H-ATLAS/GAMA: The nature and characteristics of optically red galaxies detected at submillimetre wavelengths

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**ABSTRACT**

We combine Herschel/SPIRE sub-millimeter (submm) observations with existing multi-wavelength data to investigate the characteristics of low redshift, optically red galaxies detected in submm bands. We select a sample of galaxies in the redshift range $0.01 \leq z \leq 0.2$, having $>5\sigma$ detections in the SPIRE 250 $\mu$m submm waveband. Sources are then divided into two sub-samples of red and blue galaxies, based on their UV-optical colours. Galaxies in the red sample account for $\approx 4.2$ per cent of the total number of sources with stellar masses $M_* \geq 10^{10} M_\odot$. Following visual classification of the red galaxies, we find that $\geq 30$ per cent of them are early-type galaxies and $\geq 40$ per cent are spirals. The colour of the red -spiral galaxies could be the result of their highly inclined orientation and/or a strong contribution of the old stellar population.

It is found that irrespective of their morphological types, red and blue sources occupy environments with more or less similar densities (i.e., the $\Sigma_5$ parameter). From the analysis of the spectral energy distributions (SEDs) of galaxies in our samples based on MAGPHYS, we find that galaxies in the red sample (of any morphological type) have dust masses similar to those in the blue sample (i.e. normal spiral/star-forming systems). However, in comparison to the red-spirals and in particular blue systems, red-ellipticals have lower mean dust-to-stellar mass ratios. Besides galaxies in the red-elliptical sample have much lower mean star-formation/specific-star-formation rates in contrast to their counterparts in the blue sample. Our results support a scenario where dust in early-type systems is likely to be of an external origin.

**Key words:** Galaxy

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**1 INTRODUCTION**

Galaxies display a wide variety of physical and observational properties. It is well known that the distribution of galaxy optical colours is bimodal, e.g. blue cloud versus the red sequence [Strateva et al. 2001, Baldry et al. 2004, Taylor et al. 2015]. The bimodality of the galaxy population exists at least out to $z \approx 1$ (e.g. Bell et al. 2004a, Tanaka et al. 2005, Cooper et al. 2006, Cucciati et al. 2006, Willmer et al. 2006). A number of different mechanisms (taking place in different environments) have been proposed for the observed bimodality of the galaxy population, including, but not limited to, galaxy merging (major and minor), galaxy strangulation and harassment, ram-pressure stripping as well as AGN feedback (e.g. Mulchaey 2000, Croton et al. 2006, Conselice 2014). Such mechanisms could regulate the observed optical colours of galaxies by influencing their key physical parameters such as star formation history (SFH), mean age of stellar populations, the amount of dust attenuation, dust geometry and metallicity (Bruzual et al. 2003, Burgarella et al. 2005, da Cunha et al. 2008, Conroy et al. 2009).

Besides, there are substantial differences between galaxy populations in the field and those in clusters and groups. According to Dressler (1980), galaxy morphology is a strong function of galaxy density, i.e. the morphology-density relation, and numerous studies since then have shown the dependence of galaxy properties on the local environment (Binggeli, Tammann, & Sandage 1987, Lewis et al. 2002, Balogh et al. 2004b, Ball et al. 2008). For example, the red population is substantially dominated by early-type galaxies and thus preferentially found in high-galaxy density environments, while blue galaxies are predominantly late-type systems and mostly found in low-galaxy density environments, i.e. the colour-density relation. Moreover, vast majority of galaxies in the blue cloud are actively forming stars while the red sequence consists mainly of passive galaxies with little or no ongoing star formation. There are also additional contributions to the red cloud from (a) heavily obscured star-forming or edge-on galaxies and (b) galaxies with passive disks, e.g. red spirals showing signs of low-level of star formation, which are known to be considerably redder and more massive than their blue/star-forming counterparts (van den Bergh 1976, Wolf et al. 2009, Masters et al. 2010, Cortese 2012a). It is noteworthy that the morphology-density and color-density relations evolve with redshift (e.g. Butcher & Oemler 1984, Poggianti et al. 2009, 2010).

Analyses of the dust attenuation in active/star-forming galaxies suggest that in contrast to passive galaxies, they are heavily affected by dust [Driver et al. 2007, Johnson et al. 2007, Wyder et al. 2007, Cortese et al. 2008, Tojeiro et al. 2009, Grootes et al. 2013]. It has been shown that the bulk of the dust in late-type galaxies is in the cold phase and as consequence emits at $>100 \mu$m, i.e. the far-infrared (FIR) and submm wavelengths. Such wavelengths are covered by the instruments on board the Herschel Space Observatory [Pilbratt et al. 2010]. Thus, the data collected by Herschel is uniquely suited to probe the dusty component, e.g. its characteristics and origin, in all type of galaxies, in particular early-type galaxies which contain significantly less dust than late-type systems.

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1 Herschel is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.
The existence of dust in early-type galaxies has been first reported from studying the absorption of stellar light (Bertola & Galletta 1987; Ebner et al. 1985; Goudreau et al. 1994) and since then several studies have been conducted in order to shed light on the quantitative dust content of early-type galaxies (Knapp et al. 1989; Temi et al. 2004, 2007; Lueu et al. 2004; Savoy et al. 2009). However, submm data provided by Herschel have enabled us to study dust properties, e.g. its total luminosity, mass and temperature in early-type galaxies in an unprecedented manner due to a better sensitivity, resolution and/or the long wavelength coverage necessary (Boselli et al. 2010; Davies et al. 2010; De Looze et al. 2010; Auld et al. 2013; Smith et al. 2012b; di Serego Alighieri et al. 2013).

Among various surveys, the Herschel Astrophysical Terahertz Large Area Survey (H-ATLAS; Eales et al. 2010) is the widest extragalactic survey undertaken in submm with Herschel. The large coverage of H-ATLAS helps to have a better statistic view of the dust content and its characteristic among galaxies spanning a broad range of luminosities, colours and morphologies. Results from Darwish et al. (2011) as part of the H-ATLAS Science Demonstration Phase (SDP) and based on the UV-optical colour classification, show that the majority of sources (≃ 95 per cent) with submm detections at low redshift (z < 0.2), are blue/star-forming galaxies with UV-optical colours NUV-r ≲ 4.5. This earlier study suggested that the submm-detected/optically-red galaxies (NUV-r > 4.5), with a contribution of ≲ 5 per cent to the total number of detections, are more likely to be star-forming galaxies and that their red colors are due to obscuration by dust.

