



LensWatch. I. Resolved HST Observations and Constraints on the Strongly Lensed Type Ia Supernova 2022qmx (“SN Zwicky”)

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Abstract

Supernovae (SNe) that have been multiply imaged by gravitational lensing are rare and powerful probes for cosmology. Each detection is an opportunity to develop the critical tools and methodologies needed as the sample of lensed SNe increases by orders of magnitude with the upcoming Vera C. Rubin Observatory and Nancy Grace Roman Space Telescope. The latest such discovery is of the quadruply imaged Type Ia SN 2022qmx (aka, “SN Zwicky”) at $z = 0.3544$. SN Zwicky was discovered by the Zwicky Transient Facility in spatially unresolved data. Here we present follow-up Hubble Space Telescope observations of SN Zwicky, the first from the multicycle “LensWatch (www.lenswatch.org)” program. We measure photometry for each of the four images of SN Zwicky, which are resolved in three WFC3/UVIS filters (F475W, F625W, and F814W) but unresolved with WFC3/IR F160W, and present an analysis of the lensing system using a variety of independent lens modeling methods. We find consistency between lens-model-predicted time delays ($\lesssim 1$ day), and delays estimated with the single epoch of Hubble Space Telescope colors ($\lesssim 3.5$ days), including the uncertainty from chromatic microlensing (~ 1 – 1.5 days). Our lens models converge to an Einstein radius of $\theta_E = (0.168^{+0.009}_{-0.005})''$, the smallest yet seen in a lensed SN system. The “standard candle” nature of SN Zwicky provides magnification estimates independent of the lens modeling that are brighter than predicted by $\sim 1.7^{+0.8}_{-0.6}$ mag and $\sim 0.9^{+0.8}_{-0.6}$ mag for two of the four images, suggesting significant microlensing and/or additional substructure beyond the flexibility of our image-position mass models.



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*Unified Astronomy Thesaurus concepts: Type Ia supernovae (1728); Strong gravitational lensing (1643)***1. Introduction**

Strong-gravitationally lensed supernovae (SNe) are rare events. In the strong lensing phenomenon, multiple images of a background source appear as light propagating along different paths is focused by a foreground galaxy or galaxy cluster. This requires a chance alignment along the line of sight between us the observers, the background source, and the foreground galaxy. Depending on the relative geometrical and gravitational potential differences of each path, the SN images typically appear delayed by hours to months (for galaxy-scale lenses) or years (for cluster-scale lenses).

Robust measurements of this “time delay” can constrain the Hubble constant (H_0) and the dark energy equation of state (e.g., w) in a single step (e.g., Refsdal 1964; Paraficz & Hjorth 2009; Linder 2011; Treu & Marshall 2016; Birrer et al. 2022b; Treu et al. 2022). Lensed SNe have several advantages relative to quasars, which have historically been used for this purpose (e.g., Vuissoz et al. 2008; Suyu et al. 2010; Tewes et al. 2013; Bonvin et al. 2017, 2018, 2019a; Birrer et al. 2019; Wong et al. 2020): (1) SNe fade quickly, enabling predictive experiments on the delayed appearance of trailing images, yielding more accurate models of the lens and source, as the SN (or quasar) and host fluxes are otherwise highly blended (Ding et al. 2021); (2) SNe have predictable light curves, simplifying time-delay measurements and enabling SN progenitor system constraints; (3) the impact of microlensing is somewhat mitigated, including a small (~ 0.1 day) “microlensing time delay” (Tie & Kochanek 2018; Bonvin et al. 2019b) and less-pronounced chromatic effects (Foxley-Marrable et al. 2018; Goldstein et al. 2018; Huber et al. 2019), though this can still be a significant source of uncertainty when the time delays are $\lesssim 1$ day (e.g., Goobar et al. 2017); and (4) time-delay measurements for lensed SNe require much shorter observing campaigns than lensed quasars.

While the advantages of using SNe for time-delay cosmography relative to other probes have been well documented (e.g., Refsdal 1964; Kelly et al. 2015; Goobar et al. 2017; Goldstein et al. 2018; Huber et al. 2019; Pierel & Rodney 2019; Suyu et al. 2020; Pierel et al. 2021; Rodney et al. 2021), these events have proved extremely difficult to detect. Since the first multiply imaged SN discovery by Kelly et al. (2015), there have been only four more such discoveries (Goobar et al. 2017; Rodney et al. 2021; Chen et al. 2022; Kelly et al. 2022) despite dedicated surveys to increase the sample (e.g., Petrushevskaya et al. 2016, 2018; Fremling et al. 2020; Craig et al. 2021).

SNe of Type Ia (SNe Ia), those employed for decades as “standardizable candles” to measure cosmological parameters by way of luminosity distances and the cosmic distance ladder (e.g., Garnavich et al. 1998; Riess et al. 1998; Perlmutter et al. 1999; Scolnic et al. 2018; Brout et al. 2022), are particularly valuable when strongly lensed. In addition to having a well-understood model of light-curve evolution (Hsiao et al. 2007; Guy et al. 2010; Saunders et al. 2018; Leget et al. 2020; Kenworthy et al. 2021; Pierel et al. 2022), their standardizable absolute brightness can provide additional leverage for lens modeling by limiting the uncertainty caused by mass-sheet degeneracy (Falco et al. 1985; Kolatt & Bartelmann 1998; Holz 2001; Oguri & Kawano 2003; Nordin et al. 2014;

Patel et al. 2014; Rodney et al. 2015; Xu et al. 2016; Birrer et al. 2022a), though only in cases where millilensing and microlensing are not extreme (see Goobar et al. 2017; Foxley-Marrable et al. 2018; Dhawan et al. 2020). The first such discovery was iPTF16geu (Goobar et al. 2017), which had image separations resolved using adaptive optics (AO) and the Hubble Space Telescope (HST), with very short time delays of ~ 0.25 –1.5 days (Dhawan et al. 2020). Nevertheless, the detection and analysis of objects like iPTF16geu are critical to the future of lensed SN research as unresolved, galaxy-scale lenses are expected to be relatively common among lensed SN discoveries made with the Vera C. Rubin Observatory (Collett 2015; Goldstein et al. 2019; Wojtak et al. 2019).

