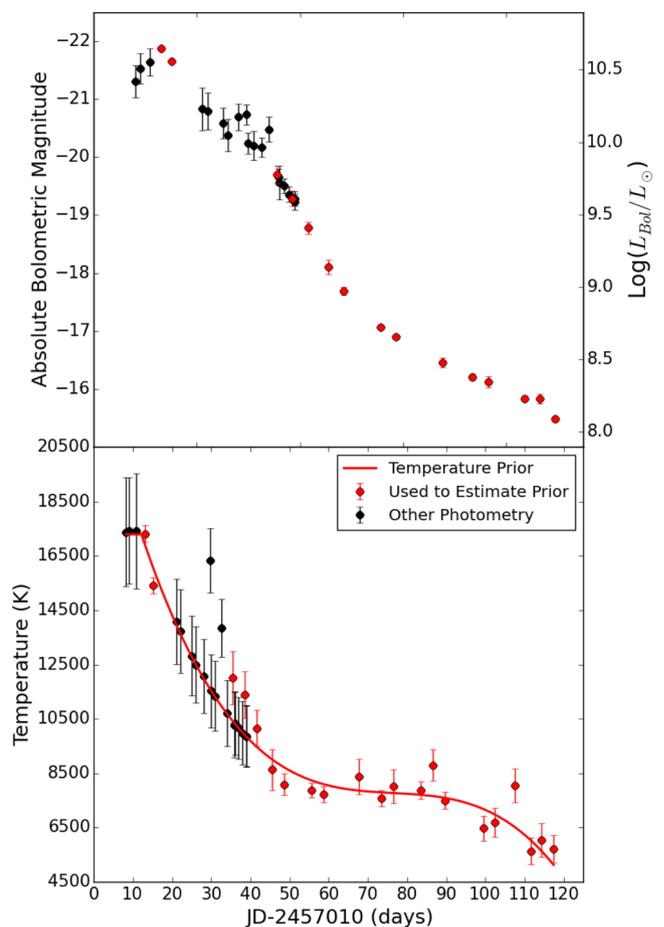


Figure 2. The bolometric light curve (top) and temperature evolution (bottom) of ASASSN-14ms. Errors are shown for all points, but can be smaller than the symbol used to represent the data. Red symbols indicate data used to build the temperature prior. The temperature prior used to calculate the bolometric light curve is shown in the lower panel. The bolometric light curve (including luminosity, temperature, and radius values) is available in the online version of the paper.

The MODS spectra (including that of the host galaxy) were reduced with a combination of the MODSCCDRED¹ PYTHON package and the MODSIDL pipeline.² We performed aperture photometry on the r' -band acquisition image to place the spectroscopic flux measurements on an absolute scale. In order to reduce the FLWO and MDM spectra we used standard techniques in IRAF to reduce the 2D images and to extract and calibrate the 1D spectra in wavelength and flux. The Supernova Identification code (SNID; Tonry & Davis 1979; Blondin & Tonry 2007), classifies all six of the ASASSN-14ms spectra as either Type Ib or Type Ic SNe. Since SNID does not include the Type Ibn subclassification this result is consistent with an SN Ibn classification because, aside from their helium emission features, SNe Ibn spectra broadly resemble those of SNe Ib/c.

The six spectra are shown in Fig. 3 along with coloured vertical lines marking the location of relevant O I, Mg II, Ca II, He I, and H Balmer lines. There is strong evidence for the presence of He I lines (the red vertical lines in Fig. 3) in all of the spectra. We see five He I lines in the +7, +11, and +12 d FAST spectra (at 3889, 4471, 5016, 5876, and 6678 Å). In the noisier +39 d OSMOS spectrum,

¹ <http://www.astronomy.ohio-state.edu/MODS/Software/modsCCDRed/>

² <http://www.astronomy.ohio-state.edu/MODS/Software/modsIDL/>

Table 3. ASASSN-14ms spectroscopic observations.

JD	Phase (d)	Telescope	Instrument	Exposure (s)	Wavelength coverage	Resolution
2457033.0	+7	FLWO-1.5m	FAST	3600	3480–7420 Å	~3 Å
2457037.0	+11	FLWO-1.5m	FAST	1800	3480–7420 Å	~3 Å
2457038.0	+12	FLWO-1.5m	FAST	1800	3480–7420 Å	~3 Å
2457043.9	+18	LBT-8.4m	MODS	1800	3400–10000 Å	~3 Å
2457065.0	+39	MDM-2.4m	OSMOS	4800	3980–6860 Å	~4 Å
2457069.8	+44	LBT-8.4m	MODS	2700	3400–10000 Å	~3 Å
2457901.7 ^a	+876	LBT-8.4m	MODS	7200	3000–9500 Å	~3 Å

^aSpectrum of the host galaxy, obtained long after ASASSN-14ms had faded.

Note that phase is calculated relative to the inferred maximum light date of JD 2457025.7.

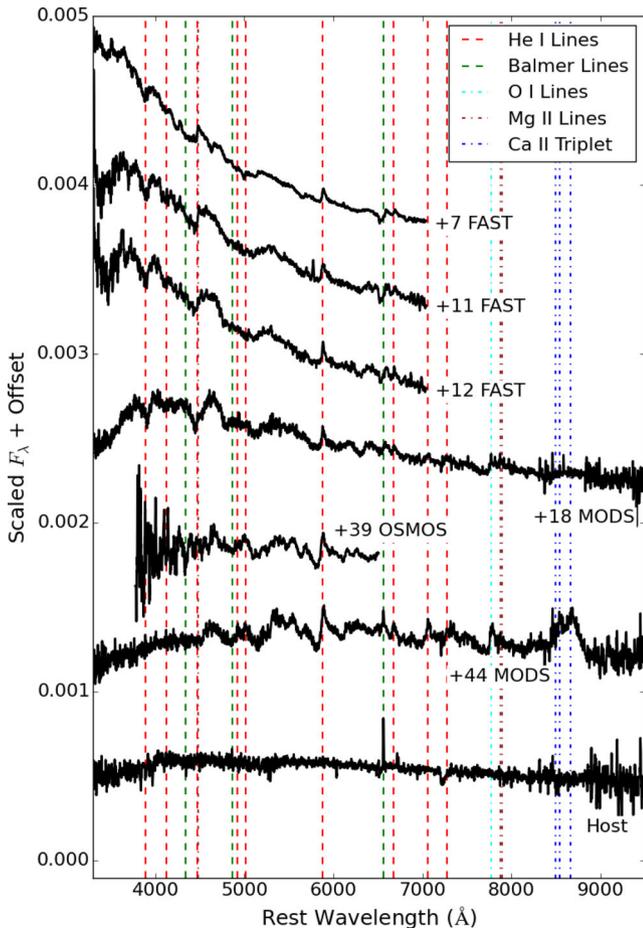


Figure 3. Rest-frame spectra of ASASSN-14ms and the spectrum of SDSS J130408.52+521846.4. No correction is applied for extinction. The locations of various atomic transitions are shown using the vertical dashed lines, and the spectra are labelled according to their phase relative to the V-band peak. The spectra have been modestly smoothed. Note the consistent presence of He I emission features throughout the spectra.

