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RADIO CONSTRAINTS ON HEAVILY OBSCURED STAR FORMATION WITHIN DARK GAMMA-RAY BURST HOST GALAXIES

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ABSTRACT

Highly dust-obscured starbursting galaxies (submillimeter galaxies and their ilk) represent the most extreme sites of star formation in the distant universe and contribute significantly to overall cosmic star formation beyond $z > 1.5$. Some stars formed in these environments may also explode as gamma-ray bursts (GRBs) and contribute to the population of “dark” bursts. Here we present Very Large Array wideband radio-continuum observations of 15 heavily dust-obscured *Swift* GRBs to search for radio synchrotron emission associated with intense star formation in their host galaxies. Most of these targets (11) are not detected. Of the remaining four objects, one detection is marginal, and for two others we cannot yet rule out the contribution of a long-lived radio afterglow. The final detection is secure, but indicates a star formation rate (SFR) roughly consistent with the dust-corrected UV-inferred value. Most galaxies hosting obscured GRBs are therefore not forming stars at extreme rates, and the amount of optical extinction seen along a GRB afterglow sightline does not clearly correlate with the likelihood that the host has a sufficiently high SFR to be radio-detectable. While some submillimeter galaxies do readily produce GRBs, these GRBs are often not heavily obscured—suggesting that the outer (modestly obscured) parts of these galaxies overproduce GRBs and the inner (heavily obscured) parts underproduce GRBs relative to their respective contributions to star formation, hinting at strong chemical or initial mass function gradients within these systems.

Key words: galaxies: starburst – gamma-ray burst: general – radio continuum: galaxies

Online-only material: color figures

1. INTRODUCTION

Long-duration gamma-ray bursts (GRBs; we exclude the physically distinct separate class of short-duration GRBs) are highly collimated and extremely luminous relativistic explosions produced during the core collapse of massive stars. These explosions are accompanied by extremely luminous multiwavelength afterglows (Meszaros & Rees 1997; Sari et al. 1998; van Paradijs et al. 2000) that pinpoint the locations of their host galaxies and often reveal their redshifts as well (e.g., Prochaska et al. 2007a, 2007b; D’Elia 2011). By virtue of their massive stellar origin, the GRB rate must also be closely tied with that of overall cosmic star formation. For these reasons, the study of GRBs and their hosts has aroused significant interest over the past 15 yr for its potential to address greater questions of galaxy evolution and cosmic history (Krumholz et al. 1998; Totani 1999; Ramirez-Ruiz et al. 2002; Lloyd-Ronning et al. 2002; Firmani et al. 2004).

While all GRBs appear to occur in star forming galaxies (e.g., Savaglio et al. 2009), it is not clear whether all types of star forming galaxies can produce GRBs. Most known GRB hosts tend to be blue, irregular, low-mass, metal-poor, and nearly dust-free; very few are spirals, have a large population of older stars, or are significantly dust-obscured (e.g., Bloom et al. 2002; Fruchter et al. 2006; Le Floch et al. 2006; Wolf & Podsiadlowski 2007; Modjaz et al. 2008; Castro Cerón et al. 2010; Levesque et al. 2010; Graham & Fruchter 2013). These trends have been interpreted as evidence that high-metallicity environments produce GRBs much less frequently or even not at all, an effect that would complicate the use of GRBs as a high- z star formation tracer. Alternatively, variation in the initial

mass function (IMF) or initial close-binary fraction might also result in a GRB population that seems to prefer certain types of star-forming galaxies while avoiding others. While GRBs undoubtedly provide a wealth of information about the high- z universe, placing this information into context will require a better empirical and physical understanding of how a galaxy’s internal properties affect its ability to produce GRBs.

While most GRBs are only mildly obscured (Jakobsson et al. 2004; Kann et al. 2006; Greiner et al. 2011), a significant minority encounter a large amount of dust along the line of sight within their host, making their optical afterglows difficult or impossible to detect. This type of event is commonly known as a “dark” GRB (Groot et al. 1998).⁴ Although the host galaxies of these events do tend to be more representative of the overall star-forming galaxy population than unobscured GRBs (Krühler et al. 2011; Perley et al. 2013), GRBs as a whole still appear to be underabundant in massive, reddened host galaxies compared to what would be expected for a perfect star formation tracer (Perley et al. 2013).

Nevertheless, the mere existence of highly dust-obscured events points toward another avenue by which to investigate the relation between GRBs and overall cosmic star formation (Djorgovski et al. 2001). A significant (if not necessarily dominant) amount of star formation beyond $z > 1.0$ – 1.5 occurs within extremely luminous galaxies with star formation rates (SFRs) several hundred times that of the Milky Way, almost all of which occurs behind a dust screen that is completely optically thick at optical and near-IR (NIR) wavelengths

⁴ In practice, the definition of a “dark” GRB is complex—the presence of dust extinction is only one of a large number of factors affecting the brightness of a GRB optical afterglow, and events are not followed up uniformly. A significantly expanded discussion of the various definitions of darkness can be found in Perley et al. (2013).

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Table 1
Summary of Previous Submm/Radio Host Detections

GRB	Redshift ^a	OA? ^b	Submm F_ν ^c (mJy)	Radio F_ν ^d (μ Jy)	Radio Freq. (GHz)	References ^e
980425 ^f	0.0085	Y		420 ± 50	4.8	M09
031203 ^f	0.105	Y		216 ± 50	5.5	S10
000210	0.8452	N	3.05 ± 0.76	18 ± 9	1.4	T04, B03
980703	0.967	Y		68 ± 6	1.4	B01
000911	1.0585	Y	2.31 ± 0.91			B03
021211	1.006	Y		330 ± 31	1.4	M08
000418	1.1185	Y	3.15 ± 0.90	59 ± 15	1.4	B03
010222	1.478	Y	3.74 ± 0.53	23 ± 8	4.9	T04, B03
000301C	2.034	Y		18 ± 7	8.5	B03
000926	2.066	Y		23 ± 9	8.5	B03
080607	3.038	Y	0.31 ± 0.09			W12
120804A ^g	~ 1.3	Y		43 ± 4	4.9	B13

Notes.

^a References: Prochaska et al. (2004); Piro et al. (2002); Djorgovski et al. (1998); Price et al. (2002); Vreeswijk et al. (2003); Bloom et al. (2003); Jha et al. (2001); Jensen et al. (2001); Fynbo et al. (2001); Berger et al. (2013).

^b Whether or not an optical afterglow was detected for this GRB.

^c At 850 μ m. The highest-confidence reported detection is given.

^d At the given frequency. The highest-confidence reported detection (at any radio frequency) is given.

^e References for submillimeter and radio flux density measurements. M09 = Michałowski et al. (2009); S10 = Stanway et al. (2010); T04 = Tanvir et al. (2004); B03 = Berger et al. (2003); B01 = Berger et al. 2001; M08 = Michałowski et al. (2008); W12 = Wang et al. (2012); B13 = Berger et al. (2013).

^f Low-redshift GRBs. Below $z \lesssim 0.2$ radio observations are sensitive to ordinary (modest-SFR, optically thin) star-forming galaxies.

^g GRB 120804A is a *Swift*-era event and has a duration of $T_{90} = 0.8$ s, consistent with a short-duration burst.

(Smail et al. 2004; Chapman et al. 2005; Dye et al. 2008; Michałowski et al. 2010; Wardlow et al. 2011; Yun et al. 2012). These galaxies—known as submillimeter galaxies (SMGs; Blain et al. 2002)—appear unassuming at UV/optical/NIR wavelengths, revealing their true nature only at long wavelengths (submillimeter/radio) where the dust screen becomes transparent.

The extreme conditions in these galaxies make them important laboratories for testing whether the GRB progenitor can form in the most extreme environments. For example, SMGs have high specific star formation rates (sSFRs; Daddi et al. 2007) but also are expected to be fairly metal-enriched (Nagao et al. 2012). The GRB host population appears overabundant in high-SFR galaxies but underabundant in high-metallicity galaxies compared to expectation for a uniform star formation tracer, so whether or not GRBs form frequently in SMGs represents an important test regarding which of these two factors is more closely associated with GRB production. There also have been suggestions that SMGs (or their possible $z = 0$ end-products, elliptical galaxies) exhibit an unusual IMF (e.g., Baugh et al. 2005; Conroy & van Dokkum 2012). Determining how often these systems host GRBs (if at all) is therefore of significant interest for understanding and applying GRBs as cosmological probes.