From a stacking analysis at submm wavelengths, Bourne et al. (2012) performed a large-scale statistical study of the submm properties of optically selected galaxies (based on the rest-frame color r - g) at z < 0.35, and found that approximately 20 per cent of low-redshift galaxies in H-ATLAS are red.

In the mean time, there have been several H-ATLAS studies trying to shed light on the existence and properties of dust in early-type galaxies. For instance Rowlands et al. (2012) used data from the H-ATLAS SDP to study dust properties and star formation histories in a sample of low redshift galaxies (z < 0.5) detected at submm wavelengths. Followed by classification of their sample based on optical morphology, Rowlands et al. (2012) found that ≃ 4.1 per cent of all detections are early-type systems and that ≃ 3.8 per cent (19 out of 496) of spiral galaxies with submm detections are passive. In another study and by using samples of early-type galaxies at low redshifts (0.013 < z < 0.06), Agius et al. (2013) found that early-type galaxies with H-ATLAS detections (based on Phase 1 Version 2.0 internal release of the H-ATLAS catalogue), are not only bluer in the UV-optical colours but also are significantly brighter in NUV than their single component counterparts.

The aim of this work is to examine in more detail the nature of submm detected red galaxies using the data of H-ATLAS. The main difference between this work and those conducted by Rowlands et al. (2012) and Agius et al. (2013) is that all sources in our sample are detected in H-ATLAS and classified by means of the UV-optical colour index. Our main objectives are: to segregate intrinsically red galaxies from heavily obscured star forming galaxies, and subsequently discuss the origin and the role of the dust in passive systems. The main improvements compared to our previous study come from:

- the inclusion of complimentary wavelengths in the mid-infrared (MIR) bands
- the extraction of various physical parameters from multi-wavelengths observations of sources by means of the SED fitting.

The paper is organized as following: In Section 2 we present the data from H-ATLAS phase 1 and select a sample of low redshift galaxies, all detected with Herschel in the SPIRE 250µm submm band. In Section 3 we select sub-samples of optically blue and red galaxies and analyze their physical characteristics such as star formation activities and dust properties as inferred from fitting their spectral energy distributions. Our main finding and conclusion are given in Section 4. Throughout the paper, we assume a concordance CDM cosmology with H₀ = 70 km s⁻¹ Mpc⁻¹, Ω_m = 0.3 and Ω_Λ = 0.7.

2 DATA

We use data from the H-ATLAS Phase 1 Version 3.0 internal release which contains the IDs of >5σ SPIRE detections at 250 µm and is reduced in a similar way to the SDP data, as described by Ibar et al. (2010), Pascale et al. (2011), Rigby et al. (2011) and Smith et al. (2011). The Phase 1 ID catalogues have been produced in a similar way to Smith et al. (2011) and will be presented in Bourne et al. (2015; in prep).

Initially observed time-line data from SPIRE and PACS instruments were processed by using the Herschel Interactive Processing Environment (HIPE) based on a custom reduction scripts. High-pass filtering was then applied to the data time-lines in order to correct the thermal drift in bolometer arrays. Cross-scan time-line observations were projected by using the naive map-making method of HIPE. For point like sources, catalogue of >5σ submm fluxes were produced from the 250 µm PSF filtered map, using the MADX algorithm (Maddox et al. in prep), as described in Rigby et al. (2011). For extended sources, larger apertures were chosen such that they match the extent of the source submm emission. For each 250 µm source, corresponding 350 and 500 µm flux densities were estimated by using the 350 and 500 µm maps (noise-weighted/beam-convolved) at the source position extracted from the 250 µm map. Finally 100 and 160 µm aperture flux densities were measured following matching each 250 µm source to the nearest PACS sources within a radius of 10 arcsec. A likelihood-ratio analysis (Sutherland & Saunders 1992) was performed by Bourne et al. (2015) to match 250 µm sources to the SDSS DR7 (Abazajian et al. 2009) sources brighter than r = 22.4 mag within a 10 arcsec radius. The probability that an optical source is associated with the submm source has been used to define the reliability of an association. According to Bourne et al. (2015), objects with reliability ≥ 0.8 are considered to be true matches to submm sources.

The H-ATLAS fields are along the celestial equator centred at RA of 9h(G09), 12h(G12) and 14.5h(G15). 144 deg² out of the 161 deg² covered by H-ATLAS overlap with the Galaxy and Mass Assembly (GAMA I) survey (Driver et al. 2009, 2011). The GAMA survey re-processes and combines optical data from the Sloan Digital Sky Survey (SDSS DR6; Adelman-McCarthy et al. 2008), NIR data from the UKIRT Infrared Deep Sky Survey (UKIDSS) Large Area Survey (LAS DR4; Lawrence et al. 2007), and UV from the Galaxy Evolution Explorer (GALEX; Morrissey et al. 2005). The pre-processing of the GAMA, SDSS and UKIDSS archive data is described in detail in Hill et al. (2011). For all galaxies with r < 19.4 mag in G09 and G15 as well as r < 19.8 mag in G12, redshifts have been measured using the Anglo Australian Telescope.
and for brighter galaxies, redshift estimates are taken from other existing redshift surveys such as SDSS, the 2dF Galaxy Redshift Survey (2dFGRS) and the Millennium Galaxy Catalogue (MGC; Liske et al. 2013; Driver et al. 2005). Furthermore, the GAMA-WISE (the Wide-field Infrared Survey Explorer; Wright et al. 2010) catalogue adds coverage in four MIR bands at 3.4\,\mu m, 4.6\,\mu m, 12\,\mu m and 22\,\mu m (Cluver et al. 2014).

In summary, we have at our disposal UV, optical and MIR data as well as redshift estimates for the submm galaxies within the H-ATLAS/GAMA-overlapping area where all submm selected sources in our sample satisfy the following criteria:

- They all have >5\,\sigma submm detected at SPIRE 250\,\mu m.
- They fall within the redshift range 0.01 < z < 0.2. We only select objects with a sufficiently reliable spectroscopic determination (i.e. $r_Q$ > 3; Driver et al. 2011).
- All submm galaxies have a reliability parameter (reliability $>0.8$) of being associated with an optical counterpart in the SDSS $r$-band catalogue, for which multi-wavelength photometry is available. As such, in addition to the 250\,\mu m emission, all sources (7131 objects) have corresponding fluxes (all corrected for Galactic extinction) via aperture matched photometry in other bands ranging from UV to MIR.
- Since a crucial aspect of our selection of red galaxies is based on the UV-optical (NUV−$r$) colour, we remove from our sample those galaxies for which their NUV fluxes as estimated in GAMA, differ by more than >0.5 magnitude from those retrieved through GALEX Gr6 Data Release based on the All-Sky Imaging survey (AIS) data products (NUV depth $\sim$ 20.8 mag). In addition, all selected sources have NUV magnitude errors, as provided by GALEX-Gr6, which are $\leq 0.2$ mag. This guarantees that all sources in our sample have enough signal-to-noise ratio in UV. The above constraints on UV fluxes, reduces our sample to 4016 sources.
- Finally since the physical parameters inferred for each galaxy are based on SED fitting techniques, an extra criterion has been applied in order to exclude sources (234 in total) with poor quality SED fits (see Sec. 3.3).