we see only the He I 5876 Å feature clearly. In the +18 and +44 d MODS spectra, we again see the 4922, 5876, 6678, and 7065 Å He I lines, but the 3889 Å feature can no longer be clearly detected. The full width at half-maximum (FWHM) of the He I emission features all lie within the range ~1000–2000 km s⁻¹, which is consistent with the narrow-/intermediate-width He I features identified in other SNe Ibn. Only weak H α signatures are seen in the spectra. This is consistent with the spectra of other SNe Ibn, as we discuss in Section 3.

The greater wavelength coverage of the two MODS spectra allows us to identify several ions not detected in the FAST and OSMOS spectra, namely Ca II, O I, and Mg II. The strength of the emission lines for these three ions is stronger in the +44 d spectrum than in the +18 d spectrum. In both MODS spectra, we see strong emission from the near-infrared Ca II triplet, as well as from the O I 7774 Å triplet. There are also emission signatures from the Mg II 7877 and 7896 Å lines, although these are not as distinct. Confident identification of Mg II is complicated by the presence of strong He I 4471 Å features, which mask the presence of the Mg II 4481 Å features. Ca II, O I, and Mg II features were also identified in the spectra of SN 2006jc (Foley et al. 2007; Anupama et al. 2009), making their detection here unsurprising.

The inferred photospheric expansion velocity, estimated by measuring the absorption minimum of P-Cygni profiles in the spectra, increases with time. From several of the strong He I features (4471, 5016, 5876, and 6678 Å), we find that the three FAST spectra (phases +7, +11, and +12 d) exhibit comparable expansion velocities of approximately 1200–1600 km s⁻¹. In the phase +18 d spectrum, we estimate the expansion velocity from the He I 4471 and 5876 Å features and find that it has increased to 2000–2300 km s⁻¹. For the +39 d OSMOS spectrum, we are limited to measuring the He I 5876 Å feature, but the 3700–4000 km s⁻¹ velocity estimate we obtain is consistent with that measured for the +44 d MODS spectrum using the O I 7774, Ca II 8498, and He I 5876 Å features. This increasing expansion velocity is likely the result of the SN ejecta shocking and accelerating the surrounding CSM.

2.4 Host galaxy SDSS J130408.52+521846.4

In order to constrain the properties of the host galaxy SDSS J130408.52+521846.4, we performed a search of archival data bases. We retrieved the archival *ugriz* model magnitudes from the Twelfth Data Release of the SDSS (Alam et al. 2015). The host was also detected in the NUV filter by the *Galaxy Evolution Explorer* (GALEX; Martin et al. 2005). The host is not significantly spatially extended on the sky, so there is no worry that the catalogued magnitudes sample drastically different spatial scales. The host is too faint to be detected in the 2MASS or WISE surveys, and there are no archival *Spitzer*, *Herschel*, *Hubble Space Telescope*, *Chandra*, or X-ray Multi-Mirror Mission (*XMM-Newton*) observations of the host galaxy. There are also no associated radio sources in the FIRST (Becker, White & Helfand 1995) or NVSS (Condon et al. 1998) catalogues.

We modelled the SDSS+GALEX host galaxy SED with the Fitting and Assessment of Synthetic Templates (FAST; Kriek et al. 2009) code. We adopted a Cardelli, Clayton & Mathis (1989) extinction law with $R_V = 3.1$ and a Galactic extinction of $A_V = 0.031$ mag

(Schlafly & Finkbeiner 2011), an exponentially declining star formation history, a Salpeter initial mass function, and the Bruzual & Charlot (2003) stellar population models. We obtained a reasonable fit ($\chi_p^2 = 0.42$) which yielded estimates of the stellar mass and star formation rates (SFRs) of $M_* = (2.6^{+0.8}_{-1.5}) \times 10^8 M_\odot$ and $\text{SFR} \sim 0.02 M_\odot \text{yr}^{-1}$.

We also obtained a late-time LBT MODS spectrum long after the transient had faded (at phase +876 d) in order to assess the spectroscopic properties of the host. The spectrum is continuum dominated, and shows only weak nebular emission features. However, we are able to cleanly detect the $H\alpha$ emission line ($\text{EW} \sim 14.8 \text{\AA}$). We determine the redshift of the host to be $z = 0.0540 \pm 0.0001$, corresponding to a luminosity distance of 242 Mpc. The $H\alpha$ luminosity ($L_{H\alpha} \sim 3.3 \times 10^5 L_\odot$) corresponds to an SFR of $\sim 9.8 \times 10^{-3} M_\odot \text{yr}^{-1}$ (Kennicutt 1998). There is a marginal detection of the $\text{N II } \lambda 6583$ line, which we use to place an approximate upper limit on the metallicity of the host galaxy. The line ratio is constrained to be $\text{N II } \lambda 6583/H\alpha \lesssim 0.075$, corresponding to an upper limit on the oxygen abundance of $12 + \log(\text{O}/\text{H}) \lesssim 8.3$ (Pettini & Pagel 2004). This places SDSS J130408.52+521846.4 near the lower end of the $12 + \log(\text{O}/\text{H})$ dex distribution for SNe Ibn host galaxies (Pastorello et al. 2015b; Taddia et al. 2015).

3 COMPARISON TO OTHER TYPE IBN SUPERNOVAE

We compare spectra of ASASSN-14ms to those of other SNe Ibn in Fig. 4. Here, the spectra are corrected for Galactic extinction using the extinction coefficient data from the Schlafly & Finkbeiner (2011) recalibration of the Schlegel, Finkbeiner & Davis (1998) dust map. The top panel shows early-phase spectra obtained approximately one week after maximum light, and the bottom panel shows later phase spectra obtained roughly 40 d after maximum light. The three other SNe Ibn presented in the early-phase comparison panel are SN 2006jc (Anupama et al. 2009), iPTF14aki (Hosseinzadeh et al. 2017), and SN 2000er (Pastorello et al. 2008a). The late phase comparison panel adds SN 2006jc (Anupama et al. 2009), SN 2002ao (Foley et al. 2007), SN 2010al (Pastorello et al. 2015a), and SN 2015G (Hosseinzadeh et al. 2017).

