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## Article

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# Photometry and spectroscopy of GRB 060526: a detailed study of the afterglow and host galaxy of a $z = 3.2$ gamma-ray burst<sup>★,★★</sup>

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## ABSTRACT

**Aims.** With this paper we want to investigate the highly variable afterglow light curve and environment of gamma-ray burst (GRB) 060526 at  $z = 3.221$ .

**Methods.** We present one of the largest photometric datasets ever obtained for a GRB afterglow, consisting of multi-color photometric data from the ultraviolet to the near infrared. The data set contains 412 data points in total to which we add additional data from the literature. Furthermore, we present low-resolution high signal-to-noise spectra of the afterglow. The afterglow light curve is modeled with both an analytical model using broken power law fits and with a broad-band numerical model which includes energy injections. The absorption lines detected in the spectra are used to derive column densities using a multi-ion single-component curve-of-growth analysis from which we derive the metallicity of the host of GRB 060526.

**Results.** The temporal behaviour of the afterglow follows a double broken power law with breaks at  $t = 0.090 \pm 0.005$  and  $t = 2.401 \pm 0.061$  days. It shows deviations from the smooth set of power laws that can be modeled by additional energy injections from the central engine, although some significant microvariability remains. The broadband spectral-energy distribution of the afterglow shows no significant extinction along the line of sight. The metallicity derived from S II and Fe II of  $[S/H] = -0.57 \pm 0.25$  and  $[Fe/H] = -1.09 \pm 0.24$  is relatively high for a galaxy at that redshift but comparable to the metallicity of other GRB hosts at similar redshifts. At the position of the afterglow, no host is detected to  $F775W(AB) = 28.5$  mag with the HST, implying an absolute magnitude of the host  $M(1500 \text{ \AA}) > -18.3$  mag which is fainter than most long-duration hosts, although the GRB may be associated with a faint galaxy at a distance of 11 kpc.

**Key words.** gamma rays bursts: individual: GRB 060526 – galaxies: high-redshift – ISM: abundances

## 1. Introduction

Gamma-ray bursts (GRBs) and their afterglows offer a powerful tool to probe the high-redshift universe, both through photometry and spectroscopy. The standard fireball model of GRB afterglows (see Zhang 2007; Mészáros et al. 2006; Gehrels et al. 2009, for recent reviews on the topic) predicts a smooth temporal evolution, and the resulting afterglow light curve can be empirically described by a joint smoothly-broken power law, the so-called Beuermann equation (Beuermann et al. 1999). For many GRBs in the pre-*Swift* era, when temporally dense afterglow photometry was obtained, the afterglow evolution was found to be smooth as, e.g., in GRB 020813 (Laursen & Stanek 2003), GRB 030226 (Klose et al. 2004) and GRB 041006

(Stanek et al. 2005) (for the complete pre-*Swift* sample see Zeh et al. 2006). Out of the total sample of 59 afterglows of that paper, only four of the GRBs analysed showed significant deviations from the expected decay (though we note that about half of the 59 afterglows were not sampled well enough to allow any conclusions). One might be explained by microlensing (GRB 000301C, Garnavich et al. 2000) and another one by assuming an inhomogeneously emitting surface (GRB 011211, Jakobsson et al. 2004). The other two GRBs which are the only ones showing long-lasting strong deviations are GRB 021004 (e.g. de Ugarte Postigo et al. 2005, and references therein) and GRB 030329 (e.g. Lipkin et al. 2004, and references therein), which incidentally also have the densest optical monitoring.

With the launch of the *Swift* satellite and its rapid localization capabilities (Gehrels et al. 2004), the number of highly variable light curves has increased considerably, though there are still examples of very smooth light curves (e.g. that of GRB 080210, which showed a smooth behaviour with a high sampling rate of 1 s, A. De Cia, priv. comm.). Some light curves show small, achromatic bumps overlying the smooth power law decay (e.g.,

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<sup>★★</sup> Table 1 is only available in electronic form at <http://www.aanda.org>

GRB 050502A, [Guidorzi et al. 2005](#); GRB 061007, [Mundell et al. 2007](#); GRBs 090323 and 090328, [McBreen et al. 2010](#)), early bumps with chromatic evolution (e.g., GRB 061126, [Perley et al. 2008a](#); GRB 071003, [Perley et al. 2008b](#)), “steps” due to energy injection episodes (e.g., GRB 070125, [Updike et al. 2008](#); GRB 071010A, [Covino et al. 2008](#); GRB 080913, [Greiner et al. 2009b](#), GRB 090926A, [Rau et al. 2010](#); [Cenko et al. 2010](#); [Swenson et al. 2010](#)) or powerful late-time rebrightenings of up to several magnitudes (e.g., GRB 050721, [Antonelli et al. 2006](#); GRB 060206, [Woźniak et al. 2006](#); [Monfardini et al. 2006](#); [Stanek et al. 2007](#); GRB 070311, [Guidorzi et al. 2007](#); GRB 071003, [Perley et al. 2008b](#)). The early time domain, which can now be routinely accessed by rapid follow-up in the *Swift* era, has yielded more types of variability, like rising afterglows (e.g., GRB 060418, [Molinari et al. 2007](#); GRB 060605, [Ferrero et al. 2009](#); GRB 060607A, [Nysewander et al. 2009](#), [Molinari et al. 2007](#); GRB 081008, [Yuan et al. 2010](#); see [Oates et al. 2009](#); [Rykoff et al. 2009](#), for further examples) and short-term variability directly linked to the prompt emission (e.g., GRB 041219A, [Vestrand et al. 2005](#); [Blake et al. 2005](#); GRB 050820A, [Vestrand et al. 2006](#); GRB 080319B, [Racusin et al. 2008](#); GRB 080129, [Greiner et al. 2009a](#)). In all these cases, dense photometric follow-up during the periods of variability was needed to characterize the phenomena involved.

In addition to their use in studying the GRB phenomenon itself, GRB afterglows can be used to study their galactic environment through absorption line spectroscopy of material in the line-of-sight towards the GRB. Above a redshift of around  $z \sim 0.2$ , resonant absorption lines from elements present in the interstellar medium (ISM) such as Mg, Zn, Fe, Si, C, and S are shifted into the optical regime and can be studied with ground-based telescopes. For 10 bursts between a redshift of  $z = 2$  to 6, the metallicity along the line-of-sight in the galaxy could be obtained so far (e.g. [Savaglio 2006](#); [Fynbo et al. 2006](#); [Price et al. 2007](#); [Prochaska et al. 2007](#); [Ledoux et al. 2009](#), and references therein). The values are usually below solar, but higher than for QSO absorbers at comparable redshifts, some of them even higher than theoretical limits for the formation of collapsars ([Woosley & Heger 2006](#)). The difference to QSO absorbers can be explained if GRB sightlines probe denser parts of the galaxy, or if GRBs reside in galaxies with higher masses and therefore higher metallicities ([Fynbo et al. 2008](#); [Pontzen et al. 2010](#)). Both for QSO and GRB absorbers, there seems to be a metallicity evolution with redshift ([Savaglio 2006](#); [Fynbo et al. 2006](#); [Price et al. 2007](#)), although the slope is different for the two samples. GRB hosts seem to show a low extinction along the line-of-sight ([Prochaska et al. 2007](#)), however, relative abundances of heavier elements indicate that some of the ions must be depleted onto dust grains and the depletion pattern resembles the one found in the warm disc and halo of the Milky Way ([MW](#); [Savaglio 2006](#)). How this can be explained together with the low extinction as also derived from the spectral energy distribution (SED) of the afterglow ([Kann et al. 2006, 2010](#); [Starling et al. 2007](#)), is still an open question. One possible solution to this problem is destruction of the dust present in the line-of-sight by the GRB and afterglow radiation (e.g., [Waxman & Draine 2000](#); [Perna & Lazzati 2002](#)), though no strong evidence has ever been found for this. Furthermore, there is clearly a strong observational bias involved, as those GRB afterglows with successful spectroscopy, especially in the case of high-resolution observations, are those which have only low extinctions and thus relatively bright afterglows ([Fynbo et al. 2009](#); [Kann et al. 2010](#)), although rapid observations with large telescopes can achieve detailed spectroscopy of highly extinguished afterglows, as in

the case of GRB 080607 ([Prochaska et al. 2009](#); [Sheffer et al. 2009](#); [Fynbo et al. 2009](#)).

GRB 060526 was detected by the *Swift* satellite on May 26.686458 (16:28:29.95 UT). The satellite slewed immediately to the burst, detecting both the X-ray and the optical afterglow ([Campana et al. 2006a](#)). The BAT instrument on *Swift* measured two emission episodes. The first one lasted 13.8 s and consisted of two FRED (fast rise exponential decay) peaks, followed by a second symmetric peak between 230 and 270 s ([Campana et al. 2006b](#)). The second peak was coincident with a giant X-ray flare followed by a softer flare at 310 s ([Campana et al. 2006c](#)) also detected in the  $\nu$  band by the UVOT telescope on-board *Swift* ([Brown et al. 2006](#)). The gamma-ray fluence was  $(4.9 \pm 0.6) \times 10^{-7}$  erg cm $^{-2}$  during the first emission episode and  $(5.9 \pm 0.6) \times 10^{-7}$  erg cm $^{-2}$  during the second, the peak flux of the second episode was however only half of the peak flux of the first one. The photon index of the two epochs changed from  $1.66 \pm 0.20$  to  $2.07 \pm 0.18$ , thus showing the typical hard-to-soft evolution ([Markwardt et al. 2006](#)). The *Watcher* telescope provided the first ground based detection with  $R \approx 15$  mag ([French & Jelínek 2006](#)) 36.2 s after the trigger. *ROTSE* observations showed a plateau for several thousand seconds after the GRB onset ([Rykoff et al. 2006](#)). A redshift of  $z = 3.21$  was determined by [Berger & Gladders \(2006\)](#) with the Magellan/Clay telescope. The brightness of the optical afterglow allowed for a dense monitoring which revealed a complex light curve structure including several flares ([Halpern et al. 2006a,b](#)) and a steepening attributed to a jet break ([Thöne et al. 2006](#)).

In this paper, we approach the analysis of GRB 060526 from two directions: Through modelling of the very detailed optical light curve and late optical imaging of the field to detect the host in Sect. 3, and analysis of low to medium resolution spectroscopic observations of absorption lines along the line of sight (Sect. 4). Throughout the paper, we follow the convention  $F_\nu(t) \propto t^{-\alpha} \nu^{-\beta}$ , and use WMAP concordant cosmology ([Spergel et al. 2003](#)) with  $H_0 = 71$  km s $^{-1}$  Mpc $^{-1}$ ,  $\Omega_M = 0.27$ , and  $\Omega_\Lambda = 0.73$ . Uncertainties are given at 68% confidence level for one parameter of interest unless stated otherwise.

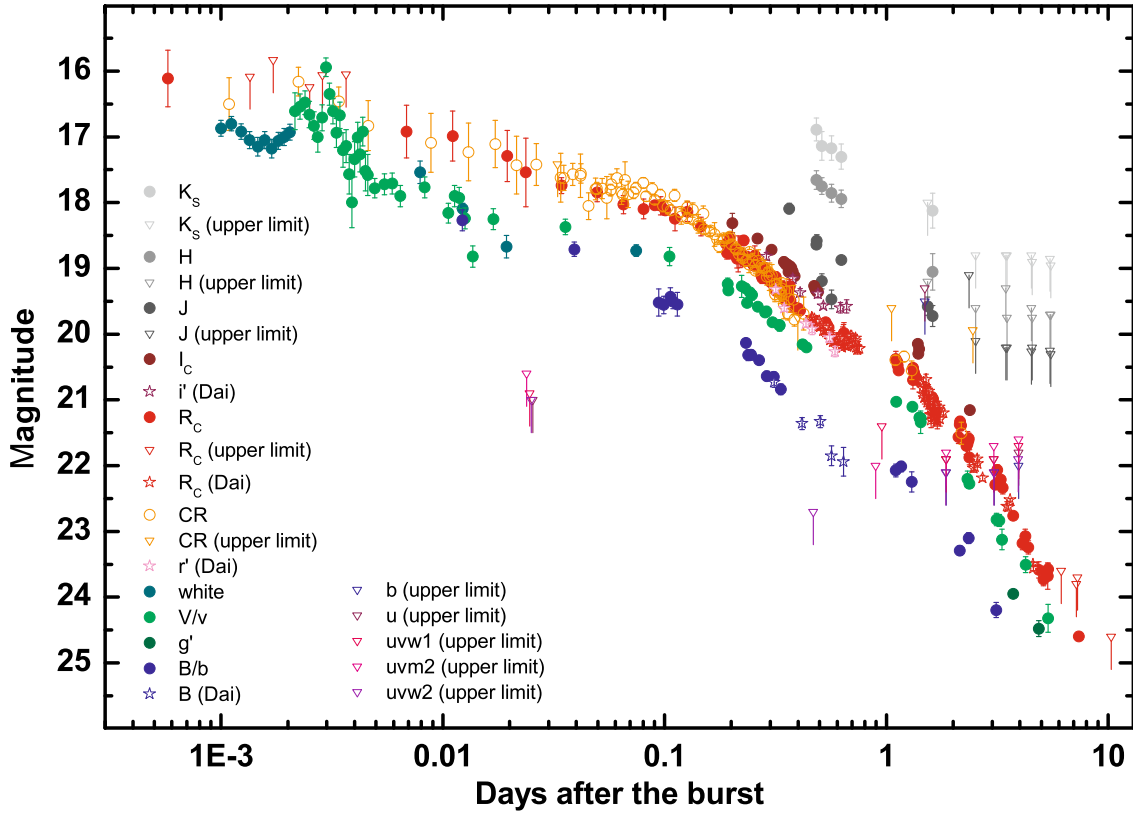
## 2. Observations

### 2.1. Photometry

In order to get a good coverage of the light curve, we obtained data using several different telescopes around the world. Our complete data set comprises a total of 412 points from the UV to  $K$ -band, one of the largest photometric samples of an optical/NIR afterglow in the *Swift* era.

