



LJMU Research Online

Vera-Casanova, A, Gómez, FA, Monachesi, A, Gargiulo, I, Pallero, D, Grand, RJJ, Marinacci, F, Pakmor, R, Simpson, CM, Frenk, CS and Morales, G

Linking the brightest stellar streams with the accretion history of Milky Way like galaxies

<http://researchonline.ljmu.ac.uk/id/eprint/19273/>

Article

Citation (please note it is advisable to refer to the publisher's version if you intend to cite from this work)

Vera-Casanova, A, Gómez, FA, Monachesi, A, Gargiulo, I, Pallero, D, Grand, RJJ, Marinacci, F, Pakmor, R, Simpson, CM, Frenk, CS and Morales, G (2022) Linking the brightest stellar streams with the accretion history of Milky Way like galaxies. Monthly Notices of the Royal Astronomical Society.

LJMU has developed [LJMU Research Online](#) for users to access the research output of the University more effectively. Copyright © and Moral Rights for the papers on this site are retained by the individual authors and/or other copyright owners. Users may download and/or print one copy of any article(s) in LJMU Research Online to facilitate their private study or for non-commercial research. You may not engage in further distribution of the material or use it for any profit-making activities or any commercial gain.

The version presented here may differ from the published version or from the version of the record. Please see the repository URL above for details on accessing the published version and note that access may require a subscription.

For more information please contact researchonline@ljmu.ac.uk

<http://researchonline.ljmu.ac.uk/>

Linking the brightest stellar streams with the accretion history of Milky Way like galaxies

Alex Vera-Casanova¹★, Facundo A. Gómez^{1,2}, Antonela Monachesi^{1,2}, Ignacio Gargiulo^{3,4},
Diego Pallero¹, Robert J. J. Grand⁵, Federico Marinacci⁶, Rüdiger Pakmor⁵, Christine
M. Simpson^{7,8}, Carlos S. Frenk⁹ and Gustavo Morales¹

¹Departamento de Astronomía, Universidad de La Serena, Avenida Juan Cisternas 1200, La Serena, Chile

²Instituto de Investigación Multidisciplinar en Ciencia y Tecnología, Universidad de La Serena, Raúl Bitrán 1305, La Serena, Chile

³Instituto de Astrofísica de La Plata (CCT La Plata, CONICET, UNLP), Paseo del Bosque s/n, B1900FWA, La Plata, Argentina

⁴Consejo Nacional de Investigaciones Científicas y Técnicas, Rivadavia 1917, C1033AAJ Buenos Aires, Argentina

⁵Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Str 1, D-85748 Garching, Germany

⁶Department of Physics & Astronomy ‘Augusto Righi’, University of Bologna, via Gobetti 93/2, I-40129 Bologna, Italy

⁷Department of Astronomy & Astrophysics, The University of Chicago, Chicago, IL 60637, USA

⁸Enrico Fermi Institute, The University of Chicago, Chicago, IL 60637, USA

⁹Institute for Computational Cosmology, Department of Physics, University of Durham, South Road, Durham DH1 3LE, UK

Accepted 2022 June 8. Received 2022 May 27; in original form 2021 March 17

ABSTRACT

According to the current galaxy-formation paradigm, mergers and interactions play an important role in shaping present-day galaxies. The remnants of this merger activity can be used to constrain galaxy-formation models. In this work, we use a sample of 30 hydrodynamical simulations of Milky Way mass haloes, from the AURIGA project, to generate surface brightness maps and search for the brightest stream in each halo as a function of varying limiting magnitude. We find that none of the models shows signatures of stellar streams at $\mu_r^{\text{lim}} \leq 25$ mag arcsec⁻². The stream detection increases significantly between 28 and 29 mag arcsec⁻². Nevertheless, even at 31 mag arcsec⁻², 13 per cent of our models show no detectable streams. We study the properties of the brightest streams progenitors (BSPs). We find that BSPs are accreted within a broad range of infall times, from 1.6 to 10 Gyr ago, with only 25 per cent accreted within the last 5 Gyrs; thus, most BSPs correspond to relatively early accretion events. We also find that 37 per cent of the BSPs survive to the present day. The median infall times for surviving and disrupted BSPs are 5.6 and 6.7 Gyr, respectively. We find a clear relation between infall time and infall mass of the BSPs, such that more massive progenitors tend to be accreted at later times. However, we find that the BSPs are not, in most cases, the dominant contributor to the accreted stellar halo of each galaxy.

Key words: methods: numerical – Galaxy: halo – galaxies: dwarf – galaxies: structure.

1 INTRODUCTION

A well-tested prediction from the current paradigm of galaxy formation establishes that galaxies grow in mass by the accretion of material from the surrounding environment (e.g. White & Frenk 1991). In addition to mass growth, the accretion of satellites plays a fundamental role in shaping the properties of the galaxies we observed at the present day. The interaction and merger with massive objects can induce a wide variety of perturbations in the central galactic regions. These can range from the destruction of pre-existing discs in the most extreme case to the excitation of non-axisymmetric perturbations such as bars, spirals, warps, and lopsidedness (e.g. Jog & Combes 2009; Quillen et al. 2009; Gómez et al. 2016; Grand et al. 2016; Gómez et al. 2017a, 2020). Smaller satellites, i.e. those with host-to-satellite mass ratios $\lesssim 1:10$, are less likely to imprint lasting and global perturbations in the inner galactic

regions. However, they significantly contribute to the formation of the outer spheroidal and extended galactic component of galaxies, known as the stellar halo (e.g. Searle & Zinn 1978; White & Rees 1978; Bullock & Johnston 2005). Low-mass satellites are not as strongly affected by dynamical friction as their more massive counterparts. As a result, these objects can spend long periods in the outer galactic region as they are tidally disrupted, leaving behind extended low surface brightness substructures known as tidal streams (e.g. Johnston, Hernquist & Bolte 1996; Majewski et al. 1999; McConnachie et al. 2009; Martínez-Delgado et al. 2010).

Substructure left in haloes by satellites of any mass are considered fossil signatures of accretion events since they can provide detailed information about the merging history of the host galaxy. As such, streams are being actively searched for not only in the Milky Way, but also in nearby galaxy using different techniques. In the Milky Way, it is possible to use measurements of the full six-dimensional, phase-space of stars. This makes it possible even to identify substructures in the inner galactic region, where the mixing times are short and

* E-mail: alex.vera@userena.cl

streams are typically well mixed in configuration space. After the pioneering work of Helmi et al. (1999), several studies have been dedicated to quantifying amount of substructures in the solar vicinity. Recently, thanks to the data from the astrometric satellite *Gaia* (Gaia Collaboration et al. 2018), the number of known streams has widely grown (e.g. Ibata et al. 2020). The combination of this information with the mapping of the outer halo with photometric and spectroscopic surveys is allowing us, for the first time, to obtain a comprehensive view of the merging history of our Galaxy (e.g. Conroy et al. 2019a, b; Helmi 2020; Naidu et al. 2020).

Information about individual stars located within a significant area of the stellar halo can only be obtained for nearby galaxies (Greggio et al. 2014; Crnojević et al. 2016; Smercina et al. 2020). Thus, for the vast majority of Milky Way type galaxies, we rely on surface brightness maps obtained from integrated photometry. A number of observational surveys have capitalized on this technique to study the merging histories of several galaxies (e.g. Martínez-Delgado et al. 2010; Atkinson, Abraham & Ferguson 2013; van Dokkum, Abraham & Merritt 2014; Kado-Fong et al. 2018; Morales et al. 2018). Thanks to very deep observations, these studies have started to conduct a census of the stellar streams in the nearby universe. A common goal is to assess the frequency with which such extended stellar streams can be observed as a function of limiting surface brightness and, thus, to constrain the merging activity these galaxies have undergone. However, in spite of the long exposure times, integrated light observations of stellar haloes have typically reached surface brightness levels of $\mu \lesssim 28$ mag arcsec⁻² (although see Merritt et al. 2016; Trujillo & Fliri 2016 for deeper observations of a few individual galaxies), and thus are typically able to detect only the brightest stellar streams.

On the theoretical side, several work have used cosmological simulations to interpret both the integrated light observations of stellar haloes (Sanderson et al. 2018; Merritt et al. 2020) as well as their individual stars (D’Souza & Bell 2018; Monachesi et al. 2019; Font, McCarthy & Belokurov 2021). The analysis of simulations of stellar haloes, in concert with observations, allows us to connect the observable quantities with the accretion history of galaxies (Amorisco 2017; D’Souza & Bell 2018; Monachesi et al. 2019) as well as to understand systematic in stellar halo properties obtained from observational methods (e.g. Sanderson et al. 2018) or to constraint the stellar mass-halo mass relation for dwarf galaxies (e.g. Rey & Starkenburg 2022). In particular, we now know that stellar haloes are primarily built up from a few massive accretion events (Deason, Mao & Wechsler 2016; Bell et al. 2017; Monachesi et al. 2019; Fattahi et al. 2020), and that the information we extract from observations within 50 kpc mainly informs us about the properties of the most massive satellite accreted.

In this work, we use a suite of state-of-the-art cosmological hydrodynamical simulations of late-type galaxies from the Auriga project (Grand et al. 2017) to analyse the information that can be extracted from the brightest stellar stream in each halo in regards to their hosts’ merging activity. Our study builds up upon previous work by Johnston et al. (2008, hereafter J08) who used cosmologically motivated simulations to study how the frequency and properties of stellar halo substructure, as a function of surface brightness, are indicators of the recent merging histories of galaxies. Here, we further study this problem by focusing on the brightest stellar streams of our simulated haloes. This is of particular interest since it allows us to make a direct link between a specific and simple observable, i.e. the brightest stream of a galaxy halo, and the accretion history of the galaxy. Unlike previous studies who have used dark matter-only simulations together with a particle tagging

technique (e.g. J08 Cooper et al. 2010), here we analyse high-resolution, fully cosmological magneto-hydrodynamical simulations of the formation and evolution of late-type galaxies that naturally account for the different distributions of the satellite dark matter and stellar components.

Our paper is organized as follows. In Section 2, we introduce the Auriga simulations and discuss their suitability for this project. In Section 3, we discuss the generation of the surface brightness maps that are used to identify the brightest, low surface brightness feature in each halo. We also quantify the number of models with observable streams as a function of limiting surface brightness. In Section 4, we study the progenitor satellites from which each brightest stream originated, and characterize their distribution of infall times and infall mass by each progenitor. We present our Summary and Conclusions in Section 5.

2 AURIGA SIMULATIONS

The Auriga project consists of a set of 50 cosmological magneto-hydrodynamic simulations of galaxies like the Milky Way (Grand et al. 2017). These are zoom-in simulations of dark matter haloes chosen from the EAGLE project (Schaye et al. 2015). The haloes analysed in this work were selected to have a narrow mass range of $1 < M_{200}/10^{12}M_{\odot} < 2$, leaving us with a subset of 30 models. Each simulation was run assuming the Λ CDM cosmology, with parameters $\Omega_m = 0.307$, $\Omega_b = 0.048$, $\Omega_{\Lambda} = 0.693$, and Hubble’s constant $H_0 = 100 h$ km s⁻¹ Mpc⁻¹, $h = 0.6777$ (Planck Collaboration et al. 2014). The multimass ‘zoom-in’ resimulations were performed in a periodic cube of side $100 h^{-1}$ Mpc using the N -body magneto-hydrodynamic moving mesh code AREPO (Springel 2010; Pakmor et al. 2016). For the simulations analysed here, the dark matter and baryonic mass is $\sim 3 \times 10^5 M_{\odot}$ and $\sim 5 \times 10^4 M_{\odot}$, respectively. The gravitational softening length for stellar and dark matter particles grows with scale factor up to a maximum of 369 pc. For the gas cells, the softening length scales with the mean radius of the cell but is not allowed to drop below the stellar softening length.

