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ASASSN-15oi: A Rapidly Evolving, Luminous Tidal Disruption Event at 216 Mpc

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4 February 2016

ABSTRACT

We present ground-based and *Swift* photometric and spectroscopic observations of the tidal disruption event (TDE) ASASSN-15oi, discovered at the center of 2MASX J20390918-3045201 ($d \simeq 216$ Mpc) by the All-Sky Automated Survey for SuperNovae (ASAS-SN). The source peaked at a bolometric luminosity of $L \simeq 1.9 \times 10^{44}$ ergs s⁻¹ and radiated a total energy of $E \simeq 5.0 \times 10^{50}$ ergs over the ~ 3.5 months of observations. The early optical/UV emission of the source can be fit by a blackbody with temperature increasing from $T \sim 2 \times 10^4$ K to $T \sim 6 \times 10^4$ K while the luminosity declines from $L \simeq 1.9 \times 10^{44}$ ergs s⁻¹ to $L \simeq 2.8 \times 10^{43}$ ergs s⁻¹, requiring the photosphere to be shrinking rapidly. The optical/UV luminosity decline is broadly consistent with an exponential decline, $L \propto e^{-t/t_0}$, with $t_0 \simeq 35$ days. ASASSN-15oi also exhibits roughly constant soft X-ray emission that is significantly weaker than the optical/UV emission. Spectra of the source show broad helium emission lines and strong blue continuum emission in early epochs, although these features fade rapidly and are not present ~ 3 months after discovery. The early spectroscopic features and color evolution of ASASSN-15oi are consistent with a TDE, but the rapid spectral evolution is unique among optically-selected TDEs.

Key words: accretion, accretion disks — black hole physics — galaxies: nuclei

1 INTRODUCTION

When a star falls deep enough into the gravitational potential of a supermassive black hole (SMBH), the tidal shear forces overpower the self-gravity of the star, resulting in a

tidal disruption event (TDE). A portion of the disrupted stellar material is ejected while a fraction of the remaining material is accreted onto the black hole, creating a luminous, short-lived ($t \lesssim 1$ yr) accretion flare (e.g., Lacy, Townes & Hollenbach 1982; Rees 1988; Evans & Kochanek

1989; Phinney 1989). When the central black hole has a mass $M_{BH} \lesssim 10^7 M_{\odot}$, the debris initially falls back onto the black hole at a super-Eddington rate, and the debris eventually returns to pericenter on a timescale that roughly follows a $t^{-5/3}$ power law for a main sequence star (Evans & Kochanek 1989; Phinney 1989). It was commonly assumed that the luminosity of the TDE flare would be proportional to this rate of return of the stellar material to pericenter and that it would peak at soft X-ray energies.

Observed TDE candidates have exhibited a large diversity in properties. TDE flares have now been discovered in hard X-ray (e.g., Bloom et al. 2011; Burrows et al. 2011; Cenko et al. 2012b; Pasham et al. 2015), soft X-ray (e.g., Bade, Komossa & Dahlem 1996; Grupe, Thomas & Leighly 1999; Komossa & Greiner 1999; Donley et al. 2002; Maksym, Ulmer & Eracleous 2010; Saxton et al. 2012), ultraviolet (e.g., Stern et al. 2004; Gezari et al. 2006, 2008, 2009), and optical (e.g., van Velzen et al. 2011; Gezari et al. 2012; Cenko et al. 2012a; Arcavi et al. 2014; Chornock et al. 2014; Holoien et al. 2014a; Vinkó et al. 2015; Holoien et al. 2016) wavelengths—see Komossa (2015) for a review. While some TDEs appear to follow the predicted $t^{-5/3}$ power law decline, many exhibit different decline rates. Many of the X-ray-bright sources do not show associated strong optical emission, and many of the recent candidates discovered by optical surveys show strong optical and UV emission without associated X-ray emission. The notable exception to this is the exceptionally well-studied TDE ASASSN-14li, which was discovered in the optical (Holoien et al. 2016) but also detected at both X-ray (Holoien et al. 2016; Miller et al. 2015) and radio wavelengths (Alexander et al. 2015; van Velzen et al. 2015).

In practice, the properties of TDE flares depend on numerous physical parameters, including the depth of the encounter, the composition of the star, the fraction of the star that is accreted, and the geometry of the accretion stream (e.g., Kochanek 1994; Lodato & Rossi 2011; Guillochon & Ramirez-Ruiz 2015; Metzger & Stone 2015; Shiokawa et al. 2015). Although the TDE flare is powered by the accretion of the stellar material onto the black hole, the observed optical/UV emission is likely dominated by a photosphere formed within the stellar debris (Evans & Kochanek 1989; Loeb & Ulmer 1997; Ulmer 1999; Strubbe & Quataert 2009; Roth et al. 2015), and is likely dependent on the viewing angle (Guillochon, Manukian & Ramirez-Ruiz 2014; Metzger & Stone 2015).

The optically discovered TDEs also exhibit a continuum of spectroscopic properties ranging from He-dominated to H-dominated (Arcavi et al. 2014). This diversity could be due to photoionization physics (e.g., Guillochon, Manukian & Ramirez-Ruiz 2014; Gaskell & Rojas Lobos 2014; Roth et al. 2015), composition variations created by stellar initial composition and evolution (Kochanek 2015) and/or require the disruption of some helium stars (Gezari et al. 2012; Strubbe & Murray 2015). The discovery and observation of additional nearby TDEs are needed to fully understand the processes governing the observed emission of these optically bright TDE flares.

Here we describe the discovery and follow-up observations of ASASSN-15oi, the third TDE discovered by the

All-Sky Automated Survey for SuperNovae (ASAS-SN¹; Shappee et al. 2014). ASAS-SN is a long-term project designed to monitor the entire sky on a rapid cadence to find nearby supernovae (e.g., Dong et al. 2016; Holoien et al. 2014b; Shappee et al. 2015) and other bright transients. In particular, ASAS-SN has already proven to be a powerful tool for discovering new TDEs, and has produced the two closest TDEs ever discovered at optical wavelengths (Holoien et al. 2014a, 2016). Our transient detection pipeline was triggered on 2015 August 14 by a new source at RA/Dec = 20:39:09.12/−30:45:20.84 (J2000) with $V = 16.2 \pm 0.1$ mag (Brimacombe et al. 2015). The source was not detected ($V \gtrsim 17.2$ mag) on 2015 July 26 or earlier. The object’s position corresponds to 2MASX J20390918-3045201, a galaxy with no previous redshift measurement. Follow-up images obtained on 2015 August 25 with the *Swift* UltraViolet and Optical Telescope (UVOT; Roming et al. 2005) and X-ray Telescope (XRT; Burrows et al. 2005) confirmed the detection of the transient in both the UV and X-rays.

A follow-up spectrum obtained by the Public ESO Spectroscopic Survey for Transient Objects (PESSTO; Smartt et al. 2015) using the ESO New Technology Telescope at La Silla showed that the transient exhibited a strong blue continuum and a broad (~ 10000 km s^{−1}) He II 4686Å emission feature (Prentice et al. 2015), consistent with helium-rich optical TDEs (e.g., Gezari et al. 2012; Arcavi et al. 2014). Although the redshift was originally reported as $z = 0.02$, our later spectra showed that the redshift of the host galaxy is $z = 0.0484$, corresponding to a luminosity distance of 216 Mpc ($H_0 = 69.6$ km s^{−1} Mpc^{−1}, $\Omega_M = 0.29$, $\Omega_{\Lambda} = 0.71$). Because of the similarities between ASASSN-15oi and previously discovered TDEs, we began a long-term follow-up campaign to monitor and characterize the object.

In §2 we describe pre-outburst data of the host galaxy as well as new observations obtained of the transient during our follow-up campaign. In §3 we analyze these data to model the transient’s luminosity and temperature evolution and compare the properties of ASASSN-15oi to those of supernovae and other TDEs to determine the nature of the transient. Finally, in §4 we discuss our findings and how they relate to the active field of TDE research.

2 OBSERVATIONS AND SURVEY DATA

Here we summarize the available archival survey data of the transient host galaxy 2MASX J20390918-3045201 as well as our new photometric and spectroscopic observations of ASASSN-15oi.

2.1 Archival Data

Because the source is located in the Southern hemisphere, there are no available archival imaging or spectroscopic data in the Sloan Digital Sky Survey (SDSS). We obtained archival near-IR JHK_s images from the Two-Micron All Sky Survey (2MASS; Skrutskie et al. 2006) and measured 5′′0 aperture magnitudes for the galaxy in these images. This aperture radius, chosen to match the PSF of the *Swift*

¹ <http://www.astronomy.ohio-state.edu/~assassin/>

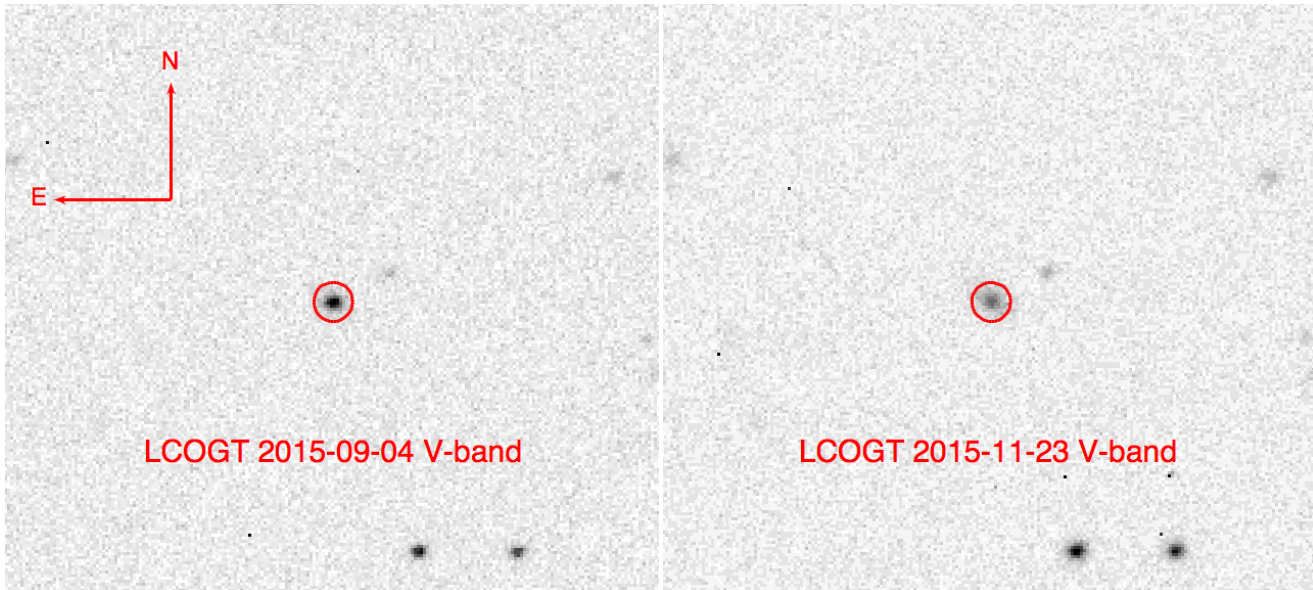


Figure 1. *V*-band images of ASASSN-15oi shortly after discovery and after the transient had faded. The left panel shows the LCOGT 1-m image from 2015 September 4 and the right panel shows the LCOGT 1-m image from 2015 November 23. The red circles have radii of $3''.0$ and are centered on the position of the transient measured in the September 4 image.

Table 1. Archival Photometry of the Host Galaxy

Filter	Magnitude	Magnitude Uncertainty
<i>V</i>	17.13	0.05
<i>J</i>	15.17	0.07
<i>H</i>	14.52	0.07
<i>K_s</i>	14.05	0.09
<i>W1</i>	14.18	0.03
<i>W2</i>	14.12	0.05

Measured $5''.0$ -radius aperture *JHK_s* magnitudes from 2MASS and a $14''.0$ -radius aperture *V*-band magnitude from ASAS-SN. The WISE *W1* and *W2* are PSF photometry magnitudes from the AllWISE source catalog.

data, was also used to measure the source flux in follow-up data. We also obtained a *V*-band aperture magnitude of the host galaxy by stacking several epochs of ASAS-SN data and measuring the magnitude in the same way using a 2-pixel aperture, roughly corresponding to a $14''.0$ aperture due to the large ASAS-SN pixel scale. These measured magnitudes were later used to model the host galaxy SED and subtract the host galaxy flux from follow-up data of the transient. We present the measured $5''.0$ aperture magnitudes from the 2MASS and ASAS-SN images in Table 1.