After applying these selection criteria, we find 3782 galaxies with detections in at least NUV + u, g, r, i, z and 250\,\mu m bands. Distributions of the SDSS $r$-band and NUV magnitudes for all galaxies as well as those qualified to be included for the subsequent data analysis are shown in panels of Fig. 1. According to the first panel, approximately $\approx 13$ per cent of the initial submm sources were excluded following the requirement of a UV detection for inclusion in the sample. But that does not seem to exclude systematically any particular type of sources as a Kolmogorov-Smirnov test (KS test) suggests a $\geq 70.0$ per cent probability that the distribution of sources detected at 250\,\mu m is similar to the one being observed simultaneously in the 250\,\mu m+NUV bands. However by limiting errors in the NUV band to $\lesssim 0.2$ mag, more sources ($\approx 31$ per cent) are excluded in particular faint objects in the NUV band.

A subset of sources have also detections in GALEX FUV, PACS (100\,\mu m, 160\,\mu m) and SPIRE (350\,\mu m, 500\,\mu m) submm bands. WISE data are available and recently have been cross-matched, with extended sources from WISE accounted for correctly, for all GAMA fields. Yet at the time of analysing galaxy SEDs in this work, WISE data were only available for the G12 and G15 fields. Thus, out of the 3782 sources, 2622 ($\approx 70$ per cent) have also aperture-matched WISE-MIR data.

3 ANALYSIS

3.1 Selection of intrinsically red objects

Though the vast majority of galaxies at low redshift with submm detection are star-forming and optically blue, a small fraction of them are red in optical bands (e.g. $u-r, g-r$). We separate blue and red galaxies in the sample using the UV-optical index. This is more robust than optical colour indices as it is more sensitive to recent star-formation activity (e.g., Kaviraj et al. 2007). Daruish et al. (2011) separate red and blue galaxies in the H-ATLAS sample at NUV−$r$ = 4.5, estimated through fitting a double Gaussian to the NUV−$r$ colour distribution of galaxies, with redshifts 0.01 $\leq z$ $\leq 0.2$ (i.e. similar to the present work), in the H-ATLAS SDP data. Hence any source with NUV−$r$ $\geq 4.5$ mag is considered as red, while blue objects are those with NUV−$r$ $< 4.5$ mag. As Fig. 1 shows, the majority of the red galaxies in our sample have apparent r-band magnitudes $\lesssim 17.5$ mag and NUV magnitudes $\gtrsim 19.0$ mag.

3.1.1 Contamination by radio AGN

In order to ensure that none of the submm emission has been contaminated by synchrotron emission from radio-jets hosted by active galactic nuclei (AGNs), we find and exclude radio AGN as follows. We cross matched the SDSS position of our sources with those from the full, unfiltered radio-source catalogue of [Virdee et al. 2012]. The radio catalogue consists of all sources detected in the H-ATLAS Phase 1 field by the NRAO VLA Sky Survey (NVSS; Condon et al. 1998) and, as such, contains 7823 sources. The outcome is 117 matches having separations of < $1.0$ arcsec. In order to determine whether the radio emission was consistent with the presence of a radio-loud AGN, we calculated $q_{250}$, defined as:

$$q_{250} = \log_{10}\left(\frac{S_{250}}{S_{1.4}}\right),$$

where $S_{250}$ and $S_{1.4}$ are fluxes at 250\,\mu m and 1.4-GHz for all matched sources respectively. If $q_{250}$ < 1.4 then part of the radio-emission is due to AGN activity (Jarvis et al. 2010). Conservatively, we exclude any source which satisfies this criterion in order to ensure none of the submm emission may be contaminated by radio AGN activity. Out of 117 sources with radio counterparts, only 13 sources (1 red and 12 blue galaxies) have $q_{250}$ < 1.4 and are thus excluded from the subsequent analysis.

3.1.2 Morphology of the red galaxies

The SDSS postage-stamp images of all red sources together with their SEDs (inferred as described in Sec. 3.3) are presented in Appendix A.

The morphology of all 117 galaxies were examined from their SDSS $r$-band images, following independent visual inspection by three team members. Galaxies were classified into three categories of elliptical (E), spiral (S) and uncertain (U). The number of sources in each morphological type is 37, 48 and 32 for the E, S and U galaxies respectively (see Fig. 2). Many of sources classified as U are too small in the SDSS images to judge their morphology and can be of any type, i.e. spiral, elliptical or merging galaxies.

In order to test the validity of this morphological classification, we compared our classification to an independent morphological classification based on the S´ersic index $n$ which we obtained from the SDSS DR7 galaxy catalogue (Simard et al. 2011). Different studies have adopted different thresholds of the S´ersic index
above/below which a galaxy is considered as early/late type. For instance, [Kavindranath et al. (2004)] adopts $n = 2.0$ to divide their sample into early and late types though Sérsic indices of $n > 2.5$ have been also used to describe early-type sources (e.g., [Bell et al. 2004b], [Barden et al. 2005]).

Fig. 3 (panel a) displays the distributions of Sérsic indices for all galaxies in our sample, i.e. the blue sample as well as the morphologically classified red galaxies. From this figure, it is clear that the distribution of Sérsic indices for the red -E sample peaks around $\approx 4$. This is larger than those estimated for the S galaxies (either blue or red ). The Sérsic index distribution of the red -U galaxies lies somewhat between those of the S and E samples.

An inspection of the ellipticity parameter $e$ of all galaxies in the sample (Fig. 3 panel b) reveals that, not surprisingly, in red sources of type S, $e \geq 0.5$ whereas in red galaxies of type E, $e \lesssim 0.5$. In fact the disk structure is extremely pronounced in highly inclined spiral galaxies and therefore the majority of galaxies in the S category are those having larger ellipticities. This is better shown in Fig. 3 where histograms of galaxy inclination angles ($i$) for blue , red -S, red -E as well as red +S+U samples are plotted. Inclinations are determined from the relation

$$
\cos^2 i = \left(\frac{b/a}{p}\right)^2 \left(1 - p^2\right)^{-1}
$$

in which $p$ is the ratio of the smallest to the largest axis of an oblate spheroid of rotation. We assume $p = 0.20$ which is an appropriate value to use for the intrinsic flattening of the distribution of the light of galactic spheroids (e.g., [Van den Bergh 1988]).