AREPO includes a comprehensive galaxy formation model (Vogelsberger et al. 2013), including baryonic processes such as primordial and metal line cooling, a prescription for a uniform background UV field for reionization, a subgrid model for star formation (Springel & Hernquist 2003), a subgrid model for two-phase interstellar medium in pressure equilibrium (Springel & Hernquist 2003), magnetic fields (Pakmor & Springel 2013; Pakmor, Marinacci & Springel 2014), gas accretion on to black holes, and energetic feedback from AGN and supernovae type II (SNII) (for more details see Springel, Di Matteo & Hernquist 2005; Vogelsberger et al. 2013; Marinacci, Pakmor & Springel 2014; Grand et al. 2017). The parameters that regulate the efficiency of each physical process were chosen by comparing the results obtained in simulations of cosmologically representative regions to a wide range of observations of the galaxy population.

In our models, each stellar particle represents a single stellar population of a given mass, age, and metallicity. Mass loss and metal enrichment from type Ia supernovae (SNIa) and asymptotic giant branch (AGB) stars are modelled by calculating, at each time step, the mass moving off the main sequence for each star particle according to a delay time distribution. Using the stellar population synthesis models from Bruzual & Charlot (2003), the luminosity of each stellar particle was estimated in multiple photometric bands. As a result, our models include detailed photometric luminosity estimates in the U , B , V , g , r , i , z , K bands, all without taking into account the effects of dust extinction. Although all our models were run in haloes of similar characteristics, the resulting galaxies present a wide variety

of properties, mostly due to the stochasticity to the diversity of merger histories possible for haloes of this type (Bullock & Johnston 2005; Cooper et al. 2010; Tumlinson 2010; Grand et al. 2017; Monachesi et al. 2019).

The Auriga simulations have been extensively analysed in the past, showing that the associated galaxy formation model can generate realistic late-type galaxy models. In addition to their star formation histories, stellar masses, sizes, rotation curves, and H I content, previous studies have carefully characterized the effect that their merging activity has on their stellar components. In particular, Monachesi et al. (2016b, 2019) studied in detail the global and radial properties of the stellar haloes in the Auriga simulations, showing that they are diverse in their masses and density profiles, mean metallicity and metallicity gradients, ages, and shapes reflecting the stochasticity inherent in their accretion and merger histories. Furthermore, a comparison with observations of nearby late-type galaxies (mainly from the GHOSTS survey, Monachesi et al. 2016a; Harmsen et al. 2017) shows very good agreement between the most observed and simulated halo properties. Simpson et al. (2018) studied the present-day satellite luminosity functions of the Auriga haloes and found that they are in excellent agreement with those observed in the MW and M31. Furthermore, they also showed that the cumulative satellite mass distribution is converged for stellar masses $\gtrsim 10^6 M_\odot$ at the resolution level used in this work. This convergence is reached at masses well below the satellite masses considered in this work. With respect to the internal properties of the Auriga satellites, Grand et al. (2021) (see also Bose et al. 2019) showed that, within the mass/luminosity range considered in this work, the satellite phase-space scaling relations obtained from the populations of simulated satellites follow reasonably well the observed scaling relations (McConnachie 2012). Moreover, they also showed that the scaling relations are well converged over 3.5 orders of magnitude in mass resolution.

Thus, the Auriga simulation suite represents a suitable simulation set to study the properties of the brightest stellar streams and their progenitors.

3 THE BRIGHTEST STELLAR STREAMS

In this section, we describe the procedure followed to identify and quantify the present-day brightest stellar stream in each Auriga model, associated with either previous or ongoing accretion events. The main steps consist of the generation of different surface brightness (SB) maps, each at a different SB limiting value, for each Auriga model. The maps are then visually inspected, one by one, to determine at what limiting SB the brightest stellar stream in each halo can be identified. Once such stream is identified, the properties of its parent satellite galaxy are extracted and analysed. We highlight that in this work, we only consider as streams low surface brightness features that are of an accreted origin. Thus, low surface brightness substructures of disc origin, such as galactic feathers, are not considered. Finally, we focus our analysis on satellites with satellite-to-host mass ratio $\leq 1/4$ that crossed the viral radius earlier than 1 Gyr ago, thus discarding very recent accretion events.

3.1 Surface brightness maps

To generate the SB maps, each simulation is first projected on two planes; i.e. it is rotated such that i) the angular momentum of the disc is aligned with the z -axis of the reference frame (edge-on) and ii) the angular momentum of the disc is aligned perpendicularly to the z -axis of the reference frame (face-on). For each projection,

the rotation is done iteratively using only young stellar particles (age < 5 Gyr) located within a cylindrical volume of 10 kpc radius and shrinking height, $h = (10, 5, 2.5)$ kpc. We choose these two particular disc orientations because they represent the two most extreme configurations to identify low surface brightness features. Whereas on a face-on configuration, stellar streams can be hidden by the presence of the much brighter disc out to larger radii (e.g. up to 50 kpc or more), on an edge-on view, the disc contamination is minimal beyond 5–10 kpc along the minor axis. We then center a 150×150 kpc, two-dimensional grid on top of both projections using grid nodes separated by 1 kpc in each direction. Centered on each grid node, we place bins of 2×2 kpc size, and calculate the total r -band magnitude by integrating over the fluxes of all enclosed stellar particles. It is worth noting that this gridding of data, as well as its bin size, mimics the smoothing done in observations to enhance diffuse structure and preserve image resolution (see Morales et al. 2018). As mentioned in Section 2, the r -band luminosity of each galaxy was modelled using Bruzual & Charlot (2003) stellar population synthesis models (Grand et al. 2017). We focus on this synthetic band since i) it is a relatively good tracer of the overall mass distribution of satellites, ii) it is a typical band employed to observe low surface brightness (LSB) features (e.g. Atkinson et al. 2013; Martínez-Delgado 2019), and iii) it is less affected by dust extinction than other bluer photometric bands. We note that the brightest streams in the r -band are also the brightest in the other simulated photometric bands. If we instead consider, e.g. the $z(B)$ -band luminosity, the limiting SB level at which the brightest stream are detected will systematically shift by $\lesssim 1$ magnitude brighter (fainter). Other than that, our overall results remain the same.

Fig. 1 shows the deepest edge-on SB maps obtained from each Auriga model, $\mu_r^{\text{lim}} = 31 \text{ mag arcsec}^{-2}$. We can see that, in most haloes, a clear stellar stream can be found. Note, however, exceptions such as Au9 and Au10, which do not show any clear LSB feature. As discussed by Hendel & Johnston (2015) (see also Karademir et al. 2019), the shape of the substructures found in these haloes is related to the mass and orbital properties of the infalling satellites, such as the impact parameter and the inclination angle with respect to the disc plane. In general, extended stellar streams, such as loops, are typically related to satellites with large impact parameters ($\alpha \geq 10^\circ$) and angular momenta. On the other hand, shell-like structures are associated with the accretion of satellites on nearly radial orbits (Cooper et al. 2011). The latter tend to have shorter lifetimes than loops (see also J08). Note as well that debris from satellites can also be found on the plane of the disc. Such streams are the result of either the accretion of satellites on low inclination infall orbits or to the tilting of the disc due to the torque exerted by a massive infalling satellite (see e.g. Gómez et al. 2017a, b).

We then generate, for each Auriga model, several SB maps reaching different limiting magnitudes, emulating different observational depths. The limiting SB magnitude, μ_r^{lim} , ranges from $22 \text{ mag arcsec}^{-2}$ to $31 \text{ mag arcsec}^{-2}$, with a step of $1 \text{ mag arcsec}^{-2}$. As discussed by Morales et al. (2018), for SB levels deeper than $28 \text{ mag arcsec}^{-2}$, the contamination from Galactic cirrus becomes very significant. We recall that, in this work, we do not account for the effects of the Galactic cirrus. We also do not model internal dust extinction or background subtraction noise, which affects detection of LSB features in photometric images. As a result, our analysis could slightly overestimate the detectability fraction at a given μ_r^{lim} . It is worth noting, however, that we are only focusing on the brightest stream on each model, and that the relatively coarse steps in SB magnitude used to generate the maps, minimize this effect. We will explore the effect of such contamination in a future work.

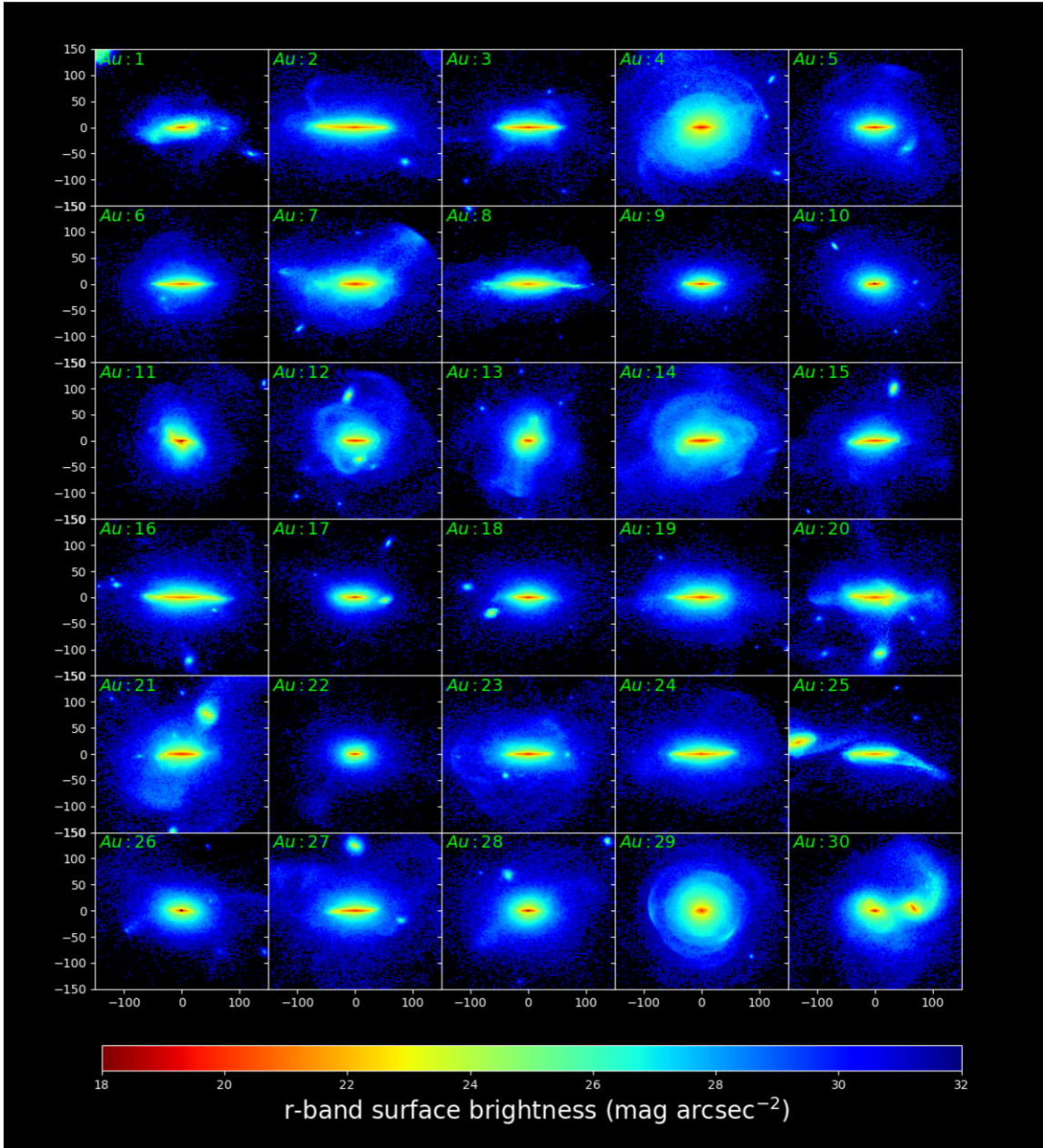


Figure 1. r -band surface brightness maps to $\mu_r^{\text{lim}} = 31 \text{ mag arcsec}^{-2}$ of the Auriga haloes at $z = 0$, seen edge-on in a square of $300 \times 300 \text{ kpc}^2$. The resolution of these maps is $1 \times 1 \text{ kpc}^2$ per pixel. In most cases, it is possible to appreciate low surface brightness features that extend well outside the host stellar disc, associated with both previous and ongoing accretion events.