It would be reasonable to expect that searches for the hosts of optically bright GRBs should yield no detections at submillimeter or radio wavelengths beyond the nearby ($z \gtrsim 0.5$) universe: observations using the previous generation of instrumentation at these wavelengths are not sensitive to ordinary galaxies at these distances and can *only* detect the most extreme star-forming galaxies (SFR $\gtrsim 200 M_\odot \text{ yr}^{-1}$). Nearly all galaxies with SFRs this high are optically thick and do not allow significant amounts of UV/optical light to escape from their star-forming regions (i.e., UV-inferred SFRs almost never significantly exceed this

value regardless of selection criterion); this opacity should stifle the optical afterglow, as well as the stellar light. Undaunted by these pessimistic expectations, a number of submillimeter/radio GRB host surveys were conducted along primarily optically selected and therefore unobscured sightlines anyway—and while a majority of the hosts targeted were indeed not detected (e.g., Barnard et al. 2003; Berger et al. 2003; Tanvir et al. 2004; Le Floc’h et al. 2006; Priddey et al. 2006; Stanway et al. 2010; Hatsukade et al. 2012; Michałowski et al. 2012), a few optically bright bursts *were* actually localized to submillimeter-bright hosts (at least three: 980703, 000418, and 010222: Berger et al. 2001, 2003; Tanvir et al. 2004; see Table 1). The very large total SFRs of these systems inferred from the submillimeter/radio observations (vastly in excess of the unobscured SFR inferred from UV/optical observations; Michałowski et al. 2008) leave little doubt that the vast majority of the star formation in these galaxies is located in regions that are completely optically thick at UV/optical wavelengths. While two of these three galaxies showed moderate obscuration, the dust columns inferred ($A_V \sim 1$ mag in the cases of GRBs 980703 and 000418 and $A_V \sim 0.1$ mag in the case of GRB 010222; Kann et al. 2006) are far less than that required to conceal the extremely luminous starbursts inferred from the radio and submillimeter data, suggesting that these GRBs were not produced by the part of the galaxy responsible for forming most of its young stars.

While this certainly indicates that the *outer* regions of SMGs form GRBs quite readily, these results are much more ambiguous about the role of the heavily obscured inner region: only a few well-localized “dark” GRBs were known at the time these early surveys were conducted, and while one of these was found to occur within an SMG as well (GRB 000210), the sample of well-localized dark GRBs was far too small (and in many cases the actual evidence of dust as the cause of the optical nondetection too uncertain) to come to any definitive

conclusions about the nature of dark GRB host galaxies at long wavelengths.

The launch of the *Swift* satellite (Gehrels et al. 2004) with its onboard X-Ray Telescope has made it possible to unambiguously identify dust-obscured GRBs and accurately determine their position, and large numbers of “dark” GRB host galaxies are now well characterized at optical and NIR wavelengths (e.g., Krühler et al. 2011; Perley et al. 2013). However, until recently similar large efforts have not been possible at long wavelengths due to the limited sensitivity of radio and submillimeter instrumentation. Fortunately, observational capabilities at these wavelengths are rapidly improving. In particular, the upgraded Karl G. Jansky Very Large Array (VLA) employs new receivers and, critically, the WIDAR digital correlator, which is capable of processing up to 8 GHz of bandwidth simultaneously, an increase by a factor of 80 (Perley et al. 2011). At submillimeter wavelengths, even larger gains in the sensitivity to high- z galaxies are becoming available with the completion of the Atacama Large Millimeter Array (ALMA), as shown by the recent detection of one GRB host at $z = 3$ (Wang et al. 2012).

In this paper we present VLA continuum observations of 15 heavily obscured GRB host galaxies, seeking to constrain the fraction of “dark” GRBs that actually originate in luminous SMGs (versus dusty regions of more ordinary galaxies) and therefore the abundance of GRBs in this type of system. We summarize the target selection, observational strategy, and data reduction in Section 2. We examine our detected systems in more detail in Section 3 to explore and examine the implications of these radio detections for the nature of those objects. Finally, we translate these observational limits into constraints on the SFRs and properties of the host galaxies of our sample in Section 4, and we discuss our overall results and their implications for the environments of GRBs and their use as high- z SFR tracers.

2. OBSERVATIONS

2.1. Targets

In this study, we exclusively examine the hosts of “dark” GRBs—or more precisely, the hosts of GRBs whose afterglows were moderately or heavily obscured by dust as determined by combined X-ray, optical, and (often) NIR observations. Our attention is focused on this group for two reasons.

First, the observation of a heavily obscured GRB immediately implies the presence of at least some obscured star formation in its host. Heavily obscured star formation (that behind columns of $A_V > 3\text{--}4$ mag) cannot be traced with ordinary UV or optical techniques since the UV light that traces the young stars is absorbed almost completely, and SFRs inferred purely from optical/UV methods alone are inherently suspect if observations of a GRB point to the presence of stars behind optically thick dust columns within the galaxy.

Second, there is accumulating evidence that dust-obscured GRB hosts (even ones with relatively modest dust extinction columns of $A_V \sim 1\text{--}2$ mag) originate in galaxies with significantly higher average SFRs than ordinary hosts. The typical host galaxy in the dark GRB host study of Perley et al. (2013) has a UV-inferred SFR of $\sim 50 M_\odot \text{ yr}^{-1}$, and several may exceed $\sim 200 M_\odot \text{ yr}^{-1}$, which would make some of these systems feasible to detect with the VLA even without a substantial additional optically thick component. Moderately to heavily dust-obscured GRBs ($A_V > 1$ mag) also tend to be hosted in significantly more massive galaxies, which are much

more likely to be extremely high SFR submillimeter sources (Michałowski et al. 2008).

The target list is a subsample of the 23 galaxies presented in Perley et al. (2013), which were selected among all *Swift* events between 2005 and 2009 with clear evidence for an afterglow extinction of at least $A_V > 1$ mag. Nineteen of those galaxies are located above $\delta > -25^\circ$ and accessible to the VLA; of these we selected 15 for observations. This (modest) down-selection was chosen based on a combination of redshift and the (UV-inferred) SFR of the host, in order to further increase the prospects for detection. Hosts at $\delta > -25^\circ$ presented in Perley et al. (2013) but not observed here are those of GRBs 060319, 061222A, 070306, and 060814. The first three of these have low (UV) SFRs and high redshifts ($\text{SFR} < 10 M_\odot \text{ yr}^{-1}$ and $z > 1$), similar to more typical GRB host galaxies at these redshifts, and seemed less promising initial targets than the other galaxies.⁵ The host of GRB 060814 (which has a very high UV SFR) was omitted from the sample on the basis of the apparent superposition of the host with a foreground galaxy at the time of the proposal, although subsequent higher-resolution observations successfully resolved the two components (Hjorth et al. 2012; Perley et al. 2013).

2.2. VLA Observations

All of our observations were carried out with the upgraded VLA between 2011 May and 2012 July. We observed using the C-band (4–8 GHz) receivers, configured to position the two central frequencies at 4.715 GHz and 5.739 GHz with 1.024 GHz of bandwidth each, providing effectively continuous coverage across 2.048 GHz of bandwidth centered at 5.227 GHz. Observations were taken with 2 s (2011) or 4 s (2012) averaging and interleaved with observations of a nearby phase calibrator approximately every 5 minutes.

The first set of observations were carried out during the summer of 2011, mostly in A configuration (resolution of $0''.45$), but with some observations conducted during the reconfiguration of the array from B to A, with most antennas in their A-configuration positions but a few on the east and west arms still in B configuration. Ten sources were observed in total during this period, with typical on-source times of approximately 1 hr producing typical rms sensitivities of 5–9 μJy .

A second set of observations were taken the following year in the B configuration (resolution $1''.4$). We re-observed four sources for which our reduction of the A-configuration data showed weak detections to significantly greater depth (an additional 3 hr per source), which confirmed three of these detections. In addition, we observed five sources that we had not previously imaged in A configuration for about 1 hr each.