There are no archival Spitzer, Herschel, Hubble Space Telescope (HST), Chandra, X-ray Multi-Mirror Mission (XMM-Newton), Galaxy Evolution Explorer (GALEX), or Very Large Array Faint Images of the Radio Sky at Twenty-cm (VLA FIRST) observations of 2MASX J20390918-3045201. Examining data from the ROSAT All-Sky Survey (Voges et al. 1999), we do not detect the host galaxy with a 3-sigma upper limit of 2.0×10^{-2} counts s^{-1} in the 0.3 – 10.0 keV band, corresponding to a limit of 1.2×10^{-13} ergs s^{-1} cm^{-2} ($L_X < 6.7 \times 10^{41}$ ergs s^{-1}), which provides evidence that the galaxy does not host a strong AGN. Detections of the host in Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010) data corroborate

this picture, as the galaxy has a mid-IR (MIR) color of $(W1 - W2) \simeq 0.06 \pm 0.06$, implying that any AGN activity in the galaxy is not strong (e.g., Assef et al. 2013).

2.2 New Photometric Observations

After the transient was classified as a TDE, we obtained a series of 26 *Swift* XRT and UVOT target-of-opportunity (ToO) observations. The UVOT observations were obtained in 6 filters: *V* (5468 Å), *B* (4392 Å), *U* (3465 Å), *UVW1* (2600 Å), *UVM2* (2246 Å), and *UVW2* (1928 Å) (Poole et al. 2008). We extracted source counts from a $5''.0$ radius region and sky counts from a $\sim 40''.0$ radius region using the UVOT software task UVOTSOURCE and converted these count rates into magnitudes and fluxes using the most recent UVOT calibrations (Poole et al. 2008; Breeveld et al. 2010).

The XRT operated in Photon Counting mode (Hill et al. 2004) for our observations. The data from all epochs were reduced with the software task XRTPIPELINE. We extracted X-ray source counts and background counts with XRTSELECT using a region with a radius of 10 pixels ($23''.6$) centered on the source position and a source-free region with a radius of 100 pixels ($235''.7$), respectively. While the source is undetected in most individual epochs, we detect X-ray emission consistent with the position of the transient after combining the signal from multiple exposures. We assume a blackbody plus power law model (see §3.1) and Galactic H I column density of 5.59×10^{20} cm^{-2} (Kalberla et al. 2005) to convert the detected counts into fluxes, with a count to flux conversion factor of 5.23×10^{-11} ergs s^{-1} (cm^2 counts s^{-1}) $^{-1}$ for the absorbed spectrum and 8.5×10^{-11} ergs s^{-1} (cm^2 counts s^{-1}) $^{-1}$ for the absorption-corrected spectrum.

We also obtained *BVI* images with the Las Cumbres Observatory Global Telescope Network (LCOGT; Brown et al. 2013) 1-m telescopes at Siding Spring, South African Astronomical, and Cerro Tololo Inter-America Observato-

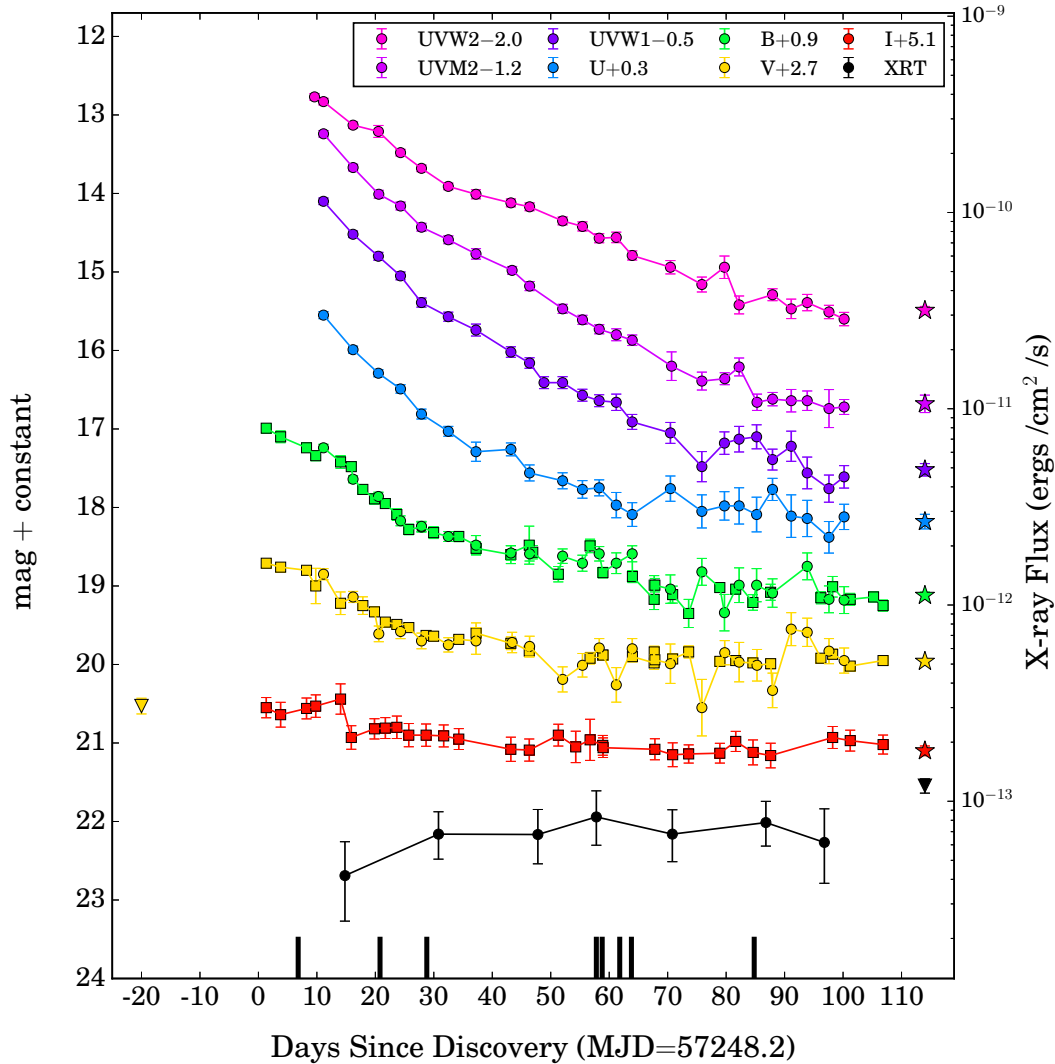


Figure 2. Light curves of ASASSN-15oi beginning on the epoch of discovery (MJD= 57248.2) and spanning 107 days. Follow-up data obtained from *Swift* (X-ray, UV, and optical) are shown as circles and follow-up data obtained by the LCOGT 1-m telescopes (optical) are shown as squares. All optical and UV magnitudes are on the Vega system (left scale), and X-ray fluxes are shown in $\text{ergs cm}^{-2} \text{s}^{-1}$ (right scale). The scales are chosen so that time variability has the same meaning for both the X-ray and optical/UV data. The X-ray data points are averages over multiple epochs. The *Swift* *B*- and *V*-band data were converted to the Johnson-Cousins system using the color corrections found at <http://heasarc.gsfc.nasa.gov/docs/heasarc/caldb/swift/docs/uvot/>. The data are not corrected for extinction and error bars are shown for all points, but can be smaller than the data points. The previous *V*-band non-detection from ASAS-SN is plotted as a triangle at -20 days, the date of our previous observation of the host galaxy. Estimated $5''$ aperture magnitudes measured by averaging the last 4 epochs in each filter (stars) and an X-ray upper limit from RASS (triangle) are shown at $+114$ days. Vertical bars at the bottom of the figure indicate the dates of spectroscopic observations. Although the rapid rise starting after our previous non-detection at -20 days was missed, the data still show that ASASSN-15oi brightened considerably in the UV and in the bluer optical filters, with the largest increase being roughly 3.4 magnitudes in the *Swift* *UVM2* band. Table A1 contains all the follow-up photometric data.

ries. We measured magnitudes using aperture photometry with a $5''$ aperture radius in order to match the host galaxy and *Swift* measurements. Photometric zero-points were determined using several stars in the field.

In order to constrain any offset between the source of the outburst and the nucleus of the host galaxy we astrometrically aligned an *I*-band image of the transient taken on 2015 September 4 with the LCOGT 1-m telescope at Sutherland, South Africa with the archival DSS image of the host galaxy. We then measured the offset between the

brightest pixel in the host galaxy in the LCOGT image and that of the DSS image, finding an offset of 0.24 ± 0.05 arcseconds (254.8 ± 53.1 parsecs). We performed the same analysis using two LCOGT *V*-band images from 2015 September 4 and 2015 November 23, after the transient had faded, and find an offset of 0.10 ± 0.06 arcseconds (106.2 ± 63.7 parsecs). Figure 1 shows the two LCOGT *V*-band images used to measure the offset. Higher resolution data are needed to improve these estimates.

As there were no archival data in many of our follow-up

Table 2. Estimated Photometry of the Host Galaxy

Filter	Magnitude	Magnitude Uncertainty
<i>UVW2</i>	17.49	0.05
<i>UVM2</i>	17.88	0.08
<i>UVW1</i>	18.02	0.09
<i>U</i>	17.88	0.11
<i>B_{UVOT}</i>	18.06	0.08
<i>B_{LCOGT}</i>	18.22	0.04
<i>V_{UVOT}</i>	17.30	0.09
<i>V_{LCOGT}</i>	17.26	0.03
<i>I</i>	16.00	0.06

Estimated $5''0$ -radius aperture magnitudes of the host galaxy measured by combining later epochs of follow-up observations in the *Swift* UVOT filters and LCOGT *BVI* filters, after the light curve had leveled off. Telescope is indicated for the *B*- and *V*-band measurements, and all magnitudes are listed in the Vega system.

filters, we lack pre-outburst magnitudes of the host. However, the light curve of the transient had begun to level-off in all filters after ~ 3 months, and transient features such as the strong blue continuum and emission lines had faded completely from follow-up spectra (see §2.3). We interpreted this to mean that ASASSN-15oi had faded completely in the optical and UV filters we were using for our follow-up observations, allowing us to obtain approximate host magnitudes by combining the later photometric observations. We averaged the measured fluxes from the final 4 epochs of observations (beginning MJD 57339.2 for the *Swift* data and MJD 57327.1 for the LCOGT data), and used this average flux to obtain host magnitudes in all six *Swift* UVOT filters and the LCOGT *BVI* filters. These approximate magnitudes are listed in Table 2.

Using the code for Fitting and Assessment of Synthetic Templates (FAST v1.0; Kriek et al. 2009), we fit stellar population synthesis (SPS) models to the estimated $5''0$ *Swift* UVOT and LCOGT *BVI* and the measured 2MASS *JHK_s* magnitudes of the host galaxy. The fit was done assuming a Cardelli, Clayton & Mathis (1988) extinction law with $R_V = 3.1$ and Galactic extinction of $A_V = 0.19$ mag based on Schlafly & Finkbeiner (2011), an exponentially declining star-formation history, a Salpeter IMF, and the Bruzual & Charlot (2003) models. We obtained a good SPS fit (reduced $\chi^2_\nu = 1.05$), with the following parameters: $M_* = (1.7^{+0.3}_{-0.3}) \times 10^{10} M_\odot$, age = $10.0^{+0.3}_{-0.3}$ Gyr, and a 1σ upper limit on the star formation rate of $\text{SFR} \leq 0.5 M_\odot \text{ yr}^{-1}$. Scaling the stellar mass of 2MASX J20390918-3045201 using the average stellar-mass-to-bulge-mass ratio from the hosts of ASASSN-14ae and ASASSN-14li (Holoien et al. 2014a, 2016) implies a bulge mass of $M_B \simeq 10^{10} M_\odot$. Using the $M_B - M_{BH}$ relation from McConnell & Ma (2013), we obtain a black hole mass of $M_{BH} = 10^{7.4} M_\odot$, which is significantly more massive than those of the hosts of the previous ASAS-SN TDEs. Fits to the transient spectral energy distribution give no indication of additional extinction related to the host galaxy. In the analyses of the event’s SED which follow, we only correct for Galactic extinction.