LensWatch³⁴ is a collaboration with the goal of finding gravitationally lensed SNe, both by monitoring active transient surveys (e.g., Fremling et al. 2020; Jones et al. 2021) and by way of targeted surveys (Craig et al. 2021). The collaboration maintains a Cycle 28 HST program,³⁵ given long-term (three-cycle) target-of-opportunity (ToO) status due to the relatively low expected lensed SN rates. The program includes three ToO triggers (two nondisruptive, one disruptive), and was designed to provide high-resolution follow-up imaging for a ground-based lensed SN discovery, which is critical for galaxy-scale multiply imaged SNe due to their small image separations (e.g., Goobar et al. 2017).

A new multiply imaged SNIa was discovered in 2022 August by the Zwicky Transient Facility (ZTF; Fremling et al. 2020),³⁶ subsequently classified and analyzed by Goobar et al. (2022; hereafter G22). The four separate images of SN 2022qmx (aka “SN Zwicky”) were spatially unresolved in ground-based imaging with separations of $\lesssim 0''.3$. In order to provide reliable photometry and the data necessary for accurate lens modeling of the system, optical space-based observations are ideal. We therefore report the first observations and results of the LensWatch collaboration, which triggered HST-GO program 16264 to schedule follow-up imaging of SN Zwicky.

This work is the first in a series of papers that utilize data from the LensWatch program. Section 2 gives an overview of SN Zwicky and presents the final HST observation characteristics including triggering, orbit design, and implementation. Section 3 details our lens modeling methodology and constraints on the lensing system, and our analysis of SN Zwicky (including photometry and measurements of time delays and magnifications) are reported in Section 4. We conclude with a discussion of implications of this new data set, as well as future observation plans, in Section 5.

2. Observing with HST

As possible discovered lensed system configurations are highly variable, it is necessary to design a custom follow-up campaign for each new discovery. We therefore give an overview of the lensing system and SN characteristics for SN Zwicky, and then the subsequent observational choices made for the LensWatch HST ToO trigger.

³⁴ <https://www.lenswatch.org>

³⁵ HST-GO-16264

³⁶ <https://www.wis-tns.org/object/2022qmx>

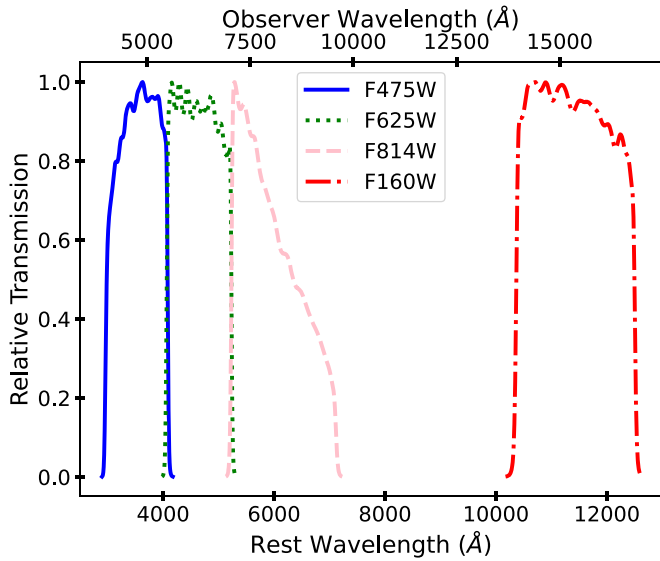


Figure 1. The HST filters used to observe SN Zwicky, with rest-frame wavelength on the lower axis and observer-frame wavelength of the upper axis. The three bluer filters are from WFC3/UVIS, while F160W is from WFC3/IR.

2.1. The Multiply Imaged SN Zwicky

The discovery, description of ground-based observations, and initial analysis of SN Zwicky are presented by G22. Briefly, the SN was discovered by ZTF at Palomar Observatory under the Bright Transient Survey (BTS; Fremling et al. 2020) on 2022 August 1 (MJD 59792). The SED Machine (SEDM) and Nordic Optical Telescope (NOT) provided spectroscopic classification of SN Zwicky as a Type Ia (SN Ia) at $z = 0.35$ and near maximum light on 2022 August 21–22 (MJD 59812–59813). Although the multiple images were not resolved by ZTF, the inferred absolute magnitude of SN Zwicky for this redshift was ~ 3 magnitudes brighter than normal, suggesting the presence of strong gravitational lensing. G22 also obtained subsequent spectroscopic observations from the Keck observatory, Hobby–Eberly Telescope, and the Very Large Telescope (VLT), which led to a final SN redshift of $z = 0.3544$ and lensing galaxy redshift of $z = 0.22615$. The multiple images of SN Zwicky were first resolved with the Keck telescope Laser Guide Star aided Adaptive Optics (LGSAO) Near-IR Camera 2 (NIRC2) on 2022 September 15 (MJD 59837; see G22 for details).

2.2. ToO: Filter Choices and Orbit Design

Roughly 12 days after the spectroscopic classification of SN Zwicky, we used a nondisruptive HST ToO trigger to obtain follow-up WFC3/UVIS and IR images of the lensing system. The average turnaround for a nondisruptive ToO trigger is $\gtrsim 21$ days, but close coordination with the HST scheduling team at STScI led to receiving our first images after 17 days on 2022 September 21 (MJD 59843), or ~ 44 observer-frame (~ 32 rest-frame) days postdiscovery and ~ 37 observer-frame (~ 27 rest-frame) days after maximum brightness.

The anticipated image separations for a galaxy-scale lens of this mass and redshift are small enough that resolving the individual images with WFC3/IR ($0''.13 \text{ pix}^{-1}$), where the point-spread function (PSF) is severely undersampled, is unlikely. For the purposes of accurate photometry and lens modeling, the highest possible resolution imaging is required,

Table 1
HST WFC3 Photometric Filter Definitions and Exposure Times

Band	Rest λ_{eff} (Å)	Obs λ_{eff} (Å)	Instrument	Exp. Time (s)
F475W	3549	4792	UVIS	126
F625W	4636	6258	UVIS	39
F814W	5965	8053	UVIS	60
F160W	11402	15392	IR	207

and we therefore turned to WFC3/UVIS ($0''.04 \text{ pix}^{-1}$) to resolve the multiple images. We selected the F814W, F625W, and F475W filters to provide nonoverlapping coverage across the full optical wavelength range (~ 3500 – 6000 Å in the rest frame; see Figure 1 and Table 1). Additionally, the ground-based follow-up campaign of SN Zwicky included (resolved) *H*-band Keck-AO imaging, and we therefore included (unresolved) WFC3/IR F160W observations to provide overall calibration and extra information about the lensing system.