Data reduction was carried out using the Astronomical Image Processing System (AIPS). Radio frequency interference was minor in all observations and generally removed by clipping outlier visibilities above a minimum flux density threshold (in all cases this step removed less than 10% of the data). In most cases we elected not to observe a primary flux calibrator and to instead calibrate using the switched power system, which injects a calibrated signal pattern into the data (Perley 2010 and work in

⁵ Of course, these remaining systems are also interesting in a different way, in that the detection of an obscured GRB from an unobscured galaxy directly indicates a very heterogeneous host; their blue colors and low-UV SFRs are also closer analogs of the pre-*Swift* submillimeter-detected hosts. Nevertheless, for the reasons previously outlined, for this study we elected to focus on the population of more massive, higher UV SFR hosts, leaving the remaining events for future work.

Table 2
VLA Observations

GRB Target	R.A.	Decl.	Configuration	Date	Int. Time (min)	rms Noise ($\mu\text{Jy beam}^{-1}$)
051008	13:31:29.550	+42:05:53.30	B	2012 Jun 11	70	3.4
051022	23:56:04.115	+19:36:24.04	AB	2011 Jun 6	30	3.2
			A	2011 Aug 13	49	
			B	2012 May 26, 2012 Jul 19	194	
060202	02:23:23.010	+38:23:03.20	B	2012 Jun 23	46	5.4
060306	02:44:22.910	-02:08:54.00	B	2012 Jun 23	48	5.9
060923A	16:58:28.160	+12:21:38.90	AB	2011 Jun 5	36	5.6
			A	2011 Aug 21	23	
070521	16:10:38.620	+30:15:22.40	AB	2011 Jun 5	36	4.0
			A	2011 Aug 21	24	
071021	22:42:34.310	+23:43:06.50	AB	2011 Jun 6	30	3.8
			A	2011 Aug 13	38	
080207	13:50:02.980	+07:30:07.40	A	2011 Jul 17/18	69	2.4
			B	2012 Jun 5	187	
080325	18:31:34.230	+36:31:24.80	AB	2011 Jun 5	36	2.6
			A	2011 Aug 21	23	
			B	2012 Jun 15, 2012 Jun 24	181	
080607	12:59:47.221	+15:55:10.86	A	2011 Jul 17/18	71	3.5
081221	01:03:10.168	-24:32:51.67	B	2012 Jun 23	46	5.9
090404	15:56:57.520	+35:30:57.50	AB	2011 Jun 5	36	2.6
			A	2011 Aug 21	23	
			B	2012 May 29, 2012 Jun 3	189	
090407	04:35:54.980	-12:40:45.50	A	2011 Jun 18	55	4.4
090417B	13:58:46.590	+47:01:05.00	B	2012 Jun 11	68	4.0
090709A	19:19:42.640	+60:43:39.30	AB	2011 Jun 5	36	5.1

preparation). Comparisons of calibrations based on this system versus a standard flux-calibration procedure using 3C 286 and 3C 48 were found to be consistent (to within 5%), and we expect the flux calibration not to be a significant source of error in our observations.

All observations were carried out at least 2 yr (and more typically 4–5 yr) after the GRB occurred, when the radio afterglow for typical GRBs has faded well below detectability, even to the upgraded VLA (Chandra & Frail 2012). However, the very brightest GRBs do remain detectable for several years (e.g., Figure 25 of that work), which must be considered when assessing our putative detections.

All observations are summarized in Table 2.

2.3. Photometry

As our images have resolution comparable to (and often finer than) the positional uncertainty of the GRB position and the intrinsic size of the host galaxy whose flux we are attempting to measure, some care must be taken in calculation of flux densities (or upper limits) for these objects. We provide these measurements by two different procedures.

First, we determine a *point-source* flux density (or limit) on any emission underlying the position of the GRB. This is calculated by measuring the flux density value of the brightest cell in the synthesized map within $0''.4$ of the afterglow location (or within the optical disk of the host galaxy if no precise afterglow position was available, using the images presented in Perley et al. 2013). Physically, this corresponds to the flux (or limiting flux) associated with any “compact” (on the spatial scale of the resolution of the map) star-forming region that may have produced the GRB, or on any late-time afterglow emission. The confidence of the detections is evaluated by calculating the peak flux density repeatedly within 1000 equivalent-sized

regions drawn randomly from blank regions across the full synthesized map.

More importantly, we also desire a measurement (or limit) of the *integrated* flux density of our hosts, corresponding to the flux emitted by all star formation within the galaxy. To calculate this, we further convolve the image using a Gaussian convolution kernel with FWHM set to the semimajor axis of the galaxy as measured in the optical/NIR images (see Figure 1). The flux density is determined by choosing the value of the maximum cell within that aperture, and its uncertainty is estimated by repeating this exercise on random regions across the map.

The flux densities of the targets, presented using both interpretations, are presented in Table 3. For most sources the estimates provide very similar results, since most of the data were taken in B configuration at resolution ($\sim 1''.4$) similar to the actual angular extent of the host galaxies, so the effect of the convolution is negligible. However, for the A-configuration-only data in particular ($0''.45$ original resolution), the difference is significant, and caution should be used in interpreting the result depending on the desired goals.

3. RESULTS

The majority of our fields (11 out of 15) resulted in no detection at the location of the GRB or its host. In three of the remaining cases, we do detect clear (but weak) emission underlying the GRB/host positions: GRBs 051022 (4.2σ detection), 080207 (7.0σ detection), and 090404 (3.9σ detection). A fourth detection (of GRB 070521) is marginal (2.7σ in the convolved image and a peak 3.1σ in the unconvolved map, although the significance in the latter case is low because of the large area searched over in relation to the beam size).

Before examining the implications of our detections, it is essential to interpret their origin. While the goal of our survey

Table 3
Host Galaxy 5.23 GHz Radio Photometry

GRB	Point Source				Integrated			
	Beam Size ^a (")	Aperture ^b (")	F_ν^c (μJy)	Conf. ^d	2σ Limit ^e (μJy)	F_ν^c (μJy)	Conf. ^d	2σ Limit ^e (μJy)
051008	1.54×1.34	0.75	7.1 ± 3.3	0.86	13.7	5.0 ± 3.6	0.69	12.2
051022	1.41×1.40	0.80	12.5 ± 3.6	0.988		13.3 ± 3.6	0.988	
060202	1.65×1.32	1.0	5.1 ± 5.8	0.25	16.7	3.4 ± 6.9	0.29	17.2
060306	1.80×1.27	1.05	13.8 ± 6.1	0.84	26.0	14.0 ± 8.6	0.79	31.2
060923A	0.56×0.45	1.1	13.1 ± 4.2	0.92	21.5	25.0 ± 12.7	0.90	50.4
070521	0.57×0.45	1.1	12.3 ± 4.0	0.87	20.3	28.0 ± 10.3	0.983	48.6
071021	0.48×0.44	0.75	10.9 ± 3.7	0.91	18.3	12.0 ± 6.7	0.75	25.4
080207	1.45×1.24	1.0	15.5 ± 2.3	0.9999		17.1 ± 2.5	0.9999	
080325	1.60×1.29	0.8	6.8 ± 2.7	0.94	12.2	6.4 ± 3.0	0.91	12.4
080607	0.47×0.43	0.8	7.5 ± 3.4	0.54	19.3	5.6 ± 6.1	0.36	17.8
081221	2.67×1.26	0.8	6.4 ± 5.4	0.57	17.2	6.4 ± 5.4	0.57	17.2
090404	1.45×1.30	0.8	10.3 ± 2.4	0.999		10.9 ± 2.7	0.998	
090407	0.61×0.40	0.75	7.5 ± 4.2	0.31	15.9	-0.1 ± 7.6	0.07	15.1
090417B	1.55×1.37	1.0	6.6 ± 4.1	0.62	14.8	7.0 ± 4.8	0.73	16.1
090709A	0.88×0.50	0.45	9.0 ± 5.0	0.61	19.0	1.2 ± 6.6	0.26	14.4

Notes.

^a Beam size of the optically weighted image (major and minor axes).

^b Effective aperture used to convolve the image to produce the integrated measurements.

^c Maximum flux density (at 5.23 GHz) in any $0''.1$ synthesized cell consistent with the position of the optical/NIR host galaxy disk. The optimally weighted image was used for the point-source measurements; the convolved image was used for integrated measurements.

^d Significance of the detection, based on placing a large number of apertures of identical size randomly across the image and calculating the maximum flux density in each one.

^e 95% confidence upper limit on the host flux density.

was to search for radio emission associated with intense, dust-enshrouded star formation, alternative possibilities must also be considered: detection of foreground/background sources, detection of AGN activity from the host, or detection of the GRB afterglow.