Figure 2 shows the X-ray, UV, and optical light curves of ASASSN-15oi. Table A1 and Table A2 give the UVOT/*VRI* magnitudes and XRT flux measurements, respectively. The photometric observations cover the period from MJD 57248.2 (the epoch of discovery) through our lat-

est epoch of observations on MJD 57355.0, a span of 106.8 days. The data shown in Figure 2 are presented without extinction correction or host flux subtraction. Also shown in Figure 2 are the host magnitudes inferred from the later epochs of follow-up data. Our observations show that ASASSN-15oi brightened considerably with respect to the host galaxy in the UV and bluer optical filters in the short period between our previous observation on 2015 July 25 and our detection of the transient on 2015 August 14. The largest increase came in the *Swift* *UVM2* band, where it brightened by $\Delta m_{UVM2} \sim -3.4$, while the *V*-band increase was weaker, at $\Delta m_V \sim -1.3$.

2.3 New Spectroscopic Observations

We obtained spectra of ASASSN-15oi spanning 69 days between UT 2015 September 04 and UT 2015 November 07. The spectrographs used for these observations were the Ohio State Multi-Object Spectrograph (OSMOS; Martini et al. 2011) mounted on the MDM Observatory Hiltner 2.4-m telescope (4200 – 6800 Å, $R \sim 4$ Å), the Fast Spectrograph (FAST; Fabricant et al. 1998) mounted on the Fred L. Whipple Observatory Tillinghast 1.5-m telescope (3700 – 9000 Å, $R \sim 3$ Å), and the Wide Field Reimaging CCD Camera (WFCCD) mounted on the Las Campanas Observatory du Pont 2.5-m telescope (3700 – 9500 Å, $R \sim 7$ Å). The spectra were reduced using standard techniques in IRAF, and we applied telluric corrections using observations of spectrophotometric standard stars from the same nights. Each spectrum was scaled to match the *V*-band photometry. Figure 3 shows a time-sequence of the flux-calibrated follow-up spectra as well as the initial spectrum obtained by PESSTO (Prentice et al. 2015). Summary information for each spectrum is listed in Table A3.

The key characteristics of the early spectra of ASASSN-15oi are a strong blue continuum, consistent with the photometric measurements, and the presence of broad helium lines in emission. The emission features are broad and asymmetric, with initial widths of $\sim 10000 - 20000 \text{ km s}^{-1}$ that narrow over time. In the initial spectrum obtained by PESSTO, the helium lines are blueshifted by approximately $6000 - 8000 \text{ km s}^{-1}$, but in later spectra little-to-no apparent shift relative to the systemic velocity is observed. The transient features fade rapidly, with spectra later than 2015 September 12 showing no blue continuum and absorption features consistent with the transient’s host galaxy. Unlike other recently discovered TDEs (e.g. Arcavi et al. 2014; Holoien et al. 2016; French, Arcavi & Zabludoff 2016), the later host-dominated spectra do not show features, such as $H\delta$ absorption with a large equivalent width, that are consistent with the host being a post-starburst galaxy. We further analyze the features of these spectra and compare them to other TDEs and supernovae in §3.

3 ANALYSIS

3.1 SED Analysis

Combining all 61.4 ks of observations, we were able to extract an X-ray spectrum of the source. The spectrum was analyzed with XSPEC 12.8.2 (Arnaud 1996) using Cash statis-

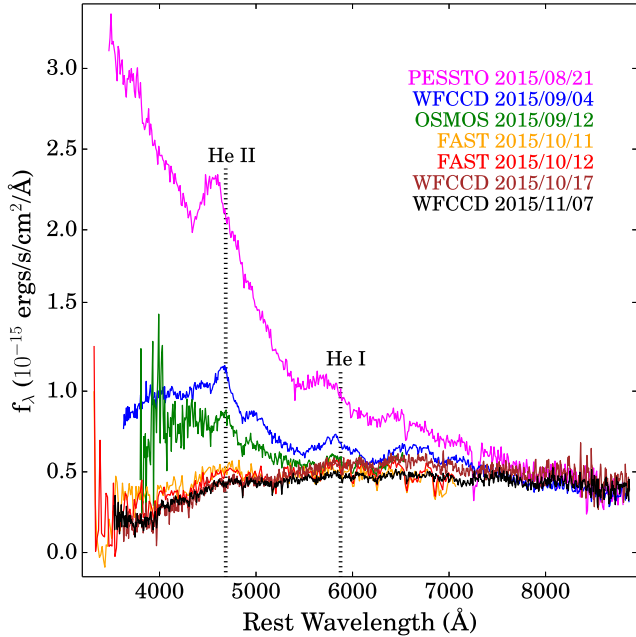


Figure 3. Spectral time-sequence for ASASSN-15oi. The dates of the observations and the instruments used are listed in matching colors to the spectra. (See Table A3 for more information.) The prominent He I 5876Å and He II 4686Å emission features are indicated by vertical dotted lines. The transient spectra show many broad emission features and blue continuum emission in early epochs which rapidly fade, with later spectra resembling that of the host galaxy rather than that of the transient.

tics (Cash 1979). A fit with a single power law model does not result in an acceptable fit. The X-ray spectral slope is very steep ($\Gamma=5.8$) with large residuals at energies above 2.0 keV. We found an acceptable fit using an absorbed blackbody plus power law model with the absorption column density fixed to the Galactic value. This results in a blackbody temperature at an energy of 49_{-9}^{+10} eV and a photon index for the hard energy component of $\Gamma = 1.76_{-0.84}^{+1.50}$. The absorbed and unabsorbed fluxes in the 0.3-10 keV band are $8.0_{-0.25}^{+0.37} \times 10^{-14}$ ergs s $^{-1}$ cm $^{-2}$ and $18.70_{-0.58}^{+0.86} \times 10^{-14}$ ergs s $^{-1}$ cm $^{-2}$, respectively. The X-ray spectrum and the model fits are shown in Figure 4.

There are two challenges to interpreting the X-ray detections. First, the archival ROSAT data are not deep enough to rule out a pre-existing source with the presently observed X-ray luminosity ($L_X < 6.7 \times 10^{41}$ versus $L_X = 4.8 \times 10^{41}$ ergs/s). Second, the X-ray emission we observe during the transient is consistent with constant luminosity. We can rule out X-ray emission dominated by sources other than an AGN. Low Mass X-ray Binaries (LMXBs) in old stellar populations, like the FAST models for the SED of 2MASX J20390918-3045201, produce X-rays. Based on the correlations of Kim & Fabbiano (2004) between the K -band luminosities of galaxies and the integrated X-ray emission by their LMXBs, however, this contribution should only be of order 10^{40} ergs s $^{-1}$, well below the observed luminosity. Thus, the X-ray emission is almost certainly dominated by AGN activity, and seems likely to be associated with the present transient, but this association is not required by the

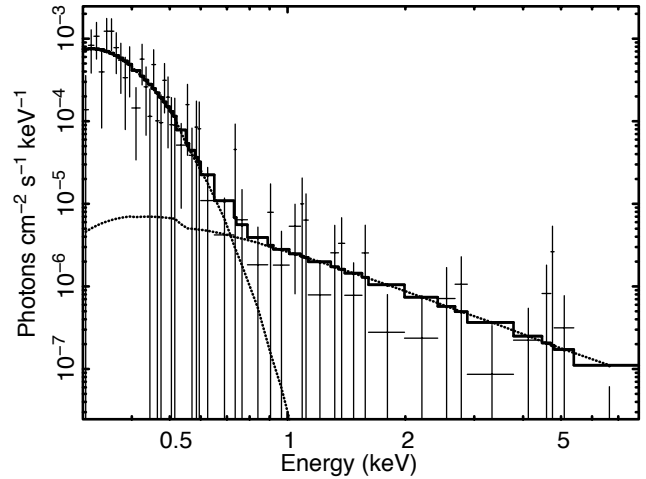


Figure 4. Unfolded X-ray spectrum (unbinned) of ASASSN 15oi fitted with an absorbed black body plus power law model, with a blackbody temperature of 49_{-9}^{+10} eV and a power law index of $\Gamma = 1.76_{-0.84}^{+1.50}$. The dotted lines display the black body and power law models and the solid line shows the combined model. The data were fitted with Cash Statistics (Cash 1979).

available data. Late-time observations of the X-ray evolution should illuminate the source of this emission.

Figure 5 shows the evolution of the UV-optical SED of ASASSN-15oi. Using the 5'0 aperture host magnitudes estimated by combining recent epochs of data, we produced host- and extinction-corrected light curves for all optical and UV filters. We then fit these host-subtracted fluxes with blackbody models using Markov Chain Monte Carlo (MCMC) methods, as was done for the previous ASAS-SN TDEs (Holoien et al. 2014a, 2016). Due to there being some uncertainty in the actual host magnitude, we only included epochs for which the transient's flux was greater than 50% of the host galaxy flux when performing these SED fits, and we did not include the later epochs (those obtained after MJD 57338) used to estimate the host magnitudes.

Unlike the previous ASAS-SN TDEs, the data for ASASSN-15oi indicate that the source's temperature is increasing in later epochs. Because of this, we fit the temperature using a changing prior based on initial, unconstrained fits. For epochs within 10 days of discovery, the temperature was fit with a prior of $T = 2 \times 10^4$ K. For later epochs, we fit the data with a prior of $T = (2 + (\Delta t - 10)/2) \times 10^4$ K, where $\Delta t = MJD - 57248.2$. These priors were applied with a log uncertainty of ± 0.05 dex if there were fewer than 4 photometric points being used for a given epoch, and with an uncertainty of ± 0.3 dex if there were 4 or more data points being used. Due to the fact that our *Swift* data do not span the peak of the SED for the later, hotter epochs, the temperatures in these phases is not well-constrained, although the prior used for these fits does capture the steepening spectral slope. Also shown in Figure 5 are the X-ray spectral model (see Figure 4) and the ionizing luminosity implied by the He II 4686Å line luminosity (see §3.2). It can be clearly seen that the blackbody model inferred from the optical/UV data is not enough to explain the observed X-ray and line emission.

The optical/UV luminosity of ASASSN-15oi, shown in

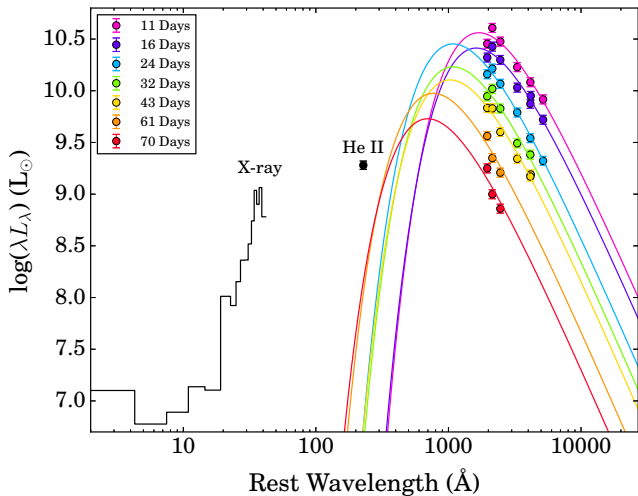


Figure 5. Evolution of the SED of ASASSN-15oi shown in different colors (points) and the corresponding best-fitting blackbody models for each epoch (lines). All data have been corrected for Galactic extinction and host contribution. Error bars are shown but can be smaller than the data points. All fits shown were made assuming a temperature prior of $T = (2 + (\Delta t - 10)/2) \times 10^4$ K, where $\Delta t = MJD - 57248.2$. The X-ray spectral model inferred from the average of all the X-ray data (see Figure 4 and the ionizing luminosity implied by the He II 4686 Å line (see §3.2) are shown in black. Neither the X-ray emission nor the ionizing luminosity for the He II line can be explained by the single blackbody model inferred from the optical/UV data, and a more complex model is likely needed.