The four filters were efficiently packed into a single orbit of observing using the 512×512 subarrays for both UVIS and IR imaging, even with three dithers per filter to reduce the impact of cosmic rays and provide optimal sampling of the (Figure 2). The four images of SN Zwicky were successfully resolved in the three UVIS filters, which provided a full-color image (Figure 3).

3. Modeling of the Lensing System

3.1. Analysis Methods

In this section, we summarize the lens modeling analysis we carried out using the HST data presented in Section 2, leading to insights into the lensing galaxy mass (see the Appendix). Given the very low number of identified strongly lensed SNe, this procedure has mainly been applied to strongly lensed quasars, e.g., by the H_0 Lenses in COSMOGRAIL’s Well-spring (H0LiCOW) collaboration (e.g., Wong et al. 2020).

For galaxy-scale lenses, the lens mass distribution is usually described by profiles such as the singular isothermal ellipsoid (SIE; Kormann et al. 1994) or the singular power-law elliptical mass distribution (SPEMD; Barkana 1998), in combination with an external shear component describing the influence from massive line-of-sight objects (McCully et al. 2017). These parameters can be constrained by the observed image positions alone (those measured in Section 4.1), and/or through the pixel intensities of the HST images. This requires a model of the lens light distribution, which is typically described by one or more stacked Sérsic profiles (De Vaucouleurs 1948; Sérsic 1963), as well as a model for the lensed SN represented by a PSF (described in Section 4.1).

Given these different potential methodologies, we used three independent software packages and five total methods to carry out independent analyses of SN Zwicky, which provide an examination of the potential modeling systematics and allow us to marginalize over them (e.g., Ertl et al. 2023; Shajib et al. 2022). The three software packages are *lfit_gui* (Shu et al. 2016), *LENSTRONOMY* (Birrer et al. 2015; Birrer & Amara 2018; Birrer et al. 2021), and the Gravitational Lens Efficient Explorer (GLEE; Suyu & Halkola 2010; Suyu et al. 2012; Ertl et al. 2023), and the five methods explored here are summarized by Table 2.

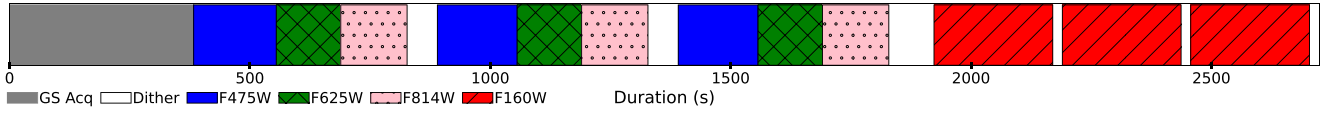


Figure 2. The layout of the orbit used for these HST observations. The dither sections (white) include other overheads such as filter changing.

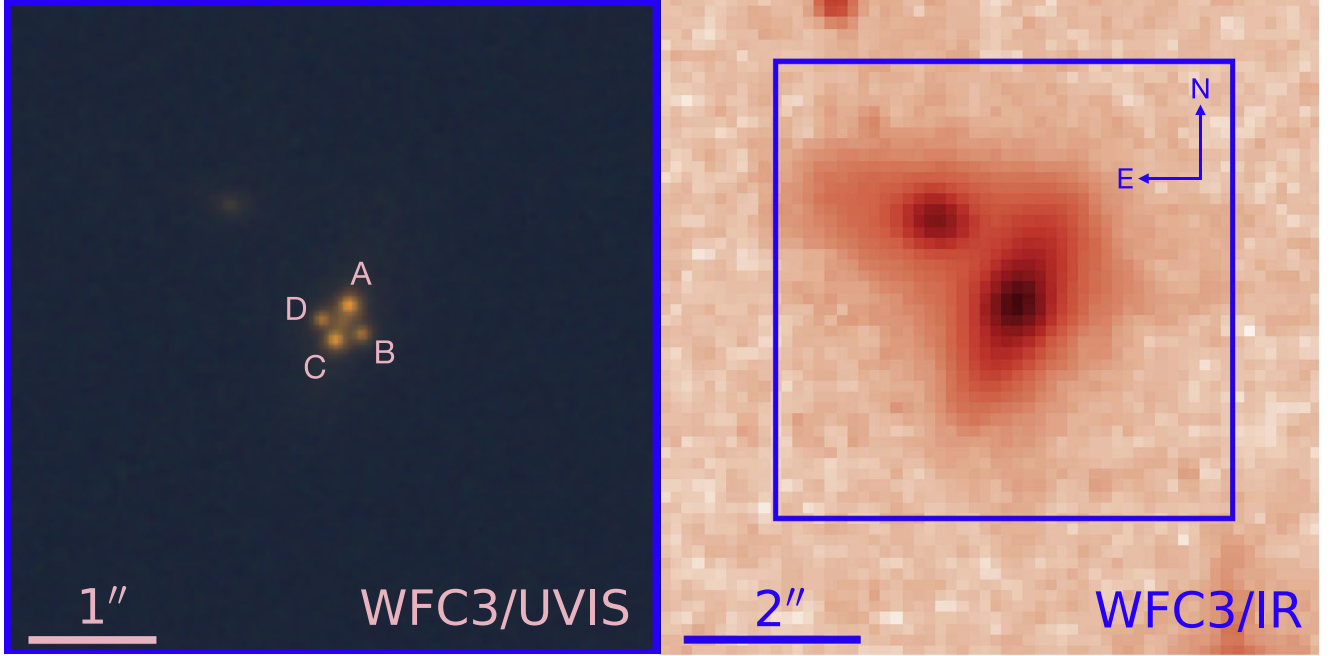


Figure 3. Left: a WFC3/UVIS combined color image ($R = F814W$, $G = F625W$, $B = F475W$) of SN Zwicky, with the multiple images labeled as A–D. Right: WFC3/IR F160W image of the galaxy lensing of SN Zwicky, where the multiple images are not resolved. The orientation of the SN host galaxy (northeast of the lens) suggests there is not much intervening matter along the line of sight apart from the lensing galaxy. The blue square is the footprint of the WFC3/UVIS image.