3.1. AGN, Afterglow, or Host Galaxy?

The probability of a given point on the sky intersecting an unassociated foreground or background radio source is very low: even at the μJy level the radio sky is mostly empty, with a source density (based on inspection of our own fields) of only two sources stronger than $10 \mu\text{Jy arcmin}^{-2}$. The probability of chance detection of an unassociated source at this level centered within $r = 1''$ of a random position on the sky is therefore approximately $P_{\text{chance}} \sim r^2 / (\pi \rho_{f > 10 \mu\text{Jy}}^2) \sim 10^{-3}$. For a sample of 15 positions, the probability of one or more chance alignments within the sample is still low ($P_{\text{chance}} = 1.5 \times 10^{-2}$), and the chance of observing three essentially negligible ($< 1 \times 10^{-7}$). Our detected sources therefore in all cases almost certainly are associated with the GRB or its host galaxy in some way.

An AGN origin is also unlikely. Statistically, most unresolved faint radio sources are star-forming galaxies (Kimball et al. 2011; Condon et al. 2012); galaxy central black holes either tend to be much brighter (radio-loud) or are not detectable at all. The chance of a given GRB occurring within a galaxy whose AGN radio flux density happens to land within the relatively narrow range expected for a star-forming galaxy is low: based on Figure 11 of Condon et al. (2012), a detected cosmological source with a measured flux density of $< 100 \mu\text{Jy}$ has about a 90% probability of being a star-forming galaxy and not an AGN based on the relative abundance of the two populations.

While all observations were taken long after the GRB event (elapsed times range from 792 to 2408 days for observations

of the four detected sources) and detection of an afterglow this late would be almost unprecedented (only two GRBs, 980703 and 030329, have reported detections after > 800 days; Chandra & Frail 2012), the large increase in sensitivity of the VLA has rendered it possible to detect afterglows for far longer than was possible in the past. To estimate an average statistical likelihood that an afterglow would remain bright enough at this epoch to be detected in our data, we acquired the database of all radio afterglow observations compiled by Chandra & Frail (2012) and calculated the flux density (or its limiting value) at late times by extrapolating the light curve following the last detection or upper limit as $F \propto \text{const}$ ($t < 30$ days) or $F \propto t^{-1}$ ($t > 30$ days), a simple model that matches most events relatively well (within a factor of 2–3). Based on this exercise, we estimate that between 10% and 20% of GRB afterglows are expected to have a flux density of $> 10 \mu\text{Jy}$ even after 1000 days—a small minority, to be sure, but as 15 GRBs were observed during this project the detection of one or even three afterglows at this level would not be surprising.

If earlier radio data are available, we can extrapolate the light curve and estimate the afterglow contribution to the late-time flux directly. Otherwise, there are two possible means of distinguishing an afterglow origin from emission from the galaxy. First, the spectral indices may differ: star-forming galaxies exhibit a relatively narrow range of spectral indices, between $\alpha = -1$ and -0.5 (Condon 1992; we define the spectral index as $F_\nu \propto \nu^\alpha$). A GRB afterglow can exhibit a much wider range of spectral indices, depending on the position of the various characteristic frequencies relative to the radio band, especially the self-absorption frequency ν_a (see, e.g., Sari et al. 1998 or Granot et al. 1999). In particular, a self-absorbed afterglow ($\nu < \nu_a$) will exhibit a steep spectral index of $\alpha = 2-2.5$; although if the afterglow is still optically thin at late times, the spectral index is less distinct, $\alpha = 0.33$

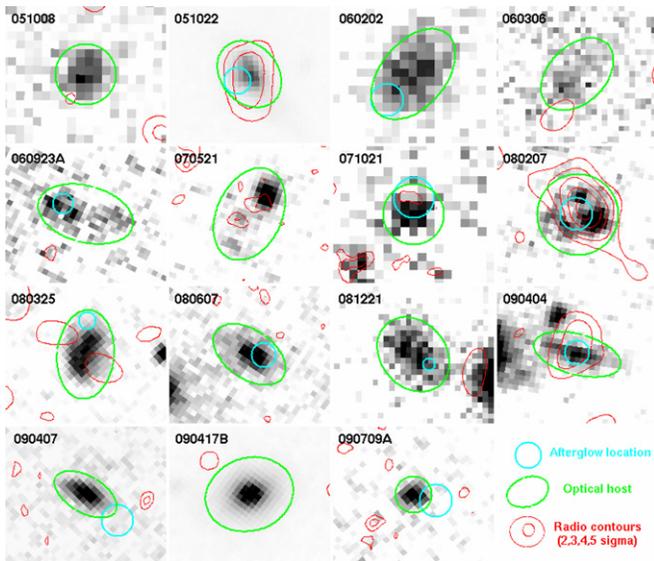


Figure 1. Radio flux density contours (2σ plus increments of 1σ) from VLA C-band observations of 15 dust-obscured GRB host galaxies (convolved with the default restoring beam), superimposed on optical or NIR imaging of the hosts from Perley et al. (2013), and uncertainty circles showing available sub-arcsecond afterglow positions (where available). Observations of 060923A, 070521, 071021, 080607, and 090407 were taken in A configuration (and/or a nonstandard A/B configuration hybrid) and have a resolution of approximately $0''.45$ FWHM; the remaining observations were taken mostly in B configuration and have a resolution of approximately $1''.4$. The 1σ sensitivity ranges from $2.5 \mu\text{Jy}$ to $6.0 \mu\text{Jy}$. Emission at the host positions of GRBs 051008, 080207, and 090404 is detected at $>3.5\sigma$; emission at GRB 070521 is detected only marginally.

(A color version of this figure is available in the online journal.)

to -1.0 (depending mostly on the location of synchrotron frequency ν_m). Alternatively, a GRB afterglow should always be unresolved (at VLA resolution) and be coincident (within astrometric accuracy) with earlier detections of the afterglow, while a host galaxy may be extended and/or offset from the early-time afterglow position.

With these factors in mind, we critically examine each of our four putative detections individually.

3.2. GRB 051022

GRB 051022 was observed by the VLA and the Westerbork Synthesis Radio Telescope (WSRT) on several occasions within the first week after the burst, so its late-time flux can be estimated directly. While the radio flux is strong at early times, it declined rapidly in subsequent observations (Cameron & Frail 2005; Rol et al. 2007; Chandra & Frail 2012); the 8.4 GHz light curve fades as approximately t^{-1} even as early as 1 day. Assuming that this decay continues, the flux density should not be more than $1 \mu\text{Jy}$ at the time of our radio observation (2053–2408 days post-burst). The more detailed broadband modeling in Rol et al. (2007) similarly predicts that even as soon as 30 days after the GRB the flux should have dropped below the level of our putative detection. As a result, the source detected in our late-time observations is almost certainly not associated with the GRB afterglow and, therefore, is likely associated with the host galaxy.

The source also appears extended: we re-weighted the data within the IMAGR task to produce images of varying resolution between $1''.5$ and $0''.4$; the flux is seen to drop (by about 2σ) at progressively higher resolution. This effect can also be seen as some limited north-south extension in the image even at the

standard resolution. This further supports a host-galaxy origin of the detected emission.

3.3. GRB 070521

No highly significant point source appears at the location of GRB 070521 in the synthesized map at the array’s native resolution (point-source limit $F < 20 \mu\text{Jy}$ at any location consistent with the optical disk), but in the convolved image a marginal detection (98% confidence) appears with an apparent flux density of $F = 28 \pm 10 \mu\text{Jy}$. GRB 070521 was not previously observed at long wavelengths, and so its late-time afterglow flux is unconstrained, although the contribution of any point source to the detection is likely minor: the flux density value at the centroid of the radio detection is only $9 \pm 4 \mu\text{Jy}$ in the unconvolved map. If the source is real at all, then it is probably also extended and therefore associated with the host galaxy. We tentatively suggest that this source also likely represents a detection, but further observations would be necessary to unambiguously establish its reality. We note that the location of the source is offset slightly from the brightest part of the host in the *Hubble Space Telescope* image, at the location of an apparent extension stretching southeast of the main disk.