In Fig. 2, we show an example of the results obtained with this procedure. The figure shows SB maps of the Au2 model, displayed on an edge-on projection. The different panels show the resulting SB maps obtained for different values of μ_r^{lim} . As expected, shallow μ_r maps ($\mu_r^{\text{lim}} \lesssim 24 \text{ mag arcsec}^{-2}$) only reveal the presence of the bright stellar disc and, if present, of the brightest satellite galaxies. In general, at $\mu_r^{\text{lim}} \sim 25 \text{ mag arcsec}^{-2}$, we reach the outer edges of all stellar discs. Deeper maps start to reveal a more extended, relatively flat stellar distribution associated with the inner stellar halo, mainly dominated by an *in-situ* component (Monachesi et al. 2019), as well as the faint and extended stellar halo. These deeper maps allow us to detect, in many cases, low surface brightness substructures mostly associated with stellar streams from ongoing or previous accretion

events. In the following section, we describe the procedure applied to identify the brightest stellar stream in each simulated halo using these SB maps.

3.2 Identification of the brightest stellar streams

Our main goal in this work is to characterize what information can be extracted from the brightest stellar streams, of accretion origin, with respect to the recent accretion history of a galaxy. Thus, the first step is to identify such streams in the stellar haloes of each Auriga model and their different projections. The automatic detection of stellar streams is a challenging task. These LSB features can show a wide variety of morphologies and are typically very extended, sampling

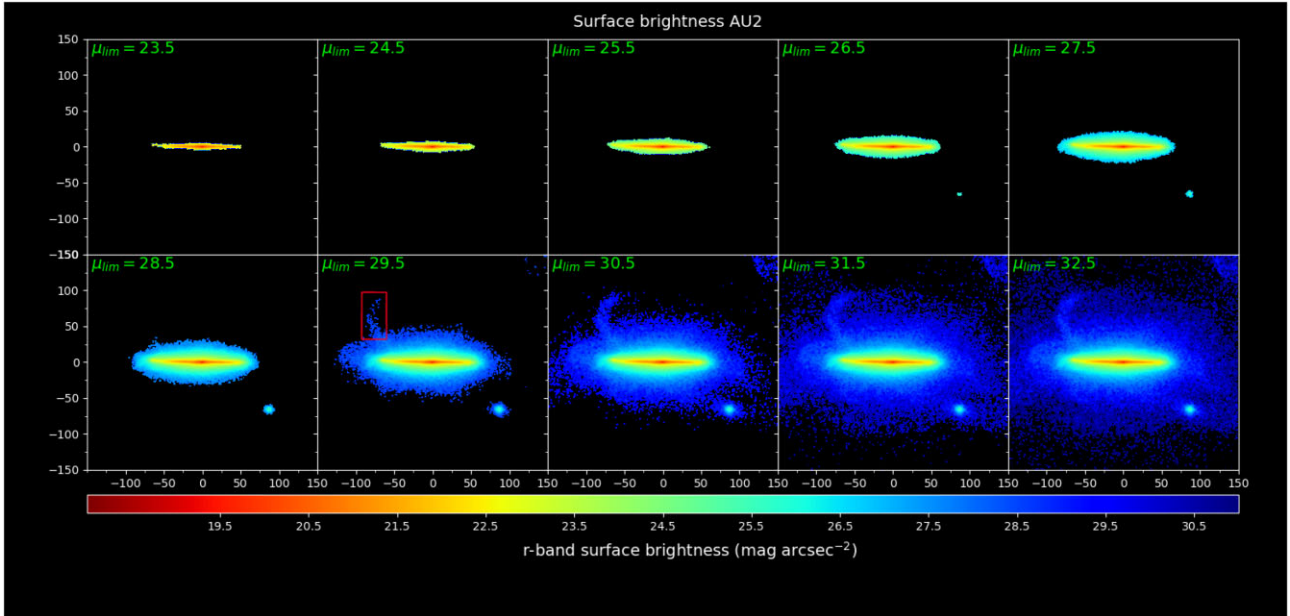


Figure 2. Surface brightness maps as a function of μ_r^{lim} for Au2. The galaxy is shown on its edge-on projection. Note that the brightest stellar stream in this galaxy can be clearly seen for the first time at $\mu_r^{\text{lim}} = 29.5 \text{ mag arcsec}^{-2}$. The corresponding stream is highlighted with a red box. The size of each projection is $300 \times 300 \text{ kpc}^2$.

galactic regions with very different SB levels. Methods to achieve this have been previously proposed, especially when dealing with large number of observations, where visual identification is neither scalable nor feasible (see e.g. Kado-Fong et al. 2018). In our work, we have 30 galactic models and, thus, visual inspection of SB maps to identify the brightest stream on each halo that can be reliably applied. We also note that visual inspection is the preferred method to identify LSB features in most observational works and, in particular, those we discuss and compare against in Section 3, such as Atkinson et al. (2013) and Morales et al. (2018).

To account for human bias on the visual inspection of SB maps, we proceed as follows. First, among the five co-authors, we distributed several SB maps with different μ_r^{lim} for each galactic model. Examples of these maps are shown in Figs 2 and 3, where each panel shows the result of reaching a progressively deeper μ_r^{lim} . These maps include all stellar particles within the given area, independently of whether they were born *in situ*, i.e. within the potential well of the main host, or in satellite galaxies, i.e. accreted particles. On each model, we start searching for LSB features on the shallowest SB map. Each person was requested to identify LSB features independently of their projected morphology. As a result, the brightest LSB features in our analysis could be associated with tidal streams, shells or plume-like structures. Note that LSB features that are associated with main galactic hosts, such as galactic feather or bound satellite galaxies, are not accounted for since we focus on features of an accreted origin. If no LSB features are identified at a given μ_r^{lim} , we proceed to the next deeper map. This is iteratively performed until the brightest stream can be first identified. Once the stream is observed, the corresponding value of μ_r^{lim} is stored on a list. In the example shown in Fig. 2, a clear tidal stream, highlighted with a red box, is first detected at $\mu_r^{\text{lim}} = 29.5 \text{ mag arcsec}^{-2}$. Note however that in some models, such as Au9, shown in Fig. 3, it was not possible to identify a well-defined LSB feature at any μ_r^{lim} .

Secondly, these five co-authors shared their independent identification list to reach a consensus. A unique and final identification list is

obtained from this procedure. Those cases where discrepancies in the identifications were found were individually discussed. A common consensus was reached after different checks, such as whether the identified features were not a low surface brightness bound satellite. Once a final list is obtained, the origin of these LSB features is asserted based on the birth location of the stellar particles that composed them. In those cases where the identified substructure is associated with debris from the main galactic host, we discarded it and proceeded to identify the following brightest LSB feature following the procedure just described.

The result of this identification process is shown in Fig. 4. Each panel shows the shallowest SB map in which the brightest stream has been clearly identified. The corresponding limiting SB is listed on the legend; the cases where no substructure is present have the SB marked with an * symbol. The brightest stream in each halo is highlighted with a red box. In some cases, such as Au12, Au20, Au21, and Au25, the brightest stream can be directly linked to the brightest satellite in the field. This shows that those satellites have been orbiting their host for a few Gyrs and are currently undergoing disruption. Such streams can be observed in relatively shallow maps, reaching $\mu_r^{\text{lim}} \lesssim 27 \text{ mag arcsec}^{-2}$. In many other cases, such as Au5, Au11, Au16, and Au22, it is necessary to reach much deeper SB levels to identify the brightest stream. These are clear examples of how low surface brightness substructures can go undetected without very deep observations. Note that, in those cases, the progenitor satellites can no longer be identified. We also find haloes, such as Au18, where the brightest satellite does not show any tidal feature, indicating that they have been very recently accreted on to their host.

3.3 Brightest stellar stream quantification

In Fig. 5, we summarize the results of the brightest stream identification process. We first focus on the results obtained from the edge-on projection. The black line shows the normalized cumulative function of a number of models with identifiable stellar streams as a

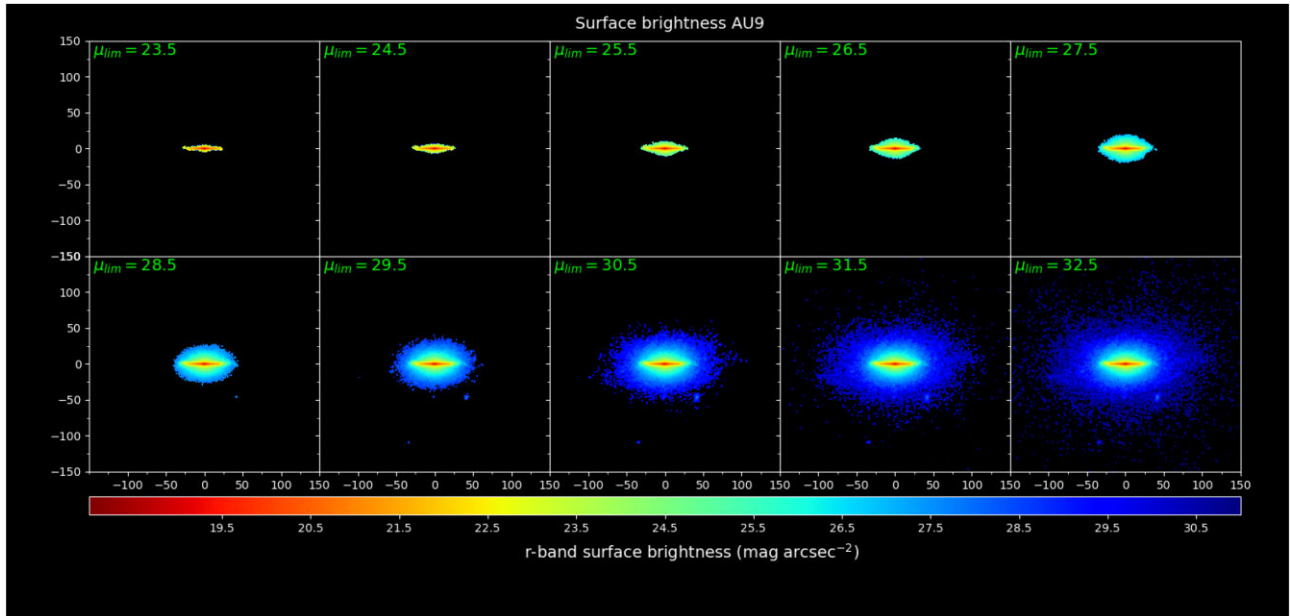


Figure 3. Surface brightness maps as a function of μ_r^{lim} , for Au9. The galaxy is shown on its edge-on projection. Note that no clear stellar stream is found at any surface brightness. The size of each projection is 300×300 kpc².

function of the limiting SB level, μ_r^{lim} , for the edge-on projections. The grey shaded area indicates the full range obtained from the five independent stellar streams identifications. To define the lower and upper bound of this area, we assigned to each model, the maximum and minimum μ_r^{lim} value reported by the five identifiers. None of the Auriga models shows stellar streams in SB maps with $\mu_r^{\text{lim}} \leq 25.5$ mag arcsec⁻². However, we find that at $\mu_r^{\text{lim}} \approx 28.5$ mag arcsec⁻², where contamination from Galactic cirrus starts to become significant, ~ 30 per cent of the models already show a detection. The cumulative function shows a steep increase in the fraction of galaxies with stream detection at this μ_r^{lim} value, and then flattens again beyond 30.5 mag arcsec⁻². Another interesting result is that, even at a very low SB limit of 31 mag arcsec⁻², 13 per cent of our models show no detectable stream. The cyan dashed line in Fig. 5 shows the deepest SB achieved by the SDSS through its IAC Stripe 82 legacy survey (Fliri & Trujillo 2016), which is at $\mu_r^{\text{lim}} \approx 28.5$ mag arcsec⁻² in the r -band. Note that, as mentioned above, at this SB limiting magnitude, we are able to detect the brightest streams on ≈ 30 per cent of late-type galaxies. For a more complete census of these LSB features, it is thus critical to obtain deeper surveys such as what will be provided by the LSST from the Vera Rubin observatory. The final 10 yr LSST catalogue is expected to provide images over a wide area (18 000 square degrees) reaching SB magnitudes of 30.5 mag arcsec⁻² in the r -band (Laine et al. 2018). This is highlighted in Fig. 5 with the purple dashed line.