The earliest dataset was obtained by the *Watcher* telescope, located at Boyden Observatory, South Africa, starting 36.2 s after the burst, followed by *Swift* UVOT starting 86 s after the trigger. Early ground based optical data were obtained with the *ROTSE-IIIc* 0.3 m telescope at the HESS site at Mt. Gamsberg, Namibia, the 1.5 m telescope on Mt. Maidanak/Uzbekistan, the 2.6 m Shajn telescope at CrAO (Crimean Astrophysical Observatory/Ukraine), the TNG (Telescopio Nazionale Galileo) on La Palma equipped with DOLoRes, the 1.2 m MIRO telescope on Mt. Abu/India, with BFOC (Bologna Faint Object Spectrograph & Camera) at the G. D. Cassini 152 cm telescope of the Bologna University under poor conditions and with the RTT150 (1.5 m Russian-Turkish telescope, Bakirlitepe, Turkey), the RTT150 data are also presented in [Khamitov et al. \(2007\)](#).

The light curve was followed up sparsely every night over nearly a week with DFOC (Danish Faint Object Spectrograph



**Fig. 1.** Light curves of the optical afterglow of GRB 060526. Filled and open circles (detections,  $Bg'VR_CCRICJHK_S$ ) and open triangles (upper limits,  $uvw2\ uvw1\ ubR_C\ CR\ JHK_S$ ) are presented in this work, and open stars ( $Br'R'$ ) are taken from Dai et al. (2007). The early light curve especially shows strong variability, with two flares seen in the UVOT  $v$  band data. The later light curve also shows evident variability, and the large  $V - R$ ,  $B - R$  colours due to the high redshift are evident. Data in this plot are not corrected for foreground extinction.

and Camera) on the Danish 1.54 m telescope on La Silla/Chile under partially photometric conditions and with MOSCA (MOSAic Camera) and ALFOSC (Andalucia Faint Object Spectrograph and Camera) at the Nordic Optical Telescope on La Palma. Furthermore, two epochs were obtained with the Tautenburg 1.34 m Schmidt telescope and two sets of images were taken several days after the GRB with Keck/LRIS simultaneously in the Kron-Cousins  $R$  and the Sloan  $g'$  bands, the second observation was performed at high airmass under bad seeing conditions. Late images were obtained with FORS2 at the VLT on Paranal/Chile on Feb. 23, 2007 and Mar. 30, 2008 in the  $R_C$  band with exposure times of 2500 and 7500 s, respectively, to look for the host galaxy. We also took observations with the *Hubble Space Telescope* on 9 August 2009, utilizing the Advanced Camera for Surveys with the F775W (roughly SDSS  $i'$ ) filter (see Fig. 6). A total of 7844 s of observations were obtained in six dithered exposures. These were reduced via *multidrizzle* in the standard fashion.

Near infrared data were collected with ANDICAM and the 1.3 m SMARTS telescope (Small and Moderate Aperture Research Telescope System) at CTIO under non-photometric conditions as well as with the robotic 1.3 m PAIRITEL telescope on Mt. Hopkins.

The UVOT data were reduced and analysed using the standard UVOT tasks within the *heasoft* package. For the photometric calibration of the ground-based data, we determined the calibrated magnitude of six comparison stars in the field using photometric zero points from DFOSC in the  $V$ ,  $R_C$  and  $I_C$  bands (see Table 3). These stars were then used to perform relative Point-Spread Function (PSF) photometry to get the calibrated

magnitude of the afterglow. For some of the late NOT images as well as the faint MIRO detections, though, we applied relative aperture photometry using a circle of 20 pixels diameter (and an annulus of 10 pixels for the sky). For the  $B$ -band, where no DFOSC data were available, we took zero points for only three comparison stars from the SDSS (see Table 3), converting them with the equations of Jester et al. (2005).  $g'$  zero points were taken from the SDSS, and they are given as  $AB$  magnitudes. For the RTT 150 data, we used one USNO-B1 star as reference that was calibrated using Landolt standard stars. The results are in full agreement with the rest of the data set. The Watcher data were analysed using a dedicated photometry pipeline (Ferrero et al. 2010). The  $J$ ,  $H$  and  $K_S$  band data from SMARTS/ANDICAM and PAIRITEL were calibrated using three and ten nearby stars, respectively, from the 2MASS catalogue. For the  $H$  and  $K_S$  band data, we used smaller apertures and applied aperture corrections to reduce the influence of the highly variable background.

All data and upper limits are given in Table 1, the data are not corrected for Galactic extinction. Note we give the  $g'$  and HST F775W magnitudes in  $AB$  magnitudes. For the final light curve fitting, we add  $Br'R_Ci'$  band data from Dai et al. (2007). We shift the  $B$  and  $R_C$  band data of Dai et al. (2007) by 0.1 mag to bring it to our zero point. The multi-color light curves are shown in Fig. 1.

## 2.2. Spectroscopy

Spectra were obtained with FORS1/VLT on May 27 from 9 to 12 h after the burst. Four different grisms cover the wavelength

**Table 2.** Spectra of GRB 060526 from FORS1/VLT.

Time (May 27 [UT])	Midtime [d]	Exptime [s]	Grism	Spectral range [Å]	Resolution [Å]
01:16:56	0.37042	600	300V	3650–8900	11.1
01:29:46	0.38105	900	600I	7000–9200	4.4
01:47:28	0.39335	900	1200B	3860–4400	2.4
02:05:27	0.40583	900	600V	4080–7200	4.5
02:24:00	0.42043	1200	300V	3650–8900	11.1
02:57:05	0.44681	1800	1200B	3860–4400	2.4
03:30:34	0.47007	1800	600V	4080–7200	4.5
04:04:08	0.49339	1800	600I	7000–9200	4.4

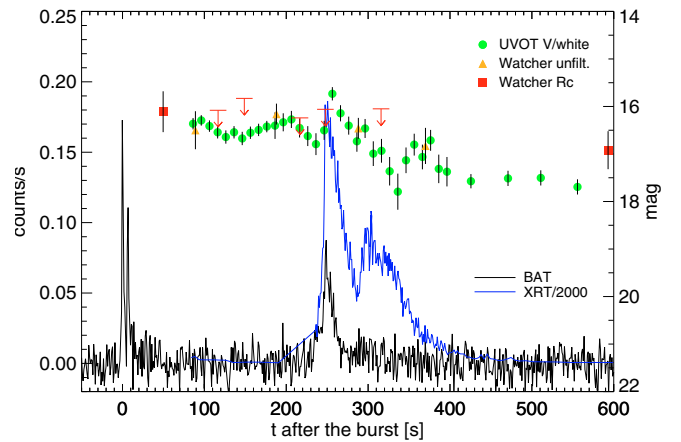
**Notes.** Times given are the start time of the observations, midexposure times are in days after the GRB trigger.

range from 3650 to 9200 Å. For all four grisms a 1'0 slit was used which provides resolutions between 2.4 and 11.1 Å. Reduction, cosmic ray removal, extraction and wavelength calibration were performed using standard tasks in IRAF<sup>1</sup>. The final spectra were then normalised as no absolute flux calibration was needed. In order to improve the S/N, we combined the datasets taken with the same grism weighted with their variance. A summary of the spectroscopic observations is given in Table 2.

### 3. Prompt emission

Up to two hours after the burst, the light curve features an optical flare contemporaneous to the XRT/BAT flare at  $\approx 250$  s and a following plateau phase (Brown et al. 2006; French & Jelínek 2006; Rykoff et al. 2006). We reduced the BAT event data using standard procedures within the software provided by HEASOFT (version 6.1). The XRT observations were reduced using the standard `xrtpipeline` (version 0.10.4) for XRT data analysis software using the most recent calibration files. The spectral data of BAT and XRT were analysed with XSPEC version 11.3 (Arnaud 1996). The X-ray Galactic column density was fixed to  $5.02 \times 10^{20} \text{ cm}^{-2}$  (Kalberla et al. 2005). We estimated the late-time extragalactic column density by fitting the XRT PC data from 517 s to  $1.34 \times 10^5$  s post trigger, where spectral evolution is negligible, using an absorbed power law model. We only find an upper limit of  $N_{\text{HX}} < 9.8 \times 10^{21} \text{ cm}^{-2}$  (see also Campana et al. 2010).

In Fig. 2 we compare the timescale of the flares at high energies (BAT and XRT) with contemporaneous optical data. The first, stronger flare is seen both by the XRT and BAT whereas the second, softer flare is only visible in the XRT data. Our earliest optical data are also coincident with the two flares. A broad early bump is observed in the optical light curve which precedes the second episode of BAT emission and the major X-ray pulse. The Watcher data only show a plateau during the high-energy flares due to a relatively low time resolution. The higher time-resolution of the UVOT data shows that there are also two significant flares in the optical. The first optical flare is contemporaneous with the peak of the BAT/XRT flare within  $2\sigma$  of the temporal error. Then, there may also be a small optical bump (significance only  $\approx 1.5\sigma$ ) at the time of the second (XRT-only) flare, whereas the second significant optical flare occurs  $\approx 90$  s after the XRT-only flare implying that these two events are probably not connected. Power-law fitting of the optical light curve (with respect to the BAT trigger time  $T_0$ ) shows the slopes are very steep. The rising slope of the first flare is  $\alpha_{r1} = -12.8 \pm 3.0$ , and the decay slope is  $\alpha_{d1} = 5.8 \pm 0.8$ , or, if one takes the bump

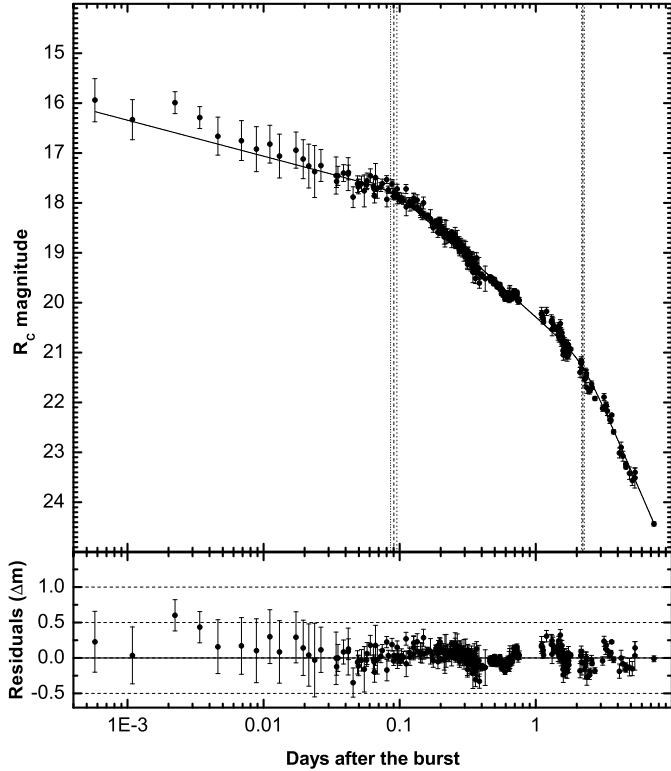


**Fig. 2.** The two early flares seen in the BAT and XRT data and temporally coincident optical observations by UVOT and Watcher. The second flare is much softer at high energies and hence it is only seen in the XRT data. Both flares are reflected in the optical data with the first optical flare being coincident with the first BAT/XRT flare while the second flare is delayed compared to the second X-ray flare. The *white* data have been shifted upward by 0.3 mag to the *v* zero point. A color version of the figure is available in the online version of the article.

as a second flare, the two decay slopes are  $\alpha_{d11} = 8.2 \pm 2.0$ ,  $\alpha_{d12} = 8.6 \pm 2.5$ . The rise and decay slopes for the second peak are similarly steep, but with larger uncertainties.

Prompt optical flashes attributed to reverse shocks, as seen in the landmark burst GRB 990123 (Akerlof et al. 1999), have been observed in only a few cases since the launch of Swift, e.g., GRB 060111B (Klotz et al. 2006) and GRB 060117 (Jelínek et al. 2006, see Kann et al. 2010 for a recent overview of light curves with probable reverse shock flashes/steep decays in the *Swift* era). Most bursts for which early optical data are available display no evidence of reverse shock emission. In a number of cases, the optical light curve is dominated by forward shock emission from very early times, e.g. GRB 050401 (Rykoff et al. 2005), GRB 060418 and GRB 060607A (Molinari et al. 2007; Nysewander et al. 2009), GRB 060605 (Ferrero et al. 2009), GRB 061007 (Mundell et al. 2007) and GRB 081008 (Yuan et al. 2010). An optical component of the emission from internal shocks has been invoked to explain the correlation between the optical and high-energy light curves observed in several bursts, e.g. GRB 041219A (Vestrand et al. 2005; Blake et al. 2005; Fan et al. 2005), GRB 050820A (Vestrand et al. 2006), and GRB 050904 (Wei et al. 2006). The very rapid variability as measured by the steep slopes and the contemporaneous first

<sup>1</sup> <http://iraf.noao.edu>



**Fig. 3.** The  $R_C$ -band light curve of GRB 060526 fitted with a double smoothly-broken power law, using data starting at 0.046 days. The residuals of the fit show an early optical flare and later strong variations of up to 0.3 mag. The dashed vertical lines mark the break times, the dotted lines the  $1\sigma$  region of uncertainty. The second break is soft ( $n \approx 4$ ) and is thus seen as a smooth rollover.

flare indicates that GRB 060526 is another case where the early optical light curve is dominated by central engine activity. Such time resolution and coincident optical and high-energy observations are still rare, however, a thorough analysis of the flares and their spectral properties goes beyond the scope of this paper.