Unlike blue and red -E galaxies, the majority of red -S galaxies are highly inclined. Note that, even in the combined red -U + red -S sample, there is still and excess of galaxies with relatively large inclination angles in comparison to the blue and red -E samples.

To illustrate this, we show in Fig. 3 the distribution of inclination angles as expected from a random sampling. The observed difference between the distribution of red -S+U galaxies in comparison to a sample of simulated inclinations, suggests that the fraction of highly inclined systems in red -S+U sample is more than one would expect for a random distribution. This shows that the inclination angle play a non-negligible role in the observed red colour of red -S systems.

The main conclusion is that the red -E sample consists of intrinsically red objects while the red -S sample contains galaxies where inclination could be a dominant factor in determining the observed red optical colours. Although these inclined sources are not the main interest of this paper, we do discuss some of their ensemble properties in Sec. 3.5.1.

### 3.2 Environmental density of red galaxies

In order to explore the environmental density of red galaxies and see if it plays an important role in shaping their observed proper-
Distributions of morphology related parameters in all blue (thick solid line) and red sources. E (red dashed line), S (green dotted line) and U (cyan line) labels represent the morphology of individual red source as explained in Sec. 3.1.2 Each histogram is normalized by its integral. Panels represent distributions of galaxy (a) Sersic index, (b) ellipticity and (c) inclination angle. In addition, the ‘black dotted line’ and ‘gray filled histogram’ in panel c represent the distribution of red-S+R galaxies and random distribution of inclination angles respectively.

Figure 2. Percentage of each morphological type in the sample of 117 red galaxies (see Fig. 2). Labels represent elliptical (E), spiral (S) and undefined (U) galaxies.

Figure 3. Distributions of morphology related parameters in all blue (thick solid line) and red sources. E (red dashed line), S (green dotted line) and U (cyan line) labels represent the morphology of individual red source as explained in Sec. 3.1.2 Each histogram is normalized by its integral. Panels represent distributions of galaxy (a) Sersic index, (b) ellipticity and (c) inclination angle. In addition, the ‘black dotted line’ and ‘gray filled histogram’ in panel c represent the distribution of red-S+R galaxies and random distribution of inclination angles respectively.

3.3 UV-to-Submm SED fitting

We derive the basic properties of galaxies by fitting their SEDs which makes use of the data (§2) going from the NUV up to all available Herschel bands. The SED of each galaxy is fitted using MAGPHYS (Multi-wavelength Analysis of Galaxy Physical Properties; da Cunha et al. 2008). MAGPHYS infers the galactic properties by matching the observed SED with a large library of calculated SEDs. These templates are constructed by considering the spectral evolution of stellar populations that are born with a Chabrier (2003) initial mass function (IMF) in combination with infrared dust spectral libraries as described in da Cunha et al. (2008). The model assumes that the energy from UV-optical radiation emitted by the stellar populations is absorbed by dust and re-radiated in the FIR. It uses also the two-component dust model of Charlot & Fall (2000) in order to account for the attenuation of starlight by dust.

The model also accounts for the enhanced attenuation of stellar radiation for stars located in star forming regions in comparison to older stars found elsewhere within the galaxy.

As the MAGPHYS analysis is based on AB magnitudes, all

$\Sigma_b (\text{Mpc}^{-2}) = \frac{5}{\pi d_5^2}$,
characteristics of red submillimetre galaxies

available photometry (aperture matched) has been converted to the AB magnitude system before estimating their associated fluxes in units of Jansky (Jy). Additional errors have been added to non-submm fluxes before running MAGPHYS to account for the total flux measurements and calibrations between the different surveys. These include adding 10 per cent of the flux values in quadrature for all optical-NIR bands and 20 per cent for the UV bands. For each output parameter, MAGPHYS produces a probability density function (PDF), in addition to the median value of each PDF. These include adding 10 per cent of the flux values in quadrature for all optical-NIR bands and 20 per cent for the UV bands. For each output parameter, MAGPHYS produces a probability density function (PDF), in addition to the median value of each PDF. The 16th and 84th percentiles of the PDF have been considered as a measure of the uncertainty.

Smith et al. (2012a) showed that it is insufficient to identify bad SED fits based on a simple $\chi^2$ threshold, instead deriving a threshold which depends on the number of bands of photometry available, above which there is < 1 per cent chance that the photometry is consistent with the MAGPHYS model. Sources exceeding this varying threshold are identified as bad fits, and excluded from the subsequent analysis. We use the H-ATLAS SED fits over the entire phase 1 area, derived using the same method as in Smith et al. (2012a), with updated PACS coverage and including data from WISE.

For the purpose of our study, we have focused on a number of galactic parameters that are inferred by fitting the observed SEDs with MAGPHYS. These are: the galactic stellar mass ($M_\ast$), the dust mass ($M_D$), the star formation rate (SFR), and the fraction of total dust luminosity contributed by the diffuse interstellar medium ($f_D$) ($0 \leq f_D \leq 1.0$). Large values of $f_D$ indicate that dust is heated by the old stellar populations while lower values suggest that ongoing star formation has a more prominent role in heating the dust. An example of a SED fit for a submm source in our red sample is shown in Fig. 5. We find that the distribution of $\chi^2$ in our sources, does not show any correlation with galaxy NUV−r colour indices. It is worth mentioning that the comparison of the results from MAGPHYS, with and without the MIR constraints from WISE, shows that including the WISE data modifies the output results from MAGPHYS. The inclusion of WISE data improves the fits of the SEDs and provides better estimates of some of the parameter, and notably of the SFR. For this reason, we include in the following sections only those galaxies for which WISE data are available (e.g. $\approx 2/3$ of the main sample). This in turn, reduces the size of our sample from 3782 to 2622 sources with 78 having NUV−r $\geq 4.5$ mag and therefore are red.

Figure 4 displays the mass distribution of galaxies in the blue and red samples (in different categories). In our sample, ~73 per cent of blue sources have stellar masses log($M_\ast/M_\odot$) $\geq$ 10.0, while the same number for the red galaxies is $\approx 97$ per cent, accounting for $\approx 4.2$ per cent of the total number of sources with log($M_\ast/M_\odot$) $\geq$ 10.0. As expected, bins associated to largest stellar masses are occupied by the red-E galaxies (see Table 1).