Figure 6, fades steadily over the ~ 3 months after the initial discovery with the temperature constrained as described above. As was the case with previous ASAS-SN TDEs, the luminosity evolution inferred from the blackbody fit is well-described by an exponential $L \propto e^{-t/t_0}$ (Holoien et al. 2014a, 2016), with $t_0 \simeq 35$ days, similar to the decay time observed for ASASSN-14ae ($t_0 \simeq 39$ days). This differs from common models of TDE emission, where the luminosity is expected to follow a power law evolution t^{-x} with $x \simeq 5/12 - 5/3$ (e.g., Strubbe & Quataert 2009; Lodato & Rossi 2011). However, this temperature and luminosity evolution is inconsistent with what would be expected if ASASSN-15oi were a supernova, which typically exhibit rapidly declining temperatures along with constant or declining luminosity (e.g., Miller et al. 2009; Botticella et al. 2010; Inserra et al. 2013; Graham et al. 2014). Normal AGN activity is ruled out by the archival host photometry, leaving a TDE as the most likely explanation for ASASSN-15oi.

In contrast to the optical/UV luminosity evolution, the X-ray luminosity remains constant for the duration of the observations. However, even in later epochs, the X-ray luminosity is nearly two magnitudes lower than the optical/UV luminosity, and even if it is associated with the TDE, it has little impact on the overall energy output of the source. As shown in Figure 5, a single blackbody model is incapable of fitting both the X-ray and optical/UV emission. Thus, if the X-ray emission is associated with the TDE, we infer that the X-ray emission likely comes from a different region of the TDE than the optical/UV emission. Including the UV luminosity implied by the He II emission, ASASSN-15oi

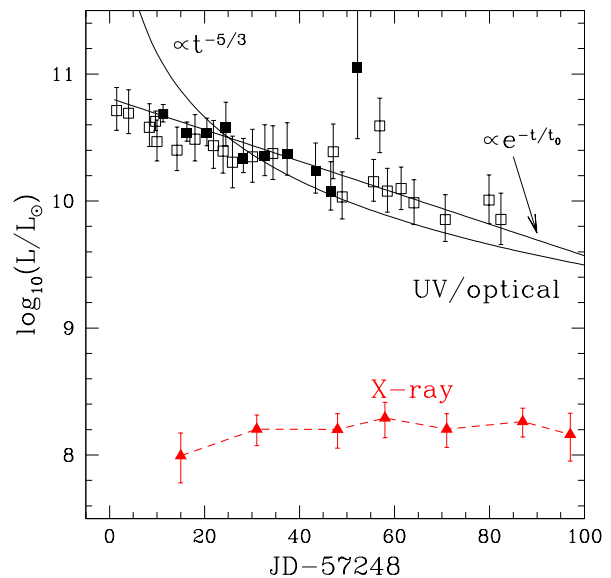


Figure 6. Evolution of ASASSN-15oi's X-ray (red triangles) and optical/UV (black squares) luminosity over time assuming the emission is consistent with a blackbody with the temperature prior described in the text. Filled points indicate those epochs where 4 or more data points were used to fit the temperature, and unfilled points indicate those epochs where fewer data points were used. The curved line shows a $L \propto t^{-5/3}$ power law fit, while the diagonal line shows an exponential fit, $L \propto e^{-t/t_0}$. The exponential model appears to be a much better fit to the data, similar to the previous ASAS-SN TDEs (Holoien et al. 2014a, 2016), though this analysis does not account for the X-ray emission, and is assuming a simple single-blackbody model.

is another piece of evidence pointing to the need for more complex models of TDE emission (e.g., Metzger & Stone 2015; Roth et al. 2015).

Integrating the combined optical/UV and X-ray luminosity curves gives a total radiated energy of $\sim 5.0 \times 10^{50}$ ergs for the ~ 3.5 months of follow-up observations. The uncertainties in this estimate are predominantly systematic due to the poorly constrained temperatures at later times.

3.2 Spectroscopic Analysis

Though the SED evolution of ASASSN-15oi is largely consistent with the transient being a tidal disruption flare, the rapid evolution and absolute magnitude are unique among TDEs discovered at optical wavelengths, and it is worth examining other possible sources of the observed emission. The most likely alternative explanation for the transient is that it was a supernova, as archival observations do not show any indication that 2MASX J20390918-3045201 hosts an AGN. While supernovae evolve on the rapid timescales observed for ASASSN-15oi, a spectroscopic comparison between ASASSN-15oi and supernovae reveals that this is not the case.

ASASSN-15oi shows broad He emission lines in its early spectra but little-to-no H emission. If ASASSN-15oi were a supernova, it would likely be a Type Ib or Type Ic-bl: the former because SNe Ib show He emission with no H emis-

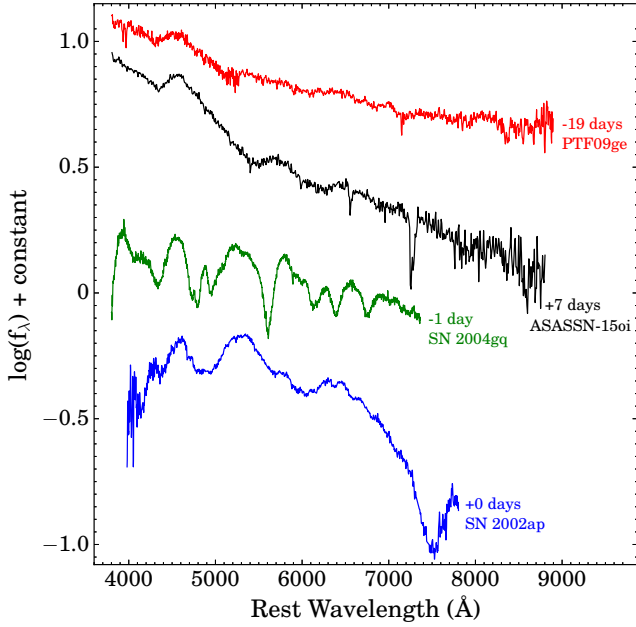


Figure 7. Comparison between spectra taken at or before maximum of ASASSN-15oi (Prentice et al. 2015), the He-rich TDE PTF09ge (Arcavi et al. 2014), the Type Ib supernova SN 2004gq (Modjaz et al. 2014), and the broad-line Type Ic supernova SN 2002ap (Gal-Yam, Ofek & Shemmer 2002). The epoch of each spectrum in days-since maximum-light is given next to each spectrum. (The epoch of ASASSN-15oi is days-since-discovery, since the peak date is not known.) Spectroscopically, ASASSN-15oi is far more similar to PTF09ge than it is to either supernova, which have different emission and absorption features and different continuum shapes.

sion, and the latter because broad-line SNe Ic exhibit line widths comparable to those observed in ASASSN-15oi. Figure 7 shows a comparison between spectra taken at or before maximum for ASASSN-15oi (Prentice et al. 2015), the He-rich TDE PTF09ge (Arcavi et al. 2014), the Type Ib supernova SN 2004gq (Modjaz et al. 2014), and the broad-line Type Ic supernova SN 2002ap (Gal-Yam, Ofek & Shemmer 2002). While SN 2004gq does show the helium emission lines present in the spectrum of ASASSN-15oi, there are also numerous absorption features that are not present in the spectrum of ASASSN-15oi. Similarly, SN 2002ap shows broad emission features, but does not have the same helium emission lines, and the shape of the continuum is very different. In contrast, the spectrum of PTF09ge is very similar to that of ASASSN-15oi, as both objects show a strong blue continuum, a broad He II 4686Å emission feature, and no hydrogen emission. Given this comparison, we conclude that it is very unlikely that ASASSN-15oi is a supernova.

Arcavi et al. (2014) noted that optically-selected TDEs span a continuum from H-rich to He-rich spectroscopic features. As indicated in Figure 7, early follow-up spectra of ASASSN-15oi are highly consistent with those of He-rich TDEs like PS1-10jh and PTF09ge (Gezari et al. 2012; Arcavi et al. 2014). The prominent features of the earliest spectra are a strong blue continuum with broad He II and He I emission features. The earliest spectrum, obtained by PESSTO on 2015 August 21, shows He II 4686Å and He I 5876Å emission features with $\text{FWHM}_{\text{He II}} \simeq 19700 \text{ km s}^{-1}$

and $\text{FWHM}_{\text{He I}} \simeq 18800 \text{ km s}^{-1}$, respectively, that are blueshifted by $\sim 6000 - 8000 \text{ km s}^{-1}$. The second spectrum, obtained with the du Pont+WFCCD on 2015 September 4, shows some interesting changes. The He II 4686Å feature is still prominent but has a smaller width ($\text{FWHM}_{\text{He II}} \simeq 9200 \text{ km s}^{-1}$) compared to the first spectrum and the peak of the line is only blueshifted by $\sim 1400 \text{ km s}^{-1}$. Conversely, the He I 5876Å line has remained broad, with $\text{FWHM}_{\text{He I}} \simeq 18800 \text{ km s}^{-1}$. There are two new broad bumps ($\text{FWHM} \sim 10000 \text{ km s}^{-1}$) that appear around the strong He II 4686Å line with peak wavelengths at $\sim 4435\text{Å}$ and $\sim 5000\text{Å}$. These features are consistent with the He I 4472Å and He I 5016Å transitions, blueshifted by $1000 - 2500 \text{ km s}^{-1}$. This spectrum also shows a new feature at $\sim 6600\text{Å}$, near the position of H α , which exhibits an integrated flux that is comparable to the He II 4686Å line flux. It is likely that this feature is due at least in part to the He II 6560Å $n = 6 \rightarrow 4$ line, which is expected to be an order of magnitude weaker than He II 4686Å in photoionized gas (Osterbrock 1989; Gezari et al. 2015).

Unlike what has been seen in previous TDEs, the spectroscopic features of ASASSN-15oi evolve rapidly. Between the classification spectrum obtained on UT 2015 August 21 and the first follow-up spectrum obtained on UT 2015 September 4, the blue continuum in the spectra of ASASSN-15oi fades and the He II line narrows considerably. The emission lines seem to have completely faded by the time of the latest follow-up spectrum obtained on UT 2015 November 7, slightly less than 3 months after discovery.

After correcting for Galactic reddening and subtracting an estimated continuum from the PESSTO classification spectrum, we estimate the luminosity of the He II 4686Å line to be $L_{\text{He II}} \simeq (1.1 \pm 0.1) \times 10^{42} \text{ ergs s}^{-1}$, where the dominant source of uncertainty is in setting the continuum level. In the UT 2015 September 4 spectrum, the line has faded to a luminosity of $L_{\text{He II}} \simeq (1.7 \pm 0.5) \times 10^{41} \text{ ergs s}^{-1}$. We are unable to obtain estimates of the line luminosity in later spectra, as the line has faded and the continuum is too uncertain to provide an accurate subtraction. Assuming case B recombination, the line luminosity from the classification spectrum implies an ionizing luminosity of $L_i \simeq (7.3 \pm 0.7) \times 10^{42} \text{ ergs s}^{-1}$, while the line luminosity from the September 4 spectrum implies an ionizing luminosity of $L_i \simeq (1.1 \pm 0.3) \times 10^{42} \text{ ergs s}^{-1}$. While this is a rapid drop, in both cases the line luminosity is greater than the peak line luminosity measured for ASASSN-14li (Holoien et al. 2016). As shown in Figure 5, the optical/UV emission is unable to provide the requisite ionizing luminosity needed to power the observed He II line, and additional ionizing flux must be coming from a different region of the TDE.

It is not uncommon for the emission features in TDE spectra to narrow and become less luminous over time (e.g., Holoien et al. 2014a, 2016), which is the opposite behavior to that observed in reverberation mapping studies of quasars, where the line width broadens as the luminosity decreases (e.g. Peterson et al. 2004; Denney et al. 2009). However, while spectroscopic evolution of ASASSN-15oi is similar to that of other optically-selected TDEs, it occurs on a significantly faster timescale. This rapid line evolution could be indicating that the “reprocessing layer” responsible for the optical/UV emission and emission lines, is becoming optically thin or is no longer in our line-of-sight to the TDE. It

is also worth noting that ASASSN-14ae, which had a similar ~ 3 week gap between the previous epoch of observation and the epoch of discovery, also showed a blueshift in its broad emission features in its early spectra which was not present in later spectroscopic observations (Holoien et al. 2014a). While ASASSN-14li did not show any blueshift in its spectroscopic observations (Holoien et al. 2016), it also had a significantly longer gap (~ 3 months) between the epoch of discovery and the previous epoch of observations due to Sun constraints. It is possible that an early blueshift in the emission lines of optically-selected TDEs is common, and that it was not seen in ASASSN-14li due to the gap in observations.