We note that two models show as brightest LSB features material directly linked to the host stellar discs. These two cases are Au11 and Au25 and both have suffered a recent and strong tidal interaction. These substructures, best known as galactic feathers, are not very common in our sample, showing in only 5 per cent of the objects. Since feathers are associated with the host pre-existing disc, they are typically very bright. In the two cases found here, they can be clearly already seen at SB brighter than ~ 27 mag/arcsec². Substructure from the perturbing satellites is typically also present, but usually is observed at fainter magnitudes. Feathers are not included in the cumulative function shown in Fig. 5. Instead, we focus on the

brightest LSB features associated to satellites, which arise at a lower SB level.

It is interesting to compare these results with those obtained from observational samples. For example, Atkinson et al. (2013), using observations from the wide-field component of the Canada–France–Hawaii telescope legacy survey, generated a sample of 1781 luminous ($M_{r'} < -19.3$ mag) galaxies in the magnitude range $15.5 \text{ mag} < r' < 17 \text{ mag}$, and in the redshift range $0.04 < z < 0.2$. The sample reaches a limiting surface brightness in the r' -band of ~ 27 mag arcsec⁻² and is visually inspected to detect LSB features. Their analysis show that 12 per cent of the galaxies in their sample present clear tidal features at the highest confidence level, but the fraction rises to about 26 per cent if systems with marginal detection are included. Note however, that the sample studied by Atkinson et al. (2013) includes galaxies lying in both the red sequence and the blue cloud. Furthermore, they find that the fraction of galaxies with detected streams is a strong function of the rest-frame colour and stellar mass, and that red galaxies are twice as likely to show tidal features than blue galaxies. In our work, we are biased towards very bright late-type galactic models, with $-23 < M_{r'} < -20$ mag, with median $M_{r'} \approx -22$ mag. Thus, our models fall within the blue cloud (see fig. 20 in Grand et al. 2017). The blue cloud subsample by Atkinson et al. (2013) shows tidal detections in 12–18 per cent of their galaxies, depending on the confidence level detection considered, which is shown in Fig. 5 as a blue shaded area. This percentage is in good agreement with our results. Around their value of $\mu_r^{\text{lim}} = 27$ mag arcsec⁻², we find that ~ 15 per cent of our models show low surface brightness features.

Similar results were obtained by Morales et al. (2018) using a sample of a post-processed Sloan digital sky survey (SDSS) images, optimized for the detection of stellar structures with low surface brightness around a volume-limited sample of nearby galaxies. Their final sample consists of images of 297 galaxies with stellar masses similar to that of the Milky Way, which are visually inspected by the authors to detect LSB features. The images sampled reach a Gaussian distributed μ_r^{lim} of mean ≈ 28.1 mag arcsec⁻² and $\sigma \approx$

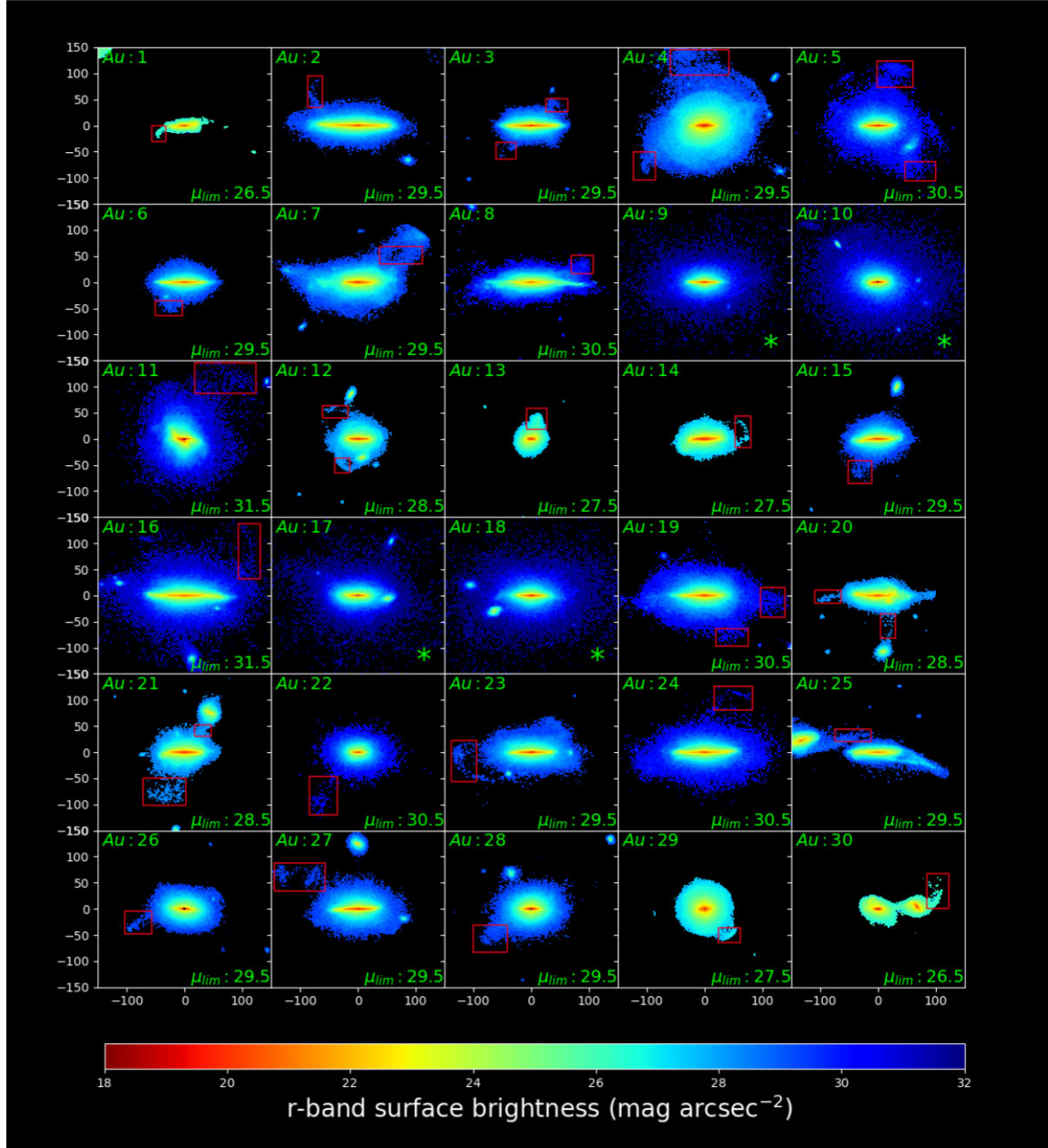


Figure 4. As in Fig. 1 but now selecting, for each halo, the SB map at the μ_r^{lim} value where the brightest stellar stream is first identified. The red boxes highlight the brightest stellar streams in each halo. Some show two boxes. This indicates that two different streams were identified at the same surface brightness. Note that some haloes (Au9, Au10, Au17, Au18) do not show identifiable stellar streams. Those haloes are highlighted with an * symbol.

$0.26 \text{ mag arcsec}^{-2}$. Within those limiting SB, they find a detection of stellar substructure in 14 per cent of the observed galaxies, whereas we find stellar streams in almost 25 per cent of the Auriga models (see Fig. 5). An important difference between the Morales et al. (2018) sample of observed galaxies and our sample of models, which most likely accounts for this mismatch, is the stellar mass distribution of galaxies. The observed sample analysed by Morales et al. (2018) has a mean stellar mass of $\log(M_*/M_\odot) = 10.37$, whereas the mean stellar mass of the Auriga models is $\log(M_*/M_\odot) = 10.82$. We note also that all our models are more massive than the median mass of the observed sample. This is rather significant, especially considering that the observed stellar substructure detections increase significantly for larger stellar masses of the host galaxy, with a detection rate of

about ~ 30 per cent for stellar masses larger than $\log(M_*/M_\odot) = 10.82$ (See fig. 7 of Morales et al. 2018).

There are several other differences between the analysis presented in this work and those based on observational samples of galaxies, which should be kept in mind while performing this comparison. It is worth recalling that, in our models, i) we are not accounting for dust extinction and background noise, which may erase the signature of faint stellar substructures, ii) we have a much smaller sample of galaxies, and iii) we count as a detection the very first time we see signs of a stream. In reality, it is likely that slightly deeper observations would be required in some of these cases to detect streams due to the above-mentioned observational effects. It is also worth recalling that in both observational studies, Atkinson et al.

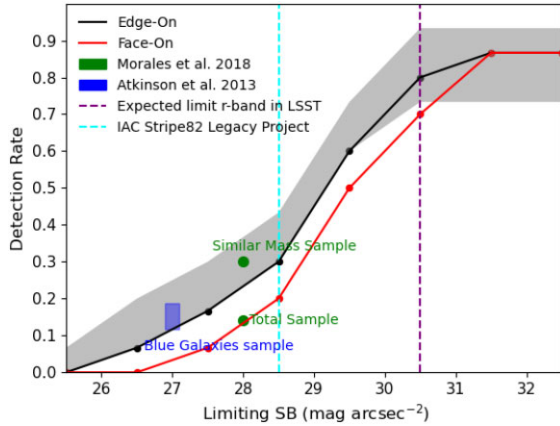


Figure 5. Cumulative fraction of the number of haloes with detected stellar streams as a function of μ_r^{lim} . The black and red lines show the results obtained when galaxies are oriented edge-on and face-on, respectively. The grey shaded area shows the range obtained from the five independent stellar streams visual identifications. The blue shaded area shows the fraction of blue galaxies sample with stellar streams, detected within the second and third highest confidence level by Atkinson et al. (2013), and green circles show the fraction of galaxies with identifiable stellar streams reported by Morales et al. (2018). The dashed lines indicate the limits in surface brightness of different surveys, in cyan the IAC Stripe82 legacy project, in purple the limit expected for the 10 years of co-added Vera Rubin observatory’s LSST images.

(2013) and Morales et al. (2018), stellar streams were also identified by visual inspection of the corresponding images, as also done in this work.

Our results discussed so far are based on edge-on projections of the galactic models. As previously mentioned, this is bound to enhance the detection of streams when compared with a sample of galaxies distributed at random inclinations. To explore this effect, we show in Fig. 5, with a red line, the cumulative fraction of models with detected streams as a function of μ_r^{lim} when projected face-on. Indeed, we find that this fraction decreases by ≈ 10 per cent within the relevant SB magnitude range.

To summarize, it is expected that the fraction of galaxies with detected streams in this work should be somewhat larger than that found in observations, which are likely to represent a lower detection limit. Nevertheless, we find a reasonably good agreement between our detection rates and those from the observed samples. In a follow-up work, we will include observational effects and biases to our models in order to perform a more quantitative and fair comparison with observational results.