3.4 Dust properties

It is important to compare the inferred parameters derived from MagPhys to other determinations. We compare the estimated dust-to-stellar mass ratio ($M_D/M_\ast$) for all sources as computed by MAGPHYS to those derived for a sample of $\approx 300$ nearby galaxies from the HRS (Herschel Reference Survey. Cortese et al. 2012b). The total dust mass of a given galaxy as estimated by MAGPHYS is the sum of the three components which includes the mass contributed by dust in thermal equilibrium in stellar birth clouds, as well as warm and cold dust components in the ambient interstellar medium (da Cunha et al. 2008).

Figure 6 displays the distribution of $M_D/M_\ast$ inferred from MAGPHYS for our sample against NUV−r for all red and blue sources. Overlaid are the $M_D/M_\ast$ estimates from the HRS using all SPIRE bands. For HRS non-detections (triangles), the submm upper-limit fluxes have been determined assuming a 3σ signal over
a circular aperture of radius 0.3×, 0.8× and 1.4× of the optical radius for the HRS E, S0 and spirals, respectively.

Note that in determining dust masses \(M_D\), both MAGPHYS and Cortese et al. (2012b) adopt a dust emissivity index \(\beta = 2.0\) for cold dust but different dust mass absorption coefficients \(\kappa_\nu\). Cortese et al. (2012b) use a dust mass absorption coefficient \(\kappa_{350}\) of 0.192 m\(^2\)kg\(^{-1}\) at 350 \(\mu\)m whereas da Cunha et al. (2008) assume \(\kappa_{850} = 0.077\) m\(^2\)kg\(^{-1}\) at 850 \(\mu\)m. Given the scaling relations \(M_D \propto \kappa_\nu^\beta\) and \(\kappa_\nu \propto \nu^{-\beta}\), one finds that \(\kappa_{850}\) in MAGPHYS can be scaled down (assuming \(\beta = 2.0\)) to 0.45 m\(^2\)kg\(^{-1}\) at 350 \(\mu\)m and that dust masses as measured by Cortese et al. (2012b) are \(\approx 2.36\) times larger than those estimated by MAGPHYS. Thus in Fig. 7 the HRS sample are scaled down for \(\approx 0.37\) dex to account for the differences between the two measurements of dust masses.

It can be seen that the \(M_D/M_*\) ratios for both the blue or red galaxies agrees reasonably well with estimates from the HRS detected objects. Furthermore, the red sources of type E exhibit, on average, \(M_D/M_*\) ratios that are noticeably lower than those of blue galaxies. This is even more clear in the right panel of Fig. 8 which displays the distributions \(M_D/M_*\) in all sources. The mean values as summarized in Table. 1 suggests that the red -E objects have values of the dust-to-stellar masses that are approximately an order of magnitude lower than those in the blue sources. This is partly because the red -Es have high stellar masses but as is visible in the left panel of Fig. 8 they also have a lower dust content in comparison to the red -S and blue systems. Note that the distribution of specific dust mass of the red -S galaxies does not match the distribution of the blue star forming galaxies. We will discuss this further in Sec. 3.5.

### 3.5 Star formation rates

In Fig. 9 we compare the MAGPHYS derived values of the star formation rates (SFRs) to those estimated based on the spectral analysis of the H\(\alpha\) lines using the Second GAMA Data Release (GAMA-DR2) catalogues (Wijesinghe et al. 2012; Hopkins et al. 2013; Gunawardhana et al. 2013; Liske et al. 2015). Galaxy SFRs in GAMA-DR2 are determined from the Kennicutt (1998) relation and based on the total aperture-corrected H\(\alpha\) luminosities observed through fibre spectroscopy. The \(r\)-band absolute magnitude of each galaxy have been used in order to correct for the Aperture and therefore recovering the total H\(\alpha\) luminosities (Hopkins et al. 2003; Gunawardhana et al. 2011). Dust corrections...
were estimated for each galaxy from the observed Balmer decrement. Finally stellar absorption corrections were applied to both Hα and Hβ fluxes which together with the Hα equivalent width (EW) allow to calculate the total aperture-corrected Hα luminosities as described in detail in Hopkins et al. (2003).

We find a strong correlation between the two estimates of SFRs such that (SFRs are in units of M⊙yr⁻¹)

\[
\log \text{SFR}_{\text{Magphys}} = 1.22 \pm 0.02 \times \log \text{SFR}_{\text{GamaDR2}} - 0.35.
\]

The distribution of SFR related parameters are displayed in Fig. 10. The first two panels, show the SFR and the specific star formation rate (SSFR) of blue and red galaxies. The mean value of the SFR in the red -E galaxies is an order of magnitude lower than in the blue galaxies with SFRblue/SFRred-E ≈13 (SFRred-S/SFRred-E ≈2.5 ; see also Table 1).

The difference between the two samples is even more pronounced when considering SFR normalized by galaxy’s stellar mass M∗ such that SSFRblue/SSFRred-E ≈100 (SSFRred-S/SSFRred-E ≈4). For both the SFR and the SSFR, the values estimated for the red -S type sources and the galaxies with uncertain morphology, lay between the red -E galaxies and the blue control sample. In comparison, Rowlands et al. (2012) (i.e. Table C1) measure ~9.99±0.03 and ~10.85±0.14 for SSFR in samples of ‘H-ATLAS spiral’ and ‘H-ATLAS elliptical’ galaxies respectively.

Figure 10 shows the normalized distributions of fμ in the blue and red populations. The red -E galaxies have an average fμ~0.92, well above the mean (~0.55) of the blue galaxies. This indicates that while about half of the observed FIR emission observed in the blue galaxies comes from dust in birth clouds, the FIR of red -E galaxies is dominated by dust in the diffuse interstellar medium (ISM). We note that the average derived fμ for the red -S systems is significantly higher than for the blue control sample and only slightly lower than for the sample of the red -E galaxies.

### 3.5.1 On the derived properties of the red -S sample

Even though the red -S galaxies are not the prime focus of this paper, this sample does display some interesting characteristics that are worth commenting on briefly. As can be derived from figures 9, 8 and 10 the deduced properties of the red -S galaxies do not match the blue galaxy properties. The red -S galaxies appear intermediate between the red -E and the blue galaxies in stellar mass, SFR and specific dust mass. This offset is primarily driven by the higher derived stellar masses and the correspondingly lower SFR. This is contrary to what one would expected if the red colours of the edge-on galaxies are only due to their high inclination.

Inclination does play a significant role in defining this sam-
ple, as can be concluded from Fig. [11]. We show in this figure the inclination of the blue + red -S for the stellar masses above $\log (M_\ast / M_\odot) \approx 10.0$, i.e. the range of stellar masses of interest. There is a definite trend of the median inclination against observed optical redness and in particular the very reddest sources are almost exclusively very inclined sources.