As discussed by French, Arcavi & Zabludoff (2016), optically-selected TDEs have shown a preference for unusual “quiescent Balmer-strong galaxies”—those galaxies whose spectra show little-to-no line emission but strong Balmer line absorption, indicating a recent period of intense star formation. While archival spectra of 2MASX J20390918-3045201 are not available, our later epochs of follow-up spectra are dominated by the host galaxy, allowing us to determine whether the host of ASASSN-15oi also falls into this rare class of galaxies. From the 2015 November 7 du Pont/WFCCD spectrum, we measure $H\delta$ and $H\alpha$ absorption features with $EW = 1.1 \text{ \AA}$ and $EW = 0.6 \text{ \AA}$, respectively. In French, Arcavi & Zabludoff (2016), a quiescent Balmer-strong galaxy is defined as having absorption with a Lick $H\delta_A$ index of $H\delta_A > 4 \text{ \AA}$ (equivalent to $H\delta \text{ EW} > 3 \text{ \AA}$) and $H\alpha \text{ EW} < 3 \text{ \AA}$. With much narrower $H\delta$ absorption, and $H\alpha$ in absorption rather than emission, 2MASX J20390918-3045201 clearly does not fall into the same class of galaxy as the TDE hosts discussed in French, Arcavi & Zabludoff (2016), and is more consistent with the old stellar population indicated by the FAST SED fits.

3.3 Comparison of ASAS-SN TDEs

ASAS-SN has now found three nearby TDEs, all of which were intensively observed with *Swift* and ground-based telescopes for months after discovery, allowing us to compare their photometric properties and the inferred temperatures and luminosities at early and late epochs. Such an analysis was not possible with earlier TDEs, due to the fact that they were fainter, and late-time observation in many cases was not possible. All data for ASASSN-14ae and ASASSN-14li shown below are taken from Holoien et al. (2014a) and Holoien et al. (2016), respectively.

From the host-subtracted light curves we produced color evolution curves for ASASSN-15oi, ASASSN-14ae, and ASASSN-14li. Figure 8 compares the *Swift* ($UVM2 - UVW2$) and ($UVM2 - U$) colors of the three TDEs. All three objects show blue UV-UV and UV-optical colors in all epochs, which is the expected for optically-selected TDEs. The ($UVM2 - UVW2$) colors of all three objects are remarkably similar and show little evolution for all three objects, but the ($UVM2 - U$) colors do show some variation: while they remain blue overall, ASASSN-14ae becomes slightly redder and ASASSN-15oi becomes slightly bluer, while ASASSN-14li shows little evolution. The blue evolution of the ($UVM2 - U$) color for ASASSN-15oi is in-line with the increasing temperature described in §3.1, and it is not unexpected that ASASSN-15oi would differ slightly

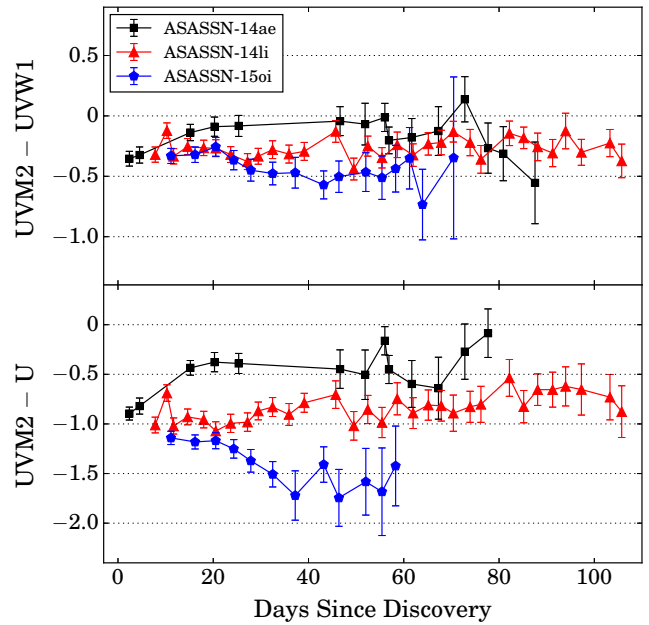


Figure 8. Comparison of ($UVM2 - UVW1$) (top panel) and ($UVM2 - U$) (bottom panel) color evolution between ASASSN-15oi (blue pentagons); ASASSN-14ae (Holoien et al. 2014a, black squares); and ASASSN-14li (Holoien et al. 2016, red triangles). Extinction correction and host flux subtraction has been applied to all objects. All three TDEs show little evolution in ($UVM2 - UVW1$) and remain quite blue in both colors for months after discovery. ASASSN-15oi does become bluer in ($UVM2 - U$), indicating an increasing temperature.

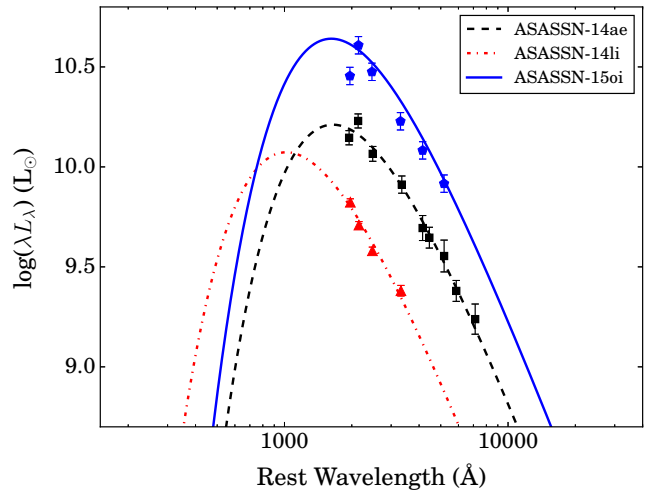


Figure 9. Comparison of the blackbody SEDs inferred from early epoch optical/UV data for ASASSN-15oi (blue solid line, 11 days after discovery), ASASSN-14ae (black dashed line, days after discovery), and ASASSN-14li (red dash-dotted line, 8 days after discovery). In all three cases, the temperature was fit to host-subtracted and extinction-corrected *Swift* and ground photometric data, which are also shown on the figure in the same colors and symbols used in Figure 8. ASASSN-14li and ASASSN-15oi were fit with a temperature prior. Though more luminous, ASASSN-15oi exhibits a blackbody temperature that is not uncommon for an optically-selected TDE.

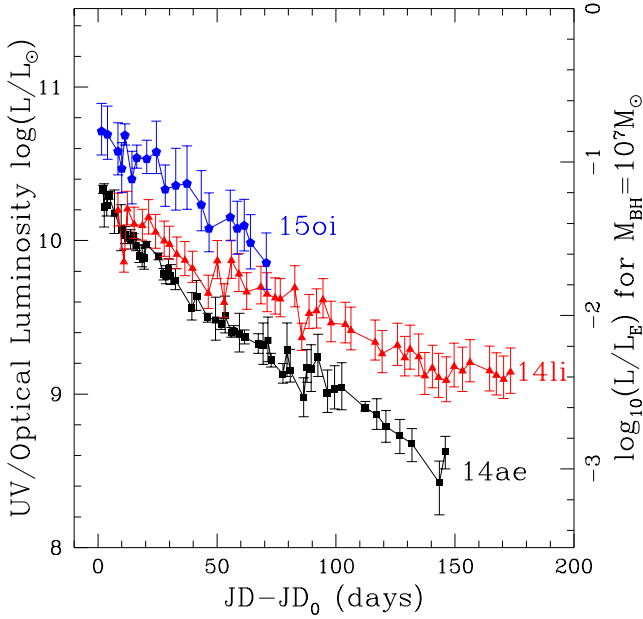


Figure 10. Comparison of the evolution of the luminosity inferred from the blackbody SED fits for ASASSN-15oi (blue pentagons), ASASSN-14ae (black squares), and ASASSN-14li (red triangles). The X-ray emission detected in ASASSN-15oi and ASASSN-14li is not included in the luminosity measurements. The scale on the right shows the luminosity in units of the Eddington luminosity for a $M_{BH} = 10^7 M_{\odot}$ black hole. ASASSN-15oi is the most luminous of the three TDEs by a clear margin, but its luminosity declines at a rate similar to that of ASASSN-14ae, and the slower-fading ASASSN-14li is of comparable luminosity roughly 70 days after discovery.

from the other two TDEs, as ASASSN-14ae and ASASSN-14li both had roughly constant temperatures in all epochs. In all cases, the magnitude of the color change is fairly minor compared to that of supernovae, which typically become redder much more rapidly (see Figure 7 of Holoien et al. (2014a) for a comparison).

Figure 9 shows a comparison of the inferred blackbody SEDs for all three objects from epochs shortly after discovery. In all three cases the fit shown used both *Swift* and ground-based data, and for ASASSN-14li and ASASSN-15oi the data were fit assuming a temperature prior. ASASSN-14ae and ASASSN-15oi exhibited similar early temperatures of roughly $T \sim 20000$ K while ASASSN-14li was hotter, with $T \sim 35000$ K. Though ASASSN-15oi is more luminous and evolved more rapidly than the other two ASAS-SN TDEs, its inferred blackbody temperature is not unusual for an optically-selected TDE.

Finally, we compare the blackbody luminosity, temperature, radius, and emission line evolution inferred for the three ASAS-SN TDEs in Figures 10, 11, 12, and 13, respectively. In all cases the bolometric luminosity was inferred from the optical/UV SED fits, as described in §3.1, and does not include the X-ray emission detected for ASASSN-15oi or ASASSN-14li. ASASSN-15oi is considerably more luminous than the other two TDEs, but its bolometric luminosity falls at a fairly similar rate to that of ASASSN-14ae. ASASSN-14li fades much more slowly than the other two; while it

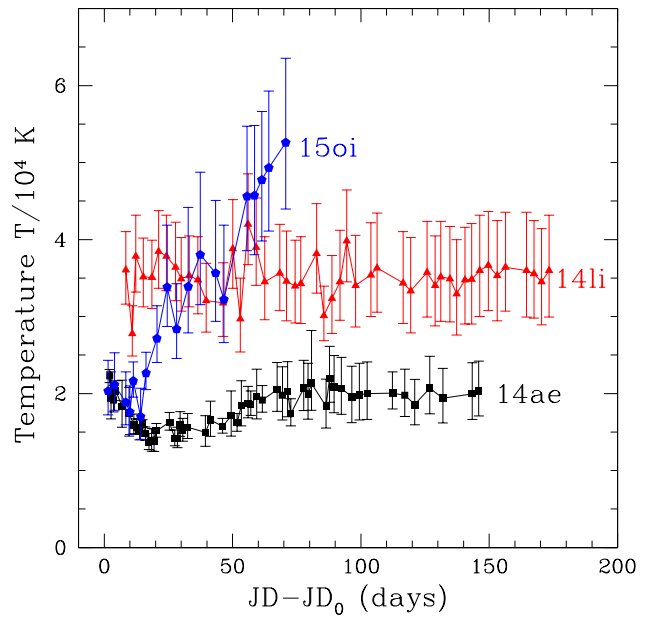


Figure 11. Comparison of the blackbody temperature evolution inferred for ASASSN-15oi (blue pentagons), ASASSN-14ae (black squares), and ASASSN-14li (red triangles). The temperature was fit with a prior for all epochs of ASASSN-15oi and ASASSN-14li and for those epochs of ASASSN-14ae that did not include *Swift* observations, as described in the text. While ASASSN-14ae and ASASSN-14li exhibit the constant temperature evolution common in TDEs, ASASSN-15oi shows significant evolution, becoming much hotter than both other ASAS-SN TDEs in the first ~ 70 days following discovery, although the maximum temperature is prior-dependent.

initially has a similar luminosity as ASASSN-14ae, by later epochs it is nearly an order of magnitude more luminous than ASASSN-14ae, and ASASSN-15oi has fallen to a comparable luminosity. This is perhaps correlated with the mass of the black hole involved in the TDE: though the three ASAS-SN TDEs have similar black hole masses given the uncertainties on the mass estimates, ASASSN-14li ostensibly had the least massive black hole, at $M_{BH} \sim 10^{6.7} M_{\odot}$, while ASASSN-14ae and ASASSN-15oi had black holes of masses $M_{BH} \sim 10^{6.8} M_{\odot}$ and $M_{BH} \sim 10^{7.4} M_{\odot}$, respectively.