There is a good deal of spectral diversity in the Ibn class at early times (Hosseinzadeh et al. 2017), and ASASSN-14ms does not seem to be an exception. The presence of He I P Cygni profiles in its spectrum and the $\sim 1300 \text{ km s}^{-1}$ FWHM of their emission components places ASASSN-14ms in the P Cygni subclass of SNe Ibn, like SN 2000er, but its continuum shape seems to more closely resemble that of SN 2006jc and iPTF14aki, both of which are in the emission subclass. At late times, however, ASASSN-14ms generally resembles other SNe Ibn. At this phase, the P Cygni absorption components of events in the P Cygni subclass have weakened, and the continuum slopes have all flattened considerably. Strong emission from the Ca II triplet (the blue vertical lines in Fig. 4) as well as moderate strength O I and Mg II emission features (the cyan and brown vertical lines in Fig. 4, respectively) are also general features of late-time SNe Ibn spectra, and ASASSN-14ms exhibits them as well. The similarity of ASASSN-14ms to SN 2006jc and SN 2002ao at this phase is particularly strong, including the development of a narrow $H\alpha$ emission line. Thus, from a spectroscopic standpoint, ASASSN-14ms is not a particularly striking outlier among SNe Ibn. It is a little unusual in the early phases when the class is broadly spectroscopically heterogeneous, and it is fairly normal in the later phases when the class becomes broadly spectroscopically homogeneous.

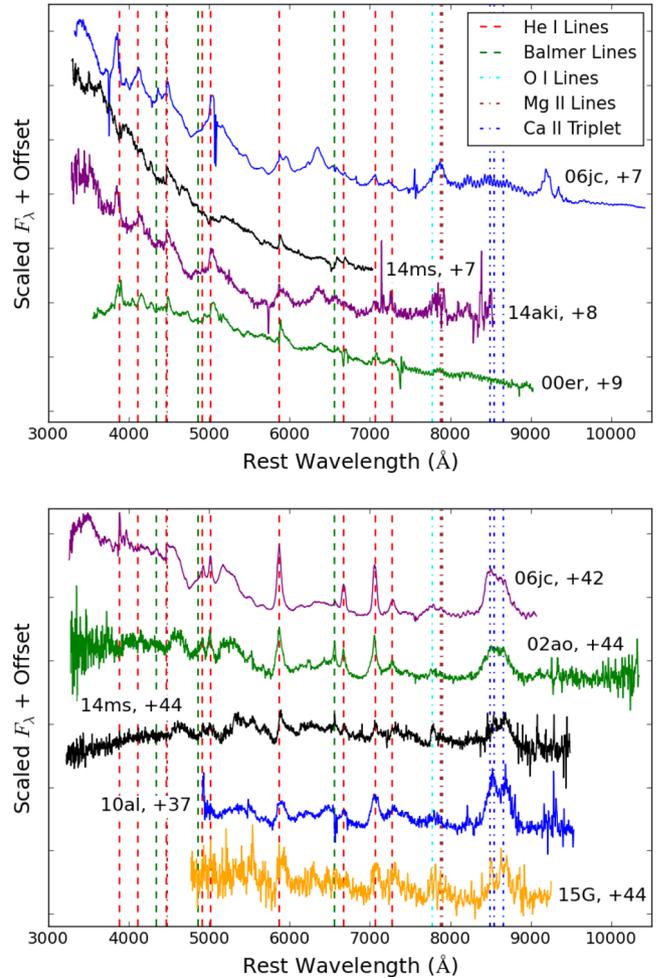


Figure 4. The upper panel shows an early (+7 d) spectrum of ASASSN-14ms compared to other Type Ibn SNe at a similarly early phase. Note that phase is defined relative to the observed date of maximum light. The lower panel shows a late (+44 d) spectrum of ASASSN-14ms as compared to other Type Ibn SNe at similar phases. Spectra for all SNe are corrected for Galactic extinction.

Its light curve, however, is notably different. The bolometric light curve of ASASSN-14ms is shown in Fig. 5 as compared to a representative sample of SNe Ibn with publicly available bolometric light curves. The four other individual events are OGLE-2012-SN-006 (Pastorello et al. 2015e), ASASSN-15ed (Pastorello et al. 2015c), SN 2010al (Moriya & Maeda 2016), and SN 2006jc (Pastorello et al. 2008a). We also include the Type Ibn light-curve template from Hosseinzadeh et al. (2017). The range of this template encompasses ~ 95 per cent of the photometric points used in their sample of 22 SNe Ibn and provides a good representation of typical SNe Ibn bolometric light-curve properties.

The ASASSN-14ms light curve is notably broader than other SNe Ibn. This is indicative of a longer diffusion time-scale and a considerably higher total radiated energy. Unlike OGLE-2012-SN-006, another unusually luminous member of this class, ASASSN-14ms does not exhibit an abnormally slow late-time decline and decays at a rate typical of other SNe Ibn. What makes ASASSN-14ms particularly unusual within the SNe Ibn class is its extreme luminosity.

The stark contrast in luminosity is clearly illustrated in Fig. 5. OGLE-2012-SN-006 and ASASSN-15ed are considered to be

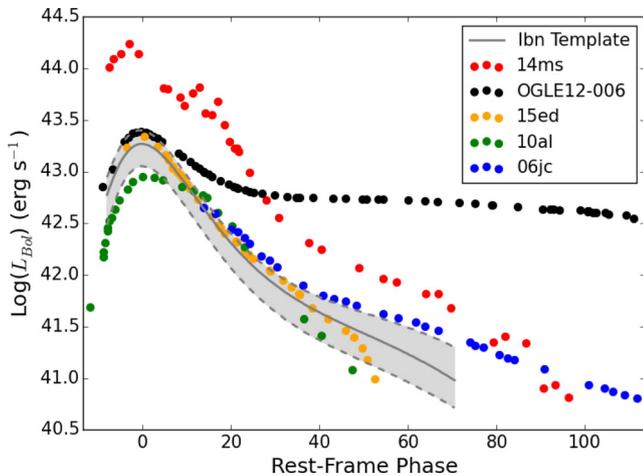


Figure 5. The rest-frame bolometric light curves of several Type Ibn SNe along with the Type Ibn template of Hosseinzadeh et al. (2017). Note the unusually high luminosity of ASASSN-14ms.

examples of unusually luminous SNe Ibn, and yet the peak luminosities for both lie within the upper limit of the Ibn template of Hosseinzadeh et al. (2017). The peak luminosity of ASASSN-14ms, on the other hand, is a factor of ~ 6 above this upper limit, and the light curve remains a factor of ~ 3 brighter than the template upper limit even at 70 d after maximum light. In terms of luminosity, ASASSN-14ms is a clear outlier among SNe Ibn.