We see two main interpretations — which could be at play simultaneously — that could explain these characteristics of the red -S sample.

(i) High inclination is a necessary, but not sufficient condition for a star-forming disk galaxy to be submm detected and very optically red. In this case the red colour would apparently select preferentially the more massive disk galaxies. Perhaps the less massive disk galaxies have enough star formation in their periphery of their disks — which would not be strongly obscured, even in the case of strong inclination — to exhibit a blue-ish optical colour. Alternatively, the red colour of those massive disks could be a direct results of a dominant old stellar population.

(ii) The galaxy parameters, derived from MAGPHYS, of the very inclined and dusty sources are systematically biased to higher stellar masses and less star formation. This is in line with the finding of da Cunha et al. (2010). These authors find that the derived SFR for edge-on galaxies is about a factor $3 \times (\approx 0.48 \text{ dex})$ below their face-on counterparts. They also find that this effect is also responsible for the lower dust masses (or dust luminosities) and higher $f_\mu$ estimated for edge-on in comparison to face-on galaxies. The amplitude of this effect is insufficient to directly explain the difference we find between the blue sample and the red -S sample. Note however that da Cunha et al. (2010) describe the effect on an inclined sample of galaxies while the red -S sample is selected to have only galaxies with very red colours. The inclined sample contains galaxies with varying degrees of hidden star-formation whereas the red -S sample contains only galaxies with very obscured star-formation. We thus would expect to find a larger offset of the derived parameters in the red -S sample than in the inclined sample.

Clearly this red disk population of nearby galaxies deserves further attention in a dedicated study.

### 3.6 Dust mass correlations with galactic properties

We show in Fig. [12] correlation plots of the derived dust mass versus a number of key parameters ($M_\ast$, SFR, and $f_\mu$) in the red -E and blue galaxies. These parameters have been chosen to elucidate the possible origin and role of the dust in the red -E galaxies. The first conclusion that can be drawn from the perusal of these diagrams is that the red -E galaxies clearly occupy a different parameter space from blue spiral galaxies.

Panel [12] shows a very different behaviour of the $M_\ast$ as a function of $M_\ast$, for the blue galaxies and the red -E sample. The blue sample shows a roughly linear correlation (with scatter) between the dust reservoir and the $M_\ast$. This relation is expected due to the $M_\ast$ -SFR relation for normal galaxies, if the $M_\ast$ is measuring the reservoir available for star formation. The red -E sample exhibits a totally different behaviour apart from being located in a distinctly different part of this diagram. While the host galaxies are all -- with one outlier -- of very similar mass ($\approx 10^{11} M_\odot$) their dust content spans more than two orders of magnitude. This complete decorrelation of stellar mass and dust content argues against a stellar origin (e.g. Cortese et al. 2012) for the dust in those galaxies. While for blue galaxies the dust mass increases with stellar mass, the dust masses found for the red -E span $\approx 2$ order of magnitudes for stellar masses that are roughly constant at $\approx 10^{11} M_\odot$ (see Table [3]).

In panel [12] we show that there is a moderate correlation in the red -E galaxies between the derived SFR and $M_\ast$ with a similar slope but offset from the blue sequence. We interpret the existence of this correlation as an indication that the star formation is probably taking place in the cold gas associated with the dust.

The observed offset between the blue control sample and the red -E sample implies that the same amount of dust in the red -E galaxies is associated with about an order of magnitude less star formation. This could be an indication that the physical state of the cold ISM phase in the red -E galaxies is significantly different perhaps due to the very different environment in which the cold gas is embedded. This interpretation is corroborated by panel [12] where we show that indeed the MAGPHYS derived fraction of the dust heating due to the interstellar radiation field, i.e. $f_\mu$ is much higher in the red -E galaxies than their blue counterparts.

### 3.7 The origin of dust in red -E

In the classical definition of galactic types, ellipticals were classified as devoid of gas and dust (Hubble 1926, de Vaucouleurs 1959, Sandage 1961). In the subsequent years, dust emission in Ellipticals has been detected from the ground (Hawarden et al. 1981, Sadler & Gerhard 1985, Sparks et al. 1985, Kormendy & Stauffer 1987, Ebneter et al. 1988, Pandey et al. 2001) and from space using the Infrared Astronomical Satellite (IRAS) (Jura et al. 1987, Knapp et al. 1989) and the Spitzer Space Telescope (Rocca-Volmerange et al. 2007). Dust lanes were observed early on along the minor axis of ellipticals (Bertola & Galleta 1978). When in some ellipticals the dust lanes and stars were observed to rotate in opposite direction, this was suggestive that this dust must have been accreted and can...
not be accounted for by mass loss from evolved stars (Kormendy & Djorgovski 1989). Kinematic information is important in order to constrain the presence of counter-rotating gas (and dust) in ellipticals in order to establish the frequency of the accretion scenario (Bertola et al. 1988).

In this study, the unresolved red ellipticals detected in the submm do not have associated kinematic information. However, we attempt to establish whether the present dust masses in our sample of elliptical galaxies can be explained with stellar sources using a model of dust formation and evolution in ellipticals. We compare the specific dust masses \( M_D/M_\ast \) with the predictions for dust mass return from a single stellar population (SSP) model and which represents an instantaneous burst of star formation. The star formation histories of the observed galaxies are more complex than that represented by a single burst of star formation. Their stellar masses and colours are however clearly dominated by the old stellar populations. Moreover, chemical evolution models of elliptical galaxies find very short timescales of their formation and high star formation efficiencies of the initial starburst (Pipino et al. 2005). The present SFR of \( 0.1 M_\odot \mathrm{yr}^{-1} \) in our sample is several orders of magnitude lower than that the SFR in the past responsible for the build-up of their stellar mass of \( 10^{11} M_\odot \). Therefore, for comparison with the dust model predictions, we assume that the entire stellar mass of each red-E galaxy is associated with a single burst with an age equal to its mass weighted age derived from the SED fitting. The observed dust mass in a galaxy is thus compared with the survived dust mass from the SSP with the same age. The model of the SSP adopted here was introduced in Zhukovska (2008) and was used to describe the chemical evolution of dust and gas in the Milky Way and dwarf galaxies (Zhukovska et al. 2008; Zhukovska 2014). For the chemical evolution aspects of the SSP model, we adopt the same ingredients as in Zhukovska (2008) except for the IMF, for which we use the Chabrier (2003) form. This is consistent with the IMF that is adopted in the SED fitting with MAGPHYS.