## 4. Afterglow photometry and light curve modelling

### 4.1. The multi-color light curve

To analyse the evolution of the light curve, we use the  $R_C/CR$  band which has the densest sampling, a total of 319 data points in all, from 50 s to more than 7 days after the GRB. For all fits, we fix the host magnitude to  $m_h = 29$  (see Sect. 4.4). A fit with a single power law to these data is very strongly rejected, with  $\chi^2/\text{d.o.f.} = 31.89$  (with 319 d.o.f.). Even if we remove the earliest data which are affected by optical flares, starting only at 400 s after the GRB, the fit is still rejected ( $\chi^2/\text{d.o.f.} = 29.46$  with 308 d.o.f.). A double smoothly-broken power law gives a much better fit, but even so, the fit is formally rejected, with  $\chi^2/\text{d.o.f.} = 3.12$  with 302 d.o.f. This is due to the strong variability in the light curve which was first found by Halpern et al. (2006a) and is also discussed in Dai et al. (2007). This fit, along with the residuals showing the strong variability, is shown in Fig. 3. The parameters we find for this fit ( $\alpha_{\text{plateau}} = 0.288 \pm 0.026$ ,  $\alpha_1 = 0.971 \pm 0.008$ ,  $\alpha_2 = 2.524 \pm 0.052$ ,  $t_{b1} = 0.090 \pm 0.005$  days,  $t_b = 2.216 \pm 0.049$  days) are concurrent with those of Dai et al. (2007) (who find  $\alpha_1 \approx 1.0$ ,  $\alpha_2 \approx 2.9$  and  $t_b \approx 2.55$  days).

Due to the high data density, we are able to let the break smoothness parameter  $n$  vary for the second break (the first break

had to be fixed to  $n = 10$ ), and our result ( $n = 4.3 \pm 0.7$ ) is in agreement with the tentative  $\alpha_1 - n$  correlation found by Zeh et al. (2006). Still, the significant improvement shows that the light curve is basically a double smoothly-broken power law, and the steep late decay indicates that this break is a jet break, as first noted by Thöne et al. (2006) and also found by Dai et al. (2007). The rest frame jet break time of 0.52 days is typical for the optical afterglows of the pre-*Swift* GRB sample (Zeh et al. 2006). Jet breaks, a common feature in well-monitored pre-*Swift* optical afterglows (e.g., Zeh et al. 2006), have not been found in many *Swift* afterglows<sup>2</sup>, especially in the X-rays (e.g., Mangano et al. 2007; Grupe et al. 2007; Sato et al. 2007; Racusin et al. 2009), and if there are breaks, then the comparison between the optical and X-ray light curves show them to often be chromatic (Panaitescu et al. 2006; Oates et al. 2007). Analysing both the optical and X-ray data of GRB 060526, Dai et al. (2007) suggest that the (jet) break is achromatic.

If we take the X-ray data from Dai et al. (2007) and fit it with the parameters of our double smoothly-broken power law fit using X-ray data from 0.06 days onward, we concur that the fit is marginally acceptable, with  $\chi^2/\text{d.o.f.} = 1.6$  for 40 degrees of freedom. Using our own X-ray reduction, we obtain a worse result, with  $\chi^2/\text{d.o.f.} = 1.7$  for 22 degrees of freedom. Similarly, with X-ray data from the *Swift* XRT repository (Evans et al. 2007, 2009), the fit is rejected ( $\chi^2/\text{d.o.f.} = 2.3$  for 33 degrees of freedom). The reason is that there are less data points, but with smaller error bars, so the outliers are weighted more strongly. A fit to the X-ray data alone results in a much earlier break time and slopes that are less steep (again in agreement with Dai et al. 2007), but we caution that the late X-ray afterglow is only sparsely sampled and shows large scatter. For the three different reductions, we derive the following values using a smoothly-broken power law (a double-broken power law yields no statistical improvement), with  $n = 10$  fixed and no host galaxy:  $\alpha_1 = 0.92 \pm 0.05$ ,  $\alpha_2 = 3.02 \pm 0.63$ ,  $t_b = 1.38 \pm 0.25$  (XRT repository light curve);  $\alpha_1 = 0.90 \pm 0.05$ ,  $\alpha_2 = 2.78 \pm 0.61$ ,  $t_b = 1.34 \pm 0.28$  (our XRT reduction);  $\alpha_1 = -0.29 \pm 0.56$ ,  $\alpha_2 = 1.69 \pm 0.17$ ,  $t_b = 0.16 \pm 0.05$  (Dai et al. 2007 data). As can be seen, the latter fit is very different from the other two (which agree fully within error bars). Even using data from 0.01 days onward (end of the X-ray flare), the fit is still significantly different:  $\alpha_1 = 0.43 \pm 0.06$ ,  $\alpha_2 = 1.69 \pm 0.17$ ,  $t_b = 0.25 \pm 0.06$ .

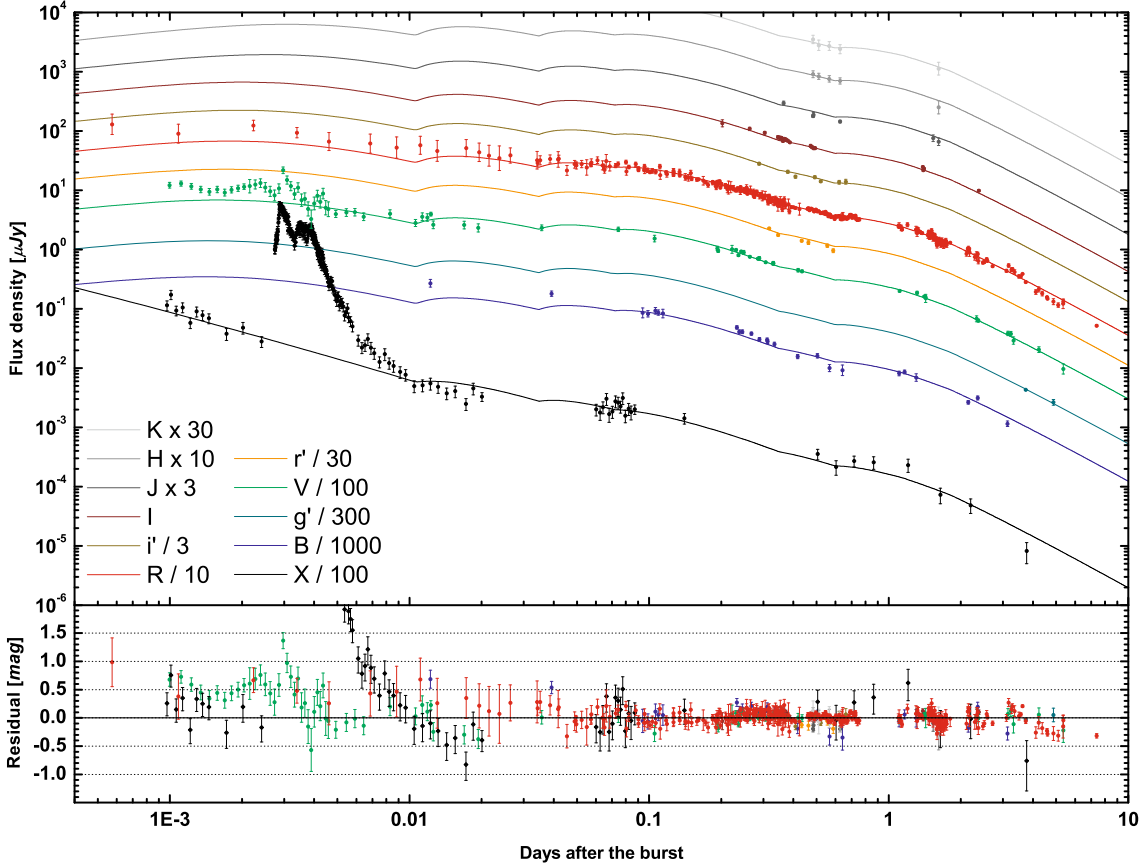
### 4.2. Modelling the light curve with energy injections

Motivated by the similarities to such highly variable light curves as that of the afterglow of GRB 021004, which was successfully modelled by multiple energy injections (“refreshed shocks”) (de Ugarte Postigo et al. 2005), we used the code of Jóhannesson et al. (2006) to model the afterglow light curve. The code numerically solves the kinematic equations of an expanding shock front and calculates the resulting synchrotron emission. Relativistic effects are fully taken into account and the code supports delayed energy injection episodes. Several energy injection episodes are applied as a possible scenario to explain the rebrightenings and shallow decay of the afterglow. Preliminary results on a smaller data set were presented in Jóhannesson et al. (2009). The number and a time range for the energy injections have to be inserted as initial guess for the fit and the fit then adjusts them to the best possible time within that range and

<sup>2</sup> Note, though, that Dai et al. (2008) argue that optical follow-up in the *Swift* era, especially with UVOT only, often does not persist long enough (and to enough depth) to find jet breaks.

**Table 3.** Magnitudes of the comparison stars used for photometry.

#	Coordinates		<i>B</i>	<i>V</i>	<i>R<sub>C</sub></i>	<i>I<sub>C</sub></i>
	RA	Dec	[mag]	[mag]	[mag]	[mag]
1	15:31:19.6	+00:16:59.5	19.90 ± 0.03	19.36 ± 0.02	19.02 ± 0.02	18.64 ± 0.02
2	15:31:22.7	+00:17:21.2	20.30 ± 0.03	19.51 ± 0.02	19.09 ± 0.02	18.71 ± 0.02
3	15:31:20.9	+00:16:39.1	19.97 ± 0.03	19.11 ± 0.02	18.55 ± 0.02	18.12 ± 0.02
4	15:31:14.4	+00:17:49.9	–	19.78 ± 0.02	19.43 ± 0.02	18.81 ± 0.03
5	15:31:16.5	+00:17:55.5	–	20.34 ± 0.03	19.96 ± 0.03	20.59 ± 0.03
6	15:31:18.3	+00:18:17.7	–	20.49 ± 0.03	19.45 ± 0.02	18.23 ± 0.02



**Fig. 4.** Fit to the light curve in X-ray and optical/NIR  $Vr'R'I'JHK_S$  bands including a total of five energy injections. Data before 500 s are excluded from the modelling, as well as the strong X-ray flare. The light curves have been offset by constant factors as given in the figure legend for better legibility. The  $Bg'$  bands are affected by additional Lyman forest absorption and were not included in the fit. For the  $Bg'$  bands, the model predictions have been multiplied by 0.50 and 0.55 respectively to match them with the data. The residuals are clearly improved in comparison to those plotted in Fig. 3, but short-timescale variations like the one at 0.7 days are still not fitted satisfactorily.

determines the magnitude of the injection. As the very early data likely contain some signature of the prompt emission, we exclude all data before 400 s as well as the very bright X-ray flare. Due to Lyman forest blanketing, data in  $B$  and  $g'$  bands were excluded from the fit. The  $V$  band is also affected by the blanketing, but was corrected with the model from Madau (1995) (brightened by 0.17 mag) and included in the fit to have a better spectral coverage at early times.

Figure 4 shows the best fit ( $\chi^2/\text{d.o.f.} = 2.0$ ) found using a model with a total of five energy injections: at 0.01, 0.04, 0.07, 0.35 and 0.60 days. Each energy injection episode adds 2 free parameters to the fit in addition to the 6 parameters needed for the standard afterglow model discussed below. The initial energy injected into the outflow is<sup>3</sup>  $E_0 = 5_{-3}^{+2} \times 10^{49}$  erg. The energy

injections then add 1.8, 4.1, 4.2, 8, and finally 14 times the initial energy release  $E_0$  to the afterglow, for a total energy release in the afterglow of  $1.6_{-1.1}^{+0.8} \times 10^{51}$  erg.

The first three injections are responsible for the shallow afterglow decay between 0.008 and 0.25 days. The quality of the data does not allow to discriminate between this three-injection scenario and a continuous injection. Using a lower number of injections does not yield a satisfactory fit to the observed data. Since there are no direct indications of injections in the light curve, the time of each of the three injections is not well determined. A direct consequence of this is that the energy of each individual injection in this phase is not well determined, while the total energy released is fairly consistent. The time and energy of the last two injections are, however, better constrained by the data.

Further results of the modelling are a high density of the circumburst medium of  $n_0 = 600_{-500}^{+3000}$  cm<sup>-3</sup> (a high value,

<sup>3</sup> The error limits presented here are  $1\sigma$  estimates, but due to degeneracy in the model parameters, the actual range of parameter values can be much larger.