3.4. GRB 080207

Unfortunately, no early-time radio observations of GRB 080207 were carried out, so we cannot directly constrain the afterglow brightness at the time of our observations. The relatively faint and fast-fading X-ray light curve—the GRB is not detected by XRT after 2 days (Evans et al. 2009)—is not suggestive of a bright late-time radio flux, but this is far from definitive.

To attempt to constrain the spectral index of the source, we split the observations into two frequency bands: 4.828–5.468 GHz and 5.980–6.620 GHz. The source is detected individually (at lower significance) in both of these windows separately, with a flux density of $18.9 \pm 6.6 \mu\text{Jy}$ at 5.15 GHz and $12.7 \pm 6.3 \mu\text{Jy}$ at 6.30 GHz. Unfortunately, this imposes only a relatively weak constraint on the spectral index of the source ($\alpha = -2.0 \pm 3.9$) and does not rule out either model.

To evaluate the spatial extent and location of the source, we created an additional map with a (non-optimal) resolution of $0''.75 \times 0''.75$ (FWHM). The point-source flux in this map is slightly less than, but statistically consistent with, the flux in the default-resolution map, suggesting a pointlike source. Its location ($\alpha = 13:50:02.96$, $\delta = +07:30:07.4$) is within the uncertainty region of the *Chandra* location (Hunt et al. 2011; Svensson et al. 2012), and roughly halfway between the brightest optical component of the host and the fainter northern component. As a result, we cannot unambiguously distinguish the origin of this source, although if it is produced by the host galaxy itself, the emission must be substantially more concentrated than the host’s optical light.

The host galaxy of GRB 080207 has been detected at $24 \mu\text{m}$ with MIPS (Hunt et al. 2011; Svensson et al. 2012). The $24 \mu\text{m}$ flux is commonly used as an independent measure of the obscured SFR of a galaxy, and the inferred value of several hundred $M_\odot \text{yr}^{-1}$ (Hunt et al. 2011; Svensson et al. 2012) is consistent with the notion of this being a heavily dust-obscured and rapidly star-forming system. Star formation therefore almost certainly contributes at least somewhat to the radio flux observed, although as the inferred radio SFR is even larger than this ($\sim 850 M_\odot \text{yr}^{-1}$; Section 3.6), an afterglow-dominated origin is not ruled out.

3.5. GRB 090404

GRB 090404 was also not observed in the radio at early times. It was, however, observed and detected in the millimeter band with the Plateau de Bure Interferometer (de Ugarte Postigo et al. 2012). The reported 87 GHz flux density from this observation is $660 \mu\text{Jy}$ at 3.4 days, which (assuming $\nu_a < \nu < \nu_m$) implies a radio flux density of $\sim 270 \mu\text{Jy}$ at this time. Assuming evolution similar to our generic radio light curve described earlier (flat evolution for the first 30 days, then decay as t^{-1}), we would expect a flux density of $\lesssim 10 \mu\text{Jy}$ at the initial observation epoch of 792 days. While the actual afterglow contribution could easily be much less than this (in particular if $\nu < \nu_a$), again we cannot unambiguously rule out an afterglow origin to the observed detection.

Higher-resolution versions of the map produced no significant change in source flux, indicating a point-like origin (although the weak detection renders this statement not strongly constraining). While the VLA detection is slightly north of the reported millimeter position, the offset is not significant given the uncertainties in both positions. Similarly, attempts to subdivide the data by frequency did not produce a useful constraint on the spectral index. As with GRB 080207, whether this source is associated with the host or an afterglow is ambiguous.

3.6. Star Formation Rates

The goal of this project was to constrain the fraction of obscured GRBs that originate from galaxies with extreme SFRs: ultra-luminous infrared galaxies (ULIRGs), SMGs, and similar classes of extreme high- z star-forming objects. To accomplish this, it is necessary to convert our flux measurements (or limits) into constraints on the host SFR.

Long-wavelength radio emission from normal galaxies is primarily generated by electrons initially accelerated by supernova remnants, but subsequently diffused into the surrounding regions of the galaxy to radiate their energy over timescales of approximately 10^7 yr via synchrotron emission. The radio flux of a galaxy can be tied to the supernova rate, and therefore the SFR (Condon 1992), allowing radio observations to serve as a star formation indicator. Specifically, our radio SFRs are estimated using Equation (17) of Murphy et al. (2011)⁶:

$$\left(\frac{\text{SFR}_{1.4\text{GHz}}}{M_\odot \text{yr}^{-1}}\right) = 6.35 \times 10^{-29} \left(\frac{L_{1.4\text{GHz}}}{\text{erg s}^{-1} \text{Hz}^{-1}}\right). \quad (1)$$

Since our observations are not at 1.4 GHz, we also have to appropriately k -correct our observations based on an assumed spectral index α . As discussed above, for none of our objects do we have a robust measurement of the actual spectral index, and so we assume a canonical average value of $\alpha = -0.75$. Incorporating this (and placing the equation in terms of the observed flux density in μJy), we employ the following relation:

$$\left(\frac{\text{SFR}_{\text{radio}}}{M_\odot \text{yr}^{-1}}\right) = 0.072 \left(\frac{F_\nu}{\mu\text{Jy}}\right) (1+z)^{1-\alpha} \left(\frac{d_L}{\text{Gpc}}\right)^{-2} \left(\frac{\nu}{\text{GHz}}\right)^{-\alpha}. \quad (2)$$

Using this equation, we convert our observed flux densities into measurements (or limits) on the SFR of each galaxy, assuming that all of the detected emission originates from the host.

⁶ Murphy et al. (2011) assume a Kroupa (2001) IMF, instead of the Chabrier IMF employed in Perley et al. (2013) to calculate the optical SFRs. However, these two forms of the IMF are very similar, and for the purposes of this paper the distinction is negligible.

Table 4
Host Galaxy Star Formation Rates

GRB	Redshift	UV SFR ^a ($M_\odot \text{yr}^{-1}$)	Radio SFR ^b ($M_\odot \text{yr}^{-1}$)	Point-source Limit ^c ($M_\odot \text{yr}^{-1}$)
051008	$2.90^{+0.28}_{-0.15}$	72^{+26}_{-54}	<1180	
051022	0.809	26^{+7}_{-7}	74 ± 20	
060202	0.785	$5.8^{+1.1}_{-2.0}$	<88	
060306	1.551	245^{+130}_{-67}	<790	<560
060923A	$2.50^{+0.58}_{-0.52}$	89^{+38}_{-30}	<4140	<1765
070521	$1.70^{+1.04}_{-0.36}$	40^{+62}_{-3}	817 ± 300	<592
071021	2.452	190^{+26}_{-20}	<1828	<1320
080207	2.086	46^{+272}_{-45}	846 ± 124	
080325	1.78	13^{+5}_{-4}	<500	
080607	3.038	19^{+7}_{-5}	<2070	<1661
081221	2.26	173^{+23}_{-30}	<1030	
090404	$3.00^{+0.83}_{-1.82}$	99^{+122}_{-99}	1230 ± 305	
090407	1.448	28^{+15}_{-10}	<325	
090417B	0.345	$0.5^{+0.3}_{-0.3}$	<12	
090709A	$1.80^{+0.46}_{-0.71}$	$8.0^{+4.1}_{-4.1}$	<653	

Notes.

^a From the SED fitting procedure of Perley et al. (2013), including dust correction.

^b Calculated using the integrated flux density from Table 3 assuming that all emission originates from the host galaxy. Uncertainties do not include the uncertainty in the photometric redshift derivations (051008, 060923A, 070521, 090404, and 090709A).

^c 2σ limit on a point-source contribution (from, e.g., nuclear starburst), calculated using the point-source flux in Table 3. Only listed if it is significantly less than that calculated from the integrated flux; otherwise, the integrated flux applies to both cases.

As discussed earlier, in the case of nondetections we provide separate limits (95% confidence) on unresolved star formation and on integrated star formation for each galaxy. These results are given in Table 4. Typical limits range from $\sim 12 M_\odot \text{yr}^{-1}$ (for GRB 090417B at $z = 0.35$) to $\sim 1000 M_\odot \text{yr}^{-1}$ for targets at $z > 2.5$. Note that our uncertainty estimates do not include the impact of uncertainty in the photometric redshifts for those sources for which a spectroscopic redshift is not yet available (effects that are particularly important for GRB 090404, which appears to be at $z \sim 3$ but could be consistent with a lower redshift).