The model includes dust production by type II supernovae (SNe) and by asymptotic giant branch (AGB) stars. Type Ia SNe are an important source of metallic iron in early type galaxies. Models of dust evolution imply that, with an assumption of high condensation efficiencies of metals into dust in their ejecta, they can dominate dust input in elliptical galaxies (e.g., Calura et al. 2008; Pipino et al. 2011). Far-infrared observational surveys of both warm and cool dust in remnants of type Ia SNe do not however find evidence of efficient dust formation, in contrast to remnants of type II SNe (Gomez et al. 2012). This is supported by theoretical models, which indicate that newly formed grains are small and are easily destroyed in shocked gas before being ejected into the ISM (Nozawa et al. 2011). Therefore, we neglect the dust input from type Ia SNe.

The net input from type II SNe is still debated. We add their contribution for completeness, as they produce dust for a limited period of time after stars have formed (\( \sim 40 \) Myr). We adopt relatively low efficiencies of dust condensation in the SNe ejecta. These are constrained by meteoritic data and the observed metallicity-dust to gas ratio relation in dwarf galaxies (Zhukovska et al. 2008; Zhukovska 2014).

The mass- and metallicity-dependent dust yields for AGB stars are taken from the work of Ferrarotti (2006) with additional models from Zhukovska et al. (2008). These dust yields were computed for stellar metallicity ranging from \( Z = 0.001 \) up to the

<table>
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<th>±-err</th>
<th>c (intercept)</th>
<th>r-value</th>
<th>p-value</th>
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<td>log(MD/⊙)</td>
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<td>0.01</td>
<td>-3.89</td>
<td>0.56</td>
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<td>-4.87</td>
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Figure 10. Distributions of (a) star formation rate, (b) specific star formation rate and (c) \( f_\mu \), e.g. the fraction of total dust luminosity contributed by the diffuse interstellar medium, in the blue (thick solid line) and the red sources. E (red dashed line), S (green dotted line) and U (cyan line) labels represent the morphology of individual red source. Each histogram is normalized by its integral. The estimated mean value associated to each histogram is given in Table 1.
We extrapolate the dust yields in the mass range \([7-8]M_\odot\). Only one galaxy in the red-E sample is old enough for stars with masses below 1 \(M_\odot\) to contribute to the dust budget. However, stars in this mass range loose a large fraction of their envelopes during Red Giant Branch evolution characterised by inefficient dust formation (Gail et al. 2009; McDonald et al. 2011a, 2015). Some amount of dust is condensed during following AGB stage, but the total dust mass returned to the ISM is very low. Estimates based on the gas mass-loss rates derived in McDonald et al. (2011b) and McDonald et al. (2015) point to \(\lesssim 10^{-3}M_\odot\) of dust per star. Given this low value, we choose not to extrapolate the dust yields down to 0.8 \(M_\odot\) and neglect dust input from these stars.

The ISM in elliptical galaxies is dominated by hot rarefied gas with temperatures of \(\sim 10^7\) K (Mathews & Brighenti 2003). Grains can be rapidly sputtered in high-temperature gas due to collisions with ions (mostly with abundant H\(^+\)) (Draine & Salpeter 1979; Itoh 1989). The time scale of destruction by thermal sputtering can be approximated as

\[
\tau_{\text{sput}} = 10^5 \left(1 + (10^7K/T)^3\right) \frac{a/0.1\mu m}{n/cm^{-3}yr}, \tag{5}
\]

where \(n\) and \(T\) are the number density and temperature of the hot gas, respectively, and \(a\) is the grain radius. The total stardust mass \(M_{\text{D}}(t)\) is reduced by thermal sputtering in the hot gas at the rate:

\[
\frac{dM_{\text{D}}(t)}{dt} = -\frac{M_{\text{D}}(t)}{\tau_{\text{sput}}}. \tag{6}
\]

The temperature and density of the hot gas are derived from observations of extended X-ray emission. For simplicity, we assume single values for the electron density and temperature of the gas of \(10^{-3}\) cm\(^{-3}\) and \(1.5 \times 10^7\) K, respectively (Mathews & Brighenti 2003) resulting in \(\tau_{\text{sput}} = 100\) Myr. Note that \(\tau_{\text{sput}}\) depends only weakly on temperature in the regime appropriate for the hot ISM of elliptical galaxies and a value of \(T = 10^6\) K results in the time scale of 200 Myr. A similarly low value of the timescale of interstellar dust destruction, only 46 Myr, is derived for ETGs detected in FIR by Spitzer observations (Clemens et al. 2010). For
a comparison, we also ran calculations of the SSP evolution with a longer dust destruction timescale of 1 Gyr which corresponds to a lower gas density of $10^{-3} \text{ cm}^{-3}$. This long timescale may also account for the fact that many early type galaxies may harbour cold gas (Mathews & Brighenti 2003; Alatalo et al. 2013; Young et al. 2014), where grains are protected for some time from the thermal sputtering and can survive longer. Another mechanism of dust destruction is inertial sputtering in SN shocks, which is thought to be the dominant mechanism of dust destruction in spiral galaxies. However, in a hot rarefied medium one SN destroys 20 times less dust compared to the local ISM conditions (McKee 1989). We therefore do not consider dust destruction by type Ia SNe and restrict our consideration to the thermal sputtering in hot gas. Dust mass in an early type galaxy can also be substantially reduced by the galactic winds (not considered in the present model). Our estimates should therefore be considered as the upper limit for the stardust mass.

Fig. 13 compares the specific dust masses we have derived for the sample of red -E to the results of the SSP models. The data are grouped in three metallicity bins of [0.5 – 1.5] $Z_{\odot}$, [1.5-2.5] $Z_{\odot}$, and > 2.5 $Z_{\odot}$ and compared to three sets of SSP models with $Z = 3 Z_{\odot}$ (left panel), $Z = 2 Z_{\odot}$ (middle panel), and $Z = 3 Z_{\odot}$ (right panel). The specific dust masses of each red -E galaxy in the sample is plotted versus the mass weighted age of its stellar populations and the metallicity of each galaxy is obtained from the SDSS DR4 (Gallazzi et al. 2005). The figure clearly shows that, as expected, SSP models with no dust destruction tend to over-predict the amount of dust in these ellipticals. On the other hand, more realistic models with dust sputtering fail to reproduce the observed $M_D/M_*$ ratio even when a relatively long dust destruction timescale of 1 Gyr is considered. The SSP models with dust destruction under-predict the ratio of $M_D/M_*$ by more than two order of magnitude. These estimates demonstrate that dust return into ISM from stellar sources is not sufficient to explain the observed $M_D/M_*$. This implies an external origin of the dust via minor mergers and/or efficient dust growth in the dense ISM.