A likely explanation is that ASASSN-14ms is hotter than other SNe Ibn. At comparable phases (~ 12 d), SN 2006jc and SN 2010al have $U - B$ colour indices of -0.7 and -0.8 , respectively (Pastorello et al. 2008b; Moriya & Maeda 2016). Although we do not have *Swift* observations of that epoch for ASASSN-14ms, making the comparison imperfect, its near-maximum $U - B$ colour index is -1.3 , indicating that it is considerably hotter than these events. Furthermore, we have g - and r -band observations for ASASSN-14ms obtained at phases (~ 15 d) comparable to those obtained for ASASSN-15ed by Pastorello et al. (2015c). At this epoch, ASASSN-14ms and ASASSN-15ed have $g - r$ colour indices of -0.03 and 0.3 , respectively, which is also consistent with ASASSN-14ms being hotter than other SNe Ibn.

4 PHYSICAL LIGHT CURVE MODELS

Semi-analytic models have proven to be a useful tool for analysing SN light curves. In this section, we compare several of these physically motivated models to ASASSN-14ms in order to probe the physics of the explosion and better understand its unusual light curve. These relatively simple models allow us to make reasonable estimates for the explosion parameters in a variety of potential light curve powering scenarios, and provide a basis for estimating the likelihood of each scenario. We consider three distinct powering sources: the radioactive decay of ^{56}Ni produced during the initial explosion, the deposition of energy into the ejecta by magnetar spin-down, and the interaction of ejecta with circumstellar material. We also consider model light curves that implement combinations of these power sources. All models are fit in the rest frame of the SN, and we exclude the final three epochs of data from our analysis because the SN is only marginally detected in the host-subtracted photometry.

We present the functional forms of the semi-analytic models we use in the following subsections. These brief descriptions are not

intended to constitute a thorough review of light-curve modelling techniques, nor are they intended to constitute derivations of the models. Rather, we present these forms to facilitate replication of our results and the implementation of similar analyses in the future. For an in-depth review of the derivations and physical assumptions of these models, please see the citations therein.

We also emphasize that we do not perform a rigorous exploration of the entire parameter space associated with these models, nor do we explore the entire range of potential power sources or power source combinations. For instance, we consider only the combination of CSM interaction and radioactive decay. In principle, any combination of these power sources could contribute to the observed light curve. The aim is only to explore physically plausible scenarios to identify parameter combinations that can explain the unusual light curve of ASASSN-14ms.

4.1 Radioactive decay model

Semi-analytic modelling techniques for the powering of SN light curves by the radioactive decay of ^{56}Ni and ^{56}Co have been in use for some time (Arnett 1982). Here, we implement a slightly modified form of the equations presented by Cano (2013) where we account for possible gamma-ray leakage with the addition of a $(1 - e^{-A_\gamma t^{-2}})$ factor. This factor is related to the gamma-ray optical depth and opacity of the ejecta via the expression $\tau_\gamma = \kappa_\gamma \rho R = A_\gamma t^{-2}$ (Chatzopoulos et al. 2013). Larger values of A_γ correspond to less gamma-ray leakage. We briefly consider radioactive decay models with complete gamma-ray trapping, but we find that such models are far more luminous (by a factor of ~ 30) at late times than the observed light curve and do not consider them further.

The luminosity of the model is determined by

$$\tau_m = \sqrt{2} \left(\frac{\kappa M_{\text{ej}}}{\beta c v_{\text{ej}}} \right)^{1/2}, \quad (1)$$

$$x = t/\tau_m, y = \tau_m/(2\tau_{\text{Ni}}), s = \tau_m/(2\tau_{\text{Co}}), \quad (2)$$

$$A(z) = 2ze^{-2zy+z^2}, B(z) = 2ze^{-2zs+z^2}, \text{ and} \quad (3)$$

$$L_{\text{RD}}(t) = M_{\text{Ni}} e^{-x^2} \left((\epsilon_{\text{Ni}} - \epsilon_{\text{Co}}) \int_0^x A(z) dz + \epsilon_{\text{Co}} \int_0^x B(z) dz \right) \times \left(1 - e^{-A_\gamma (x\tau_m)^{-2}} \right), \quad (4)$$

where $\beta \simeq 13.8$ is a constant of integration (Arnett 1982), c is the speed of light, M_{ej} is the ejecta mass, κ is the opacity of the ejecta, v_{ej} is the ejecta velocity and τ_m is the effective diffusion time. The decay times of ^{56}Ni and ^{56}Co are $\tau_{\text{Ni}} = 8.77$ d and $\tau_{\text{Co}} = 111.3$ d (Taubenberger et al. 2006; Martin 1987), and the rates of energy release for ^{56}Ni and ^{56}Co are $\epsilon_{\text{Ni}} = 3.90 \times 10^{10}$ erg $\text{s}^{-1} \text{g}^{-1}$ and $\epsilon_{\text{Co}} = 6.78 \times 10^9$ erg $\text{s}^{-1} \text{g}^{-1}$ (Sutherland & Wheeler 1984; Cappellaro et al. 1997). Like Cano (2013) and Karamehmetoglu et al. (2017), we adopt an opacity of $\kappa = 0.07 \text{ cm}^2 \text{ g}^{-1}$, which is a reasonable value for electron scattering in hydrogen-free ejecta (Chugai 2000; Chevalier 1992).

We also include a time offset parameter to account for the uncertainty in our explosion date estimates from Section 2.2. Since we cannot measure the ejecta velocity directly in our observed spectra due to the presence of the CSM, we adopt a value of $v_{\text{ej}} = 7000 \text{ km s}^{-1}$, consistent with observations of ASASSN-15ed (Pastorello et al. 2015c). By assuming this value for v_{ej} , we reduce