4. DISCUSSION

4.1. Comparing Optical and Radio Star Formation Rates

The primary goal of our study is to determine the contribution of optically thick star formation to the total. Our radio continuum observations measure the *total* (optically thick + optically thin) SFR of each galaxy, which can then be compared to the UV/optical values that (even with dust attenuation included in the modeling) probe *only* optically thin regions. A significant discrepancy between the two estimates (as seen in SMGs) would indicate the presence of substantial additional optically thick extinction.

Our sole unambiguous host-galaxy detection, GRB 051022, has a radio SFR of $74 \pm 20 M_\odot \text{yr}^{-1}$. This is higher, but not much higher, than our optically derived estimates: our UV-based estimate for this source (from spectral energy distribution (SED) fitting including dust absorption; Perley et al. 2013) is $26 \pm 7 M_\odot \text{yr}^{-1}$. The two values are marginally inconsistent

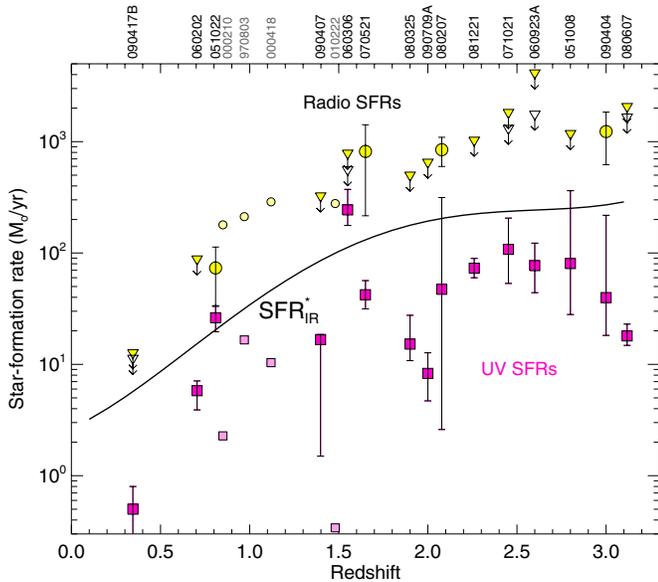


Figure 2. SFRs of targeted galaxies as a function of redshift. Magenta squares indicate reddening-corrected UV/optical SFRs from the SED fitting procedure of Perley et al. (2013). Yellow points indicate our VLA measurements of the SFR of these galaxies (including the contribution from optically thick regions). For comparison, we also show the four probable submillimeter-galaxy hosts known from pre-*Swift* work (values from Savaglio et al. 2009 and Michałowski et al. 2008) as smaller, pale symbols. The x -axis positions of a few targets have been offset slightly from their actual redshift values for clarity. The solid curve shows the SFR of an L_* galaxy (from the IR luminosity functions of Pérez-González et al. 2005 and Lapi et al. 2011) as a function of redshift.

(A color version of this figure is available in the online journal.)

(by 2.3σ), hinting at the presence of modest additional, optically thick star formation present within this galaxy not apparent in optical observations, even with the effects of dust reddening considered. Given systematic uncertainties in both SFR estimators, it is not certain that this marginal excess is real, although since the GRB itself exploded in a highly optically thick region, it seems likely that at least some optically thick star formation must be present. In any case, the optically thick SFR is modest in comparison to SMGs and previous radio-detected GRB hosts, where the optically thick star-forming regions dominate the optically thin ones by an order of magnitude or more.

If interpreted as host-galaxy emission (and not afterglow), the detections of GRBs 080207 and 090404 correspond to SFRs of $850 M_\odot \text{ yr}^{-1}$ and $1230 M_\odot \text{ yr}^{-1}$, respectively (assuming $z = 3$ for GRB 090404). These are, indeed, greatly in excess of UV/optical estimates and would be indicative of large optically thick star-forming regions that presumably also produced the GRB: no NIR afterglow was detected in either case. Of course, our limits would be consistent with the optically inferred SFRs in the event of an afterglow origin. Similarly, the possible detection of GRB 070521 would imply an SFR ($\sim 850 M_\odot \text{ yr}^{-1}$) vastly in excess of the optical value if confirmed by further observations.

The remaining upper limits are generally well in excess of the optically determined SFRs (by approximately a factor of 10 in most cases; see Figure 2)—consistent with the absence of a significant optically thick component, but only formally ruling out models highly dominated by such an obscured component.

We have also compared the radio SFRs (or limits) inferred here with dust-*uncorrected* SFRs from the UV–optical SED fitting, by taking the Perley et al. (2013) values and re-extinguishing them using the Calzetti extinction law (the same extinction law used in the original fits). Among our

putative detections, we find $\text{SFR}_{\text{radio}}/\text{SFR}_{\text{UV}} = 12, 500, 2000,$ and 2200 for GRBs 051022, 090404 (at $z = 3$), 070521, and 080207, respectively, clearly emphasizing that these are extremely dusty galaxies (typical ratios are 1–10 for “normal” low- z galaxies, 10–500 for local LIRGs, and 1000 for Arp 220; Howell et al. 2010). Upper limits for our nondetections range from 100 (GRB 051008) to 2000 (GRB 060923A).

4.2. Implications of a Low Detection Fraction

Regardless of interpretation for the source of the emission seen in the cases of GRBs 080207 and 090404, it is clear that we do not detect most objects in the sample. There is no suggestion of a detection (at 95% confidence, based on photometry of random blank locations) from either afterglow or host galaxy for 11 out of 15 fields, with typical flux density limits of $F_\nu < 15\text{--}20 \mu\text{Jy}$ (for a compact source; A-configuration-only observations have limits 2–3 times this for an extended object).

To some extent, a low radio detection fraction is to be expected: even after its upgrade, the VLA is only sensitive to the most luminous galaxies at any given redshift beyond about $z > 0.2$. To quantify these expectations, we used the observed luminosity functions for cosmological galaxy populations selected on the basis of obscured star formation (mid-IR and far-IR field surveys) to calculate an *anticipated* detection fraction, under the assumption that obscured GRBs trace obscured star formation uniformly.

Infrared luminosity functions are drawn from the Schechter (1976) function fits of Pérez-González et al. (2005) (*Spitzer* rest-frame $12 \mu\text{m}$, for $z < 1.5$) and from Lapi et al. (2011) (*Herschel/ATLAS* rest-frame $100 \mu\text{m}$, for $z > 1.5$; we performed our own Schechter fits to their data with α fixed to -1.2). We then convert the L^* parameter of the $12 \mu\text{m}$ or $100 \mu\text{m}$ luminosity function to an equivalent SFR at each redshift using the $F_{\text{IR-to-SFR}}$ conversions as detailed in each of the two reference works.⁷ (This SFR^* is plotted as the solid curve in Figure 2.) We then convert this SFR^* to an equivalent observed 5 GHz flux density $F_{\nu,\text{radio}}^*$ at each redshift using the Murphy et al. (2011) relation. The value of the Schechter parameter α is unaltered by the conversions (since the scaling is linear to a good approximation) and the normalization is unimportant, so this converts the mid-IR/far-IR luminosity Schechter function, $\phi_{L(\text{IR})}(z) \propto (L/L^*)^\alpha e^{-L/L^*}$, to a predicted radio flux density function, $\phi_F(z) \propto (F/F^*)^\alpha e^{-F/F^*}$.

Galaxies with brighter radio fluxes have proportionally higher SFRs, so the relative fraction of obscured star formation at a given redshift occurring in all galaxies brighter than our typical flux density limit $F_{(\text{obs})}$ is given by a normalized integral of ϕ_F weighted by the radio flux density:

$$\text{frac}_{\text{SFR}}(F > F_{\text{obs}}) = C^{-1} \int_{F_{\text{obs}}}^{+\infty} F \phi_F dF$$

(C is a normalization constant equal to the integral of the term on the right from $-\infty$ to $+\infty$).

This value is redshift dependent, though only by a factor of a few over the range of our sample: at higher z the increase in importance of highly star-forming galaxies partially compensates for flux dimming. We find that only a small, albeit non-negligible, fraction of obscured SFR (2%–5% at

⁷ The conversions employed by Pérez-González et al. (2005) assume a Salpeter IMF, so we convert to Chabrier by scaling the SFR down by a correction factor of 1.6.