The amount of dust in the sub-mm detected galaxies as well as its correlation with the present day star-formation rate (Fig. 12 panel b) suggests a connection between the dust and the dense ISM in agreement with Alatalo et al. (2013), who find that the distribution of the CO and dust in nearby ETG is spatially correlated. The timescale for dust growth in molecular clouds is short and of the order of a few to several 10^{7} yrs (Hirashita 2000). We estimate an upper bound on the dust mass that may result from dust growth in the dense ISM in the following manner. Assuming a specific mass of molecular gas $M_{HI}/M_*$ of 0.01 and a value of 0.06 for the specific mass of the atomic gas $M_{HII}/M_*$ (these are the observed upper limits in Young et al. 2014), a dust-to-hydrogen mass ratio of 0.018 (i.e., about 3 times the solar value), and a complete condensation of heavy elements into dust in the molecular gas, this yields a specific dust mass $M_D/M_*$ of $0.07 \times 0.018 \approx 1.3 \times 10^{-3}$ which is only slightly higher than the largest specific dust masses measured for the sample of red Ellipticals that are displayed in Fig. 13. This means that it is difficult, but not impossible, to explain the measured dust masses as resulting from grain growth in the dense gas inside the elliptical galaxies. It should be noted that dust growth does not preclude the role of minor mergers because the molecular gas may have an external origin (Davis et al. 2011).

Characteristics of red submillimetre galaxies

Figure 12. Distribution of dust mass $M_D$ against (a) stellar mass $M_*$, (b) star formation rate SFR and (c) $S_\mu$ in blue (blue square), red -E (red circle). In addition, horizontal and vertical histograms show the distributions of data points along x and y axes with blue/thick and red/dashed lines representing blue and red -E. Each histogram is normalized by its integral. Typical errors associated to various parameters are indicated on the bottom-left corner. Results of linear regression analysis to blue and red -E observed data points in panels ‘a’ and ‘b’ are given in Table 3.

4 Value of the solar metallicity adopted here is $Z_{\odot} = 0.014$ Asplund et al. 2009.
4 CONCLUSIONS

In this work, we examine the properties of low redshift galaxies detected in 250µm (>5σ) using H-ATLAS DR1 catalogue. We define two sub-samples of red and blue galaxies based on NUV−r colours. Our aim is to understand the nature of the red subset in comparison to those in the blue sub-sample. We can summarize our findings as follows:

- Within the redshift range 0.01 ≤ z ≤ 0.2 of our sample, red sources with the UV-optical colour indices of NUV−r≥4.5, constitute ≈4.2 per cent of the total number of systems in H-ATLAS. The fraction of red sources increases with the galaxy stellar mass such that in ≥97 per cent of the red sample, M_* ≥ 10^{10}M_☉.

- Following the visual inspection of galaxies, sources in the red sample were grouped into three categories of elliptical (E), spiral (S) and uncertain (U). We find that at least ≥30 per cent of the red sources are of type E and more than ≥40 per cent of sources belong to type S.

- Both blue and red sources, seem to occupy environments with similar densities (e.g. having similar log(Σ_*)) distributions though in comparison to blue and red objects of type S and U, a slightly larger fraction of red -E sources are in relatively denser regions with log(Σ_*/(Mpc^{-2})) ≥ 1.5.

- The SED analysis of galaxies in our sample based on MAGPHYS, reveals that the red galaxies (either type S or E) span a similar range of dust masses but different dust-to-stellar mass ratios in comparison to the blue galaxies. The specific dust masses in the blue and red -S galaxies are, on average, larger than those found for the red -E sample by a factor of 7× and 2× respectively. Similarly, galaxies of type E have lower levels of mean SFR and SSFR in contrast to sources in the blue and red -S samples. Furthermore, analysis of f_ν shows that unlike blue galaxies where star-forming regions have the main contribution to the observed submm fluxes, FIR emission in the red systems of type E is mainly from the dust in the ISM.

- The UV-optical colours of the red -S sample could be the result of their highly inclined orientation and/or a strong contribution of the old stellar population. However, in the current work we did not further investigate the contribution of each factor to the observed colour of the red -S sources.

- Finally, the comparison of specific dust masses (MD/M_* ) of the red elliptical galaxies to the dust evolution in single stellar populations models excludes that the origin of the dust is from internal stellar sources. Dust growth in molecular clouds and/or gas and dust accretion through minor mergers provide more realistic and appealing alternatives (e.g., [Gomez et al. 2010], [Smith et al. 2012b]).

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The Herschel-ATLAS is a project with Herschel: which is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA. The H-ATLAS website is http://www.h-atlas.org/. GAMA is a joint European-Australasian project based around a spectroscopic campaign using the Anglo-Australian Telescope. The GAMA input catalogue is based on data taken from the Sloan Digital Sky Survey and UKIRT Infrared Deep Sky Survey. Complementary imaging of the GAMA regions is being obtained by a number of independent survey programs including GALEX MIS, VST KIDS, VISTA VIKING, WISE, Herschel -ATLAS, GMRT and ASKAP providing UV to radio coverage. GAMA is funded by the STFC (UK), the ARC (Australia), the AAO, and the participating institutions. The GAMA website is http://www.gama-survey.org/. MAGPHYS is available via http://www.iap.fr/magphys/magphys/MAGPHYS.html

REFERENCES

APPENDIX A: SDSS POSTAGE-STAMP IMAGES OF RED GALAXIES AND THEIR SED FITS
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<td>8h41m39.55</td>
<td>+1h53m46.57</td>
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<td>8h34m44.02</td>
<td>-0h12m43.98</td>
<td>4.67</td>
<td>0.035</td>
<td>73.62</td>
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Table A1. List of red galaxies detected in HATLAS.
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<th>SDSS OBJID</th>
<th>SDSS Ra</th>
<th>SDSS Dec</th>
<th>NUV−r</th>
<th>\log(\Sigma_{5})</th>
<th>i</th>
<th>type</th>
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<td>588889004842562600</td>
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Table A2. Table A1 Continued.
Table A3. MAGPHYS output parameters for the red galaxies having WISE observed photometric data.
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<th>( \log (SFR) ) ( [M_\odot \text{yr}^{-1}] )</th>
<th>( \log (SFR/M_*) )</th>
<th>( \log (M_D/M_\odot) )</th>
<th>( \log (M_D/M_*) )</th>
<th>( f_\mu )</th>
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Table A4. Table A3 Continued.