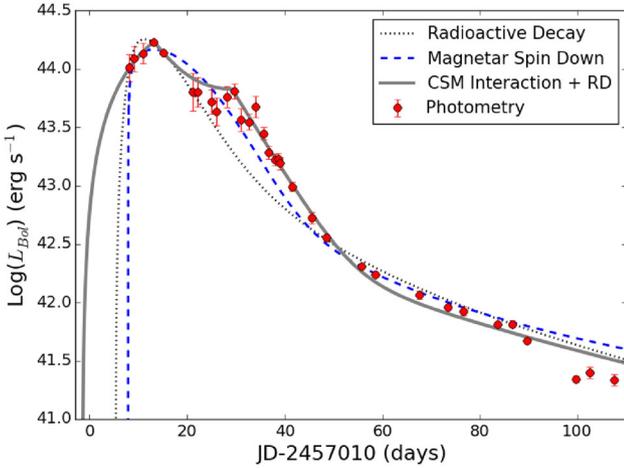


Figure 6. The bolometric light curve of ASASSN-14ms compared with several physical models. The grey curve is a model powered solely by the radioactive decay of ^{56}Ni and ^{56}Co . The blue curve is a model powered solely by the spin-down of a magnetar. The grey curve is a model powered by a combination of CSM interaction, which dominates at early times, and radioactive decay, which dominates at late times. Note that the radioactive decay model requires the unphysical condition that $M_{\text{Ni}} > M_{\text{ej}}$. The other two models utilize physically plausible parameters. Note that, while all models are fit in the rest frame of the SN, they are presented here in the observed frame. The fits did not include the last three epochs because of uncertainties in correcting for the flux of the host galaxy.

the complexity of this model to four fitting parameters: M_{ej} , M_{Ni} , A_{γ} , and the explosion date. A best fit is then obtained by calculating a grid of models in this 4D parameter space and finding the model with the lowest χ^2 value.

The result is shown by the grey curve in Fig. 6. The physical parameters are $M_{\text{ej}} = 0.85 M_{\odot}$, $M_{\text{Ni}} = 25.7 M_{\odot}$, and $A_{\gamma} = 1.66 \times 10^{12} \text{ s}^2$. The inferred explosion date of JD 2457015.3 agrees poorly with our crude estimate of JD 2456998.0 in Section 2.2. Of the physical models we explored, this pure radioactive decay model produces the worst fit. Furthermore, the only way this model can reproduce the observed high peak luminosity is with $M_{\text{Ni}} > M_{\text{ej}}$, an unphysical condition. We also considered opacities up to $0.2 \text{ cm}^2 \text{ g}^{-1}$ and values of v_{ej} up to $20\,000 \text{ km s}^{-1}$ but find that such models still cannot produce good fits with physically reasonable parameters. Thus, although we know that some ^{56}Ni must have been produced in the SN explosion, its radioactive decay can be confidently ruled out as the primary power source for the light curve.

4.2 Magnetar spin-down model

In this work, we implement the magnetar spin-down model of Inserra et al. (2013). We use simplified forms of the equations therein that have been informed by those presented by Chatzopoulos et al. (2013). Under the simplifying assumption of uniform ejecta density $E_k = 3M_{\text{ej}}v_{\text{ej}}^2/10$, where E_k is the SN’s kinetic energy, M_{ej} is the ejecta mass, and v_{ej} is the ejecta velocity. The luminosity of the model is then described by

$$\tau_m = (6/5)^{1/4} \left(\frac{\kappa}{\beta c} \right)^{1/2} M_{\text{ej}}^{3/4} E_k^{-1/4}, \quad \tau_p = 4.7 B_{14}^{-2} P_{\text{ms}}^2, \quad (5)$$

$$x = t/\tau_m, \quad \tilde{y} = \tau_m/\tau_p, \quad \text{and} \quad (6)$$

$$L_{\text{Mag}}(t) = \left(4.9 \times 10^{46} \text{ erg s}^{-1} \right) B_{14}^2 P_{\text{ms}}^{-4} e^{-x^2} \times \left(\int_0^x \frac{2z e^{-z^2}}{(1+y z)^2} dz \right). \quad (7)$$

The variables τ_m , β , and κ are the same as in the radioactive decay model, and we again adopt $\kappa = 0.07 \text{ cm}^2 \text{ g}^{-1}$. The magnetar is described by its characteristic spin-down time-scale, τ_p , the strength of its magnetic field, B_{14} , in units of 10^{14} G , and its initial rotational period, P_{ms} , in units of ms. We again include an explosion date offset parameter in our implementation of this model.

While the model fits improve with shorter initial rotation periods, we choose to limit our models to $P_{\text{ms}} = 1.0 \text{ ms}$, in order to be consistent with the theoretical limits on magnetar periods calculated by Metzger et al. (2015). For this model, we fit v_{ej} as a fitting parameter with values constrained to be between 5000 and $20\,000 \text{ km s}^{-1}$, leading to a model with four fitting parameters: M_{ej} , v_{ej} , B_{14} , and the explosion date. We again calculate a grid of models in this 4D parameter space and obtain a best fit by selecting the model with the lowest χ^2 value.

The result for this model is shown by the blue curve in Fig. 6. The physical parameters are $M_{\text{ej}} = 3.29 M_{\odot}$, $v_{\text{ej}} = 14\,000 \text{ km s}^{-1}$, $B = 17.2 \times 10^{14} \text{ G}$, and the imposed $P = 1.0 \text{ ms}$ value. The model reproduces the observed light curve fairly well, and unlike the pure radioactive decay model, none of the parameters are obviously unphysical.

Although this model produces a fit in reasonable agreement with the observed light curve, there is little other evidence to support the hypothesis that ASASSN-14ms is powered by a magnetar. Additionally, the inferred explosion date of JD 2457017.9 is in poor agreement with our crude estimate of JD 2456998.0 in Section 2.2. Furthermore, although it is not completely unphysical, a 1 ms initial spin period approaches theoretical limits, and models with considerably slower initial rotation do not match the bolometric light curve as well. As such, while we cannot rule out this model, we will focus instead on considering the contribution of CSM interaction, for which we have compelling spectroscopic evidence.

4.3 Circumstellar material interaction model

We use the CSM-interaction model of Chatzopoulos, Wheeler & Vinko (2012) and Chatzopoulos et al. (2013), which is based on the work of Chevalier (1982) and Chevalier & Fransson (1994). This particular formulation is a hybrid model combining radioactive decay plus circumstellar interactions (hence the two integrals in the expression for L_{CSM} below), but it can trivially be reduced to a purely CSM-interaction model by setting $M_{\text{Ni}} = 0$.

This model utilizes a CSM density profile of the form $\rho_{\text{CSM}} = qr^{-s}$. The exponent s of the power law is commonly set either to 0 (constant density CSM) or 2 (steady-wind CSM). The scaling factor q is defined by $q = \rho_{\text{CSM},1} R_p^s$, where R_p is the radius of the progenitor star and $\rho_{\text{CSM},1}$ is the density of the CSM immediately beyond this stellar radius.

This model assigns a double power-law density profile to the SN ejecta of the form $\rho_{\text{SN}} = g^n t^{n-3} r^{-n}$. The power-law exponent of the steeper outer component of the SN ejecta is given by n (with values of 6–14 being typical, as used by Chevalier 1982), while that of the inner component is given by δ (and is typically set to either 0 or 2).