4 PROPERTIES OF SATELLITE PROGENITORS

In the previous section, we have identified and quantified the brightest stellar streams in each Auriga halo by inspecting SB maps at different μ_r^{lim} and projections. The goal of this section is to identify and characterize the main properties of the satellite progenitors of these detected low surface brightness features, in particular their infall times and masses. The progenitor satellite of a given stream is identified by searching among the satellites that contributed the largest number of particles to a small area surrounding the brightest stream. These areas for each Auriga model are highlighted with a red square in Fig. 4. This procedure is summarized in Fig. 6. The top panels shows, as an example, the SB maps of five Auriga models reaching $\mu_r^{\text{lim}} = 31 \text{ mag arcsec}^{-2}$. The middle panels show stellar particle scatter plots of the same haloes. With red dots we highlight

the stellar particles that belong to the brightest stream progenitor detected in each halo. Some models, such as Au21 (rightmost panel), show more than one stream at the corresponding μ_r^{lim} where substructures were first revealed (see also Fig. 4). In those cases where the two streams are associated to two different progenitors, both satellites are highlighted with red and blue dots. Note that, while some satellite progenitors can still be identified, i.e. are surviving satellites (e.g. Au2 and Au21-blue dots), in other cases, they have been fully disrupted (e.g. Au11 and Au21-red dots). We will further explore this in what follows.

4.1 Progenitor infall times

Once a progenitor satellite is identified, we proceed to trace it back in time. This is done by following the corresponding merger trees. Note that in Auriga every satellite is assigned a unique identification number. This ID is assigned to every stellar particle born within the potential well of the corresponding satellite. As a result, it is possible to track a satellite’s stellar particles even after its full disruption. The bottom panels of Fig. 6 show with red and blue lines, the evolution of the satellite galactocentric distance, R_{sat} , as a function of lookback time. For comparison, we also show with a black line the time evolution of the host virial radius, R_{vir} .

To estimate the brightest stream progenitor (BSP) lookback infall time, t_{infall} , which is the time at which the satellite first crosses the host virial radius, we search for the snapshot where the first minimum of $|R_{\text{sat}} - R_{\text{vir}}|$ is reached. The t_{infall} associated with those snapshots are listed in Table 1. In Fig. 7 we show the overall t_{infall} distribution. Interestingly, we find that the satellites that give rise to the brightest stream in each simulation are accreted in a very wide range of times, with t_{infall} values as high and low as 10 and 1.6 Gyr, respectively and a median value of $t_{\text{infall}} = 6.39$ Gyr. It is worth highlighting that 50 per cent of the BSPs were accreted within the range $5 \text{ Gyr} \leq t_{\text{infall}} \leq 7.5 \text{ Gyr}$, as shown by the striped box area. As a result, only 25 per cent of the BSPs correspond to a very recent accretion event, with $t_{\text{infall}} < 5 \text{ Gyr}$ (e.g. Au11). Conversely, the BSPs in 25 per cent of the cases are related to satellites that were accreted as early as 8 to 10 Gyr ago (e.g. Au26).

The fifth column of Table 1 indicates whether the BSPs have survived to the present-day or not. The surviving BSPs are those satellites that are still identified by the sub-find routine, at $z = 0$, as an independent sub-halo from the main host halo. Interestingly, ~ 34 per cent of the BSPs can still be identified at $z = 0$ as an independent sub-halo. The bottom panel of Fig. 6 seems to suggest that, typically, surviving BSPs have crossed R_{vir} at later times than their disrupted counterparts (see e.g. the blue and red line in the last panel of Fig. 6 for an example of a surviving and destroyed satellite, respectively). Indeed, we find that the mean t_{infall} for the surviving and disrupted BSPs are 5.1 Gyr and 6.7 Gyr, respectively. In this context, it is worth taking into account the results of Fattahi et al. (2020, hereafter F20), who analysed the Auriga simulations to examine the build-up of the MW’s stellar halo. Their analysis focused on the comparison between the properties of the surviving and destroyed dwarf galaxies that are accreted by these haloes over cosmic time. However, they did not explore the correlation between these accretion events and the low surface brightness substructure left behind. F20 showed that, on average, destroyed dwarfs have early infall times, $t_{\text{infall}} \gtrsim 7.5 \text{ Gyr}$, whereas the majority of dwarfs accreted at $t_{\text{infall}} < 4 \text{ Gyr}$ survive to the present-day, in agreement with what we find for the BSPs. Moreover, they also find a dependence between survivability, t_{infall} , and satellite mass at infall. For surviving satellites, the typical t_{infall} are ~ 8 and ~ 4 Gyr for satellites with

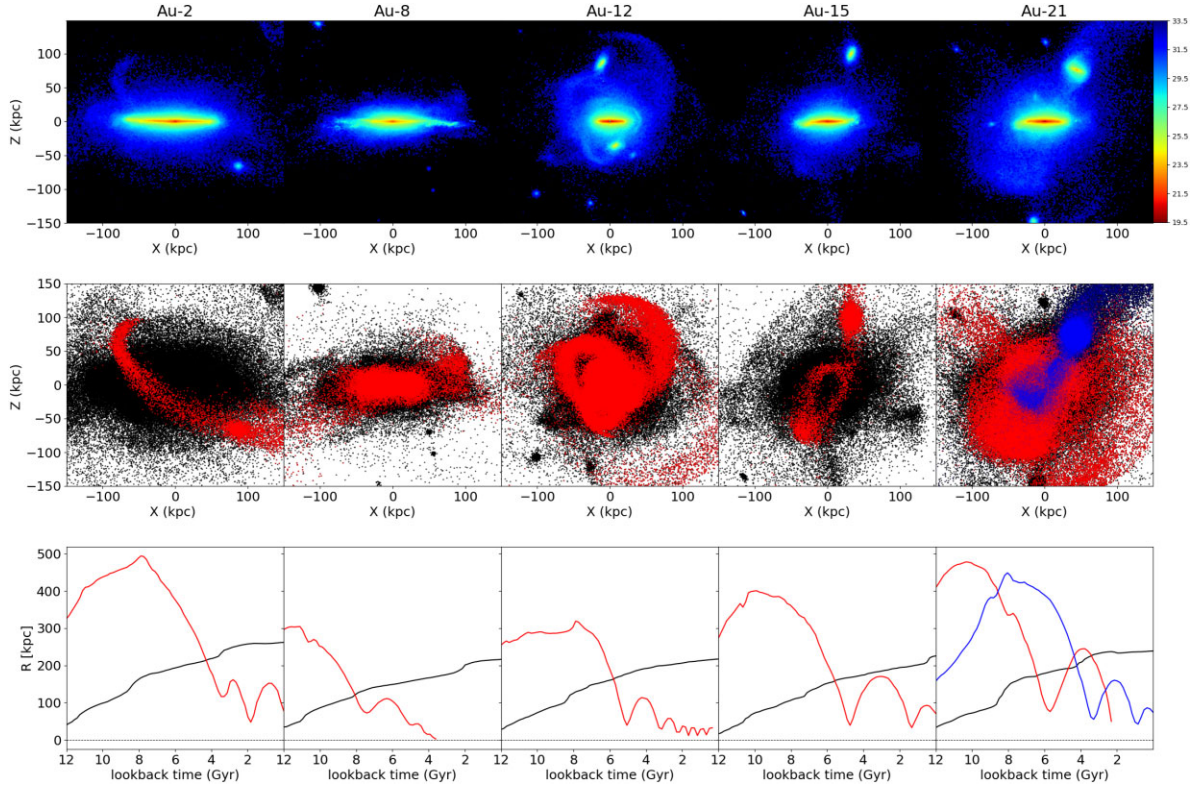


Figure 6. Top panels: surface brightness maps of five different Auriga galaxies, shown edge-on in a $300 \times 300 \text{ kpc}^2$ area. Middle panels: the black dots show the stellar particle distributions of these haloes. Red and blue dots show the stellar particles associated with the brightest stream progenitors on each halo. The different colours in Au21 correspond to particles from two different satellites that contributed with streams at the same μ_r^{lim} . Bottom panels: the black line shows the time evolution of the host virial radius. The red and blue lines show the time evolution of the galactocentric distance of the brightest stream progenitors. The dashed line indicates $R = 0 \text{ kpc}$.

infall stellar masses of 10^6 and $10^9 M_\odot$, respectively. Instead, for disrupted satellites, the infall times are ~ 11.5 and ~ 9 Gyr for infall stellar masses of 10^6 and $10^9 M_\odot$, respectively. Interestingly, one could naively expect the brightest LSB feature to be associated with the most massive accretion event a galaxy has recently experienced. However, as we will discuss later in Section 4.2, the correlation between satellite infall mass and survivability rate discussed in F20, translates into BSPs not necessarily being the most significant mass contributors to the present-day accreted stellar haloes mass distribution.

We now examine the relation between BSPs infall times and the μ_r^{lim} at which the brightest streams are first identified. For this analysis, we divide the sample of BSPs into two subsamples based on the limiting SB at which the stream was first identified: fainter or brighter than $28.5 \text{ mag arcsec}^{-2}$. The division at this limiting SB is motivated by the following. First, it is at about this SB magnitude where contamination from Galactic cirrus becomes very significant. Second, and related to the first point, current large observational surveys that have tried to quantify LSB features on nearby galaxies have not reached deeper (in general) than $\mu_r^{\text{lim}} = 28.5 \text{ mag arcsec}^{-2}$. Future surveys, such as those provided by the LSST, will reach much fainter SB levels and we want to show what information will be possible to obtain from these fainter features. The top and bottom panels of Fig. 8 show the t_{infall} distribution for the BSPs which streams were first detected at values of $\mu_r^{\text{lim}} \leq 28.5$ and $\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$, respectively. These two subsamples present different distributions. Brightest stellar streams, first detected on SB maps at $\mu_r^{\text{lim}} < 28.5 \text{ mag arcsec}^{-2}$, represent accretion events

that typically took place 5 Gyr ago, while those streams found at $\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$ are, on average, related to accretion events that took place about 7 Gyr ago. Note, however, that both distributions show significant dispersion in t_{infall} . For this calculation, we have removed Au11 whose BSP has just been accreted and, thus, presents a very faint shell-like substructure that has not yet had the time to develop.

The relation between brightest stream brightness and the progenitor infall times is not surprising. If, for simplicity, we assume that the BSP infall mass distribution is similar for both samples, streams that have more time to phase-mix should undoubtedly look fainter at $z = 0$ (see also Helmi et al. 1999; J08; Gómez & Helmi 2010). We explore this further in the next section.

4.2 Mass of progenitor satellites

In this section, we explore the BSPs mass distribution. In particular, we focus on the mass of each BSP at infall, i.e. at the time of R_{vir} crossing (see Section 4.1). We search for correlations with other properties such as their t_{infall} and first pericentric distance. As we just showed, the μ_r^{lim} at which the brightest stellar streams are detected provides information about the BSP infall time. In Fig. 9, we show the infall mass distribution for the BSPs which streams were first detected at values of $\mu_r^{\text{lim}} < 28.5$ (top panel) mag arcsec^{-2} and $\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$ (bottom panel). The figure shows that BSPs with brighter LSB features are typically more massive than their fainter counterparts. The difference is of approximately 1 dex in stellar mass.

Table 1. Properties of the brightest stream satellite progenitors. The columns are (1) the Auriga galaxy halo number; (2) the satellite lookback infall time, or time when the satellite first crossed their host virial radius; (3) the logarithm of the total mass of the satellite at infall time: $\log(M_{\text{tot}}/M_{\odot})$; (4) the logarithm of the stellar mass of the satellite at infall time: $\log(M_{\star}/M_{\odot})$; (5) flag that indicates whether satellites survive at $z = 0$ or not; (6) satellite’s ranking, $\#_{\text{prog}}$, based on the stellar mass contributed to the $z = 0$ host stellar halo mass; (7) the logarithm of the total mass of the host galaxy at $z = 0$, $\log(M_{\text{tot}}^{\text{Host}}/M_{\odot})$. Galaxies that show multiple streams, above the surface brightness threshold, associated with different satellites are highlighted with an (*a*).