Table 2
Summary of the Lens Modeling Methodologies

Name	Code	Lens Model Components	Fitted Filters	Modeling Team
lfit_gui	lfit_gui	SIE	All	Y.S.
LS1	LENSTRONOMY	SIE	All	N.A., A.J.S.
LS2	LENSTRONOMY	Power law + γ_{ext}	All	L.M., S.B.
GLEE	GLEE	Power law + γ_{ext}	F814W	S.E., S.S., S.H.S.
SALT+LS ^a	LENSTRONOMY	Power law + γ_{ext}	All	X.H., W.S., E.S., S.A.

Note.

^a The SALT+LS model also includes constraints from the SN Ia’s absolute magnitude (see Section 3.1).

Using the first four models reported in Table 2, we found a significantly higher reduced- χ^2 when fitting both image positions and fluxes compared to fitting only image positions, indicating the presence of substructure and/or microlensing not captured by the lens modeling process. Additionally, the SALT+LS modeling method first performs PSF photometry to obtain individual image fluxes, which are fit directly with the SALT2.4 model (Betoule et al. 2014). This initial fit is used to provide a prior on the magnifications given the standardizability of SN Zwicky as an SN Ia (see Section 4.2 for details), then the lens model parameters are constrained with the image positions. As is apparent in the following section, the SALT+LS model magnification estimates agree with the models constrained with positions only for images B and D but not A and C, further supporting our assumption that images A and C are impacted by additional factors.

As a result of the above, we rely only on the positions of the multiple images to constrain the initial four lens models. Our interpretation is that, pending updated difference image photometry, there is some combination of additional microlensing and/or millilensing impacting images A and C. We discuss this more in Section 5, and will wait for the upcoming template image to improve both our lens models and photometry. For the remainder of this work, we refer to the models used to constrain the lensing system parameters (lfit_gui, LS1, LS2, and GLEE) as the “primary” models, and we refer to the SALT+LS model by name.

3.2. Lens Model Constraints

Each of the primary lens modeling methods described in Section 3.1 was used to constrain the magnifications and time

Table 3
Lens Modeling Constraints on Key Parameters

Parameter	Unit	lfit_gui	LS1	LS2	GLEE	Final	SALT+LS
θ_E	"	$0.166^{+0.0010}_{-0.0019}$	$0.167^{+0.0005}_{-0.0005}$	$0.173^{+0.0086}_{-0.0071}$	$0.168^{+0.0046}_{-0.0037}$	$0.168^{+0.009}_{-0.005}$	0.155 ± 0.0004
μ_A	...	$-2.05^{+0.17}_{-0.26}$	$-2.46^{+0.12}_{-0.11}$	$-1.45^{+0.49}_{-1.05}$	$-1.26^{+0.56}_{-1.43}$	$-1.81^{+0.90}_{-0.89}$	-5.70 ± 0.42
μ_B	...	$3.96^{+0.32}_{-0.20}$	$4.41^{+0.12}_{-0.14}$	$3.75^{+1.17}_{-0.65}$	$2.78^{+1.71}_{-0.67}$	$3.72^{+1.04}_{-1.24}$	3.44 ± 0.42
μ_C	...	$-3.47^{+0.25}_{-0.43}$	$-3.99^{+0.16}_{-0.16}$	$-1.94^{+0.67}_{-1.45}$	$-2.07^{+0.95}_{-2.29}$	$-2.87^{+1.51}_{-1.50}$	-4.57 ± 0.42
μ_D	...	$4.36^{+0.37}_{-0.23}$	$4.85^{+0.14}_{-0.15}$	$4.19^{+1.33}_{-0.78}$	$3.10^{+1.92}_{-0.73}$	$4.12^{+1.19}_{-1.36}$	3.72 ± 0.42
Δt_{BA}	Days	$-0.48^{+0.05}_{-0.03}$	$-0.41^{+0.01}_{-0.02}$	$-0.50^{+0.15}_{-0.12}$	$-0.59^{+0.22}_{-0.20}$	$-0.50^{+0.15}_{-0.21}$	-0.26 ± 0.13
Δt_{CA}	Days	$-0.24^{+0.03}_{-0.02}$	$-0.21^{+0.01}_{-0.01}$	$-0.15^{+0.04}_{-0.11}$	$-0.27^{+0.09}_{-0.09}$	$-0.22^{+0.10}_{-0.10}$	0.02 ± 0.13
Δt_{DA}	Days	$-0.41^{+0.04}_{-0.03}$	$-0.36^{+0.01}_{-0.02}$	$-0.41^{+0.12}_{-0.11}$	$-0.50^{+0.18}_{-0.16}$	$-0.42^{+0.12}_{-0.18}$	-0.17 ± 0.13

Note.

^a See the [Appendix](#) for the κ and γ results for each lens model.

delays for each SN image independently, as well as the Einstein radius (θ_E) of the lensing galaxy. These results are summarized in Table 3, where Δt_{iA} refers to the relative delay between the i th image and image A (i.e., $t_i - t_A$), and we also report the “Final” combined measurement for each parameter. This final value is the equally weighted average of each lens modeling result, and the uncertainty is a combination of the standard deviation of the primary model results and the statistical scatter in the average posterior distribution. As the models are equally weighted to obtain our final values for each parameter, we show normalized posterior distributions simply for visual comparison in Figures 4–6. Note that the models from Table 2 with fewer parameters also have narrower posterior distributions. We also include the results of the SALT+LS model, which reveals the impact of including information about the SN Ia standardizability. The additional specific lensing model parameters measured by each method, as well as more details about the modeling processes, are given in the [Appendix](#).