The scaling factor g^n , the total radius of the CSM, R_{CSM} , and the fixed photospheric radius of the CSM, R_{ph} , are given by

$$g^n = \frac{1}{4\pi(\delta - n)} \frac{(2(5 - \delta)(n - 5)E_{\text{SN}})^{(n-3)/2}}{((3 - \delta)(n - 3)M_{\text{ej}})^{(n-5)/2}}, \quad (8)$$

$$R_{\text{CSM}} = \left(\frac{M_{\text{CSM}}(3 - s)}{4\pi q} + R_{\text{p}}^{3-s} \right)^{1/(3-s)}, \quad \text{and} \quad (9)$$

$$R_{\text{ph}} = \left(-\frac{2(1 - s)}{3\kappa q} + R_{\text{CSM}}^{1-s} \right)^{1/(1-s)}. \quad (10)$$

Here, E_{SN} denotes the total SN energy, M_{CSM} is the mass of circumstellar material, and M_{ej} is the ejecta mass. The values of the constants A (which differs from the gamma-ray leakage parameter A_γ), β_{F} , and β_{R} are dependent on the values of n and s and can be found in table 1 of Chevalier (1982), where $\beta_{\text{F}} = R_1/R_c$ and $\beta_{\text{R}} = R_2/R_c$. We adopt the assumptions used by Karamehmetoglu et al. (2017) when modelling the SN Ibn OGLE-2014-SN-131 and take $s = 0$ and $n = 12$. The values of A , β_{F} , and β_{R} are then $A = 0.19$, $\beta_{\text{F}} = 1.121$, and $\beta_{\text{R}} = 0.974$.

The luminosity of this model is then set by

$$F(t') = \frac{2\pi}{(n - s)^3} g^n \frac{\beta_{\text{F}}^{5-n}}{q^{n-5}} q^{\frac{n-5}{n-s}} (n - 3)^2 (n - 5) \beta_{\text{F}}^{5-s} A^{\frac{5-s}{n-s}} (t' + t_i)^\gamma, \quad (11)$$

$$R(t') = 2\pi \left(\frac{A g^n}{q} \right)^{\frac{5-n}{n-s}} \beta_{\text{R}}^{5-n} g^n \left(\frac{3 - s}{n - s} \right)^3 (t' - t_i)^\gamma, \quad \text{and} \quad (12)$$

$$\begin{aligned} L_{\text{CSM}}(t) &= \frac{M_{\text{Ni}}}{t'_0} e^{-\frac{t}{t'_0}} \int_0^t e^{\frac{t'}{t'_0}} \left[(\epsilon_{\text{Ni}} - \epsilon_{\text{Co}}) e^{-\frac{t'}{t'_{\text{Ni}}}} + \epsilon_{\text{Co}} e^{-\frac{t'}{t'_{\text{Co}}}} \right] dt' \\ &\times \left(1 - e^{-A_\gamma t'^{-2}} \right) + \frac{e^{-\frac{t}{t'_0}}}{t'_0} \int_0^t e^{\frac{t'}{t'_0}} \\ &\times \left[F(t') \cdot \theta(t_{\text{FS,BO}} - t') + R(t') \cdot \theta(t_{\text{RS,*}} - t') \right] dt'. \end{aligned} \quad (13)$$

The first integral in this expression corresponds to the pure radioactive decay model described in Section 4.1 as $L_{\text{RD}}(t)$, including the gamma-ray leakage term. The time of initial interaction between the CSM and the SN ejecta is $t_i = R_{\text{p}}/v_{\text{SN}}$, where the characteristic velocity of the SN ejecta is $v_{\text{SN}} = [10(n - 5)E_{\text{SN}}/(3(n - 3)M_{\text{ej}})]^{1/2}/x_0$. Here, x_0 is the dimensionless radius where the SN ejecta density profile transitions from the inner profile ($\delta = 2$) to the outer profile ($n = 12$), and we use $x_0 = 0.7$. $\theta(t_{\text{FS,BO}} - t')$ and $\theta(t_{\text{RS,*}} - t')$ are Heaviside theta functions, and $t_{\text{FS,*}}$ and $t_{\text{RS,*}}$ are the forward and reverse shock termination time-scales, respectively. We adopt the assumption of Chatzopoulos et al. (2012) that $t_{\text{FS,BO}}$ can be reasonably approximated by $t_{\text{FS,*}}$ when the optically thick component of the CSM mass ($M_{\text{CSM,th}}$) is used in the calculation. The shock termination times are then given by

$$M_{\text{CSM,th}} = \frac{4\pi q}{3 - s} \left(R_{\text{ph}}^{3-s} - R_{\text{p}}^{3-s} \right), \quad (14)$$

$$\xi = \frac{n - s}{(n - 3)(3 - s)}, \quad (15)$$

$$t_{\text{FS,BO}} \approx t_{\text{FS,*}} = \left| \frac{3 - s}{4\pi \beta_{\text{F}}^{3-s}} q^{\frac{3-n}{n-s}} (A g^n)^{\frac{s-3}{n-s}} \right|^\xi M_{\text{CSM,th}}^\xi, \quad \text{and} \quad (16)$$

$$t_{\text{RS,*}} = \left| \frac{v_{\text{SN}}}{\beta_{\text{R}} (A g^n / q)^{1/(n-s)}} \left(1 - \frac{(3 - n)M_{\text{ej}}}{4\pi v_{\text{SN}}^{3-n} g^n} \right)^{1/(3-n)} \right|^{\frac{n-s}{s-3}}. \quad (17)$$

As in the other models, $\beta \simeq 13.8$ is a constant of integration from Arnett (1982), c is the speed of light, and κ is the opacity. We adopt the assumption of Karamehmetoglu et al. (2017) that the opacity is $0.20 \text{ cm}^2 \text{ g}^{-1}$. The diffusion time-scale expressions used in the pure radioactive decay and magnetar spin-down models are valid only for the SN ejecta. In this model, we adopt more general expressions, necessary for considering diffusion through the CSM. The diffusion time-scale through the optically thick component of the CSM, and that through the sum of the ejecta mass and the optically thick CSM are given by

$$t_0 = \frac{\kappa M_{\text{CSM,th}}}{\beta c R_{\text{ph}}} \quad \text{and} \quad t'_0 = \frac{\kappa (M_{\text{ej}} + M_{\text{CSM,th}})}{\beta c R_{\text{ph}}}, \quad (18)$$

respectively. This model is considerably more complex than the others. As such, we do not attempt to calculate a true best fit and instead simply present an optimized manual fit. We varied the numerous model parameters until we obtained a fit that reproduces the observed light curve reasonably well with physically justifiable input parameters. Excluding the three density profile exponents, there are seven fitting parameters: M_{ej} , M_{CSM} , M_{Ni} , A_γ , E_{SN} , $\rho_{\text{CSM},1}$, and the explosion date. We generate a grid of models in this 7D parameter space around the values used to produce the initial manual fit and then select the model with the lowest χ^2 value.