Au	$t_{\text{infall}}(\text{Gyr})$	$M_{\text{tot}}^{\text{Sat}}$	M_{\star}^{Sat}	Surv.	$\#_{\text{prog}}$	$M_{\text{tot}}^{\text{Host}}$
1	4.56	10.93	9.57	NO	1	11.98
2	4.24	10.30	8.53	YES	5	12.3
3	9.42	10.34	9.13	NO	2	12.18
4 ^a	3.11	11.49	10.37	NO	1	12.17
	6.32	10.70	9.03	NO	3	–
5 ^a	8.49	10.27	8.46	NO	3	12.1
	6.78	10.29	8.45	YES	4	–
6	8.96	10.43	8.70	NO	1	12.04
7	2.79	11.04	9.80	NO	1	12.07
8	8.03	10.84	9.38	NO	1	12.04
9	–	–	–	–	–	12.04
10	–	–	–	–	–	12.04
11	0.82	11.47	10.51	YES	–	12.23
12	5.68	10.67	9.28	NO	3	12.06
13	7.25	10.82	9.33	NO	1	12.1
14	7.25	10.92	9.45	NO	6	12.24
15	5.52	10.51	9.10	YES	6	12.1
16	8.65	9.73	7.68	YES	10	12.19
17	–	–	–	–	–	12.04
18	–	–	–	–	–	12.11
19 ^a	7.10	10.61	9.14	YES	2	12.1
	5.99	10.41	8.89	NO	3	–
20 ^a	5.99	11.28	9.98	NO	1	12.11
	5.04	10.49	9.05	YES	4	–
21 ^a	6.47	10.80	9.16	NO	3	12.18
	4.24	11.00	9.65	YES	4	–
22	5.52	10.23	8.44	NO	1	11.99
23	6.63	10.38	8.70	YES	3	12.22
24	9.88	10.33	8.48	NO	4	12.19
25	1.65	11.32	10.21	YES	1	12.1
26	8.96	10.47	8.86	NO	2	12.22
27	6.94	10.69	8.93	NO	3	12.26
28	7.25	9.94	8.73	YES	4	12.23
29	5.04	11.56	10.48	NO	1	12.21
30	3.60	11.15	9.92	YES	1	12.06

As previously discussed, one would naively expect that the BSPs are among the most massive satellites accreted by each individual host. However, satellites more massive than $10^8 M_{\odot}$ in stellar mass are severely affected by dynamical friction and very rapidly disrupted, as discussed by F20. The more massive the satellite, the more efficient this process is. Such massive satellites tend to sink rapidly to the host galactic center, typically leaving behind shell-like low surface brightness substructures. As discussed by Hendel & Johnston (2015) (see also Amorisco 2015a; Pop et al. 2018; Karademir et al. 2019), this type of low surface brightness substructures tend to have shorter lifetimes than other stream types, such as loops, associated with satellites less massive than $10^8 M_{\odot}$ in stellar mass on less eccentric orbits. Thus, BSPs may not necessarily be associated with the most massive accretion event.

In Fig. 10, we show the distribution of BSPs infall mass against their corresponding t_{infall} . The left and middle panels show the total

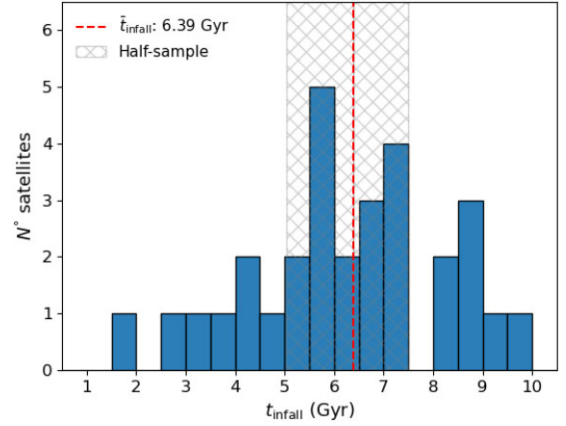


Figure 7. Lookback infall time distribution of all brightest stream progenitors. The distribution has a median value of 6.39 Gyr. The cross-hatched area, centred on the median, encloses 50 per cent of the sample and ranges from 5 to 7.5 Gyr.

and stellar mass, respectively. The symbols in both panels have been colour coded according to the μ_r^{lim} at which the brightest stream was detected. The first thing to notice is that the population of BSPs shows a wide range of satellite masses. Values span $9.5 \lesssim \log_{10}(M_{\text{tot}}/M_{\odot}) \lesssim 11.5$ in total mass and $7.5 \lesssim \log_{10}(M_{\star}/M_{\odot}) \lesssim 10.5$ in stellar mass. This represents a variation of 2 and 3 dex in total and stellar mass, respectively. In general, we find brighter streams to be associated with more massive progenitors. However, a clear relation between the mass of the progenitor and its infall time can also be seen. In addition to producing the brighter streams, more massive BSPs typically have more recent infall time. Similar results were found by J08, using dark matter-only cosmologically motivated simulations of the formation of MW stellar haloes. However, unlike J08 who studied all satellites that have contributed streams at any μ_r^{lim} , we are only focusing here on the main properties of the BSPs. Therefore, our results could be potentially compared to those from the high-mass end ($>10^{7.5} M_{\odot}$) of the satellite mass distribution analysed in J08, which contains satellites with stellar masses even smaller than $10^5 M_{\odot}$ (see their Fig. 4). Nevertheless, a direct comparison is not possible because we cannot distinguish the brightest stream of each model from all the streams shown in J08, neither the properties of its corresponding progenitor.

As before, we subdivide the sample of BSPs by the SB at which the stream was first identified $\mu_r^{\text{lim}} = 28.5 \text{ mag arcsec}^{-2}$. For $\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$, we find BSPs with median values of $(M_{\text{tot}}, M_{\star}) \approx (10^{10.4}, 10^{8.8}) M_{\odot}$. On the other hand, for $\mu_r^{\text{lim}} \leq 28.5 \text{ mag arcsec}^{-2}$, we find $(M_{\text{tot}}, M_{\star}) \approx (10^{10.9}, 10^{9.6}) M_{\odot}$ for the corresponding BSPs. The right-hand panel of Fig. 10 shows the total mass – t_{infall} relation, but now symbols are colour coded according to whether the satellites have survived to the present day (yellow dots) or not (blue dots). This panel clearly shows that surviving BSPs are both i) less massive at any given t_{infall} and ii) have been accreted later at any given M_{tot} . The size of the symbols in this panel indicates the first satellite pericentric distance. Note that late accreted satellites that, at a given mass, have been fully disrupted show very small first infall pericentric distances. Indeed, the BSPs in Au22 and Au7 have very small first pericentric distances with values of ≈ 12 and 22 kpc, respectively. As a result, they were rapidly disrupted by strong tidal forces associated with these inner galactic regions.

It is interesting to highlight that in most cases BSPs are not the dominant contributor to the overall galactic stellar haloes, even

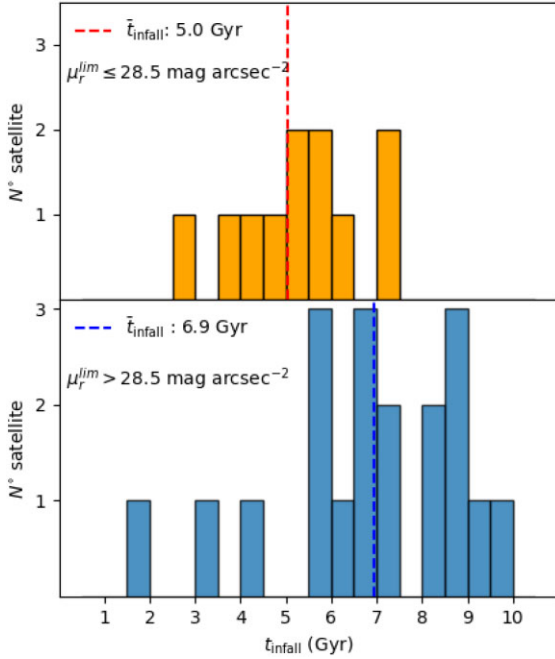


Figure 8. Top panel: distribution of lookback infall times for the BSPs of streams first identified in SB maps with $\mu_r^{\text{lim}} \leq 28.5 \text{ mag arcsec}^{-2}$. The red line shows the median $t_{\text{infall}} = 5 \text{ Gyr}$. Bottom panel: same as top panel but for the BSPs of brightest streams first identified in SB maps with $\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$. The blue line shows the median $t_{\text{infall}} = 6.9 \text{ Gyr}$.

though they are among the significant progenitors. As shown by Monachesi et al. (2019), the number of satellites that contribute to 90 per cent of the accreted halo mass (significant progenitors) range from 1 to 14, with a median of 6.5 (see also Cooper et al. 2010, F20). In Table 1, we list the rank of each BSP according to its mass contributed to the overall stellar halo. Only 35 per cent of the BSPs correspond to the most significant progenitor of an Auriga halo. Another 32 per cent of the BSPs are distributed among the 2nd and 3rd most significant contributors. The remaining 32 per cent have been ranked above the 4th significant contributor. Thus, most of these haloes have accreted more massive satellites than the BSPs over time.

This is closely related to the mass– t_{infall} relation shown in Fig. 10. As previously discussed, BSPs accreted earlier are typically less massive than those recently accreted. The difference can be as large as two orders of magnitude in total mass. This is not purely due to the fact that at later times galaxies are more likely to accrete more massive substructures. To illustrate this, in the right-hand panel of Fig. 10, we show the five most massive satellites accreted by two Auriga haloes, Au5 and Au28, with red and green triangles, respectively. It is clear that both haloes have accreted satellites more than 1 dex more massive than their corresponding BSPs, and at similar times. However, as recently discussed in Panithanpaisal et al. (2021), the infall mass of the accreted satellites plays a key role in determining the mixing time-scales of their streams (see also Johnston 1998; Helmi et al. 1999). More massive progenitors generate warmer stellar streams due to their higher internal velocity dispersions (Amorisco 2015b). Indeed, as discussed in Helmi, White & Springel (2003), streams originating in smaller haloes are narrower, more clearly defined, and, typically, they phase-mix over longer time-scales. Note however that this is not the only process at play. As shown by F20, for satellites with $M_* > 10^8 M_\odot$ (i.e. within the BSPs range), their survival time strongly depends on their t_{infall} . Due to dynamical

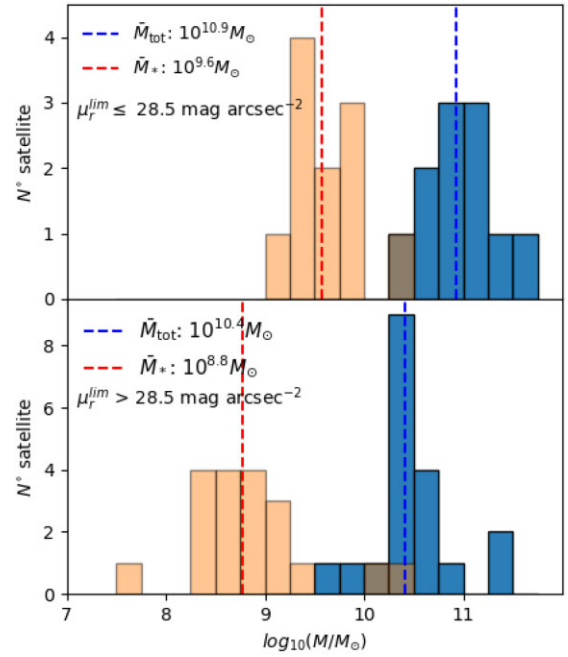


Figure 9. The histograms show the total mass and stellar mass distributions, in blue and orange, respectively, for all the brightest stream progenitors. Similar to Fig. 8, the top panel shows the sample for the BSPs of brightest streams first identified at $\mu_r^{\text{lim}} \leq 28.5 \text{ mag arcsec}^{-2}$ and the bottom panel shows the sample for the BSPs of the brightest streams first identified at $\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$. The dashed lines indicate the median infall mass for each distribution.

friction, these more massive and luminous haloes are more rapidly disrupted compared to lower mass haloes accreted at similar times. In such cases, the brightest streams can be associated with less massive (but still luminous) galaxies that can continue orbiting their host, and releasing debris, for longer periods. As t_{infall} gets closer to $z = 0$, the chances of finding coherent stellar streams arising from the more massive satellites grows. This is more clearly illustrated in Fig. 11, where we show the total mass distribution of accreted satellites (not only BSPs) as a function t_{infall} for six particular Auriga models. Based on the infall mass distribution of the SPs (Fig. 9), we only focus on accreted satellites with total infall masses $> 10^{9.5} M_\odot$. The BSPs on each model are highlighted in red. Note that Au3 and Au19 (middle column) show two BSPs. The reason is these haloes show, at the same μ_r^{lim} , two LSB features associated with different progenitors (see Section 3.2). In all these examples the host has accreted satellites more massive than the BSPs, with later t_{infall} . Indeed, we find that 52 per cent of the models had a satellite more massive than the corresponding BSPs infalling at a later time. As previously discussed, we find that obvious debris features observed around galaxies today are not typically associated with the most recent and most luminous accretion events.