While we expect improved constraints following the Lens-Watch template image scheduled for 2023 as we will be able to disentangle the SN and lensing galaxy flux more reliably, the level of agreement between the primary modeling methods gives us confidence in the final constraints and uncertainties. We also find good agreement between these key parameters with the modeling of [G22](#), who used only the resolved near-IR Keck data. We note that our measured Einstein radius of $\theta_E = (0.168^{+0.009}_{-0.005})''$ is the smallest detected value for a multiply imaged SN thus far, and corresponds to a lens mass of $\sim 8 \times 10^9 M_\odot$. The similar lensed SN iPTF16geu had a measured $\theta_E = 0''.29$ (More et al. 2017; Mörtzell et al. 2020), with time delays of ~ 1 day (Dhawan et al. 2020). Here we predict time delays of ~ 0.2 – 0.5 days for each of the images of SN Zwicky, which are well below the predicted time-delay measurement uncertainty for even a resolved and well-sampled lensed SN Ia due to the impacts of microlensing (e.g., Pierel et al. 2021; Huber et al. 2022).

We also note that the SALT+LS method, which uniquely uses the measured photometry to infer a standardized absolute magnitude measurement of SN Zwicky and sets a prior on the image magnifications (see Section 3.1), is generally in good agreement with other methods apart from the predicted magnifications for images A and D and a slightly lower θ_E . The method also significantly reduced the plausible model parameter space (see the [Appendix](#)), which lends weight to the claims that SN Ia standardization can significantly improve lens

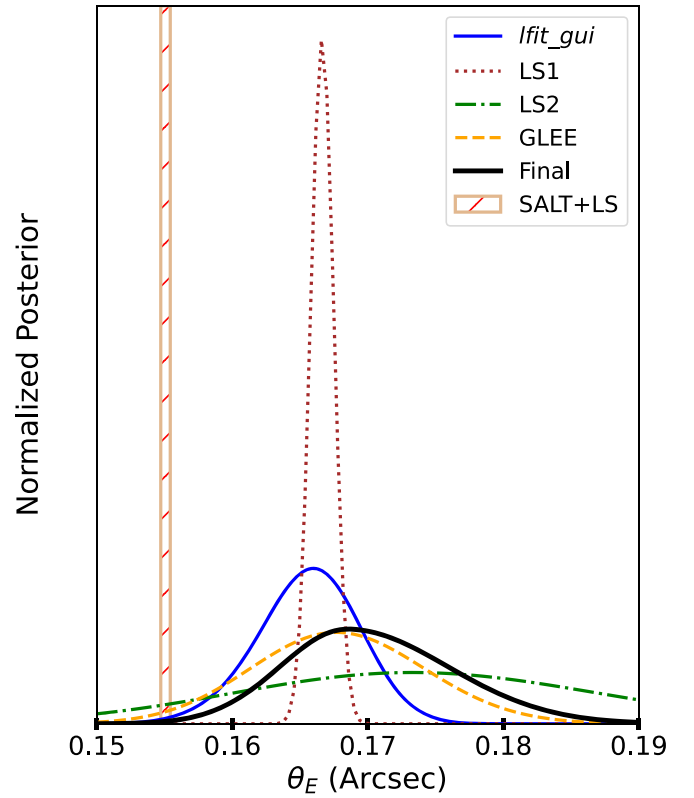


Figure 4. Normalized posterior distributions for θ_E of the four primary lens models, and their combined constraint (dark gray mean, light gray uncertainty). The SALT+LS model constraint is also shown for comparison (red slashed).

modeling efforts when microlensing is minimal. By implementing models that did not include this extra step alongside SALT+LS, the relative agreement (or disagreement) between methods was a useful indicator of additional substructure/microlensing beyond the primary lens modeling flexibility.

4. Analysis of SN Zwicky

4.1. HST Photometry

Due to the compact nature of the lensing system and difficulty in disentangling the SN and lens galaxy fluxes, an identical “template” epoch has been scheduled for ~ 6 – 12 months after the first, once the SN has long faded. This will provide more precise measurements and constraints for SN

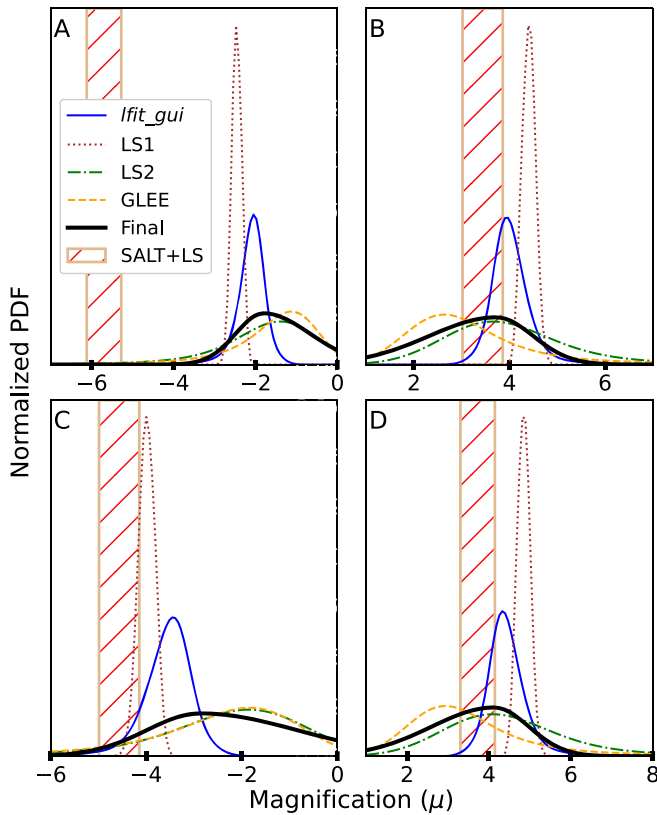


Figure 5. Normalized posterior distributions of the four primary lens models for the four image magnifications, and their combined constraints (dark gray mean, light gray uncertainty). The SALT+LS model constraints are also shown for comparison (red slashed). Note that images A, C have negative magnification and B, D have positive magnification, indicating all lens models agree on the parity of each image in addition to the absolute value of the magnification.

Zwicky and the lensing system in general. In the meantime, we have used PSF photometry to measure optimally the brightness of each SN image.