This fit is shown by the grey curve in Fig. 6. The inferred explosion date of JD 2457008.5 is significantly earlier than those of the other two models and better approximates our crude estimate of JD 2456998.0 in Section 2.2. The physical parameters are $M_{\text{ej}} = 4.28 M_\odot$, $M_{\text{CSM}} = 0.51 M_\odot$, $M_{\text{Ni}} = 0.23 M_\odot$, $A_\gamma = 1.97 \times 10^{13} \text{ s}^2$, $E_{\text{SN}} = 8.75 \times 10^{50} \text{ erg}$, and $\rho_{\text{CSM},1} = 3.18 \times 10^{-13} \text{ g cm}^{-3}$. Such a large M_{Ni} would be unusual for an SN with $E_{\text{SN}} < 10^{51} \text{ erg}$, but it is not dramatically unphysical. We note that the photospheric radius of this fit ($\sim 10^{15} \text{ cm}$) is consistent with the early-phase radius estimates from our bolometric light curve (out to $\sim 40 \text{ d}$ beyond maximum light) and is broadly comparable to R_{ph} values found using this physical model to fit other highly luminous SNe (Chatzopoulos et al. 2013).

This model is able to reproduce the observed light curve extremely well across all epochs, although this may not be surprising given the number of fitting parameters. Still, the considerable spectral evidence for CSM interaction provides ample justification for its use. Considering both the spectral evidence and the quality fit of this physically reasonable model, we conclude that the light curve of ASASSN-14ms can be plausibly explained by a combination of CSM interaction (which dominates at early times) and radioactive decay (which dominates at late times).

Since this model reproduces the observed light curve well, we used it to make an estimate of the total radiated energy. Integrating this model over the rest-frame time period for which we have photometric observations yields a total radiated energy of $2.41 \times 10^{50} \text{ erg}$, which is slightly larger than our more direct estimate of $(2.1 \pm 0.3) \times 10^{50} \text{ erg}$. We can then expand this integral to calculate the total energy radiated from the time of initial explosion to 500 d later (when the SN has faded to a negligible luminosity). Doing so, we obtain $2.66 \times 10^{50} \text{ erg}$.

5 DISCUSSION AND CONCLUSIONS

We discuss the discovery of the highly luminous type Ibn SN ASASSN-14ms and present follow-up observations spanning over 100 d. This includes optical and UV photometry as well as six epochs of spectroscopy. We have also presented a spectrum of the

host galaxy SDSS J130408.52+521846.4, which previously had no publicly available spectroscopic observations.

Spectroscopically, ASASSN-14ms is a broadly normal member of the Type Ibn SN class. Its spectral features differ slightly from those of other SNe Ibn at early times, but the class as a whole is fairly heterogeneous at this stage. By late times, ASASSN-14ms settles into the typical spectroscopic behaviour of an SN Ibn. However, its light curve is broader and considerably more luminous than that of a typical SNe Ibn, prompting us to investigate potential explanations for this peculiar feature.

After fitting three semi-analytic light-curve models to the observed bolometric light curve, we conclude that it cannot be powered purely by the radioactive decay of ^{56}Ni . Such models require an unphysically large nickel mass ($M_{\text{Ni}} > M_{\text{ej}}$). We conclude that a light curve powered purely by magnetar spin-down is unlikely because such models require initial spin periods that are borderline unphysical and because we do not detect any observational signatures that indicate the formation of a magnetar. We note that the observed late-time behaviour could likely be fit comparably well by magnetar spin-down, but in the absence of other evidence for such a remnant we elected not to include this model.

A model light curve powered primarily by CSM interaction that includes the additional luminosity produced by the radioactive decay of a moderate amount of nickel ($\sim 0.25 M_{\odot}$) can reproduce the observed light curve of ASASSN-14ms quite well with physically reasonable input parameters. Although we cannot rule out a light curve powered by magnetar spin-down, the CSM interaction model is well motivated by the presence of intermediate-width He I emission features in the spectra of SNe Ibn (ASASSN-14ms included). Thus, we conclude that, although ASASSN-14ms is an unusually luminous member of the SNe Ibn class, it can plausibly be explained by CSM interaction and radioactive decay with no need to invoke more exotic sources of energy deposition.

ACKNOWLEDGEMENTS

We thank the referee for helpful comments. We thank the Las Cumbres Observatory and its staff for its continuing support of the ASAS-SN project. We also thank Swift PI, the Observation Duty Scientists, and the science planners for promptly approving and executing our observations.

ASAS-SN is supported by the Gordon and Betty Moore Foundation through grant GBMF5490 to the Ohio State University and NSF grant AST-1515927. Development of ASAS-SN has been supported by NSF grant AST-0908816, the Mt. Cuba Astronomical Foundation, the Center for Cosmology and AstroParticle Physics at the Ohio State University, the Chinese Academy of Sciences South America Center for Astronomy (CAS- SACA), the Villum Foundation, and George Skestos.

The LT is operated on the island of La Palma by Liverpool John Moores University in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias with financial support from the UK Science and Technology Facilities Council. This research uses data products produced by the OIR Telescope Data Center, supported by the Smithsonian Astrophysical Observatory.

This research uses data obtained through the Telescope Access Program, which has been funded by ‘the Strategic Priority Research Program-The Emergence of Cosmological Structures’ of the Chinese Academy of Sciences (grant no. 11 XDB09000000) and the Special Fund for Astronomy from the Ministry of Finance. This

research was made possible through the use of the APASS, funded by the Robert Martin Ayers Sciences Fund.

Support for JLP is provided in part by FONDECYT through the grant 1151445 and by the Ministry of Economy, Development, and Tourism’s Millennium Science Initiative through grant IC120009, awarded to The Millennium Institute of Astrophysics, MAS. KZS, CSK, and TAT are supported by NSF grants AST-1515876 and AST-1515927. TS is partly supported by NSF PHY-1404311 to John Beacom. PC and SD acknowledge Project 11573003 supported by NSFC. MDS is supported by a research grant (13261) from VIL LUM FONDEN.