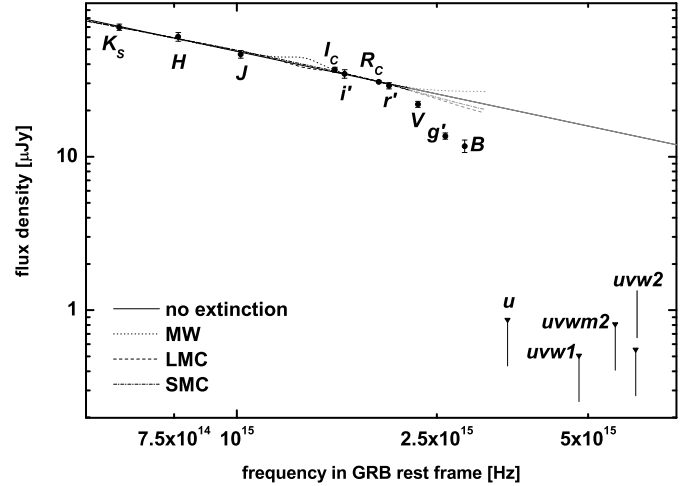
comparable to the high- $z$  GRB 050904, Frail et al. 2006), a rather low opening angle of  $\theta_0 = 2.2_{-1.8}^{+1.5}$  and an electron index of  $p = 2.1_{-0.3}^{+0.2}$  (host galaxy extinction is assumed to be negligible, see Sect. 4.3). The peak frequency  $\nu_m$  passes through the optical/NIR at very early times, while the data are most likely affected by the prompt emission. The cooling break  $\nu_c$  is between the optical and X-rays up to 6 days after the burst, i.e., over the whole data span. Furthermore, we find  $\varepsilon_e = 2_{-1.5}^{+5} \times 10^{-2}$  and  $\varepsilon_B = 7_{-6}^{+40} \times 10^{-5}$ , with  $\varepsilon_e$  being the fraction of the energy in the electron population and  $\varepsilon_B$  the fraction of the energy in the magnetic field. Note that the definition of  $\varepsilon_e$  has been changed in the model from Jóhannesson et al. (2006) to the definition of Panaitescu & Kumar (2001) to allow for  $p < 2$  in the model. This model reproduces the global properties of the optical/NIR light curves well, with the rather high  $\chi^2/\text{d.o.f.} = 2.0$  resulting from rapid variability in the late time afterglow. This is incompatible with the homogeneous shock front assumed in the numerical code. The X-ray light curve is also fitted reasonably well.

#### 4.3. The spectral energy distribution and host extinction

Following the procedures outlined in Kann et al. (2006), we derive the optical spectral energy distribution (SED) of the GRB 060526 afterglow and fit it with several dust models (Milky Way, MW; Large Magellanic Cloud, LMC; and Small Magellanic Cloud, SMC; Pei 1992) to derive the line-of-sight extinction in the host galaxy. Due to the strongly variable light curve, we choose the approach Kann et al. (2006) used for the SED of GRB 030329, and shift the other bands to the  $R_C$ -band zero point to derive the colours. With this method, we can also look for colour changes. There may be marginal variations in  $B - R$ , but this colour remains constant within conservative errors, and is more sensitive to Lyman forest blanketing, as different filters will suffer a different amount of blanketing.

The SED clearly shows the decreasing flux in the  $B$ ,  $g'$  and  $V$  bands, and especially in the  $uvw2$   $uvm2$   $uvw1$   $u$  bands, where only upper limits are found, due to the Lyman forest blanketing as well as the Lyman cutoff (see Fig. 5), and we thus do not include these filters in our fit. Given the size of the errors, all fits, even with no extinction, are acceptable (see Table 4) and we are thus unable to prefer one dust model over another. The lack of  $z$  band data does not allow us to constrain the existence of a 2175 Å bump, therefore we have no evidence in favor of or ruling out MW and LMC dust. We do note, though, that MW dust leads to slightly negative extinction, while the LMC dust fit yields large errors. Thus, we henceforth use the SMC dust fit, as this is the most common dust type found in GRB host galaxies, both in pre-*Swift* (Kann et al. 2006; Starling et al. 2007) and *Swift*-era (Schady et al. 2007; Kann et al. 2010) results. Note that for SMC dust, the extinction is 0 within errors as well.

Assuming the cooling break  $\nu_c$  to lie blueward of the optical bands (highly likely considering the X-ray spectral slope  $\beta_X \approx 1$ , Dai et al. 2007, and the result of the numerical modelling), the standard fireball model gives for the electron power law index  $p = 2\beta + 1 = 2.10 \pm 0.40$ . This result (albeit with large errors) is in agreement with the result from the broadband modelling (Sect. 4.2) and with the canonical  $p = 2.2$ , and very similar to many other GRB afterglows (e.g. Kann et al. 2006; Starling et al. 2007). Our results are in contrast to Dai et al. (2007), who derive a very steep slope  $\beta_0 = 1.69_{-0.49}^{+0.53}$  from  $Br'i'$  data only (correcting for Lyman absorption) and conclude that the optical and the X-ray data lie on the same slope.



**Fig. 5.** Spectral energy distribution of the afterglow of GRB 060526 in  $uvw2$   $uvm2$   $uvw1$   $uBg'Vr'Rc'i'IcJHKs$ , and fits with no extinction (straight black line), MW extinction (dotted line), LMC extinction (dashed line) and SMC extinction (thick dash-dotted line). The fit with MW dust finds (unphysical) negative extinction, causing an “emissive component” instead of a 2175 Å absorption bump. Data beyond  $2.2 \times 10^{15}$  Hz ( $Vg'Bu$   $uvw1$   $uvm2$   $uvw2$ ) were not included in the fit due to Lyman forest blanketing, the grey curves represent extrapolations. The extinction curves of Pei (1992) are not correctly defined beyond  $3.2 \times 10^{15}$  Hz. The UVOT UV filters are upper limits only, showing the strong flux decrease beyond the Lyman cutoff. The flux density scale is measured at the break time.

**Table 4.** Fits to the spectral energy distribution of GRB 060526.

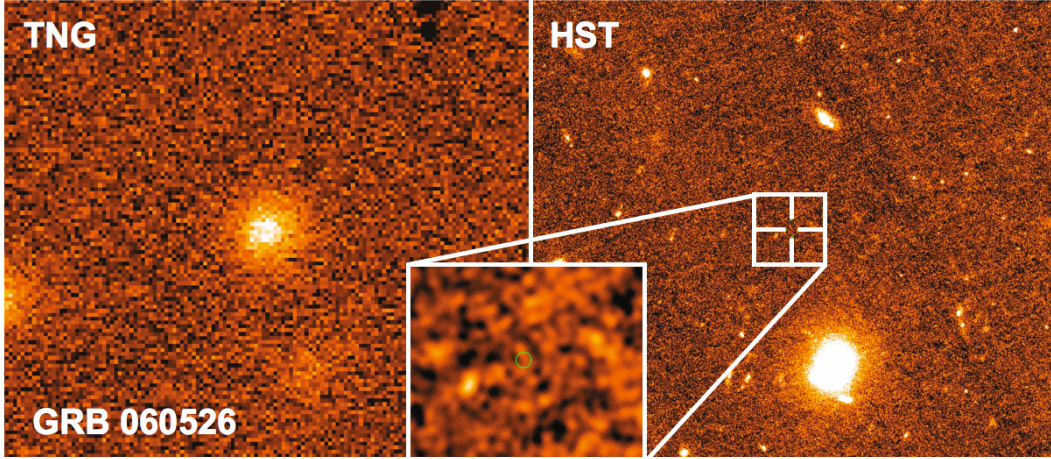
Dust	$\chi^2/\text{d.o.f.}$	$\beta$	$A_V$
none	0.425	$0.695 \pm 0.035$	...
MW	0.309	$0.833 \pm 0.149$	$-0.126 \pm 0.132$
LMC	0.466	$0.494 \pm 0.397$	$0.113 \pm 0.223$
SMC	0.396	$0.552 \pm 0.198$	$0.055 \pm 0.075$

**Notes.** Columns are the dust model, the goodness of the fit, the spectral slope and the derived extinction in the rest-frame  $V$  band.

#### 4.4. Host search

In the late-time VLT images, we do not detect any source at the position of the afterglow down to a limit of  $R_C > 27.1$  mag (which transforms into an absolute magnitude limit of  $M_R > -20.1$ ). There is one source present at  $\sim 1''.5$  South-East of the afterglow position with  $R_C = 26.4 \pm 0.2$ , which at a redshift of  $z = 3.221$  would transform into a physical offset between GRB and host galaxy of  $\sim 11.5$  kpc. A long-slit spectrum covering the afterglow position and this galaxy does not show any trace at these positions. If this galaxy was associated with the GRB, it would be one of the largest offsets between a long-duration GRB and its host galaxy known (Bloom et al. 2002). The strong absorption lines from the ISM seen in the afterglow spectrum (see Sect. 5) as well as the high circumburst density inferred from the numerical modeling (see Sect. 4.2) do in fact not favour a large distance from the host galaxy. A position outside of their host galaxies had been inferred for GRB 070125 (Cenko et al. 2008) and GRB 071003 (Perley et al. 2008b), however, those spectra showed very weak absorption lines, contrary to what we observe for GRB 060526. Concluding, we have no direct spectroscopic evidence for or against an association of the GRB with the galaxy.

We also observed the field with ACS on *HST*. Astrometry was performed relative to our TNG image, yielding a position



**Fig. 6.** *Left panel:* an image from the TNG with the afterglow present. *Right panel:* the same region observed with *HST/ACS* and a zoom-in around the region of the afterglow. There is no source detected at the afterglow position itself down to a limit of  $F775W(AB) > 28.5$ . The brighter spot to the South-East in the zoomed *HST* image corresponds to the host candidate from the VLT mentioned in the text. A color version of the figure is available in the online version of the article.

accurate to  $\sim 0.1''$ . At the location of the burst we do not find any evidence for an underlying host galaxy. To estimate the limits we place 50 apertures randomly on the sky and measure the standard deviation in their count rates. This implies that the underlying host galaxy of GRB 060526 is fainter than  $F775W(AB) > 28.5$  ( $3\sigma$ ) or an absolute magnitude of  $M(1500 \text{ \AA}) > -18.3$  mag which is fainter than  $0.5 L^*$  according to Gabasch et al. (2004). The flux density at the host position is  $0.004 \pm 0.005 \mu\text{Jy}$ .

Host galaxies of long-duration GRBs have often been found to be faint irregular galaxies (Fruchter et al. 2006; Christensen et al. 2004) which are difficult to detect at higher redshifts. So far, there are only eight bursts with  $z > 3$  where the detection of a host galaxy has been published, namely GRB 971214 ( $z = 3.418$ , Kulkarni et al. 1998), GRB 000131 ( $z = 4.500$ , Andersen et al. 2000; Fruchter et al. 2006), GRB 030323 ( $z = 3.3718$ , Vreeswijk et al. 2004), GRB 060206 ( $z = 4.04795$ , Fynbo et al. 2006; Thöne et al. 2008; Chen et al. 2009), GRB 060210 ( $z = 3.9133$ , Fynbo et al. 2009; Perley et al. 2009), GRB 060605 ( $z = 3.773$ , Ferrero et al. 2009), GRB 090205 ( $z = 4.6503$ , D’Avanzo et al. 2010), and GRB 090323 ( $z = 3.568$ , Cenko et al. 2010; McBreen et al. 2010). The host galaxies of GRB 020124 (Berger et al. 2002; Chen et al. 2009), GRB 050730, GRB 050908, GRB 060607A, GRB 070721B (all Chen et al. 2009), GRB 050904 (Berger et al. 2007) and GRB 060510B (Perley et al. 2009), on the other hand, were not detected to very deep limits in the optical and the NIR.

The distribution of pre-*Swift*  $R_C$  band host galaxy magnitudes peaks at  $R_C(AB) = 25$  (Fruchter et al. 2006) but extends out to 29 with a typical redshift  $z \approx 1.4$ . *Swift* GRBs (and thus their hosts) however have a higher mean redshift of  $z = 2.8$  (Jakobsson et al. 2006a)<sup>4</sup>, so the distribution will be shifted out to even fainter magnitudes. Ovaldsen et al. (2007) also find a higher magnitude for *Swift* hosts than for pre-*Swift* bursts by comparing the expected detection rate from pre-*Swift* hosts with detections and upper limits derived from imaging the fields of 24 *Swift* and *HETE II* bursts from 2005–2006.

<sup>4</sup> An updated version of the redshift distribution can be found under <http://raunvis.hi.is/~pja/GRBSample.html> – as of July 1, 2010, the mean redshift has become lower,  $z = 2.19$ , see also Fynbo et al. (2009), who find a mean and median of  $z = 2.2$ .

## 5. Spectroscopy results

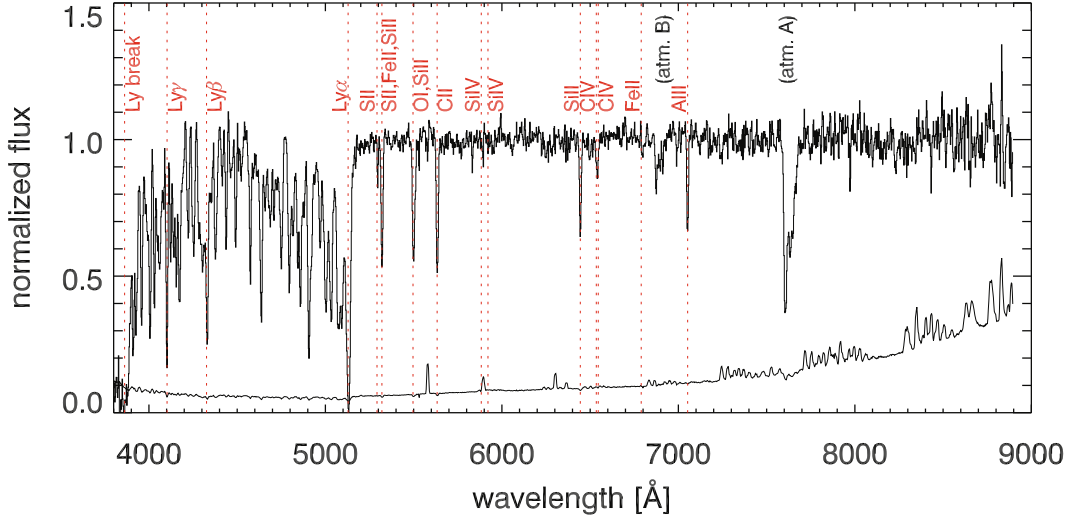
### 5.1. Line identification

We detect a range of metal absorption lines as well as a Lyman limit system (LLS) originating in the host galaxy of GRB 060526. A redshift of  $z = 3.221$  was determined in Jakobsson et al. (2006b) from a number of these absorption lines using the spectra taken with the 600V grism presented in this article.