Including both apparently thermal optical/UV and X-ray emission (where observed), we estimate total radiated energies of 1.7×10^{50} ergs, 7.0×10^{50} ergs, and 5.0×10^{50} ergs for ASASSN-14ae, 14li, and 15oi, respectively. Converted to an equivalent rest mass, these energies correspond to 0.001, 0.004, and 0.003 M_{\odot} , compared to a typical TDE stellar mass of $M_* = 0.3 M_{0.3} M_{\odot}$ (Kochanek 2016). If the radiated energy is $E = \eta f M_* c^2$ where $\eta \simeq 0.1$ is the radiative efficiency of accretion and f is the fraction of the stellar mass accreted ($f \simeq 0.5$ is bound to the black hole), then we must have $\eta f \simeq 0.003 M_{0.3}^{-1}$, $0.013 M_{0.3}^{-1}$ and $0.010 M_{0.3}^{-1}$ for these three events. All three events require either that TDE activity is far less radiatively efficient than normal disk accretion ($\eta \ll 0.1$) or that a very small fraction of the bound mass ($f \ll 0.5$) is accreted. Variants of these possibilities appear in recent theoretical models (e.g., Metzger & Stone 2015;

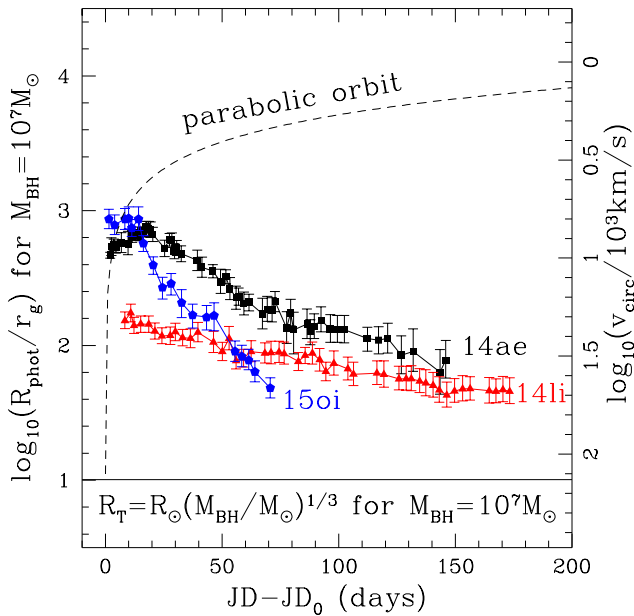


Figure 12. Comparison of the evolution of the photospheric radius inferred from the blackbody SED fit for ASASSN-15oi (blue pentagons), ASASSN-14ae (black squares), and ASASSN-14li (red triangles). The left scale gives the radius in units of the gravitational radius of a $M_{BH} = 10^7 M_{\odot}$ black hole, and the right scale gives the implied circular velocity. For comparison, the horizontal line gives the tidal disruption radius R_T of the Sun, and the dashed line shows the radial evolution of a parabolic orbit. The photospheric radii are roughly bounded by the parabolic expansion at peak and then shrink back towards the tidal radius. ASASSN-15oi shrinks fastest, going from the largest to the most compact in roughly 70 days.

Piran et al. 2015; Strubbe & Murray 2015; Svirski, Piran & Krolik 2015).

ASASSN-15oi stands out very clearly from the other two ASAS-SN TDEs in its temperature evolution. While ASASSN-14ae and ASASSN-14li evolve at nearly constant temperatures, ASASSN-15oi increases in temperature. The exact increase is poorly constrained because the data do not sample the peak of the SED, but it is consistent with a sample model where the temperature linearly increases from $T \sim 20000$ K to $T \sim 55000$ K over roughly 70 days. ASASSN-15oi was fit with the temperature prior described in §3.1 above and ASASSN-14li was fit with a prior of $\log T/K = 4.55 \pm 0.05$. As the UV data for ASASSN-14ae captured the peak of the transient’s SED, its temperature was not fit with a prior for those epochs where *Swift* data were available, and was fit with a prior that tracks the UV data for those epochs where *Swift* data was not available. While there are systematic uncertainties in these temperature fits, particularly for ASASSN-15oi and ASASSN-14li, it is clear that ASASSN-15oi has an increasing temperature, and that the other two TDEs do not.

Figure 12 shows the evolution of the apparent photospheric radii of the ASAS-SN TDEs. All three show the same decreasing trend, but the rate at which the radius falls differs significantly between the three TDEs, with ASASSN-15oi standing out from the other two events. To put the ra-

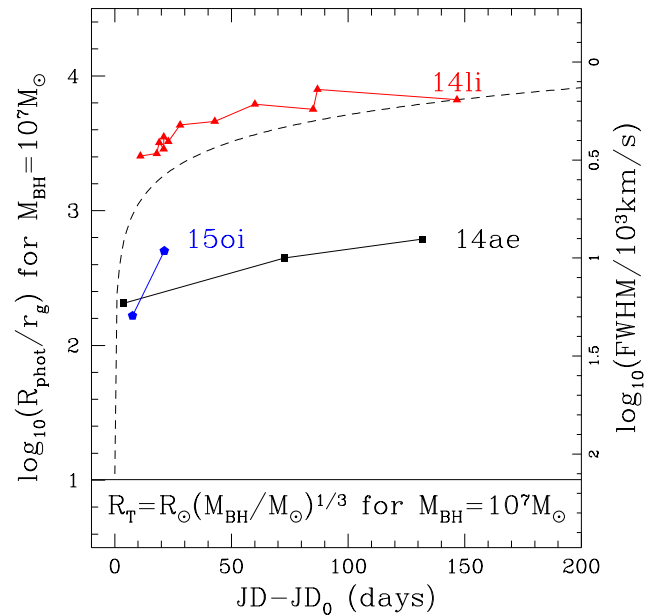


Figure 13. Comparison of the evolution of the FWHM (right scale, km s⁻¹). The left scale converts it to a radius on the same scale as in Figure 12 (i.e., assuming $v_c = FWHM$). Also as in Figure 12, the horizontal line gives the tidal disruption radius R_T of the Sun, and the dashed line shows the radial evolution of a parabolic orbit. Results are shown for the most prominent emission line in each ASAS-SN TDE. H α is shown for ASASSN-14ae (black squares) and ASASSN-14li (red triangles), while He II 4686Å is shown for ASASSN-15oi (blue pentagons). To the extent that the velocity can be mapped into a radius, the line emission regions lie outside the photospheric radii (Figure 12), inside a parabolic orbit, and evolve to larger radii. ASASSN-15oi exhibits a much more rapid narrowing in its emission lines than ASASSN-14ae or ASASSN-14li.

dius estimates on a more physical scale, Figure 12 shows the radii in units of the gravitational radius $r_g = GM_{BH}/c^2 = 1.5 \times 10^{12}$ cm of a $M_{BH} = 10^7 M_{\odot}$ black hole. The mass estimates of $M_{BH} \sim 10^{6.8}$, $10^{6.7}$ and $10^{7.4} M_{\odot}$ for ASASSN-14ae, 14li and 15oi are sufficiently close to this mass to use a common scaling given the overall dynamic range. For comparison, we also show the tidal disruption radius of the Sun, $R_T = R_{\odot}(M_{BH}/M_{\odot})^{1/3} \simeq 1.5 \times 10^{13}$ cm $\simeq 10r_g$, and the radius of a parabolic orbit initiated at the origin at $t = 0$. To facilitate comparisons to velocities, we also convert the radial scale on the left to the corresponding circular velocity scale, $v_c = c(r/r_g)^{1/2}$, on the right.

Broadly speaking, the initial photospheric radii are close to the radius that would be achieved by free parabolic expansion from the disruption radius R_T rather than being comparable to R_T or smaller. Given the uncertainties in the black hole masses and the temporal start of the transients, this is mainly a qualitative observation. The apparent photospheres then shrink, but are still well outside R_T at the end of the phase where they can be well-estimated. That the photospheric radii are all large compared to R_T favors models where the optical/UV emission is reprocessed by debris on large scales rather than direct emission from a disk. Svirski, Piran & Krolik (2015) propose generating the luminosity by shocks at apocenter, rather than by reprocessing

disk emission, but this would imply a much smaller emitting area than $4\pi r^2$ at the apocentric radius. If the covering fraction of the shock emission region is f^2 , then the radius has to increase to r/f to produce the observed luminosity and this would initially lie well outside the radius corresponding to a parabolic orbit if $f \sim 0.1$. If the dominant source of line emission was closely associated with the photosphere, we would expect the typical line widths to steadily increase with time. As we noted earlier, this is also the pattern observed in AGN, where line widths increase with decreasing luminosity.

For comparison to the photospheric evolution, Figure 13 shows the evolution of the FWHM of the most prominent emission line in each TDE ($H\alpha$ for ASASSN-14ae and ASASSN-14li, and $He\ II\ 4686\text{\AA}$ for ASASSN-15oi). The scales and limits of Figure 13 are the same as for Figure 12 except for relabeling the circular velocity to be the FWHM (i.e. just taking $v_c = FWHM$) to allow qualitative comparisons of the radial and velocity scales. For ASASSN-14ae and ASASSN-14li, the initial line formation region is of order the photospheric radius. For ASASSN-14li, the two scales are quite different. The simplest explanation for the mild conflict with parabolic expansion would be to shift the start of the expansion to be 10-20 days earlier. In all three cases, the FWHM diminish with time, moving outwards in radius if we can interpret the velocities kinematically. At late times, the dominant source of line emission appears to arise from radii far outside the (continuum) photosphere. Though we are only able to obtain two clear measurements of the $He\ II\ 4686\text{\AA}$ line for ASASSN-15oi, the data show that the line is narrowing much more rapidly than $H\alpha$ did for the previous two ASAS-SN TDEs. The primary caveat for the line evolution is that continuum subtraction is challenging for all these events, particularly at late times, and an additional (but sub-dominant) broad emission component might be hard to identify.

The extensive optical and UV data sets obtained for the three ASAS-SN TDEs allows us to perform comparisons between them that are largely impossible with other, fainter TDEs. While the three TDEs show unique characteristics, their colors, temperatures, and luminosity evolution are all broadly similar.

4 CONCLUSIONS

ASASSN-15oi was discovered by ASAS-SN on 2015 August 14 and had a peak absolute V -band magnitude of $M_V \sim -20.5$. Follow-up data indicate that it is consistent both photometrically and spectroscopically with previously discovered He-rich, optically-selected TDEs, and is inconsistent with either a supernova or an AGN event. Unlike other optically discovered TDEs, ASASSN-15oi faded rapidly in the optical and UV, and spectra obtained ~ 3 months after discovery indicate that the transient features had almost completely disappeared. If we fit the optical and UV emission with a single blackbody model, the best-fit temperature is $T \sim 20000\text{ K}$, and the temperature becomes hotter in later epochs. The luminosity declines steadily at a rate best fit by an exponential decay, $L \propto e^{-t/t_0}$, with $t_0 \simeq 35$ days. Like ASASSN-14li (Holoien et al. 2016), ASASSN-15oi exhibits soft X-ray emission, though this emission is weaker than

previous limits and cannot be conclusively tied to the TDE flare. Late-time X-ray observations to see if the X-ray emission starts to fade are needed to confirm the association of the X-ray emission with the transient. ASASSN-15oi is the third nearby TDE discovered by ASAS-SN, and possibly the second to exhibit both X-ray and optical/UV emission.