$0.6 < z < 2$; more at lower redshifts and less at higher redshifts) is in radio-detectable galaxies to an average limit of $\sim 15 \mu\text{Jy}$. Detecting only one galaxy in the sample would be entirely consistent with a host population that traces obscured star formation. Detection of several (accepting the detection of 070521 as genuine and interpreting 090404 and 080207 as galaxy, not afterglow, emission) would actually indicate a population skewed in *favor* of the most luminous galaxies.

4.3. On the Relative Detection Fractions of Dark and Non-dark Bursts

Interestingly, non-dark GRBs are also sometimes found in very luminous hosts dominated by optically thick star formation, even though very little unobscured star formation in the universe occurs in such systems. Even if we conservatively consider only those host galaxies with secure detections at both submillimeter and radio wavelengths,⁸ there are at least two highly secure, IR ultra-luminous (inferred $L_{\text{IR}} \gtrsim 10^{12} L_{\odot}$) host galaxies belonging to GRBs that were not heavily obscured: GRBs 000418 and 010222. Considering the fact that the control sample of pre-*Swift* GRBs for which radio/submillimeter host searches have been conducted was only a few dozen, this corresponds to a fraction of $\sim 10\%$. While still a minority, this fraction is much larger than would be expected given that these events were optically selected, as the fraction of *optically thin* star formation in the universe occurring within SMGs is almost negligible (less than 1%). These galaxies would therefore appear to be extremely effective (optically thin) GRB producers given their (optically thin) SFRs.

Yet, it cannot be the case that the *entire* galaxy experiences an elevated GRB rate. SMGs are completely dominated by optically thick star formation, so if the GRB rate is the same throughout this type of galaxy, then for every mildly obscured GRB localized within an SMG, there should be many more dark GRBs occurring within similar systems. Our new results show clearly that this is not the case: heavily obscured dark GRBs are already relatively uncommon ($\sim 15\%$ of the population), and only a few of them at most occur within bright SMGs. In SMGs, therefore, *only* the optically thick outer regions seem to show a large amplification of the GRB rate relative to star formation, providing evidence that the GRB rate relative to that of overall star formation varies within galaxies as well as between them.

Alternative interpretations are possible, although they appear unlikely. If the gas and dust densities within SMGs were sufficiently high, the column would become Compton-thick, suppressing the X-ray afterglow completely and preventing the afterglows and host galaxies from ever being localized. However, since the prompt gamma-rays would generally not be absorbed, these events would be manifest as a large “X-ray-dark” population, the presence of which appears to be ruled out: almost every *Swift* GRB ($\sim 96\%$) that can be followed rapidly with the XRT is successfully detected in the X-ray band (Burrows et al. 2007).

It is also possible that the unobscured pre-*Swift* GRBs from submillimeter host galaxies actually *did* occur inside highly obscured regions of their hosts, but were able to destroy the dust along the line of sight by photoevaporation and/or X-ray grain shattering (Waxman & Draine 2000; Fruchter

et al. 2001; Draine & Hao 2002; Perna & Lazzati 2002; Perna et al. 2003), allowing the afterglow to emerge mostly unimpeded. In this case, our results would be much less surprising. However, for this to be possible nearly all the obscuring dust would have to be concentrated quite close to the GRB (dust destruction is only thought to be effective within ~ 10 pc). It is difficult to understand how the dust in these galaxies would be entirely concentrated within such a small distance of the massive stars and yet have such a high covering fraction (up to 99%). Furthermore, evidence of dramatic early-time extinction variations from rapid follow-up of GRB afterglows is so far lacking (Oates et al. 2009; Perley et al. 2010; although see also Morgan et al. 2013).

We are left, then, with the likelihood that the inferred trend is real, and that optically thin star-forming regions within SMGs are indeed much more likely to produce GRBs than the optically thick regions. While the physical cause of this trend is not clear, almost any interpretation would have interesting implications for both the formation mechanism of GRBs and the structure of SMGs.

The most popular interpretation for the apparent variation of the GRB rate between galaxies is that it is a metallicity effect: high-metallicity environments underproduce GRBs and low-metallicity environments overproduce them (Modjaz et al. 2008; Levesque et al. 2010; Graham & Fruchter 2013). SMGs and ULIRGs do indeed seem to show significant differences in metallicity between their optically thick inner regions and optically thin outer regions (Swinbank et al. 2004; Rupke et al. 2008; Caputi et al. 2008; Santini et al. 2010; Nagao et al. 2012), which could explain our observations. Nevertheless, the contribution of optically thin regions of SMGs to metal-poor cosmic star formation is (like the contribution of these regions to star formation overall) likely to be very low, so it is still not clear whether this effect alone can explain the observations or whether some other condition more specific to SMGs—such as a non-standard IMF (Baugh et al. 2005; Conroy & van Dokkum 2012)—might also be required to explain the observations.

These different possibilities—any of which would be quite significant—clearly demonstrate the need for continued submillimeter and radio follow-up of GRB host galaxies for understanding the environmental dependences of these enigmatic objects. Our observations only skirt the tip of the luminosity function (typical upper limit of $3L_*$)—and while pre-*Swift* observations were similarly sensitive only to these depths, deeper observations might unveil more highly starbursting hosts. Our interpretations also depend critically on observations of a relatively small number of pre-*Swift* hosts conducted over a decade ago. Very little radio and submillimeter follow-up of GRB hosts has been conducted since 2004, most of which has targeted galaxies at $z < 1$ —a period when optically thick star formation played only a minor role in the overall cosmic story.

With the upgraded VLA, SCUBA-2, and ALMA all now available, renewed investigation of the “dark” side of GRB host galaxies in these dust-unbiased long-wavelength windows is now more practical than ever before, even at $z > 2$ and beyond. Future long-wavelength follow-up of the host galaxies of both obscured and “ordinary” GRBs will help to further illuminate the role of different types of environments in producing the GRB phenomenon.

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⁸ A number of other possible detections have been reported in radio or submillimeter wavelengths alone (Berger et al. 2003; Tanvir et al. 2004; see Table 1). But since submillimeter measurements are often confusion limited and radio observations are subject to afterglow contamination, it is not clear if all of these represent genuine host detections.