Most of the lines were fitted from the combined 600V spectrum which covers all metal absorption lines detected longwards of Ly- $\alpha$ , but provides a higher resolution than the 300V spectra. The 600I spectra only covers the AlIII  $\lambda$  1670 line at the same resolution as the 600V grism. The range of 1200B is entirely within the Lyman- $\alpha$  forest but does not have a high enough resolution to deblend metal transitions from absorption caused by the Ly- $\alpha$  forest lines. We do, however, detect Ly- $\beta$  and Ly- $\gamma$  absorption in the 1200B grism. In the blue end of the 300V spectrum, one can clearly see the 915 Å Lyman break at the redshift of the host galaxy. In Fig. 7, we show the combined spectrum of the 300V grism with the identified lines indicated. We note that we do not detect any intervening system in the sightline towards GRB 060526, which is rather unusual for a GRB sightline, in particular at that redshift (see, e.g., Prochter et al. 2006). Taking the strong Mg II  $\lambda$  2796, 2803 doublet, the redshift path probed for any intervening system is between  $z \sim 0.8$  and 2.2. The  $3\sigma$  limits on the non-detection of the Mg II doublet vary between  $0.17 \text{ \AA}$  ( $z = 0.8$ ) and  $0.82 \text{ \AA}$  ( $z = 2.2$ ) (rest frame).

In order to determine the equivalent width ( $EW$ ) of the strong absorption lines, we fitted the continuum around the lines in regions that were free of absorption and summed over the absorption contained within two times the full-width at half-maximum (FWHM) of the lines. For weak lines, we obtained better results due to the low S/N by fitting Gaussians. For this fit, we used a modified version of the gaussfit procedure provided in IDL<sup>5</sup> which is more reliable in determining the continuum and fitting the actual line even if it is slightly blended with a neighbouring line. The upper limits on the  $EW$ s for a range of ions noted in Table 5 was determined from the spectra taken with the 300V grism due to the better S/N of those spectra. Between the

<sup>5</sup> Available at <http://www.pa.iasf.cnr.it/~nicastro/IDL/Lib/gfit.pro>



**Fig. 7.** The spectrum taken with 300V grism, the two exposures have been combined to improve the S/N. The identified lines as well as the main atmospheric absorption bands are labelled. The line at the bottom shows the error spectrum. A color version of the figure is available in the online version of the article.

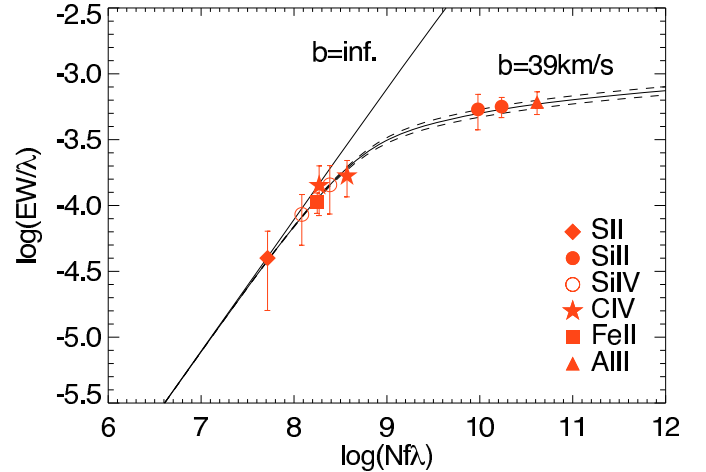
individual spectra, we do not find any variability in the  $EW$  of the individual absorption lines.

S II  $\lambda$  1259, Si II  $\lambda$  1260 and Fe II  $\lambda$  1260 are blended and cannot be fitted separately. We therefore cannot consider them for the derivation of the column density from the curve of growth fit as described below and only give the total  $EW$  in Table 5. In contrast to what is noted in Jakobsson et al. (2006b), we cannot reliably detect any fine-structure lines and only give an upper limit for Si II\*. Fine-structure lines would be a clear indication that the detected absorption lines indeed originate in the host galaxy of the GRB, as they are assumed to be produced by UV pumping from the afterglow (e.g. Prochaska et al. 2006; Vreeswijk et al. 2007; D’Elia et al. 2009). The redshift derived is therefore to be strictly taken as a lower limit only, the detection of the Lyman  $\alpha$  forest redward of the proposed redshift, however, excludes a significantly higher redshift for the burst.

## 5.2. Column densities from curve of growth analysis

Some of the strong absorption lines are saturated, which is a problem in low resolution spectra as the damping wings are not resolved and Voigt profile (VP) fitting cannot be adopted to derive a reliable column density. Furthermore, high resolution spectra of GRBs (Prochaska 2006) have shown that the strong metal absorption lines unresolved in low resolution spectra usually consist of a number of narrow, unsaturated components that would allow an accurate determination of the column density by fitting the different components separately.

If only low resolution spectra are available, one has to adopt a curve of growth (CoG) analysis (Spitzer 1978; Savage & Sembach 1996) which directly relates the  $EW$  to the column density on the linear part of the CoG where the lines are not saturated (optical depth  $\tau_0 < 1$ ), but depends on the Doppler  $b$  parameter of the medium on the flat part of the curve which applies for mildly saturated lines. For GRBs, one usually has to do a multiple-ion single-component CoG (MISC-CoG) analysis, adopting the same effective Doppler parameter for all ions. Here, we used all unblended ionic lines that were not heavily saturated, namely S II, Si II, Fe II, Al II as well as C IV and Si IV and calculate the  $\chi^2$  minimum going through the parameter space for the column densities of each ion and a range of  $b$  parameters. We then find the best fit for  $b = 39 \pm 3 \text{ km s}^{-1}$  (see Fig. 8) and the



**Fig. 8.** Multi-ion single-component curve of growth fit for the absorption systems in GRB 060526 using 6 different ions. The CoG is expressed in units of the column density “N”, the oscillator strength of the transition “f” and the wavelength of the transition. As comparison,  $b = \infty$  is plotted, the dotted line marks the  $1 \sigma$  deviation from the fit. A color version of the figure is available in the online version of the article.

column densities as noted in Table 5. Most of the ions lie near the linear part of the CoG and are therefore rather independent of  $b$ , Si II and Al II lie on the flat part of the CoG and the column densities are only noted as lower limits as derived from the linear part of the CoG. We excluded the saturated O I and C II transitions from the fit and only list lower limits for these two ions for which we take the column density resulting from the linear part of the CoG. We also discarded the blended absorption lines of Si II  $\lambda$  1260, S II and Fe II even though Si II is the dominant contribution to the absorption since we detect two other unblended Si II transitions at  $\lambda$  1304 and  $\lambda$  1526 Å.

There are several problems connected with the use of the CoG that have to be considered. Different ionisation levels should actually be treated in separate analyses as they might occur in different regions in the absorbing system. However, we assume that the absorption takes place in a relatively small region of the host galaxy and the resolution of the spectrum does resolve different components of the absorbing material in the

**Table 5.** *EW*s of detected absorption lines and  $2\sigma$  upper limits on some undetected lines, the *EW* for the blended systems include the contributions from all lines.

$\lambda_{\text{obs}}$ [Å]	$\lambda_{\text{rest}}$ [Å]	ID	$z$	$EW_{\text{rest}}$ [Å]	$\log N/\text{cm}^{-2}$
4102.26	972.54	Ly $\gamma$	3.218	$2.39 \pm 0.10$	–
4329.70	1025.72	Ly $\beta$	3.221	$2.92 \pm 0.06$	–
5131	1215.67	Ly $\alpha$	3.221	–	$20.00 \pm 0.15$
(5254)	1242.80	N V	(3.221)	$<0.09$	$<13.66$
(5276)	1250.58	S II	(3.221)	$<0.08$	$<14.71$
5291.01	1253.81	S II	3.220	$0.05 \pm 0.02$	$14.58 \pm 0.25$
5320.00	1259.52	S II	3.221	$(1.32 \pm 0.07)$	–
	1260.42	Si II	3.221	<i>blended</i>	–
	1260.53	Fe II	3.221		–
(5335)	1264.74	Si II*	(3.221)	$<0.10$	$<12.76$
5497.30	1302.17	O I	3.222	$1.07 \pm 0.05$	$>15.15$
5506.82	1304.37	Si II	3.222	$0.70 \pm 0.21$	$15.87 \pm 0.16$
(5559)	1317.22	Ni II	(3.221)	$<0.11$	$<13.59$
5634.24	1334.53	C II	3.222	$1.51 \pm 0.10$	$>15.55$
5884.87	1393.76	Si IV	3.222	$0.20 \pm 0.08$	$13.50 \pm 0.14$
5923.05	1402.77	Si IV	3.222	$0.12 \pm 0.05$	$13.50 \pm 0.14$
6444.50	1526.71	Si II	3.221	$0.86 \pm 0.04$	$15.87 \pm 0.16$
6539.58	1548.20	C IV	3.223	$0.26 \pm 0.08$	$14.10 \pm 0.13$
6546.39	1550.78	C IV	3.221	$0.22 \pm 0.09$	$14.10 \pm 0.13$
6790.24	1608.45	Fe II	3.221	$0.17 \pm 0.03$	$14.28 \pm 0.24$
6800.38	1611.20	Fe II	3.221	$<0.01$	–
7052.62	1670.79	Al II	3.221	$1.02 \pm 0.05$	$15.12 \pm 0.18$
(7863)	1862.79	Al III	(3.221)	$<0.35$	$<12.71$
(8552)	2026.13	Zn II	(3.221)	$<0.45$	$<12.73$
(8679)	2056.25	Cr II	(3.221)	$<0.54$	$<13.47$

**Notes.** The column densities were derived from CoG fitted for S II, Si II, Fe II, C IV, Si IV and Al II. Si II\* denotes a fine-structure line. Upper and lower limits were determined by assuming the ions to lie on the linear part of the CoG. The Ly $\alpha$  column density is taken from [Jakobsson et al. \(2006b\)](#) and is based on the 600V grism.

host galaxy. Therefore, we also fit the higher ionisation levels of Si IV and C IV in the same CoG. Furthermore, they lie close to the linear part of the CoG and excluding them from the common fit would not change the derived  $b$  parameter very much. Another problem with doing multiple-ion CoG analysis using strong lines in low-resolution spectra has been noted by [Prochaska \(2006\)](#), who compared column densities derived from low resolution spectra and CoG with high resolution data and VP fitting from the same bursts. He found that when including saturated lines in the fit, column densities are generally underestimated. Indications for that are if an effective  $b$  parameter of  $\gg 20 \text{ km s}^{-1}$  is found, since strong lines actually consist of a range of components with  $b < 20 \text{ km s}^{-1}$ . [Savaglio \(2006\)](#) however performed a similar analysis using CoG and the apparent optical depth (AOD) method as described in [Pettini et al. \(2002\)](#), which can be applied for medium resolution high S/N spectra, and found a good agreement between the two methods. The column densities of the saturated lines Al II and Si II are in fact rather sensitive to the adopted  $b$  parameter. These lines very likely consist of a number of unresolved weaker components which would lie on the linear part of the CoG, the real errors should therefore be larger but are difficult to estimate.

### 5.3. Metallicity and relative abundances

Absorption lines that are likely not affected by dust depletion can be used to derive a metallicity of the medium in the line of sight to the GRB. The least dust-depleted element is Zn, which is, however, undetected in our spectra. We then use the relatively weak lines S II and Fe II derive relative metallicities compared to the hydrogen density with  $[M/H] = \log(N_M/N_H) - \log(N_M/N_H)_\odot$

using solar abundances from [Asplund et al. \(2005\)](#). Here we derive metallicities of  $[S/H] = -0.57 \pm 0.25$  and  $[Fe/H] = -1.22 \pm 0.24$ .

Fe is usually affected by dust depletion ([Savage & Sembach 1996](#)) and corrections have to be adopted. Using the relation between the Zn and Fe abundance in [Savaglio \(2006\)](#), we find a metallicity of  $[Fe/H] = -1.09 \pm 0.24$  which marginally agrees within the errors with the value derived from sulphur. The only detected and unblended S II line at  $\lambda 1253$ , however, is only marginally detected and therefore the *EW* has large errors. The Fe II doublet taken for the CoG fit, in contrast, is also slightly blended, but the fit of the stronger component can be considered as reliable. Despite the dust depletion, the metallicity derived from Fe might be the most reliable one in this case and we therefore assume a metallicity of  $[Fe/H] = -1.09$  for the host of GRB 060526.

Independent of the ion used, the metallicity is rather high compared to other galaxies at redshift  $z \approx 3$ , but among the typical metallicities derived for other GRB hosts ([Fynbo et al. 2006](#)) at that redshift. For those measurements, different ions have been used depending on the quality and the wavelength coverage of the spectra. Our results show that caution is required when comparing the metallicities derived from different elements, as they might be differently affected by dust depletion and/or evolution. This is especially true when saturated lines in low resolution spectra are used to derive the column densities and hence the abundances as it is the case for, e.g., GRB 000926 ([Savaglio et al. 2003](#)) and GRB 011211 ([Vreeswijk et al. 2006](#)). For GRB 050401 ([Watson et al. 2006](#)) and GRB 050505 ([Berger et al. 2006](#)), the authors themselves note that due to saturation the reported metallicity is indeed a lower limit.