HST photometry for SN Zwicky was measured using PSF photometry on the WFC3/UVIS “FLC” images, which are individual exposures that have been bias subtracted, dark subtracted, and flat fielded but not yet corrected for geometric distortions. The UVIS data processing also includes a charge transfer efficiency (CTE) correction, which results in FLC images instead of the FLT images used for WFC3/IR. The WFC3/UVIS2 pixel area map (PAM) for the corresponding subarray was also applied to each exposure to correct for pixel area variations across the images.³⁷

In most cases, the individual exposures for each filter are “drizzled” together to create a single image (e.g., Figure 3). Here, we primarily are concerned with precisely measuring the position and brightness of each SN image, which (without a template image) requires accurate fitting of a PSF model. Drizzled images can introduce inconsistencies into the modeling of a PSF, and so we restrict ourselves to the FLCs to preserve the PSF structure. We use the standard HST PSF models³⁸ to represent the PSF, which also take into account spatial variations across the detector.

³⁷ <https://www.stsci.edu/hst/instrumentation/wfc3/data-analysis/pixel-area-maps>

³⁸ <https://www.stsci.edu/hst/instrumentation/wfc3/data-analysis/psf>

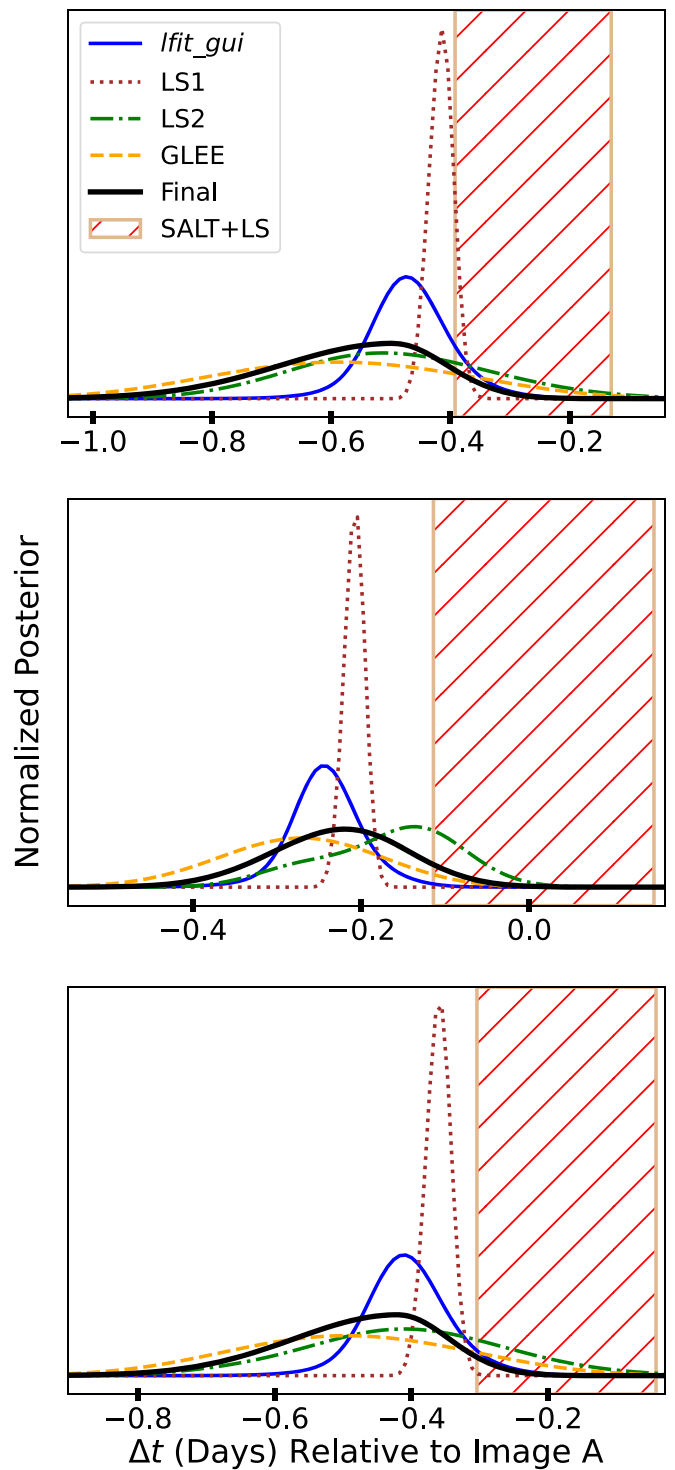


Figure 6. Normalized posterior distributions of the four primary lens models for the time delays relative to image A, and their combined constraints (dark gray mean, light gray uncertainty). The SALT+LS model constraints are also shown for comparison (red slashed).

For each UVIS filter, we implement a Bayesian nested sampling routine³⁹ to constrain the (common) SN flux and relative position simultaneously in all three FLCs for all four SN images (Figure 7). Each PSF was fit to the multiple SN images within a 5×5 pixel square in an attempt to limit the

³⁹ NESTLE: <http://kylebarbary.com/nestle>.

Table 4
Photometry and Colors Measured for Each Image of SN Zwicky in AB Magnitudes

Image	F475W	F625W	F814W	F475W – F625W	F475W – F814W	F625W – F814W
A	23.22 ± 0.04	21.67 ± 0.02	20.67 ± 0.01	1.55 ± 0.04	2.55 ± 0.04	1.00 ± 0.02
B	24.31 ± 0.07	22.65 ± 0.03	21.71 ± 0.02	1.66 ± 0.08	2.60 ± 0.08	0.94 ± 0.04
C	23.35 ± 0.04	21.90 ± 0.02	20.88 ± 0.02	1.44 ± 0.04	2.47 ± 0.04	1.02 ± 0.03
D	24.26 ± 0.07	22.72 ± 0.04	21.60 ± 0.02	1.53 ± 0.08	2.65 ± 0.07	1.12 ± 0.04