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Facility: VLA

REFERENCES

- Barnard, V. E., Blain, A. W., Tanvir, N. R., et al. 2003, *MNRAS*, **338**, 1
- Baugh, C. M., Lacey, C. G., Frenk, C. S., et al. 2005, *MNRAS*, **356**, 1191
- Berger, E., Cowie, L. L., Kulkarni, S. R., et al. 2003, *ApJ*, **588**, 99
- Berger, E., Kulkarni, S. R., & Frail, D. A. 2001, *ApJ*, **560**, 652
- Berger, E., Zauderer, B. A., Levan, A., et al. 2013, *ApJ*, **765**, 121
- Blain, A. W., Smail, I., Ivison, R. J., Kneib, J.-P., & Frayer, D. T. 2002, *PhR*, **369**, 111
- Bloom, J. S., Berger, E., Kulkarni, S. R., Djorgovski, S. G., & Frail, D. A. 2003, *AJ*, **125**, 999
- Bloom, J. S., Kulkarni, S. R., & Djorgovski, S. G. 2002, *AJ*, **123**, 1111
- Burrows, D. N., Kennea, J. A., Abbey, A. F., et al. 2007, *Proc. SPIE*, **6686**, 668607
- Cameron, P. B., & Frail, D. A. 2005, GCN Circular 4154
- Caputi, K. I., Lilly, S. J., Aussel, H., et al. 2008, *ApJ*, **680**, 939
- Castro Cerón, J. M., Michałowski, M. J., Hjorth, J., et al. 2010, *ApJ*, **721**, 1919
- Chandra, P., & Frail, D. A. 2012, *ApJ*, **746**, 156
- Chapman, S. C., Blain, A. W., Smail, I., & Ivison, R. J. 2005, *ApJ*, **622**, 772
- Condon, J. J. 1992, *ARA&A*, **30**, 575
- Condon, J. J., Cotton, W. D., Fomalont, E. B., et al. 2012, *ApJ*, **758**, 23
- Conroy, C., & van Dokkum, P. G. 2012, *ApJ*, **760**, 71
- Daddi, E., Dickinson, M., Morrison, G., et al. 2007, *ApJ*, **670**, 156
- de Ugarte Postigo, A., Lundgren, A., Martín, S., et al. 2012, *A&A*, **538**, A44
- D'Elia, V. 2011, *AN*, **332**, 272
- Djorgovski, S. G., Frail, D. A., Kulkarni, S. R., et al. 2001, *ApJ*, **562**, 654
- Djorgovski, S. G., Kulkarni, S. R., Bloom, J. S., et al. 1998, *ApJL*, **508**, L17
- Draine, B. T., & Hao, L. 2002, *ApJ*, **569**, 780
- Dye, S., Eales, S. A., Aretxaga, I., et al. 2008, *MNRAS*, **386**, 1107
- Evans, P. A., Beardmore, A. P., Page, K. L., et al. 2009, *MNRAS*, **397**, 1177
- Firmani, C., Avila-Reese, V., Ghisellini, G., & Tutukov, A. V. 2004, *ApJ*, **611**, 1033
- Fruchter, A., Krolik, J. H., & Rhoads, J. E. 2001, *ApJ*, **563**, 597
- Fruchter, A. S., Levan, A. J., Strolger, L., et al. 2006, *Natur*, **441**, 463
- Fynbo, J. U., Gorosabel, J., Dall, T. H., et al. 2001, *A&A*, **373**, 796
- Gehrels, N., Chincarini, G., Giommi, P., et al. 2004, *ApJ*, **611**, 1005
- Graham, J. F., & Fruchter, A. S. 2013, *ApJ*, **774**, 119
- Granut, J., Piran, T., & Sari, R. 1999, *ApJ*, **527**, 236
- Greiner, J., Krühler, T., Klose, S., et al. 2011, *A&A*, **526**, A30
- Groot, P. J., Galama, T. J., van Paradijs, J., et al. 1998, *ApJL*, **493**, L27
- Hatsukade, B., Hashimoto, T., Ohta, K., et al. 2012, *ApJ*, **748**, 108
- Hjorth, J., Malesani, D., Jakobsson, P., et al. 2012, *ApJ*, **756**, 187
- Howell, J. H., Armus, L., Mazzarella, J. M., et al. 2010, *ApJ*, **715**, 572
- Hunt, L., Palazzi, E., Rossi, A., et al. 2011, *ApJL*, **736**, L36
- Jakobsson, P., Hjorth, J., Fynbo, J. P. U., et al. 2004, *ApJL*, **617**, L21
- Jensen, B. L., Fynbo, J. U., Gorosabel, J., et al. 2001, *A&A*, **370**, 909
- Jha, S., Pahre, M. A., Garnavich, P. M., et al. 2001, *ApJL*, **554**, L155
- Kann, D. A., Klose, S., & Zeh, A. 2006, *ApJ*, **641**, 993
- Kimball, A. E., Kellermann, K. I., Condon, J. J., Ivezić, Ž., & Perley, R. A. 2011, *ApJL*, **739**, L29
- Kroupa, P. 2001, *MNRAS*, **322**, 231
- Krühler, T., Greiner, J., Schady, P., et al. 2011, *A&A*, **534**, A108
- Krumholz, M., Thorsett, S. E., & Harrison, F. A. 1998, *ApJL*, **506**, L81
- Lapi, A., González-Nuevo, J., Fan, L., et al. 2011, *ApJ*, **742**, 24
- Le Floch, E., Charmandaris, V., Forrest, W. J., et al. 2006, *ApJ*, **642**, 636
- Levesque, E. M., Kewley, L. J., Graham, J. F., & Fruchter, A. S. 2010, *ApJL*, **712**, L26
- Lloyd-Ronning, N. M., Fryer, C. L., & Ramirez-Ruiz, E. 2002, *ApJ*, **574**, 554
- Meszáros, P., & Rees, M. J. 1997, *ApJ*, **476**, 232
- Michałowski, M., Hjorth, J., & Watson, D. 2010, *A&A*, **514**, A67
- Michałowski, M. J., Hjorth, J., Castro Cerón, J. M., & Watson, D. 2008, *ApJ*, **672**, 817
- Michałowski, M. J., Hjorth, J., Malesani, D., et al. 2009, *ApJ*, **693**, 347
- Michałowski, M. J., Kamble, A., Hjorth, J., et al. 2012, *ApJ*, **755**, 85
- Modjaz, M., Kewley, L., Kirshner, R. P., et al. 2008, *AJ*, **135**, 1136
- Morgan, A. N., Perley, D. A., Cenko, S. B., et al. 2013, arXiv:1305.1928
- Murphy, E. J., Condon, J. J., Schinnerer, E., et al. 2011, *ApJ*, **737**, 67
- Nagao, T., Maiolino, R., De Breuck, C., et al. 2012, *A&A*, **542**, L34
- Oates, S. R., Page, M. J., Schady, P., et al. 2009, *MNRAS*, **395**, 490
- Pérez-González, P. G., Rieke, G. H., Egami, E., et al. 2005, *ApJ*, **630**, 82
- Perley, D. A., Bloom, J. S., Klein, C. R., et al. 2010, *MNRAS*, **406**, 2473
- Perley, D. A., Levan, A. J., Tanvir, N. R., et al. 2013, arXiv:1301.5903
- Perley, R. A. 2010, EVLA Memo 145
- Perley, R. A., Chandler, C. J., Butler, B. J., & Wrobel, J. M. 2011, *ApJL*, **739**, L1
- Perna, R., & Lazzati, D. 2002, *ApJ*, **580**, 261
- Perna, R., Lazzati, D., & Fiore, F. 2003, *ApJ*, **585**, 775
- Piro, L., Frail, D. A., Gorosabel, J., et al. 2002, *ApJ*, **577**, 680
- Price, P. A., Berger, E., Kulkarni, S. R., et al. 2002, *ApJ*, **573**, 85
- Priddey, R. S., Tanvir, N. R., Levan, A. J., et al. 2006, *MNRAS*, **369**, 1189
- Prochaska, J. X., Bloom, J. S., Chen, H.-W., et al. 2004, *ApJ*, **611**, 200
- Prochaska, J. X., Chen, H.-W., Bloom, J. S., et al. 2007a, *ApJS*, **168**, 231
- Prochaska, J. X., Chen, H.-W., Dessauges-Zavadsky, M., & Bloom, J. S. 2007b, *ApJ*, **666**, 267
- Ramirez-Ruiz, E., Trentham, N., & Blain, A. W. 2002, *MNRAS*, **329**, 465
- Rol, E., van der Horst, A., Wiersema, K., et al. 2007, *ApJ*, **669**, 1098
- Rupke, D. S. N., Veilleux, S., & Baker, A. J. 2008, *ApJ*, **674**, 172
- Santini, P., Maiolino, R., Magnelli, B., et al. 2010, *A&A*, **518**, L154
- Sari, R., Piran, T., & Narayan, R. 1998, *ApJL*, **497**, L17
- Savaglio, S., Glazebrook, K., & Le Borgne, D. 2009, *ApJ*, **691**, 182
- Schechter, P. 1976, *ApJ*, **203**, 297
- Smail, I., Chapman, S. C., Blain, A. W., & Ivison, R. J. 2004, *ApJ*, **616**, 71
- Stanway, E. R., Davies, L. J. M., & Levan, A. J. 2010, *MNRAS*, **409**, L74
- Svensson, K. M., Levan, A. J., Tanvir, N. R., et al. 2012, *MNRAS*, **421**, 25
- Swinbank, A. M., Smail, I., Chapman, S. C., et al. 2004, *ApJ*, **617**, 64
- Tanvir, N. R., Barnard, V. E., Blain, A. W., et al. 2004, *MNRAS*, **352**, 1073
- Totani, T. 1999, *ApJ*, **511**, 41
- van Paradijs, J., Kouveliotou, C., & Wijers, R. A. M. J. 2000, *ARA&A*, **38**, 379
- Vreeswijk, P., Fruchter, A., Hjorth, J., & Kouveliotou, C. 2003, GCN Circular 1785
- Wang, W.-H., Chen, H.-W., & Huang, K.-Y. 2012, *ApJL*, **761**, L32
- Wardlow, J. L., Smail, I., Coppin, K. E. K., et al. 2011, *MNRAS*, **415**, 1479
- Waxman, E., & Draine, B. T. 2000, *ApJ*, **537**, 796
- Wolf, C., & Podsiadlowski, P. 2007, *MNRAS*, **375**, 1049
- Yun, M. S., Scott, K. S., Guo, Y., et al. 2012, *MNRAS*, **420**, 957