Our spectra do not allow us to determine a dust depletion pattern from the relative abundances of heavier elements, since at least four elements out of Zn, Si, Mn, Cr, Fe, Ni or S are necessary to do such a fit. The difference in the relative abundances between S and Fe might suggest the presence of some dust in the line-of-sight towards the GRB, however, large extinction is excluded from the afterglow SED. On the other hand, this difference could also be due to an enhancement in  $\alpha$  element production<sup>6</sup> which is likely to happen in the young star-forming host of a GRB. Generally, GRB hosts have higher  $\alpha$ /Fe ratios than QSO-DLAs (Prochaska et al. 2007), which can either be interpreted as large dust depletion consistent with the higher metallicities of GRB sightlines or as  $\alpha$  element enhancement. There are indications that most of the  $\alpha$ /Fe ratio is due to dust depletion traced by a large  $[\text{Zn}/\text{Fe}]$  or  $[\text{Ti}/\text{Fe}] < 0$ . The latter is assumed to be a clear indicator of dust depletion ( $[\text{Ti}/\text{Fe}] < 0$ ) vs.  $\alpha$  enhancement ( $[\text{Ti}/\text{Fe}] > 0$ ) (Dessauges-Zavadsky et al. 2002). Both elements are, however, not detected and the limit on Zn does not lead to a strong constraint on the  $[\text{Zn}/\text{Fe}]$  ratio to settle this issue in the case of GRB 060526.

The ratio between high and low ionisation species in the spectrum clearly shows that most of the material in the line of sight is in a low ionization state. Si IV and C IV have rather low column densities and lie close to the linear part of the CoG in contrast to their low ionization species Si II and C II which are both saturated, whereas Al III is not even detected in our spectra. We then derive column density ratios of  $\log(\text{Si IV}/\text{Si II}) = -2.37$ ,  $\log(\text{C IV}/\text{C II}) < -1.45$  and  $\log(\text{Al III}/\text{Al II}) < -2.41$ . From the large sample of *Swift* long GRB afterglow spectra (Fynbo et al. 2009), this seems fairly normal for an average GRB sightline. However, the few GRBs occurring in LLSs ( $\log N_{\text{HI}}/\text{cm}^{-2} < 20.3$ ) usually show a higher fraction of ionized material compared to GRB-DLA sightlines. This might either be due to decreased shielding of the highly ionizing afterglow flux by the lower hydrogen column density, by a rather special arrangement of the GRB inside the host galaxy or simply by a small host galaxy. Since our spectra show a low fraction of highly ionized material this implies that the absorbing material has a relatively large distance from the GRB.

We do not detect NV in our spectra but provide an upper limit of  $\log N < 13.66$ . NV was detected in only four sightlines towards GRBs (Prochaska et al. 2008) and likely traces the immediate environment of the GRB as it has a high ionization potential and requires a strong radiation field. Our upper limit is lower than the column densities for those GRB sightlines where NV could be detected which again implies that the absorbing gas probed by our spectra are most likely not close to the GRB itself. Furthermore, we do not detect any fine structure lines and only derive an upper limit on Si II\* of  $\log N/\text{cm}^{-2} < 12.76$ . Fine structure lines are assumed to be pumped by the UV radiation field of the GRB (Vreeswijk et al. 2007) which also indicates that the gas is likely very far from the GRB itself. Si II\* requires a less strong radiation field and has been detected in Lyman break galaxies (Pettini et al. 2002) where the UV radiation from young stars provides the necessary radiation field.

## 6. Discussion and conclusions

GRB 060526 had a relatively bright afterglow that allowed us to obtain a solid dataset, both photometrically and

spectroscopically. We achieved a dense light curve coverage over several days which allowed a detailed study of the afterglow properties, and obtained a series of low resolution but high signal-to-noise spectra to study the host environment.

The optical light curve can be fitted with a double smoothly-broken power law with breaks at  $t_{b1} = 0.090 \pm 0.005$  and  $t_b = 2.216 \pm 0.049$  days, and decay slopes of  $\alpha_{\text{plateau}} = 0.288 \pm 0.026$ ,  $\alpha_1 = 0.971 \pm 0.008$ , and  $\alpha_2 = 2.524 \pm 0.052$ . The dense sampling of especially the  $R_C$ -band light curve reveals additional variability on top of the power laws. These features could be explained either by extended activity of the central engine or through interactions of the shock with the interstellar medium. For the case that the variability arises from external shocks, several mechanisms have been considered. In GRB 021004, both density variations of the external medium into which the GRB jet plows and angular inhomogeneities of the jet surface were considered (Nakar et al. 2003). However, Nakar & Granot (2007) show that density variations would cause much smaller fluctuations than those observed in GRB afterglows and can therefore be ruled out. Another possibility is the injection of additional energy into the shock by slower shells that catch up with the shocked region as it decelerates, this model was used successfully to describe GRB 021004 (de Ugarte Postigo et al. 2005) and also works better than two other models (double jet and density fluctuations) to describe the highly complex light curve of GRB 030329 (Huang et al. 2006). Thus, variability can give either information on the medium surrounding the GRB or on the activity of the central engine. A more intriguing possibility is that the flares may be emitted from another region closer to the central engine, resulting from late internal shocks. Powerful X-ray flares that are attributed to late central engine activity have been observed in about 50% of all *Swift* GRBs (e.g. Burrows et al. 2005; Chincarini et al. 2007; Krimm et al. 2007), and strong optical/NIR flaring contemporaneous with the GRB prompt emission may also occur (Vestrand et al. 2005, 2006; Blake et al. 2005; Racusin et al. 2008), thus making optical flares from late central engine activity an interesting prospect (Kann 2007; Malesani et al. 2007).

Indeed, Dai et al. (2007) have suggested that the optical variability of the afterglow of GRB 060526 is due to flares from late internal shocks (the very early rapid optical variability we present here is very probably due to central engine activity, as it is seen contemporaneously in gamma and X-rays). Khamitov et al. (2007), on the other hand, conclude that the short timescale of the variabilities requires the jet to be non-relativistic already at  $\sim 1$  day and could then be explained by external density fluctuations. Our analysis lends tentative support to the notion of flares from internal shocks, finding decay slopes for two flares that exceed what should be possible from external shocks. But we caution that the errors of these fits are large due to a low amount of data in the decaying parts. Furthermore, globally, a model using refreshed external shocks is able to account for the light curve variations, although microvariability remains. This creates the intriguing possibility of reverberation effects (see Vestrand et al. 2006, for a case of reverberation between gamma-rays and optical emission). Short flares in the X-ray or optical bands signal internal shocks from long-term central engine activity, and when these shells catch up with the forward shock front, they re-energise the external forward shock. The detection of such behaviour would probably require dense multi-band observations of a bright afterglow to search for SED changes at high time resolution combined with detailed modelling of the data. This way, one could discern between internal shocks (which are expected to have a different spectral index from the

<sup>6</sup>  $\alpha$  elements are produced in massive, metal-poor stars through the  $\alpha$  process and include elements with integer multiples of the He nucleus mass such as O, Si, S, Ca, Mg and Ti.

forward shock afterglow) and refreshed, external shocks (which are achromatic). Our data set of the afterglow of GRB 060526 does not allow us such a detailed decomposition.

From the analysis of our low resolution spectra with different resolutions, we detect a LLS and a number of metal absorption lines that all lie at a redshift of  $z = 3.221$ . The low resolution only allows us to derive column densities from measuring the *EWs* of the absorption lines and adopt a MISC-CoG analysis where we exclude the most saturated as well as blended transitions. We find a best fit for the Doppler parameter of  $b = 39 \pm 3 \text{ km s}^{-1}$  and most of the ions used for the fit lie on the linear part of the CoG which allows a relatively reliable determination of the column densities. The relative abundances of different metals in the spectra indicate some dust extinction, but an intrinsic difference due to enhancement of the production of certain elements cannot be excluded. The very low amount of dust detected in the afterglow SED may indicate that the latter might be the favored possibility.

The column density of neutral hydrogen is rather low compared to other GRBs. We derive a metallicity for the host of  $[\text{Fe}/\text{H}] = -1.09$  which is slightly higher than metallicities determined from other GRB afterglow spectra. According to the definition of QSO absorbers, the host of GRB 060526 is classified as a LLS ( $19 < \log N_{\text{HI}}/\text{cm}^{-2} < 20.3$ ), which seem to have on average higher metallicities than damped Lyman  $\alpha$  systems (DLA; Peroux et al. 2007) and a steeper evolution towards lower redshifts. Around redshift 3, however, the metallicities of both samples are within the same range. Also, GRB hosts show a trend towards increasing metallicity with lower redshifts (Fynbo et al. 2006; Savaglio 2006). Taking into account that most of the sample used only low-resolution spectra to derive the metallicity (which only gives lower limits for the column densities and the metallicity) this evolution might, however, not be as pronounced as for DLAs and LLS. This might imply that the enrichment of the ISM in the early universe had taken place at earlier times than assumed. The absorbing material along the line-of-sight is mostly in the neutral state, as usually observed for long GRB-DLAs, while sightlines with lower  $\log N_{\text{HI}}$  often contain more ionized material. This might either imply that we have a very small host galaxy or that the GRB is placed somewhere in the outskirts of its host. In general, the low ionization points to a relatively large distance of the absorbing material from the GRB itself.

There is no underlying host galaxy of GRB 060526 detected down to a deep limit of 28.5 mag (in F775W AB) in HST/ACS data. At that redshift, this means the host has an absolute magnitude  $M(1500 \text{ \AA}) > -18.3$  mag, fainter than an  $0.5 L^*$  galaxy at that redshift. Long GRBs have been found to occur in actively star-forming galaxies and star formation is assumed to shift towards smaller and fainter galaxies over time (e.g. Cowie et al. 1996) while massive galaxies prove to be rather unchanged throughout the history of the universe (e.g. Abraham et al. 1999; Heavens et al. 2004). One would therefore expect that GRB hosts should also have higher luminosities towards higher redshifts. Fynbo et al. (2008) concluded that the observed metallicity distribution of GRB hosts (as well as QSO absorbers) at  $z \approx 3$  can be explained by the luminosity function of galaxies at that redshift and assuming a luminosity-metallicity relation as derived for other high-redshift samples (Ledoux et al. 2006; Erb et al. 2006). The non-detection of the host of GRB 060526 down to deep limits, however, would not support this suggested evolution. The data neither strongly support nor allow us to rule out that the GRB is associated with a nearby galaxy at  $\sim 11.5$  kpc. While the offset would be very large compared to typical long

GRB offsets, it is possible the burst occurred in a locally dense star-forming region which is not detected even in our very deep imaging.

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**Table 1.** Broad band observations of GRB 060526, times given are the midpoint of the observations.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
0.0250277	509.94	UVOT	<i>uvw2</i>	> 21.0	...
0.4692036	6180.78	UVOT	<i>uvw2</i>	>22.7	...
1.8554459	2306.75	UVOT	<i>uvw2</i>	>22.1	...
3.0522168	2212.79	UVOT	<i>uvw2</i>	>22.1	...
3.9406911	1778.93	UVOT	<i>uvw2</i>	>21.9	...
0.0238351	471.09	UVOT	<i>uwm2</i>	>20.6	...
0.8961627	3319.91	UVOT	<i>uwm2</i>	>22.0	...
1.8577054	2306.9	UVOT	<i>uwm2</i>	>21.8	...
3.0550826	2212.9	UVOT	<i>uwm2</i>	>21.7	...
3.9434308	1775.78	UVOT	<i>uwm2</i>	>21.6	...
0.0246621	490.49	UVOT	<i>uvw1</i>	>20.9	...
0.9543864	1030.99	UVOT	<i>uvw1</i>	>21.4	...
1.8588450	2306.88	UVOT	<i>uvw1</i>	>21.9	...
3.0565271	2192.57	UVOT	<i>uvw1</i>	>21.9	...
3.9448082	1777.08	UVOT	<i>uvw1</i>	>21.7	...
0.0254789	490.45	UVOT	<i>u</i>	>21.0	...
1.4941792	40.13	UVOT	<i>u</i>	>19.3	...
1.8599737	2169.1	UVOT	<i>u</i>	>21.9	...
3.0579600	2210.83	UVOT	<i>u</i>	>21.9	...
3.9461762	1773.99	UVOT	<i>u</i>	>21.8	...
0.0122205	58.12	UVOT	<i>b</i>	18.270	0.160
0.0392111	216.07	UVOT	<i>b</i>	18.710	0.110
0.09394	180	Maidanak	<i>B</i>	19.518	0.207
0.09898	180	Maidanak	<i>B</i>	19.557	0.135
0.10625	180	Maidanak	<i>B</i>	19.432	0.136
0.10925	180	Maidanak	<i>B</i>	19.505	0.107
0.11421	900	Maidanak	<i>B</i>	19.546	0.179
0.23242	100	RTT150	<i>B</i>	20.133	0.069
0.23813	100	RTT150	<i>B</i>	20.321	0.083
0.24525	100	RTT150	<i>B</i>	20.316	0.081
0.26704	300	RTT150	<i>B</i>	20.395	0.064
0.28883	300	RTT150	<i>B</i>	20.636	0.074
0.31108	300	RTT150	<i>B</i>	20.649	0.073
0.33371	300	RTT150	<i>B</i>	20.839	0.085
1.10542	900	RTT150	<i>B</i>	22.072	0.097
1.16521	900	Maidanak	<i>B</i>	22.013	0.079
1.30279	900	RTT150	<i>B</i>	22.246	0.154
1.4946505	30.99	UVOT	<i>b</i>	>19.5	...
1.8609920	1909.33	UVOT	<i>b</i>	>22.1	...
2.14626	2880	Maidanak	<i>B</i>	23.290	0.072
2.35104	900	MOSCA	<i>B</i>	23.104	0.082
3.0592114	1810.67	UVOT	<i>b</i>	>22.1	...
3.13135	2700	Maidanak	<i>B</i>	24.198	0.115
3.9474369	1536.37	UVOT	<i>b</i>	>22.0	...
3.73439	660	Keck/LRIS	<i>g'</i>	23.952	0.022
4.87114	660	Keck/LRIS	<i>g'</i>	24.480	0.120
0.0021561	10	UVOT	<i>v</i>	16.610	0.280
0.0022719	10	UVOT	<i>v</i>	16.540	0.190
0.0023876	10	UVOT	<i>v</i>	16.480	0.180
0.0025034	10	UVOT	<i>v</i>	16.660	0.200
0.0026191	10	UVOT	<i>v</i>	16.830	0.210
0.0027348	10	UVOT	<i>v</i>	17.000	0.230
0.0028506	10	UVOT	<i>v</i>	16.710	0.200
0.0029663	10	UVOT	<i>v</i>	15.940	0.140
0.0030821	10	UVOT	<i>v</i>	16.350	0.170
0.0031978	10	UVOT	<i>v</i>	16.610	0.190
0.0033135	10	UVOT	<i>v</i>	16.940	0.220
0.0034293	10	UVOT	<i>v</i>	16.670	0.200
0.0035450	10	UVOT	<i>v</i>	17.200	0.260
0.0036608	10	UVOT	<i>v</i>	17.140	0.250
0.0037765	10	UVOT	<i>v</i>	17.570	0.300
0.0038923	10	UVOT	<i>v</i>	18.000	0.380
0.0040080	10	UVOT	<i>v</i>	17.340	0.270
0.0041237	10	UVOT	<i>v</i>	17.010	0.230
0.0042395	10	UVOT	<i>v</i>	17.270	0.260
0.0043552	10	UVOT	<i>v</i>	16.920	0.220