Early follow-up spectra show a strong blue continuum and broad $He\ II\ 4686\text{\AA}$ and $He\ I\ 5725\text{\AA}$ emission features, and while the spectroscopic features evolve rapidly, their evolution is consistent with that of a TDE, and not that of a supernova or quasi-steady state AGN. These spectral features are characteristic of He-rich TDEs like PS1-10jh and PTF09ge (Gezari et al. 2012; Arcavi et al. 2014), making ASASSN-15oi another member of this intriguing class of objects which exhibit strong helium emission lines but no corresponding hydrogen emission. That He-dominated spectra are so common essentially rules out the possibility that they are produced by the disruption of helium stars. While it might be possible to explain a single event such as PS1-10jh with this scenario (Gezari et al. 2012; Strubbe & Murray 2015), it seems improbable that He stars can represent $\sim 1/3$ of the overall optical/UV TDE rate.

In Holoien et al. (2016), we found that a 90% confidence interval on the implied TDE rate given the ASAS-SN discoveries was $r = (2.2-17.0) \times 10^{-5}\text{ yr}^{-1}$ per galaxy, a range that is consistent with theoretical estimates (e.g., Stone & Metzger 2016; Kochanek 2016) while also significantly higher but marginally consistent with the rates inferred from other optical surveys (e.g., van Velzen & Farrar 2014). The discovery of a third ASAS-SN TDE less than a year after the previous one implies that the relatively high (compared to previous observational studies) rate from (Holoien et al. 2016) is correct, which helps to alleviate some of the tension between observed and theoretical rates. With the addition of this third TDE, ASAS-SN is finding roughly 1 TDE for every 70 type Ia supernovae (3 and 211, respectively, at the time of writing), a rate that is significantly higher than that of other surveys (see Holoien et al. (2016) for a detailed comparison). This suggests that the TDE sample discovered by ASAS-SN is significantly more complete than previous surveys, and consequently that our rate estimate may be more accurate.

ASAS-SN has now discovered the three nearest optically discovered TDEs, each of which has been extensively observed over long periods of time in the optical, UV, and X-rays. Due to the nearby and bright nature of objects discovered by ASAS-SN, future TDEs will be similarly easy to observe with a wide variety of instruments, which will allow the creation of a catalog of well-studied TDEs that can be used to study the early and late-time behaviors of these transients, which cannot be done with higher-redshift objects. Moreover, due to the rapid evolution of the black hole mass function, the TDE rate should decline rapidly with increasing redshift (Kochanek 2016), making TDEs creatures of the local universe probed by ASAS-SN. Given the success of the ASAS-SN TDE search to-date, we expect it will remain a powerful tool for finding and studying TDEs and other bright transients in the future.

ACKNOWLEDGMENTS

The authors thank PI Neil Gehrels and the *Swift* ToO team for promptly approving and executing our observations. We thank LCOGT and its staff for their continued support of ASAS-SN.

ASAS-SN is supported by NSF grant AST-1515927. Development of ASAS-SN has been supported by NSF grant AST-0908816, the Center for Cosmology and AstroParticle Physics at the Ohio State University, the Mt. Cuba Astronomical Foundation, and by George Skestos.

TW-SH is supported by the DOE Computational Science Graduate Fellowship, grant number DE-FG02-97ER25308. Support for JLP is in part provided by FONDECYT through the grant 1151445 and by the Ministry of Economy, Development, and Tourism's Millennium Science Initiative through grant IC120009, awarded to The Millennium Institute of Astrophysics, MAS. PC and SD are supported by "the Strategic Priority Research Program-The Emergence of Cosmological Structures" of the Chinese Academy of Sciences (Grant No. XDB09000000) and Project 11573003 supported by NSFC. BJS is supported by NASA through Hubble Fellowship grant HST-HF-51348.001 awarded by the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., for NASA, under contract NAS 5-26555. JFB is supported by NSF grant PHY-1404311. EJ acknowledges support from the Marie Curie Actions of the European Commission (FP7-COFUND).

This research has made use of the XRT Data Analysis Software (XRTDAS) developed under the responsibility of the ASI Science Data Center (ASDC), Italy. At Penn State the NASA *Swift* program is support through contract NAS5-00136.

This research was made possible through the use of the AAVSO Photometric All-Sky Survey (APASS), funded by the Robert Martin Ayers Sciences Fund.

This research has made use of data provided by Astrometry.net (Barron et al. 2008).

This paper uses data products produced by the OIR Telescope Data Center, supported by the Smithsonian Astrophysical Observatory.

This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation.

This publication makes use of data products from the Wide-field Infrared Survey Explorer, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration.

This research has made use of the NASA/IPAC Extragalactic Database (NED), which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

This work is based in part on observations collected at the European Organisation for Astronomical Research in the Southern Hemisphere, Chile as part of PESSTO, (the Public

ESO Spectroscopic Survey for Transient Objects Survey) ESO program 188.D-3003, 191.D-0935.

This publication makes use of data obtained from the Weizmann interactive supernova data repository (WISE-EREP; Yaron & Gal-Yam 2012).

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APPENDIX A: FOLLOW-UP PHOTOMETRY

All follow-up photometry are presented in Table A1 and Table A2, while information about follow-up spectroscopic observations is given in Table A3. All optical and UV photometry is presented in the Vega system.

Table A1. Photometric data of ASASSN-15oi.

MJD	Magnitude	Filter	Telescope	MJD	Magnitude	Filter	Telescope
57249.532	15.45 0.129	<i>I</i>	LCOGT	57327.129	17.26 0.045	<i>V</i>	LCOGT
57252.002	15.54 0.158	<i>I</i>	LCOGT	57329.844	17.25 0.055	<i>V</i>	LCOGT
57256.428	15.46 0.133	<i>I</i>	LCOGT	57332.791	17.28 0.051	<i>V</i>	LCOGT
57258.008	15.43 0.143	<i>I</i>	LCOGT	57335.816	17.29 0.065	<i>V</i>	LCOGT
57262.251	15.34 0.193	<i>I</i>	LCOGT	57344.415	17.22 0.060	<i>V</i>	LCOGT
57264.083	15.83 0.150	<i>I</i>	LCOGT	57346.420	17.17 0.066	<i>V</i>	LCOGT
57268.052	15.72 0.129	<i>I</i>	LCOGT	57349.423	17.32 0.056	<i>V</i>	LCOGT
57269.908	15.71 0.133	<i>I</i>	LCOGT	57355.035	17.25 0.052	<i>V</i>	LCOGT
57271.882	15.70 0.143	<i>I</i>	LCOGT	57259.334	16.12 0.071	<i>V</i>	<i>Swift</i>
57273.924	15.80 0.148	<i>I</i>	LCOGT	57264.392	16.42 0.071	<i>V</i>	<i>Swift</i>
57276.874	15.80 0.142	<i>I</i>	LCOGT	57268.788	16.88 0.1	<i>V</i>	<i>Swift</i>
57279.895	15.81 0.140	<i>I</i>	LCOGT	57272.514	16.86 0.091	<i>V</i>	<i>Swift</i>
57282.476	15.85 0.133	<i>I</i>	LCOGT	57276.096	16.98 0.1	<i>V</i>	<i>Swift</i>
57291.383	15.98 0.154	<i>I</i>	LCOGT	57280.688	17.03 0.091	<i>V</i>	<i>Swift</i>
57294.522	15.99 0.141	<i>I</i>	LCOGT	57285.415	16.99 0.17	<i>V</i>	<i>Swift</i>
57299.495	15.80 0.139	<i>I</i>	LCOGT	57291.590	17.01 0.13	<i>V</i>	<i>Swift</i>
57302.481	15.95 0.200	<i>I</i>	LCOGT	57294.588	17.06 0.13	<i>V</i>	<i>Swift</i>
57304.921	15.86 0.262	<i>I</i>	LCOGT	57300.236	17.47 0.16	<i>V</i>	<i>Swift</i>
57307.106	15.93 0.126	<i>I</i>	LCOGT	57303.628	17.29 0.15	<i>V</i>	<i>Swift</i>
57307.140	15.96 0.125	<i>I</i>	LCOGT	57306.494	17.08 0.12	<i>V</i>	<i>Swift</i>
57316.018	15.98 0.134	<i>I</i>	LCOGT	57309.413	17.54 0.22	<i>V</i>	<i>Swift</i>
57319.059	16.05 0.152	<i>I</i>	LCOGT	57312.137	17.09 0.13	<i>V</i>	<i>Swift</i>
57321.796	16.04 0.116	<i>I</i>	LCOGT	57318.671	17.28 0.25	<i>V</i>	<i>Swift</i>
57327.133	16.03 0.126	<i>I</i>	LCOGT	57324.051	17.82 0.36	<i>V</i>	<i>Swift</i>
57329.849	15.88 0.128	<i>I</i>	LCOGT	57327.961	17.16 0.15	<i>V</i>	<i>Swift</i>
57332.796	16.02 0.159	<i>I</i>	LCOGT	57330.420	17.26 0.25	<i>V</i>	<i>Swift</i>
57335.821	16.06 0.157	<i>I</i>	LCOGT	57333.480	17.30 0.2	<i>V</i>	<i>Swift</i>
57346.424	15.83 0.140	<i>I</i>	LCOGT	57336.141	17.61 0.22	<i>V</i>	<i>Swift</i>
57349.419	15.87 0.133	<i>I</i>	LCOGT	57339.334	16.84 0.21	<i>V</i>	<i>Swift</i>
57355.039	15.92 0.121	<i>I</i>	LCOGT	57342.069	16.88 0.18	<i>V</i>	<i>Swift</i>
57249.533	16.01 0.053	<i>V</i>	LCOGT	57345.793	17.12 0.16	<i>V</i>	<i>Swift</i>
57252.004	16.06 0.055	<i>V</i>	LCOGT	57348.384	17.23 0.16	<i>V</i>	<i>Swift</i>
57256.429	16.1 0.053	<i>V</i>	LCOGT	57249.531	16.09 0.055	<i>B</i>	LCOGT
57258.009	16.3 0.225	<i>V</i>	LCOGT	57252.001	16.2 0.075	<i>B</i>	LCOGT
57262.252	16.52 0.144	<i>V</i>	LCOGT	57256.426	16.34 0.042	<i>B</i>	LCOGT
57266.065	16.55 0.112	<i>V</i>	LCOGT	57258.007	16.44 0.065	<i>B</i>	LCOGT
57268.053	16.63 0.047	<i>V</i>	LCOGT	57262.250	16.52 0.082	<i>B</i>	LCOGT
57269.909	16.76 0.050	<i>V</i>	LCOGT	57264.081	16.58 0.062	<i>B</i>	LCOGT
57271.883	16.79 0.062	<i>V</i>	LCOGT	57266.062	16.87 0.063	<i>B</i>	LCOGT
57273.926	16.83 0.058	<i>V</i>	LCOGT	57268.051	16.99 0.068	<i>B</i>	LCOGT
57276.875	16.93 0.057	<i>V</i>	LCOGT	57269.906	17.05 0.060	<i>B</i>	LCOGT
57278.172	16.94 0.062	<i>V</i>	LCOGT	57271.880	17.19 0.074	<i>B</i>	LCOGT
57282.475	16.98 0.050	<i>V</i>	LCOGT	57273.923	17.38 0.067	<i>B</i>	LCOGT
57285.424	16.9 0.130	<i>V</i>	LCOGT	57276.872	17.26 0.069	<i>B</i>	LCOGT
57291.382	17.03 0.081	<i>V</i>	LCOGT	57278.169	17.42 0.076	<i>B</i>	LCOGT
57294.521	17.13 0.074	<i>V</i>	LCOGT	57282.474	17.47 0.065	<i>B</i>	LCOGT
57304.919	17.22 0.074	<i>V</i>	LCOGT	57285.423	17.62 0.083	<i>B</i>	LCOGT
57307.104	17.18 0.047	<i>V</i>	LCOGT	57291.381	17.7 0.115	<i>B</i>	LCOGT
57307.138	17.18 0.056	<i>V</i>	LCOGT	57294.520	17.58 0.242	<i>B</i>	LCOGT
57312.170	17.2 0.068	<i>V</i>	LCOGT	57295.130	17.67 0.085	<i>B</i>	LCOGT
57315.877	17.28 0.082	<i>V</i>	LCOGT	57299.493	17.95 0.098	<i>B</i>	LCOGT
57315.894	17.14 0.073	<i>V</i>	LCOGT	57304.918	17.59 0.086	<i>B</i>	LCOGT
57316.016	17.24 0.075	<i>V</i>	LCOGT	57307.137	17.93 0.067	<i>B</i>	LCOGT
57319.057	17.23 0.059	<i>V</i>	LCOGT	57312.168	17.98 0.077	<i>B</i>	LCOGT
57321.794	17.14 0.077	<i>V</i>	LCOGT	57315.873	18.27 0.130	<i>B</i>	LCOGT