5 SUMMARY AND CONCLUSIONS

In this work, we have searched for the brightest stellar streams in 30 fully cosmological magneto-hydrodynamical simulations of Milky Way mass galaxies from the Auriga project. Our main goal was to quantify the number of haloes with clear tidal streams as a function of limiting surface brightness, and to characterize the main properties of the BSPs. The properties of the satellites within this mass range are reasonably well numerically converged at the

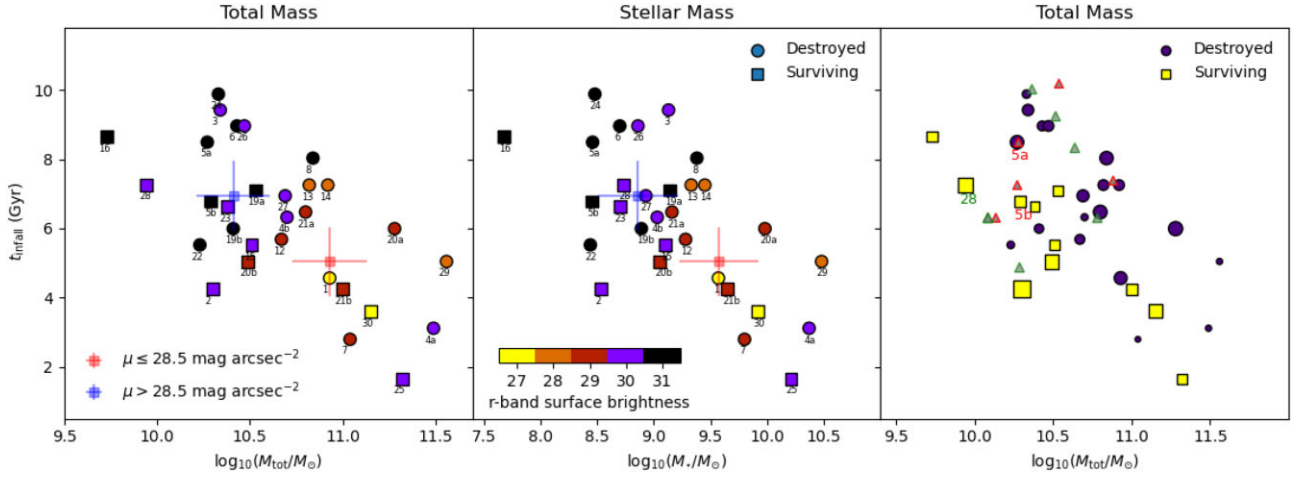


Figure 10. Left-hand panel: distribution of the brightest stream progenitor (BSP) lookback infall time as a function of their total mass at infall time. The colour coding indicates the μ_r^{lim} at which the brightest stream was first identified. Circles and squares indicate destroyed and surviving BSPs, respectively. The light red and blue crosses show the median (t_{infall} , M_{tot}) values for the population of BSPs with stellar streams first identified at $\mu_r^{\text{lim}} \leq 28.5$ and >28.5 mag arcsec $^{-2}$, respectively. Middle panel: as the left-hand panel, but for (t_{infall} , M_*) distribution. Right-hand panel: as in the left-hand panel but now symbols are sized according to their first pericentric distance (larger symbols indicate greater distances) and colour coded according to whether the BSP survived to the present day or not. The red and green triangles show the five most massive satellites accreted by haloes Au5 and Au28, respectively. This shows that the BSP is not necessarily the most massive nor the most recent accretion event.

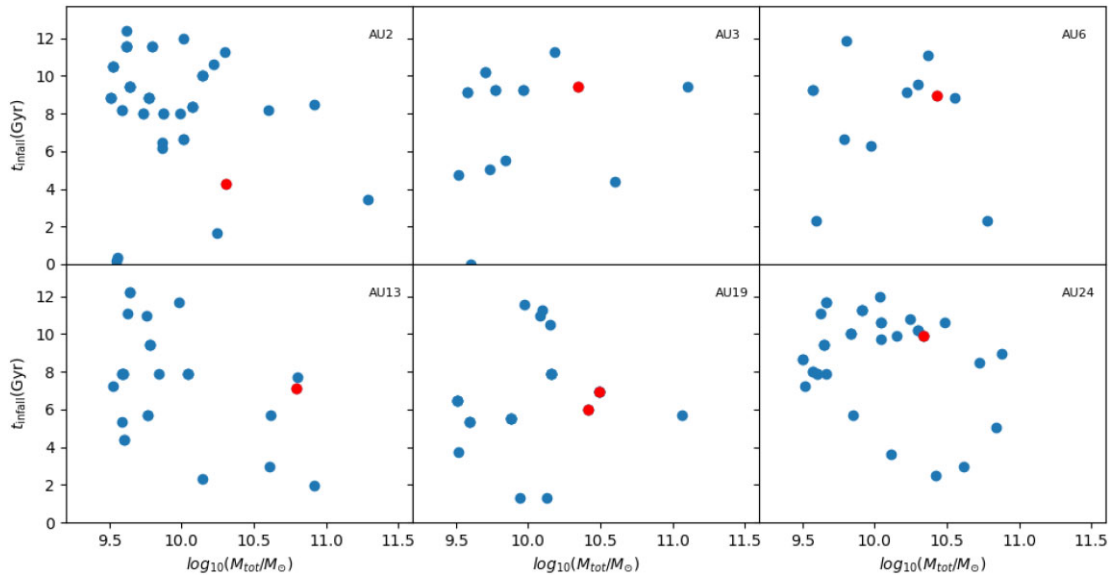


Figure 11. Each panel shows the distribution of lookback infall times and total infall mass of satellites in different Auriga haloes, similar to the right-hand panel of Fig. 10. Only satellites with total infall mass $>10^{9.5} M_{\odot}$ are considered. The brightest stream progenitors (BSPs) on these haloes are highlighted with red dots. Note that, in all cases, the BSPs are neither the most massive nor the most recent accretion event.

considered resolution level. This allowed us to link, for the first time, using fully cosmological hydrodynamical simulations, a clear observable property – the brightest stream surface brightness – with the accretion history of Milky Way-like galaxies.

For each halo, we generated several surface brightness (SB) maps reaching different limiting surface brightness levels. Starting from the shallower SB map, we have searched for the first clear signature of a stellar stream. This was done iteratively by increasing the limiting surface brightness of the maps at each step. To minimize the effects that dust and the Galactic cirrus would have on our detections, we have focused our analysis on the model r -band photometry. Considering other photometric

bands does not affect our main results. To take into account the effect of different galaxy inclinations on the identification of the brightest streams as a function of μ_r^{lim} , we have considered edge-on and face-on disc projections. To mimic the smoothing performed in observations to enhance diffuse structure and preserve image resolution, the distribution of stellar particles was assigned to bins of 2×2 kpc, and fluxes were integrated within each bin.

None of our models show signatures of streams for $\mu_r^{\text{lim}} \leq 25.5$ mag arcsec $^{-2}$. At the typical SB, limiting magnitude reached by current surveys, $\mu_r^{\text{lim}} \approx 28.5$ mag arcsec $^{-2}$, we show that the brightest stream can be detected in ≈ 30 per cent of our models.

Independently of the projection, we find that the cumulative function of the detected brightest stream strongly rises at values of $28.5 \text{ mag arcsec}^{-2}$, and then flattens again beyond $30.5 \text{ mag arcsec}^{-2}$. Furthermore, 13 per cent of our models show no detectable streams up to $\mu_r^{\text{lim}} = 31 \text{ mag arcsec}^{-2}$. Varying the projected orientation of our models has a significant impact in the detectability of bright stellar streams. With respect to the face-on projection, the cumulative function shows an increase of about 10 per cent of brightest stream detections at all limiting SB levels when the galaxies are projected edge-on. Our results show that, to obtain a representative census of the brightest LSB features in the nearby universe, it is thus critical to conduct surveys that reach fainter SB limiting magnitude than current surveys ($\mu_r^{\text{lim}} \approx 28 \text{ mag arcsec}^{-2}$, see Atkinson et al. 2013; Kado-Fong et al. 2018; Morales et al. 2018). Such surveys will be provided by the Vera Rubin-LSST as a result of its 10-yr campaign, which will reach $\mu_r^{\text{lim}} \approx 30 \text{ mag arcsec}^{-2}$. To achieve this goal, it will be key to develop accurate methods to systematically and efficiently perform sky background subtraction, as well as to characterize sources of contamination such as Galactic cirrus (see e.g. Román, Trujillo & Montes 2020), which become relevant for observation deeper than $\mu_r^{\text{lim}} \approx 28.5 \text{ mag arcsec}^{-2}$.

In general, we find our results are in good agreement with previous observational studies. We note, however, that the comparison between our results and those presented in, e.g. Atkinson et al. (2013) and Morales et al. (2018), is not straightforward, and there are several differences to take into account. First, our models do not account for dust extinction and background noise, both likely to conceal signatures of faint stellar substructures. We also have a much smaller sample of galaxies which, on an average, are more luminous than the observed ones. Because of these differences, it is expected that the fraction of galaxies with streams detected in the models to be slightly larger than those found in observations, which are likely to represent a lower detection limit. Finally, it is worth noting that, even though the properties of the simulated satellite populations are generally in good agreement with those observed in the MW and M31 satellites, our results are still limited by certain aspects of the model, such as numerical resolution and stellar feedback. For example, while the satellite phase-space scaling relation is well reproduced, the detailed internal structure of our simulated satellites, specially those with lower mass, are likely to be unresolved. In addition, diffusion of stellar streams in phase-space could also be enhanced due to the limited numerical resolution. Note that since we are focusing on the brightest stellar streams on each model, this limitation should be minimized. We also note that, as reported in Monachesi et al. (2019), Auriga stellar haloes are slightly more massive than some of the observed haloes in the local universe. This is partially the result of the accretion of slightly more massive satellites than observed (Simpson et al. 2018).

We have identified and characterized the main properties of BSPs, focusing in particular on their infall times (the time at which they first cross the virial radius of the host galaxy) and their infall mass. We find that BSPs can be accreted in a very wide range of t_{infall} , with values that can range from 10-Gyr ago to very recent accretion events at $t_{\text{infall}} = 1.6 \text{ Gyr}$. Interestingly, only 25 per cent of the BSPs have been recently accreted, within the last 5 Gyr. Thus, most BSPs correspond to relatively early accretion events (i.e. $5 < t_{\text{infall}} \lesssim 10 \text{ Gyr}$). As expected, BSPs associated with brighter stellar streams ($\mu_r^{\text{lim}} < \text{mag arcsec}^{-2}$) were typically accreted later than those with fainter substructures ($\mu_r^{\text{lim}} > 28.5 \text{ mag arcsec}^{-2}$). We also find that only 37 per cent of the BSPs can still be identified at the present day. The median t_{infall} for the surviving and the disrupted BSPs are $t_{\text{infall}} = 5.2$ and 6.7 Gyr , respectively.