REFERENCES

- Alam S. et al., 2015, *ApJS*, 219, 12
 Alard C., 2000, *A&AS*, 144, 363
 Alard C., Lupton R. H., 1998, *ApJ*, 503, 325
 Anupama G. C., Sahu D. K., Gurugubelli U. K., Prabhu T. P., Tominaga N., Tanaka M., Nomoto K., 2009, *MNRAS*, 392, 894
 Arnett W. D., 1982, *ApJ*, 253, 785
 Becker R. H., White R. L., Helfand D. J., 1995, *ApJ*, 450, 559
 Blondin S., Tonry J. L., 2007, *ApJ*, 666, 1024
 Breeveld A. A., Landsman W., Holland S. T., Roming P., Kuin N. P. M., Page M. J., 2011, in McEnery J. E., Racusin J. L., Gehrels N., eds, *AIP Conf. Ser.*, Vol. 1358, *Gamma Ray Bursts*. Am. Inst. Phys., New York, p. 373
 Brown T. M. et al., 2013, *PASP*, 125, 1031
 Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
 Cano Z., 2013, *MNRAS*, 434, 1098
 Cappellaro E., Mazzali P. A., Benetti S., Danziger I. J., Turatto M., della Valle M., Patat F., 1997, *A&A*, 328, 203
 Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245
 Chatzopoulos E., Wheeler J. C., Vinko J., 2012, *ApJ*, 746, 121
 Chatzopoulos E., Wheeler J. C., Vinko J., Horvath Z. L., Nagy A., 2013, *ApJ*, 773, 76
 Chevalier R. A., 1982, *ApJ*, 258, 790
 Chevalier R. A., 1992, *ApJ*, 394, 599
 Chevalier R. A., Fransson C., 1994, *ApJ*, 420, 268
 Chugai N. N., 2000, *Astron. Lett.*, 26, 797
 Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693
 Fabricant D., Cheimets P., Caldwell N., Geary J., 1998, *PASP*, 110, 79
 Filippenko A. V., 1997, *ARA&A*, 35, 309
 Flewelling H. A. et al., 2016, preprint ([arXiv:e-prints](https://arxiv.org/abs/1608.07447))
 Foley R. J., Smith N., Ganeshalingam M., Li W., Chornock R., Filippenko A. V., 2007, *ApJ*, 657, L105
 Godoy-Rivera D. et al., 2017, *MNRAS*, 466, 1428
 Gorbikov E. et al., 2014, *MNRAS*, 443, 671
 Henden A. A., Levine S., Terrell D., Welch D. L., 2015, *American Astronomical Society Meeting Abstracts*, Vol. 225, p. 336.16
 Holoien T. W.-S. et al., 2017, *MNRAS*, 464, 2672
 Hosseinzadeh G. et al., 2017, *ApJ*, 836, 158
 Inserra C. et al., 2013, *ApJ*, 770, 128
 Karamehmetoglu E. et al., 2017, *A&A*, 602, A93
 Kennicutt R. C., Jr, 1998, *ARA&A*, 36, 189
 Kiyota S. et al., 2014, *ATel*, 6853
 Kriek M., van Dokkum P. G., Labbé I., Franx M., Illingworth G. D., Marchesini D., Quadri R. F., 2009, *ApJ*, 700, 221
 Martin M. J. E., 1987, 58, 67
 Martin D. C. et al., 2005, *ApJ*, 619, L1
 Martini P. et al., 2011, *PASP*, 123, 187
 Matheson T., Filippenko A. V., Chornock R., Leonard D. C., Li W., 2000, *AJ*, 119, 2303
 Metzger B. D., Margalit B., Kasen D., Quataert E., 2015, *MNRAS*, 454, 3311
 Moriya T. J., Maeda K., 2016, *ApJ*, 824, 100

- Nakano S., Itagaki K., Puckett T., Gorelli R., 2006, Central Bureau Electronic Telegrams, 666
- Nugent P. E. et al., 2011, *Nature*, 480, 344
- Pastorello A. et al., 2007, *Nature*, 447, 829
- Pastorello A. et al., 2008a, *MNRAS*, 389, 113
- Pastorello A. et al., 2008b, *MNRAS*, 389, 131
- Pastorello A. et al., 2015a, *MNRAS*, 449, 1921
- Pastorello A. et al., 2015b, *MNRAS*, 449, 1954
- Pastorello A. et al., 2015c, *MNRAS*, 453, 3649
- Pastorello A. et al., 2015d, *MNRAS*, 454, 4293
- Pastorello A. et al., 2015e, *MNRAS*, 449, 1941
- Pettini M., Pagel B. E. J., 2004, *MNRAS*, 348, L59
- Pogge R. W. et al., 2010, in McLean I. S., Ramsay S. K., Takami H., eds, *Proc. SPIE Conf. Ser.*, Vol. 7735, *Ground-based and Airborne Instrumentation for Astronomy III*. SPIE, Bellingham, p. 77350A
- Prieto J. L. et al., 2014, *ATel*, 6293
- Riess A. G. et al., 1999, *AJ*, 118, 2675
- Roming P. W. A. et al., 2005, *Space Sci. Rev.*, 120, 95
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Schlegel E. M., 1990, *MNRAS*, 244, 269
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
- Shappee B. J. et al., 2014, *ApJ*, 788, 48
- Shivvers I. et al., 2016, *MNRAS*, 461, 3057
- Smartt S. J. et al., 2016, *ApJ*, 827, L40
- Smith N., Mauerhan J. C., Silverman J. M., Ganeshalingam M., Filippenko A. V., Cenko S. B., Clubb K. I., Kandrashoff M. T., 2012, *MNRAS*, 426, 1905
- Steele I. A. et al., 2004, in Oschmann J. M. Jr, ed., *Proc. SPIE Conf. Ser.*, Vol. 5489, *Ground-based Telescopes*. SPIE, Bellingham, p. 679
- Sutherland P. G., Wheeler J. C., 1984, *ApJ*, 280, 282
- Taddia F. et al., 2015, *A&A*, 580, A131
- Taubenberger S. et al., 2006, *MNRAS*, 371, 1459
- Tokarz S. P., Roll J., 1997, in Hunt G., Payne H., eds, *ASP Conf. Ser.*, Vol. 125, *Astronomical Data Analysis Software and Systems VI*. Astron. Soc. Pac., San Francisco, p. 140
- Tonry J., Davis M., 1979, *AJ*, 84, 1511
- Tsvetkov D. Y., Volkov I. M., Pavlyuk N. N., 2015, *Inf. Bull. Var. Stars*, 6140
- Wright E. L., 2006, *PASP*, 118, 1711

SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](#) online.

Table 1. Ground-based photometric observations.

Table 2. *Swift* UVOT photometric observations.

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.