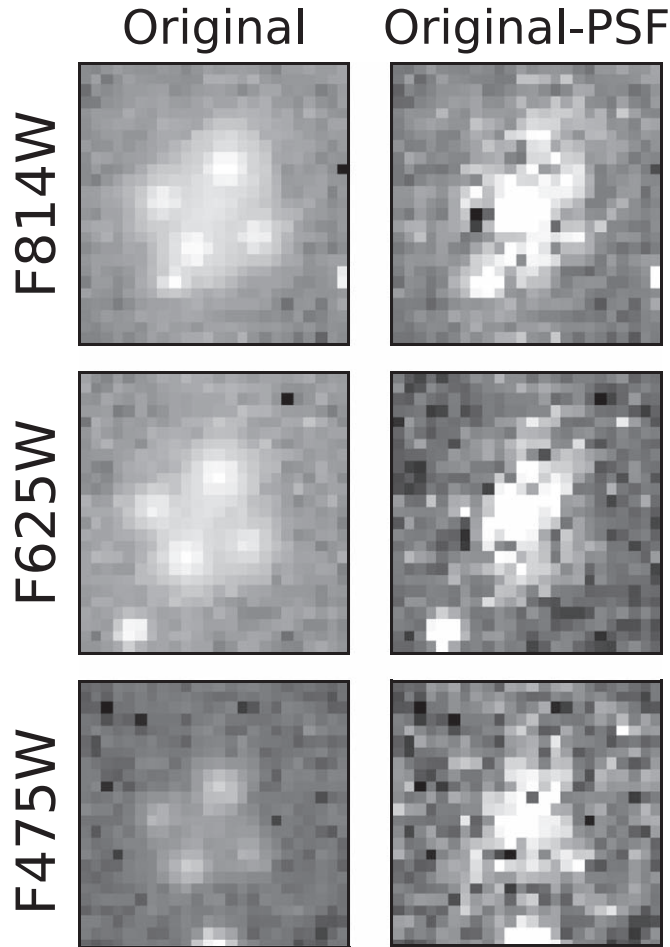


Figure 7. Results (right) of subtracting the best-fit PSF models from a single FLC for each WFC3/UVIS filter (left). Some residuals remain, particularly for the brightest images, but our planned template epoch will significantly improve the measured fluxes.

contamination of both the lensing galaxy (as we assume a constant background in the fitted region) and other SN images. The PSF FWHM for WFC3/UVIS is <2 pixels, so this PSF size should include $\sim 99\%$ of the total SN flux and not be contaminated by significant flux from the other images (each $\gtrsim 5$ pixels away).

The final measured flux is the integral of each full-fitted PSF model, which is 101×101 pixels and large enough to contain approximately all of the SN flux. These corrected fluxes were converted to AB magnitudes using the time-dependent inverse sensitivity and filter pivot wavelengths provided with each data file. The final measured magnitudes and colors are reported in Table 4.

4.2. Single Epoch Time Delays and Magnifications from HST Photometry

We use the single epoch of HST photometry from Table 4 to constrain the time delays and magnifications for the multiple images of SN Zwicky in the manner of Rodney et al. (2021). As measuring the difference in time of peak brightness for each image directly (e.g., Rodney et al. 2016) is not possible with a single epoch, we instead constrain the age of each SN image given a single light-curve model. The relative age difference for each image is also a measure of the time delay, though we note this method is only possible because we have a reliable model for the light (and color) curve evolution as SN Zwicky is of Type Ia. As we use some information from the unresolved light curve in G22, we also fit the data with the commonly used Spectral Adaptive Lightcurve Template (SALT2; Guy et al. 2010), which provides a simple parameterization of SN Ia normalization (x_0), shape or stretch (x_1), and color (c) used for light-curve standardization. The remaining SALT2 parameters are the time of peak B -band brightness (t_{pk}) and the SN redshift.

We fit the photometry of the multiple images simultaneously using the SNTD software package (Pierel & Rodney 2019), where we also include the known effects of Milky Way dust ($E(B - V) = 0.16$, $R_V = 3.1$) based on the maps of Schlafly & Finkbeiner (2011) and the extinction curve of Fitzpatrick (1999). Additionally, we use the simulations of Goldstein et al. (2018) to estimate the additional uncertainty introduced by chromatic microlensing. Using the time of peak estimate from G22 our observations are ~ 46 days post explosion, which corresponds to ~ 0.05 , 0.05 , and 0.11 mag of additional color uncertainty in rest-frame $U - B$ ($\sim F475W - F625W$), $B - V$ ($\sim F625W - F814W$), and $U - V$ ($\sim F475W - F814W$), respectively (95% confidence; see Figure 5 in Goldstein et al. 2018). We add these uncertainties in quadrature to the color uncertainties shown in Table 4 for the fitting process.

We follow the methods outlined by Rodney et al. (2021) to measure the time delays for SN Zwicky, who performed a similar analysis with the single epoch of SN Requiem. This process uses the SN color curves to constrain the time delay (with the SNTD “Color” method), and then fits for relative magnifications (with the SNTD “Series” method). Unlike the analysis of SN Requiem, an unresolved light curve exists for SN Zwicky, which in G22 was analyzed to give $t_{\text{pk}} = 59808.6$, $c = 0.005$, and $x_1 = 1.16$. While our single epoch of photometry should constrain the color parameter, it will be unable to constrain the x_1 parameter and there will be significant degeneracies between the time delays and t_{pk} (as seen in Rodney et al. 2021). We therefore allow the t_{pk} parameter, which here describes the time of peak for image A (see Figure 3 for the naming convention), to vary only within fifteen days of 59808.6. We also fix x_1 to the parameter derived