Looking at the BSPs infall mass, we find a wide range of masses with values $9.5 \lesssim \log_{10} M_{\text{tot}}/M_{\odot} \lesssim 11.5$ in total mass and $7.5 \lesssim \log_{10} M_{*}/M_{\odot} \lesssim 10.5$ in stellar mass. This represents a variation of 2 and 3 dex in total and stellar mass, respectively. A comparison between surviving and destroyed BSPs shows that surviving BSPs are less massive at any given t_{infall} and have been accreted later at any given M_{tot} . We find that brighter streams tend to be associated with more massive BSPs. Indeed, there is a correlation between the BSPs infall mass and t_{infall} , such that more massive progenitors tend to be accreted at later times. However, we showed that this is not simply due to the fact that close to $z = 0$, galaxies are more likely to accrete more massive substructures. Indeed, we find that haloes that have relatively low-mass BSPs accreted at earlier times have accreted other satellites up to 1 dex more massive than the BSPs, and at similar times. Due to their larger internal velocity dispersion, debris from these more massive and luminous satellites are more rapidly phase-mixed. Furthermore, dynamical friction acts more efficiently on these more massive and luminous haloes. As a result, they are more rapidly disrupted compared to lower mass haloes accreted at similar times. In such cases, the brightest streams can be associated with less massive (but still luminous) galaxies that can continue orbiting their host and realising debris, for longer periods. Finally, we also show that, for most of the cases, BSPs are not the dominant contributors to the accreted stellar halo, even though they are always significant contributors; i.e. part of the subset of satellites that contribute to over 90 per cent of the accreted stellar halo mass. Our work thus complements the information that can be obtained from stellar haloes regarding the accretion history of late-type MW-mass galaxies (e.g. Deason et al. 2016; D’Souza & Bell 2018; Monachesi et al. 2019), namely the dominant accretion event, by adding extra information provided by the stellar streams.

ACKNOWLEDGEMENTS

We wish to thank the anonymous reviewers for thoughtful and constructive reports that helped us improve our manuscript. AV and FAG acknowledges support from ANID FONDECYT Regular 1211370. AV acknowledges support from DIDULS PTE192137. AM acknowledges support from ANID FONDECYT Regular 1212046. FAG and AM gratefully acknowledge support by the ANID BASAL project FB210003. FAG, AM, DP, and IG acknowledge funding from the Max Planck Society through a ‘Partner Group’ grant. IG acknowledges financial support from CONICYT Programa de Astronomía, Fondo ALMA-CONICYT 2017 31170048. FM acknowledges support through the Program ‘Rita Levi Montalcini’ of the Italian MUR. CSF acknowledges support from European Research Council (ERC) Advanced Investigator grant DMIDAS (GA 786910). This work was also supported by the Consolidated Grant for Astronomy at Durham (ST/L00075X/1). This work made use of the DiRAC Data Centric system at Durham University, operated by the Institute for Computational Cosmology on behalf of the STFC DiRAC HPC Facility (www.dirac.ac.uk). This equipment was funded by BIS National E-infrastructure capital grant ST/K00042X/1, STFC capital grants ST/H008519/1 and ST/K00087X/1, STFC DiRAC Operations grant ST/K003267/1 and Durham University. DiRAC is part of the National E-Infrastructure.

DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

REFERENCES

- Amorisco N. C., 2015a, *MNRAS*, 450, 575
 Amorisco N. C., 2015b, *MNRAS*, 450, 575
 Amorisco N. C., 2017, *MNRAS*, 469, L48
 Atkinson A. M., Abraham R. G., Ferguson A. M. N., 2013, *ApJ*, 765, 28
 Bell E. F., Monachesi A., Harmsen B., de Jong R. S., Bailin J., Radburn-Smith D. J., D'Souza R., Holwerda B. W., 2017, *ApJ*, 837, L8
 Bose S. et al., 2019, *MNRAS*, 486, 4790
 Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
 Bullock J. S., Johnston K. V., 2005, *ApJ*, 635, 931
 Conroy C. et al., 2019a, *ApJ*, 883, 107
 Conroy C., Naidu R. P., Zaritsky D., Bonaca A., Cargile P., Johnson B. D., Caldwell N., 2019b, *ApJ*, 887, 237
 Cooper A. P. et al., 2010, *MNRAS*, 406, 744
 Cooper A. P. et al., 2011, *ApJ*, 743, L21
 Crnojević D. et al., 2016, *ApJ*, 823, 19
 D'Souza R., Bell E. F., 2018, *Nat. Astron.*, 2, 737
 Deason A. J., Mao Y.-Y., Wechsler R. H., 2016, *ApJ*, 821, 5
 Fattahi A. et al., 2020, *MNRAS*, 497, 4459
 Fliri J., Trujillo I., 2016, *MNRAS*, 456, 1359
 Font A. S., McCarthy I. G., Belokurov V., 2021, *MNRAS*, 505, 783
 Gaia Collaboration et al., 2018, *A&A*, 616, A1
 Gómez F. A. et al., 2017b, *MNRAS*, 472, 3722
 Gómez F. A. et al., 2020, preprint ([arXiv:2011.12323](https://arxiv.org/abs/2011.12323))
 Gómez F. A., Helmi A., 2010, *MNRAS*, 401, 2285
 Gómez F. A., White S. D. M., Marinacci F., Slater C. T., Grand R. J. J., Springel V., Pakmor R., 2016, *MNRAS*, 456, 2779
 Gómez F. A., White S. D. M., Grand R. J. J., Marinacci F., Springel V., Pakmor R., 2017a, *MNRAS*, 465, 3446
 Grand R. J. J. et al., 2017, *MNRAS*, 467, 179
 Grand R. J. J. et al., 2021, *MNRAS*, 507, 4953
 Grand R. J. J., Springel V., Gómez F. A., Marinacci F., Pakmor R., Campbell D. J. R., Jenkins A., 2016, *MNRAS*, 459, 199
 Greggio L., Rejkuba M., Gonzalez O. A., Arnaboldi M., Iodice E., Irwin M., Neeser M. J., Emerson J., 2014, *A&A*, 562, A73
 Harmsen B., Monachesi A., Bell E. F., de Jong R. S., Bailin J., Radburn-Smith D. J., Holwerda B. W., 2017, *MNRAS*, 466, 1491
 Helmi A., 2020, *ARA&A*, 58, 205
 Helmi A., White S. D. M., de Zeeuw P. T., Zhao H., 1999, *Nature*, 402, 53
 Helmi A., White S. D. M., Springel V., 2003, *MNRAS*, 339, 834
 Hendel D., Johnston K. V., 2015, *MNRAS*, 454, 2472
 Ibata R. et al., 2020, *ApJ*, 914, 22
 Jog C. J., Combes F., 2009, *Phys. Rep.*, 471, 75
 Johnston K. V., 1998, *ApJ*, 495, 297
 Johnston K. V., Hernquist L., Bolte M., 1996, *ApJ*, 465, 278
 Johnston K. V., Bullock J. S., Sharma S., Font A., Robertson B. E., Leitner S. N., 2008, *ApJ*, 689, 936
 Kado-Fong E. et al., 2018, *ApJ*, 866, 103
 Karademir G. S., Remus R.-S., Burkert A., Dolag K., Hoffmann T. L., Moster B. P., Steinwandel U. P., Zhang J., 2019, *MNRAS*, 487, 318
 Laine S. et al., 2018, preprint ([arXiv:1812.04897](https://arxiv.org/abs/1812.04897))
 Majewski S. R., Siegel M. H., Kunkel W. E., Reid I. N., Johnston K. V., Thompson I. B., Landolt A. U., Palma C., 1999, *AJ*, 118, 1709
 Marinacci F., Pakmor R., Springel V., 2014, *MNRAS*, 437, 1750
 Martínez-Delgado D. et al., 2010, *AJ*, 140, 962
 Martínez-Delgado D., 2019, in Montesinos B., Asensio Ramos A., Buitrago F., Schödel R., Villaver E., Pérez-Hoyos S., Ordóñez-Etxeberria I., eds, Highlights on Spanish Astrophysics X, Proceedings of the XIII Scientific Meeting of the Spanish Astronomical Society held on July 16-20, 2018. Sociedad Española de Astronomía, Salamanca, Spain, p. 146
 McConnachie A. W. et al., 2009, *Nature*, 461, 66
 McConnachie A. W., 2012, *AJ*, 144, 4
 Merritt A., van Dokkum P., Abraham R., Zhang J., 2016, *ApJ*, 830, 62
 Merritt A., Pillepich A., van Dokkum P., Nelson D., Hernquist L., Marinacci F., Vogelsberger M., 2020, *MNRAS*, 495, 4570
 Monachesi A. et al., 2019, *MNRAS*, 485, 2589
 Monachesi A., Bell E. F., Radburn-Smith D. J., Bailin J., de Jong R. S., Holwerda B., Streich D., Silverstein G., 2016a, *MNRAS*, 457, 1419
 Monachesi A., Gómez F. A., Grand R. J. J., Kauffmann G., Marinacci F., Pakmor R., Springel V., Frenk C. S., 2016b, *MNRAS*, 459, L46
 Morales G., Martínez-Delgado D., Grebel E. K., Cooper A. P., Javanmardi B., Miskolczy A., 2018, *A&A*, 614, A143
 Naidu R. P., Conroy C., Bonaca A., Johnson B. D., Ting Y.-S., Caldwell N., Zaritsky D., Cargile P. A., 2020, *ApJ*, 901, 48
 Pakmor R., Springel V., 2013, *MNRAS*, 432, 176
 Pakmor R., Marinacci F., Springel V., 2014, *ApJ*, 783, L20
 Pakmor R., Springel V., Bauer A., Mocz P., Muñoz D. J., Ohlmann S. T., Schaal K., Zhu C., 2016, *MNRAS*, 455, 1134
 Panithanpaisal N., Sanderson R. E., Wetzel A., Cunningham E. C., Bailin J., Faucher-Giguère C.-A., 2021, *ApJ*, 920, 10
 Planck Collaboration et al., 2014, *A&A*, 571, A16
 Pop A.-R., Pillepich A., Amorisco N. C., Hernquist L., 2018, *MNRAS*, 480, 1715
 Quillen A. C., Minchev I., Bland-Hawthorn J., Haywood M., 2009, *MNRAS*, 397, 1599
 Rey M. P., Starkenburg T. K., 2022, *MNRAS*, 510, 4208
 Román J., Trujillo I., Montes M., 2020, *A&A*, 644, A42
 Sanderson R. E. et al., 2018, *ApJ*, 869, 12
 Schaye J. et al., 2015, *MNRAS*, 446, 521
 Searle L., Zinn R., 1978, *ApJ*, 225, 357
 Simpson C. M., Grand R. J. J., Gómez F. A., Marinacci F., Pakmor R., Springel V., Campbell D. J. R., Frenk C. S., 2018, *MNRAS*, 478, 548
 Smercina A. et al., 2020, *ApJ*, 905, 60
 Springel V., 2010, *MNRAS*, 401, 791
 Springel V., Hernquist L., 2003, *MNRAS*, 339, 289
 Springel V., Di Matteo T., Hernquist L., 2005, *MNRAS*, 361, 776
 Trujillo I., Fliri J., 2016, *ApJ*, 823, 123
 Tumlinson J., 2010, *ApJ*, 708, 1398
 van Dokkum P. G., Abraham R., Merritt A., 2014, *ApJ*, 782, L24
 Vogelsberger M., Genel S., Sijacki D., Torrey P., Springel V., Hernquist L., 2013, *MNRAS*, 436, 3031
 White S. D. M., Rees M. J., 1978, *MNRAS*, 183, 341
 White S. D. M., Frenk C. S., 1991, *ApJ*, 379, 52

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.