Table 1. continued.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
0.0044710	10	UVOT	<i>v</i>	17.520	0.300
0.0045867	10	UVOT	<i>v</i>	17.580	0.310
0.0049254	50	UVOT	<i>v</i>	17.780	0.150
0.0054498	40	UVOT	<i>v</i>	17.720	0.160
0.0059132	40	UVOT	<i>v</i>	17.710	0.160
0.0064320	50	UVOT	<i>v</i>	17.900	0.160
0.0082840	41.68	UVOT	<i>v</i>	17.770	0.160
0.0105975	70	UVOT	<i>v</i>	18.160	0.150
0.0112959	50	UVOT	<i>v</i>	17.900	0.160
0.0118748	50	UVOT	<i>v</i>	17.930	0.160
0.0125663	70	UVOT	<i>v</i>	18.240	0.160
0.0136547	120	UVOT	<i>v</i>	18.820	0.160
0.0169180	66.88	UVOT	<i>v</i>	18.250	0.160
0.0356456	210.99	UVOT	<i>v</i>	18.370	0.120
0.1052804	311.39	UVOT	<i>v</i>	18.820	0.140
0.19284	60	DOLORES	<i>V</i>	19.235	0.087
0.19443	60	DOLORES	<i>V</i>	19.338	0.065
0.22210	120	MIRO	<i>V</i>	19.274	0.167
0.23067	100	RTT150	<i>V</i>	19.302	0.046
0.23583	100	RTT150	<i>V</i>	19.527	0.047
0.24254	100	RTT150	<i>V</i>	19.364	0.045
0.24818	1200	BFOSC	<i>V</i>	19.403	0.081
0.26313	300	RTT150	<i>V</i>	19.584	0.040
0.28487	300	RTT150	<i>V</i>	19.669	0.040
0.28809	1200	BFOSC	<i>V</i>	19.655	0.078
0.30667	300	RTT150	<i>V</i>	19.821	0.040
0.32942	300	RTT150	<i>V</i>	19.861	0.048
0.33005	1200	BFOSC	<i>V</i>	19.870	0.083
0.41910	600	DFOSC	<i>V</i>	20.155	0.073
0.43507	600	DFOSC	<i>V</i>	20.200	0.057
1.10946	900	RTT150	<i>V</i>	21.027	0.055
1.31288	900	RTT150	<i>V</i>	21.105	0.060
1.40937	600	DFOSC	<i>V</i>	21.267	0.037
1.41685	600	DFOSC	<i>V</i>	21.273	0.040
1.42464	600	DFOSC	<i>V</i>	21.276	0.038
1.42725	11382.11	UVOT	<i>v</i>	21.340	0.170
2.32642	900	RTT150	<i>V</i>	22.203	0.123
2.36561	600	DFOSC	<i>V</i>	22.279	0.062
3.14483	3000	RTT150	<i>V</i>	22.818	0.094
3.23404	3000	RTT150	<i>V</i>	22.834	0.094
3.32054	3000	RTT150	<i>V</i>	23.123	0.154
4.24058	7200	RTT150	<i>V</i>	23.508	0.121
5.35387	8400	RTT150	<i>V</i>	24.322	0.211
0.0009974	10	UVOT	<i>white</i>	16.870	0.120
0.0011133	10	UVOT	<i>white</i>	16.800	0.110
0.0012292	10	UVOT	<i>white</i>	16.920	0.120
0.0013451	10	UVOT	<i>white</i>	17.050	0.130
0.0014609	10	UVOT	<i>white</i>	17.150	0.140
0.0015767	10	UVOT	<i>white</i>	17.050	0.130
0.0016925	10	UVOT	<i>white</i>	17.180	0.140
0.0018083	10	UVOT	<i>white</i>	17.060	0.130
0.0019241	10	UVOT	<i>white</i>	17.000	0.130
0.0020399	9.76	UVOT	<i>white</i>	16.930	0.120
0.0078848	9.77	UVOT	<i>white</i>	17.540	0.170
0.0122853	107.98	UVOT	<i>white</i>	18.090	0.080
0.0193600	29.28	UVOT	<i>white</i>	18.670	0.170
0.0744556	196.63	UVOT	<i>white</i>	18.730	0.090
0.0005741	10	Watcher	<i>R<sub>C</sub></i>	16.110	0.430
0.0013513	10	Watcher	<i>R<sub>C</sub></i>	>16.080	...
0.0017219	10	Watcher	<i>R<sub>C</sub></i>	>15.830	...
0.0025092	10	Watcher	<i>R<sub>C</sub></i>	>16.240	...
0.0028681	10	Watcher	<i>R<sub>C</sub></i>	>16.060	...
0.0036553	10	Watcher	<i>R<sub>C</sub></i>	>16.050	...
0.0068646	60	Watcher	<i>R<sub>C</sub></i>	16.920	0.400
0.0110924	60	Watcher	<i>R<sub>C</sub></i>	16.990	0.380
0.0195091	60	Watcher	<i>R<sub>C</sub></i>	17.290	0.390
0.0237226	60	Watcher	<i>R<sub>C</sub></i>	17.540	0.520

Table 1. continued.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
0.0342566	120	Watcher	$R_C$	17.740	0.120
0.0497130	240	Watcher	$R_C$	17.847	0.140
0.0651205	240	Watcher	$R_C$	18.030	0.140
0.0805210	240	Watcher	$R_C$	18.097	0.150
0.09030	180	Maidanak	$R_C$	18.039	0.062
0.09539	180	Maidanak	$R_C$	18.042	0.064
0.095935	240	Watcher	$R_C$	18.047	0.140
0.10079	180	Maidanak	$R_C$	18.117	0.067
0.1113202	240	Watcher	$R_C$	18.247	0.180
0.1267018	240	Watcher	$R_C$	18.137	0.150
0.1457669	360	Watcher	$R_C$	18.363	0.140
0.18919	60	DOLORES	$R_C$	18.627	0.109
0.19079	60	DOLORES	$R_C$	18.763	0.097
0.19165	180	MIRO	$R_C$	18.616	0.066
0.19390	180	MIRO	$R_C$	18.646	0.068
0.19615	180	MIRO	$R_C$	18.546	0.094
0.19838	180	MIRO	$R_C$	18.775	0.046
0.20914	600	BFOSC	$R_C$	18.723	0.087
0.21225	180	MIRO	$R_C$	18.851	0.143
0.21449	180	MIRO	$R_C$	18.778	0.273
0.21897	180	MIRO	$R_C$	18.756	0.232
0.22144	900	BFOSC	$R_C$	18.779	0.083
0.22679	150	RTT150	$R_C$	18.576	0.076
0.22892	150	RTT150	$R_C$	18.892	0.061
0.23417	150	RTT150	$R_C$	18.831	0.060
0.23428	900	BFOSC	$R_C$	18.821	0.078
0.24058	150	RTT150	$R_C$	18.888	0.066
0.24829	150	RTT150	$R_C$	18.881	0.071
0.25054	150	RTT150	$R_C$	18.935	0.071
0.25279	150	RTT150	$R_C$	18.960	0.070
0.25504	150	RTT150	$R_C$	18.805	0.070
0.25530	180	MIRO	$R_C$	18.815	0.155
0.25729	150	RTT150	$R_C$	18.910	0.071
0.25954	150	RTT150	$R_C$	18.917	0.073
0.27050	150	RTT150	$R_C$	18.994	0.075
0.27275	150	RTT150	$R_C$	19.105	0.081
0.27468	900	BFOSC	$R_C$	19.017	0.074
0.27500	150	RTT150	$R_C$	19.147	0.081
0.27721	150	RTT150	$R_C$	18.993	0.074
0.27946	150	RTT150	$R_C$	19.014	0.073
0.28171	150	RTT150	$R_C$	18.986	0.075
0.29208	150	RTT150	$R_C$	19.073	0.080
0.29438	150	RTT150	$R_C$	19.131	0.081
0.29662	150	RTT150	$R_C$	19.170	0.079
0.29888	150	RTT150	$R_C$	19.129	0.077
0.30113	150	RTT150	$R_C$	19.206	0.079
0.30342	150	RTT150	$R_C$	19.183	0.074
0.31462	150	RTT150	$R_C$	19.144	0.078
0.31599	900	BFOSC	$R_C$	19.332	0.080
0.31692	150	RTT150	$R_C$	19.128	0.075
0.31917	150	RTT150	$R_C$	19.292	0.082
0.32142	150	RTT150	$R_C$	19.349	0.083
0.32371	150	RTT150	$R_C$	19.289	0.087
0.32596	150	RTT150	$R_C$	19.303	0.088
0.33725	150	RTT150	$R_C$	19.290	0.083
0.33950	150	RTT150	$R_C$	19.275	0.086
0.34179	150	RTT150	$R_C$	19.402	0.089
0.34404	150	RTT150	$R_C$	19.336	0.089
0.34629	150	RTT150	$R_C$	19.341	0.089
0.34854	150	RTT150	$R_C$	19.454	0.101
0.35083	150	RTT150	$R_C$	19.466	0.103
0.35308	150	RTT150	$R_C$	19.505	0.104
0.35533	150	RTT150	$R_C$	19.364	0.105
0.35758	150	RTT150	$R_C$	19.383	0.106
0.35983	150	RTT150	$R_C$	19.417	0.106
0.36208	150	RTT150	$R_C$	19.434	0.107
0.36433	150	RTT150	$R_C$	19.462	0.111

Table 1. continued.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
0.36658	150	RTT150	$R_C$	19.270	0.101
0.36883	150	RTT150	$R_C$	19.492	0.109
0.37108	150	RTT150	$R_C$	19.542	0.122
0.39230	600	DFOSC	$R_C$	19.603	0.051
0.40064	600	DFOSC	$R_C$	19.607	0.039
0.40800	600	DFOSC	$R_C$	19.644	0.037
0.53319	600	DFOSC	$R_C$	19.820	0.044
0.54030	600	DFOSC	$R_C$	19.868	0.041
0.54841	600	DFOSC	$R_C$	19.899	0.042
0.62960	600	DFOSC	$R_C$	20.058	0.090
0.63685	300	DFOSC	$R_C$	20.030	0.142
0.64096	300	DFOSC	$R_C$	20.001	0.145
0.64468	300	DFOSC	$R_C$	19.982	0.158
1.10432	540	Shajn 2.6 m	$R_C$	20.382	0.120
1.11983	900	RTT150	$R_C$	20.426	0.069
1.12784	540	Shajn 2.6 m	$R_C$	20.478	0.113
1.13576	1080	Maidanak	$R_C$	20.551	0.041
1.31530	900	BFOSC	$R_C$	20.700	0.133
1.31725	900	RTT150	$R_C$	20.512	0.068
1.33046	900	BFOSC	$R_C$	20.582	0.108
1.35328	600	DFOSC	$R_C$	20.648	0.043
1.36091	600	DFOSC	$R_C$	20.668	0.040
1.37108	600	DFOSC	$R_C$	20.686	0.042
1.37829	600	DFOSC	$R_C$	20.699	0.038
1.49769	600	DFOSC	$R_C$	20.864	0.038
1.50549	600	DFOSC	$R_C$	20.836	0.039
1.51330	600	DFOSC	$R_C$	20.871	0.037
1.60597	600	DFOSC	$R_C$	20.954	0.047
1.61369	600	DFOSC	$R_C$	21.031	0.059
1.62117	600	DFOSC	$R_C$	20.986	0.067
2.11423	1260	Maidanak	$R_C$	21.565	0.100
2.14083	900	RTT150	$R_C$	21.373	0.080
2.15346	900	RTT150	$R_C$	21.330	0.080
2.16637	900	RTT150	$R_C$	21.389	0.082
2.29040	2400	TLS 1.34 m	$R_C$	21.700	0.130
2.32674	900	MOSCA	$R_C$	21.671	0.188
2.33858	900	RTT150	$R_C$	21.619	0.091
2.35204	900	RTT150	$R_C$	21.587	0.083
2.36504	900	RTT150	$R_C$	21.875	0.115
3.10104	1800	Maidanak	$R_C$	22.286	0.060
3.16792	3000	RTT150	$R_C$	22.062	0.072
3.25538	3000	RTT150	$R_C$	22.239	0.077
3.28438	1800	MOSCA	$R_C$	22.213	0.052
3.34992	3000	RTT150	$R_C$	22.338	0.095
3.73440	660	Keck/LRIS	$R_C$	22.758	0.042
4.11354	2700	Maidanak	$R_C$	23.181	0.100
4.23200	7200	RTT150	$R_C$	23.072	0.108
4.37326	1800	MOSCA	$R_C$	23.244	0.099
4.87336	900	Keck/LRIS	$R_C$	23.587	0.123
5.10382	3420	Maidanak	$R_C$	23.733	0.096
5.34483	9000	RTT150	$R_C$	23.680	0.201
5.35173	3600	ALFOSC	$R_C$	23.571	0.088
6.14534	3420	Maidanak	$R_C$	>23.6	...
7.17834	3600	Maidanak	$R_C$	>23.8	...
7.26820	12000	TLS 1.34m	$R_C$	>23.7	...
7.37882	3600	ALFOSC	$R_C$	24.602	0.039
10.3198	3600	ALFOSC	$R_C$	>24.6	...
272.654	2500	FORS2	$R_C$	>27.1	...
672.332	7500	FORS2	$R_C$	(combined)	...
0.0010848	10	Watcher	CR	16.500	0.400
0.0022314	10	Watcher	CR	16.160	0.220
0.0033890	10	Watcher	CR	16.460	0.220
0.0046149	60	Watcher	CR	16.830	0.380
0.0088341	60	Watcher	CR	17.090	0.450
0.0130493	60	Watcher	CR	17.230	0.440
0.0172518	60	Watcher	CR	17.110	0.360
0.0214654	60	Watcher	CR	17.430	0.440