MJD	Magnitude	Filter	Telescope	MJD	Magnitude	Filter	Telescope		
57316.015	18.09	0.086	<i>B</i>	LCOGT	57339.331	17.81	0.271	<i>U</i>	<i>Swift</i>
57319.055	18.21	0.108	<i>B</i>	LCOGT	57342.066	17.84	0.231	<i>U</i>	<i>Swift</i>
57321.792	18.45	0.179	<i>B</i>	LCOGT	57345.787	18.08	0.201	<i>U</i>	<i>Swift</i>
57327.124	18.12	0.058	<i>B</i>	LCOGT	57348.377	17.82	0.161	<i>U</i>	<i>Swift</i>
57329.839	18.14	0.049	<i>B</i>	LCOGT	57259.325	14.60	0.042	<i>W1</i>	<i>Swift</i>
57332.786	18.31	0.099	<i>B</i>	LCOGT	57264.379	15.02	0.042	<i>W1</i>	<i>Swift</i>
57335.812	18.18	0.102	<i>B</i>	LCOGT	57268.719	15.30	0.05	<i>W1</i>	<i>Swift</i>
57344.411	18.25	0.077	<i>B</i>	LCOGT	57272.503	15.55	0.05	<i>W1</i>	<i>Swift</i>
57346.415	18.11	0.131	<i>B</i>	LCOGT	57276.085	15.89	0.058	<i>W1</i>	<i>Swift</i>
57349.414	18.27	0.076	<i>B</i>	LCOGT	57280.675	16.07	0.058	<i>W1</i>	<i>Swift</i>
57353.420	18.24	0.064	<i>B</i>	LCOGT	57285.411	16.24	0.076	<i>W1</i>	<i>Swift</i>
57355.030	18.35	0.071	<i>B</i>	LCOGT	57291.384	16.52	0.067	<i>W1</i>	<i>Swift</i>
57259.329	16.33	0.045	<i>B</i>	<i>Swift</i>	57294.579	16.66	0.067	<i>W1</i>	<i>Swift</i>
57264.384	16.73	0.045	<i>B</i>	<i>Swift</i>	57297.048	16.91	0.076	<i>W1</i>	<i>Swift</i>
57268.724	16.95	0.054	<i>B</i>	<i>Swift</i>	57300.226	16.91	0.076	<i>W1</i>	<i>Swift</i>
57272.507	17.26	0.063	<i>B</i>	<i>Swift</i>	57303.618	17.07	0.076	<i>W1</i>	<i>Swift</i>
57276.089	17.33	0.063	<i>B</i>	<i>Swift</i>	57306.483	17.14	0.076	<i>W1</i>	<i>Swift</i>
57280.679	17.46	0.063	<i>B</i>	<i>Swift</i>	57309.407	17.16	0.104	<i>W1</i>	<i>Swift</i>
57285.413	17.57	0.122	<i>B</i>	<i>Swift</i>	57312.128	17.41	0.095	<i>W1</i>	<i>Swift</i>
57291.388	17.67	0.092	<i>B</i>	<i>Swift</i>	57318.666	17.55	0.133	<i>W1</i>	<i>Swift</i>
57294.582	17.68	0.092	<i>B</i>	<i>Swift</i>	57324.047	17.98	0.192	<i>W1</i>	<i>Swift</i>
57300.230	17.71	0.092	<i>B</i>	<i>Swift</i>	57327.909	17.68	0.143	<i>W1</i>	<i>Swift</i>
57303.622	17.80	0.102	<i>B</i>	<i>Swift</i>	57330.416	17.63	0.163	<i>W1</i>	<i>Swift</i>
57306.487	17.68	0.092	<i>B</i>	<i>Swift</i>	57333.412	17.60	0.153	<i>W1</i>	<i>Swift</i>
57309.409	17.80	0.132	<i>B</i>	<i>Swift</i>	57336.132	17.89	0.133	<i>W1</i>	<i>Swift</i>
57312.132	17.68	0.102	<i>B</i>	<i>Swift</i>	57339.330	17.72	0.192	<i>W1</i>	<i>Swift</i>
57318.668	18.13	0.181	<i>B</i>	<i>Swift</i>	57342.064	18.06	0.202	<i>W1</i>	<i>Swift</i>
57324.049	17.91	0.171	<i>B</i>	<i>Swift</i>	57345.784	18.26	0.173	<i>W1</i>	<i>Swift</i>
57327.911	18.44	0.231	<i>B</i>	<i>Swift</i>	57348.374	18.11	0.143	<i>W1</i>	<i>Swift</i>
57330.418	18.08	0.221	<i>B</i>	<i>Swift</i>	57259.335	14.44	0.042	<i>M2</i>	<i>Swift</i>
57333.414	18.08	0.211	<i>B</i>	<i>Swift</i>	57264.394	14.87	0.042	<i>M2</i>	<i>Swift</i>
57336.135	18.17	0.181	<i>B</i>	<i>Swift</i>	57268.789	15.21	0.05	<i>M2</i>	<i>Swift</i>
57342.066	17.84	0.171	<i>B</i>	<i>Swift</i>	57272.515	15.36	0.05	<i>M2</i>	<i>Swift</i>
57345.788	18.26	0.171	<i>B</i>	<i>Swift</i>	57276.098	15.63	0.05	<i>M2</i>	<i>Swift</i>
57348.378	18.27	0.171	<i>B</i>	<i>Swift</i>	57280.689	15.79	0.05	<i>M2</i>	<i>Swift</i>
57259.328	15.25	0.045	<i>U</i>	<i>Swift</i>	57285.416	15.97	0.067	<i>M2</i>	<i>Swift</i>
57264.383	15.69	0.045	<i>U</i>	<i>Swift</i>	57291.591	16.18	0.05	<i>M2</i>	<i>Swift</i>
57268.722	15.99	0.045	<i>U</i>	<i>Swift</i>	57294.589	16.38	0.058	<i>M2</i>	<i>Swift</i>
57272.506	16.19	0.054	<i>U</i>	<i>Swift</i>	57300.237	16.67	0.058	<i>M2</i>	<i>Swift</i>
57276.088	16.51	0.063	<i>U</i>	<i>Swift</i>	57303.629	16.81	0.058	<i>M2</i>	<i>Swift</i>
57280.678	16.73	0.063	<i>U</i>	<i>Swift</i>	57306.495	16.93	0.058	<i>M2</i>	<i>Swift</i>
57285.412	16.99	0.122	<i>U</i>	<i>Swift</i>	57309.413	17.00	0.076	<i>M2</i>	<i>Swift</i>
57291.387	16.96	0.082	<i>U</i>	<i>Swift</i>	57312.138	17.07	0.067	<i>M2</i>	<i>Swift</i>
57294.581	17.26	0.102	<i>U</i>	<i>Swift</i>	57318.869	17.40	0.182	<i>M2</i>	<i>Swift</i>
57300.229	17.36	0.102	<i>U</i>	<i>Swift</i>	57324.052	17.59	0.114	<i>M2</i>	<i>Swift</i>
57303.621	17.47	0.112	<i>U</i>	<i>Swift</i>	57327.962	17.56	0.076	<i>M2</i>	<i>Swift</i>
57306.486	17.45	0.102	<i>U</i>	<i>Swift</i>	57330.421	17.41	0.114	<i>M2</i>	<i>Swift</i>
57309.408	17.67	0.161	<i>U</i>	<i>Swift</i>	57333.481	17.86	0.104	<i>M2</i>	<i>Swift</i>
57312.131	17.79	0.151	<i>U</i>	<i>Swift</i>	57336.142	17.82	0.085	<i>M2</i>	<i>Swift</i>
57318.667	17.46	0.161	<i>U</i>	<i>Swift</i>	57339.334	17.84	0.143	<i>M2</i>	<i>Swift</i>
57324.048	17.75	0.211	<i>U</i>	<i>Swift</i>	57342.069	17.84	0.124	<i>M2</i>	<i>Swift</i>
57327.910	17.68	0.181	<i>U</i>	<i>Swift</i>	57345.794	17.94	0.242	<i>M2</i>	<i>Swift</i>
57330.418	17.68	0.231	<i>U</i>	<i>Swift</i>	57348.385	17.92	0.095	<i>M2</i>	<i>Swift</i>
57333.413	17.79	0.221	<i>U</i>	<i>Swift</i>	57257.794	14.77	0.042	<i>W2</i>	<i>Swift</i>
57336.134	17.47	0.141	<i>U</i>	<i>Swift</i>	57259.33	14.83	0.042	<i>W2</i>	<i>Swift</i>

MJD	Magnitude	Filter	Telescope	MJD	Magnitude	Filter	Telescope
57268.725	15.21 0.076	W2	<i>Swift</i>	57312.133	16.79 0.058	W2	<i>Swift</i>
57272.508	15.48 0.042	W2	<i>Swift</i>	57318.668	16.94 0.085	W2	<i>Swift</i>
57276.091	15.68 0.042	W2	<i>Swift</i>	57324.049	17.16 0.095	W2	<i>Swift</i>
57280.681	15.91 0.042	W2	<i>Swift</i>	57327.911	16.94 0.143	W2	<i>Swift</i>
57285.413	16.01 0.058	W2	<i>Swift</i>	57330.418	17.42 0.114	W2	<i>Swift</i>
57291.389	16.12 0.05	W2	<i>Swift</i>	57336.136	17.29 0.076	W2	<i>Swift</i>
57294.583	16.17 0.05	W2	<i>Swift</i>	57339.332	17.47 0.124	W2	<i>Swift</i>
57300.231	16.35 0.05	W2	<i>Swift</i>	57342.067	17.39 0.104	W2	<i>Swift</i>
57303.623	16.42 0.058	W2	<i>Swift</i>	57345.789	17.51 0.085	W2	<i>Swift</i>
57306.488	16.57 0.058	W2	<i>Swift</i>	57348.379	17.60 0.085	W2	<i>Swift</i>
57309.41	16.56 0.067	W2	<i>Swift</i>				

All magnitudes and uncertainties are presented in the Vega system. Uncertainties are given next to the magnitude measurements. Data are not corrected for Galactic extinction.

Table A2. Swift XRT photometry of ASASSN-15oi.

MJD	Flux	Lower Uncertainty	Upper Uncertainty	Epochs Combined
57263.0	4.18	2.04	1.73	01 – 04
57279.0	6.80	2.04	1.73	05 – 09
57296.0	6.77	2.31	1.97	10 – 13
57306.0	8.32	2.98	2.35	14 – 16
57319.0	6.80	2.25	1.88	17 – 20
57335.0	7.79	2.20	1.88	21 – 24
57345.0	6.17	2.98	2.35	25 – 27

All X-ray fluxes and uncertainties are given in units of 10^{-14} ergs s^{-1} cm^{-2} . Data are not corrected for Galactic extinction. The “Epochs Combined” column indicates which *Swift* epochs were combined to obtain the listed measurement.

Table A3. Spectroscopic Observations of ASASSN-15oi.

UT Date	MJD	Telescope/Instrument	Exposure (s)
2015 September 04.22	57269.22	du Pont-2.5m/WFCCD	1 × 900
2015 September 12.20	57277.20	MDM-2.4m/OSMOS	1 × 600
2015 October 11.82	57306.82	Tillinghast-1.5m/FAST	1 × 2400
2015 October 12.82	57307.82	Tillinghast-1.5m/FAST	1 × 2100
2015 October 15.75	57310.75	Tillinghast-1.5m/FAST	1 × 1320
2015 October 17.12	57312.12	du Pont-2.5m/WFCCD	2 × 900
2015 November 07.06	57333.06	du Pont-2.5m/WFCCD	2 × 1200