Table 1. continued.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
0.0263665	120	Watcher	CR	17.420	0.320
0.0328664	38.6	ROTSE	CR	>17.412	...
0.0340538	120	Watcher	CR	17.620	0.370
0.0351278	143.89	ROTSE	CR	17.625	0.192
0.0385665	143.28	ROTSE	CR	17.570	0.170
0.0417402	120	Watcher	CR	17.590	0.330
0.0419979	143.28	ROTSE	CR	17.558	0.147
0.0454286	143.23	ROTSE	CR	18.054	0.206
0.0488545	143.18	ROTSE	CR	17.792	0.164
0.0494170	120	Watcher	CR	17.783	0.090
0.0522827	143.23	ROTSE	CR	17.785	0.153
0.0548713	67.49	ROTSE	CR	17.927	0.321
0.0571028	120	Watcher	CR	17.810	0.090
0.0574228	143.84	ROTSE	CR	17.723	0.155
0.0608539	143.43	ROTSE	CR	17.624	0.139
0.0639448	113.58	ROTSE	CR	17.841	0.188
0.0647885	120	Watcher	CR	17.883	0.100
0.0662157	38.6	ROTSE	CR	17.659	0.282
0.0684630	148	ROTSE	CR	17.876	0.200
0.0724741	120	Watcher	CR	17.850	0.090
0.0742371	343.17	ROTSE	CR	17.780	0.088
0.0801480	120	Watcher	CR	17.700	0.080
0.0822853	342.81	ROTSE	CR	17.905	0.074
0.0878566	120	Watcher	CR	17.790	0.090
0.0903340	342.96	ROTSE	CR	17.959	0.064
0.0955536	120	Watcher	CR	17.887	0.090
0.0983836	342.71	ROTSE	CR	18.049	0.091
0.1032274	120	Watcher	CR	18.070	0.100
0.1064264	342.4	ROTSE	CR	18.114	0.081
0.1109128	120	Watcher	CR	17.890	0.090
0.1179262	335.45	ROTSE	CR	18.141	0.079
0.1185981	120	Watcher	CR	18.210	0.120
0.1262346	342.96	ROTSE	CR	18.232	0.078
0.1262834	120	Watcher	CR	18.073	0.100
0.1339918	120	Watcher	CR	18.113	0.110
0.1342816	342.96	ROTSE	CR	18.255	0.069
0.1416771	120	Watcher	CR	18.323	0.140
0.1423343	343.47	ROTSE	CR	18.354	0.066
0.1493624	120	Watcher	CR	18.170	0.120
0.1503864	343.01	ROTSE	CR	18.407	0.070
0.1584325	342.25	ROTSE	CR	18.440	0.070
0.1611113	975	Watcher	CR	18.423	0.040
0.1682310	342.61	ROTSE	CR	18.484	0.180
0.1724372	975	Watcher	CR	18.560	0.050
0.1762647	342.3	ROTSE	CR	18.666	0.093
0.1843071	343.27	ROTSE	CR	18.578	0.082
0.1844514	1098	Watcher	CR	18.540	0.050
0.1923501	342.46	ROTSE	CR	18.541	0.070
0.1971843	1097	Watcher	CR	18.517	0.050
0.2003860	342.56	ROTSE	CR	18.637	0.075
0.2084307	342.61	ROTSE	CR	18.698	0.096
0.2106114	1220	Watcher	CR	18.707	0.060
0.2164729	342.96	ROTSE	CR	18.782	0.101
0.2245117	342.41	ROTSE	CR	18.789	0.102
0.2254456	1342	Watcher	CR	18.847	0.060
0.2325434	342.35	ROTSE	CR	18.881	0.084
0.2405816	342.46	ROTSE	CR	18.941	0.123
0.2410154	1343	Watcher	CR	18.753	0.040
0.2486180	342.51	ROTSE	CR	18.877	0.105
0.2566562	342.82	ROTSE	CR	18.958	0.095
0.2572685	1464	Watcher	CR	19.033	0.050
0.2647067	343.17	ROTSE	CR	18.888	0.170
0.2727542	342.96	ROTSE	CR	18.954	0.109
0.2749340	1587	Watcher	CR	18.987	0.050
0.2807917	342.3	ROTSE	CR	18.995	0.118
0.2888233	342.46	ROTSE	CR	19.122	0.132
0.2940183	1709	Watcher	CR	19.053	0.060

Table 1. continued.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
0.2968695	342.96	ROTSE	CR	19.125	0.210
0.3049105	342.61	ROTSE	CR	19.108	0.169
0.3129479	342.56	ROTSE	CR	19.318	0.159
0.3145154	1831	Watcher	CR	19.397	0.110
0.3209870	342.45	ROTSE	CR	19.300	0.132
0.3290197	342.4	ROTSE	CR	19.327	0.173
0.3330194	689.53	ROTSE	CR	19.360	0.174
0.3364198	1952	Watcher	CR	19.400	0.100
0.3370540	342.5	ROTSE	CR	19.305	0.163
0.3450996	342.86	ROTSE	CR	19.551	0.213
0.3531434	342.45	ROTSE	CR	19.353	0.196
0.3597482	2075	Watcher	CR	19.683	0.100
0.3651768	689.38	ROTSE	CR	19.464	0.201
0.3812533	690.3	ROTSE	CR	19.480	0.220
0.3844837	2197	Watcher	CR	19.777	0.100
0.3982605	770.22	ROTSE	CR	> 19.741	...
0.4241566	1462.69	ROTSE	CR	19.687	0.252
1.0566720	1037.58	ROTSE	CR	>19.602	...
1.0986892	8554	Watcher	CR	20.397	0.070
1.1985530	8688	Watcher	CR	20.343	0.080
1.3005326	8923	Watcher	CR	20.550	0.140
2.1764170	22125	Watcher	CR	21.507	0.170
2.4570745	1040.76	ROTSE	CR	>19.937	...
0.20221	180	MIRO	$I_C$	18.311	0.149
0.26296	900	BFOSC	$I_C$	18.549	0.070
0.30365	900	BFOSC	$I_C$	18.715	0.055
0.34481	900	BFOSC	$I_C$	18.901	0.070
0.35175	360	ANDICAM	$I_C$	18.940	0.090
0.35683	360	ANDICAM	$I_C$	18.937	0.062
0.36191	360	ANDICAM	$I_C$	19.052	0.056
0.36692	360	ANDICAM	$I_C$	18.967	0.043
0.37107	600	DFOSC	$I_C$	19.009	0.048
0.37200	360	ANDICAM	$I_C$	18.987	0.036
0.37888	600	DFOSC	$I_C$	19.050	0.045
0.38632	600	DFOSC	$I_C$	19.118	0.042
0.47240	360	ANDICAM	$I_C$	19.266	0.059
0.47737	360	ANDICAM	$I_C$	19.310	0.050
0.48238	360	ANDICAM	$I_C$	19.338	0.052
0.48747	360	ANDICAM	$I_C$	19.361	0.060
0.49254	360	ANDICAM	$I_C$	19.359	0.048
1.38701	600	DFOSC	$I_C$	20.152	0.039
1.39420	600	DFOSC	$I_C$	20.294	0.042
1.40178	600	DFOSC	$I_C$	20.223	0.049
2.37793	900	MOSCA	$I_C$	21.156	0.041
0.36436	360	ANDICAM	$J$	18.090	0.080
0.48350	1177	PAIRITEL	$J$	18.638	0.054
0.48494	360	ANDICAM	$J$	18.580	0.090
0.51112	2260	PAIRITEL	$J$	19.196	0.114
0.56496	683	PAIRITEL	$J$	19.472	0.141
0.62604	2190	PAIRITEL	$J$	18.872	0.054
1.53323	706	PAIRITEL	$J$	19.575	0.138
1.61412	2237	PAIRITEL	$J$	19.725	0.157
2.36277	360	ANDICAM	$J$	>19.1	...
2.52278	2237	PAIRITEL	$J$	>20.1	...
3.46425	2213	PAIRITEL	$J$	>20.2	...
4.50410	2266	PAIRITEL	$J$	>20.3	...
5.52048	2260	PAIRITEL	$J$	>20.3	...
0.48350	1177	PAIRITEL	$H$	17.693	0.141
0.51112	2260	PAIRITEL	$H$	17.784	0.127
0.56496	683	PAIRITEL	$H$	17.890	0.136
5.52048	2260	PAIRITEL	$H$	>19.7	...
0.62604	2190	PAIRITEL	$H$	17.978	0.136
1.53323	706	PAIRITEL	$H$	>19.2	...
1.61412	2237	PAIRITEL	$H$	19.090	0.278
2.52278	2237	PAIRITEL	$H$	>19.6	...
3.50562	2237	PAIRITEL	$H$	>19.8	...

**Table 1.** continued.

$\Delta t$ [d]	Exp [s]	Instrument	Filter	Mag	Error
4.54525	2260	PAIRITEL	<i>H</i>	>19.8	...
0.48350	1177	PAIRITEL	<i>K<sub>S</sub></i>	16.913	0.174
0.51112	2260	PAIRITEL	<i>K<sub>S</sub></i>	17.162	0.217
0.56496	683	PAIRITEL	<i>K<sub>S</sub></i>	17.197	0.185
0.62604	2190	PAIRITEL	<i>K<sub>S</sub></i>	17.326	0.191
1.53323	706	PAIRITEL	<i>K<sub>S</sub></i>	>18	...
1.61412	2237	PAIRITEL	<i>K<sub>S</sub></i>	18.148	0.265
2.52278	2237	PAIRITEL	<i>K<sub>S</sub></i>	>18.8	...
3.46425	2213	PAIRITEL	<i>K<sub>S</sub></i>	>18.8	...
4.54525	2260	PAIRITEL	<i>K<sub>S</sub></i>	>18.9	...
5.52048	2260	PAIRITEL	<i>K<sub>S</sub></i>	>19.0	...
1169.785	7864	HST ACS	<i>F775W<sub>AB</sub></i>	>28.5	...

**Notes.** Data have not been corrected for Galactic extinction. All data except the Keck/LRIS *g'* and the HST *F775W* data points (in *AB* magnitudes) are in Vega magnitudes. UVOT is the 30 cm UltraViolet and Optical Telescope onboard the *Swift* satellite.

Maidanak is the 1.5 m telescope of the Maidanak observatory in Uzbekistan.

RTT150 is the 1.5 m Russian-Turkish Telescope at TÜBİTAK National Observatory on Mount Bakyrlytepe, Antalya, Turkey.

MOSCA is the MOSaic CAmera on the 2.5 m Nordic Optical Telescope, La Palma, Canary Islands, Spain.

Keck/LRIS is the Low Resolution Imaging Spectrograph on the 10 m Keck I telescope, Mauna Kea, Hawaii, USA.

DOLORES is the Device Optimized for the LOw REsOLUTION detector on the 3.6 m TNG (Telescopio Nazionale Galileo) telescope on La Palma, Canary Islands, Spain.

MIRO is the 1.2 m telescope of the Mt. Abu Infrared Observatory, India.

BFOSC is the Bologna Faint Object Spectrograph & Camera at the G. D. Cassini 152 cm telescope of the Bologna University, Loiano, Italy.

DFOSC is the Danish Faint Object Spectrograph and Camera on the Danish 1.54 m telescope on La Silla, Chile.

Watcher is the 0.3 m Watcher robotic telescope at Boyden Observatory, South Africa.

Shajn is the 2.6 m Shajn telescope of the Crimean Astrophysical Observatory (CrAO), Ukraine.

TLS 1.34 m is the 1.34 m Schmidt telescope of the Thüringer Landessternwarte Tautenburg, Germany.

ALFOSC is the Andalucia Faint Object Spectrograph and Camera on the 2.5 m Nordic Optical Telescope, La Palma, Canary Islands, Spain.

FORS2 is the FOcal Reducer and low dispersion Spectrograph 2 on the 8.2 m Very Large Telescope, Paranal Observatory, Chile.

ROTSE is the 0.3 m Robotic Optical Transient Search Experiment III-B telescope at the HESS site, Mt. Gamsberg, Namibia.

ANDICAM is the A Novel Double-Imaging CAMera detector on the 1.3 m Small and Moderate Aperture Research Telescope System (SMARTS) telescope at the Cerro Tololo Interamerican Observatory in Chile.

PAIRITEL is the 1.3 m Peters Automatic InfraRed Imaging TELEscope on Mt. Hopkins, Arizona, USA.

HST ACS is the Hubble Space Telescope equipped with the Advanced Camera for Surveys.