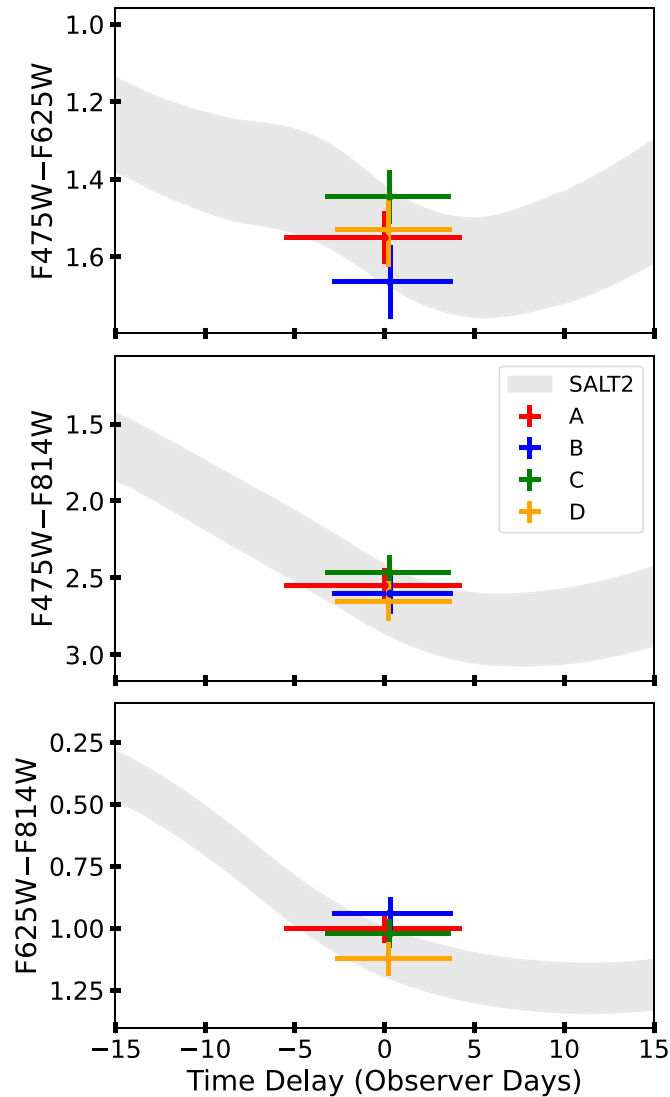


Figure 8. Measurements of the time delay and color for each image of SN Zwicky, with image A the reference image (i.e., $\Delta t = 0$). The vertical error bars are the photometric precision based on the work in Section 4.1 combined with an additional microlensing uncertainty, while the horizontal error bars are the 16th and 84th quantiles for the time-delay posterior of images B–D (columns 2 and 3–5 in Figure 9), and t_{pk} for image A. The gray shaded region is the best-fit SALT2 model from the SNTD Color method.

by G22 ($x_1 = 1.16$), mainly to ensure an accurate light-curve standardization. We repeated the fitting first following the choice in Rodney et al. (2021) to set $x_1 = 0$ and second allowing x_1 to vary within 3σ of the value measured in G22, and found these varied the time delays by $\lesssim 0.5$ days, well within our measurement error bars. We also checked the difference in the measured time delays when fixing the value of t_{pk} to 59808.6 and found a difference of $\lesssim 0.5$ days. Finally, we note that the additional uncertainty added due to chromatic microlensing changes the measured time delay by $\lesssim 0.5$ days as well, but increases the measurement uncertainties by ~ 1 –1.5 days.

SNTD finds a common value for c among all SN Zwicky images while varying the time delays of images B–D relative to the value of t_{pk} , which describes image A, within relative bounds of $[-15, 15]$ days. Figure 8 shows the measured colors and time delays for SN Zwicky with the best-fit SALT2 model

Table 5
Summary of the SALT2 Parameters Used in Fitting SN Zwicky

SNTD Method	Parameter	Varied?	Bounds	Value
...	z	No	...	0.3544
...	x_1	No	...	1.16
Color	t_{pk} (MJD)	Yes	$[-15, 15] + 59808.6$	$59808.24^{+4.30}_{-5.59}$
Color	c	Yes	$[-0.3, 0.3]$	$0.03^{+0.14}_{-0.06}$
Series	x_0	Yes	$[0, 1]$	$m_B = 19.62^{+0.02}_{-0.03}$

Table 6
Measured Time Delays and Magnifications Compared to the Predictions from the Lens Models from Section 3

Image	$(\Delta t_{IA})_{\text{meas}}$ Days	$(\Delta t_{IA})_{\text{pred}}$ Days	$ \mu_{\text{meas}} $	$ \mu_{\text{pred}} $
A	$8.31^{+4.16}_{-1.43}$	$1.81^{+0.90}_{-0.89}$
B	$0.30^{+3.51}_{-3.22}$	$-0.50^{+0.15}_{-0.21}$	$3.24^{+1.69}_{-0.57}$	$3.72^{+1.04}_{-1.24}$
C	$0.30^{+3.40}_{-3.59}$	$-0.22^{+0.10}_{-0.10}$	$6.73^{+3.38}_{-1.16}$	$2.87^{+1.51}_{-1.50}$
D	$0.19^{+3.53}_{-2.97}$	$-0.42^{+0.12}_{-0.18}$	$3.39^{+1.65}_{-0.62}$	$4.12^{+1.19}_{-1.36}$

overlaid. While all colors are used in the fit simultaneously, the photometric and model uncertainties mean $F625W - F814W$ provides the most constraining power, followed by the colors that include the rest-frame ultraviolet. After these time delays have been measured with the SNTD Color method, where we fix all the best-fit parameters and use the SNTD Series method to estimate the magnification ratios for images B–D (within bounds of $[0.01, 100]$) relative to image A. The fitting procedure for the light-curve parameters is summarized in Table 5.

As mentioned above, SNTD measures an overall normalization (x_0) and relative magnifications, and so we convert the combination of x_0 and magnification ratios to absolute magnitudes by assuming SN Zwicky is a perfect standardizable candle. Specifically, we apply light-curve corrections based on Table 5 for the stretch ($x_1 = 1.16$, with luminosity coefficient $\alpha = 0.14$) and color ($c = 0.03$, with a luminosity coefficient of $\beta = 3.1$) in the manner of Scolnic et al. (2018) to obtain absolute magnitude estimates. We then compare the distance modulus of each image to the value predicted by a flat Λ CDM model (with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_m = 0.3$) for an average SN Ia ($M_B = -19.36$; Richardson et al. 2014) at $z = 0.3544$, which results in a measure of the absolute magnifications. We combine the statistical uncertainties on each measured magnification with a systematic uncertainty based on the intrinsic scatter of SN Ia absolute magnitudes (0.1 mag; Scolnic et al. 2018). The measured time delays and magnifications (with subscript “meas”) are shown in Table 6 compared with lens-model-predicted values from Section 3.2 (with subscript “pred”). The posterior distributions for all parameters fit with SNTD (using the conversions listed above) are shown in Figures 9 and 10. While the relative time-delay uncertainties are too large to provide a useful direct cosmological constraint, these results are a valuable check on our lens modeling predictions. The agreement also supports the plausibility of measuring time delays in a single epoch, at least when the lensed SN is of Type Ia and there is some constraint on the overall explosion date.

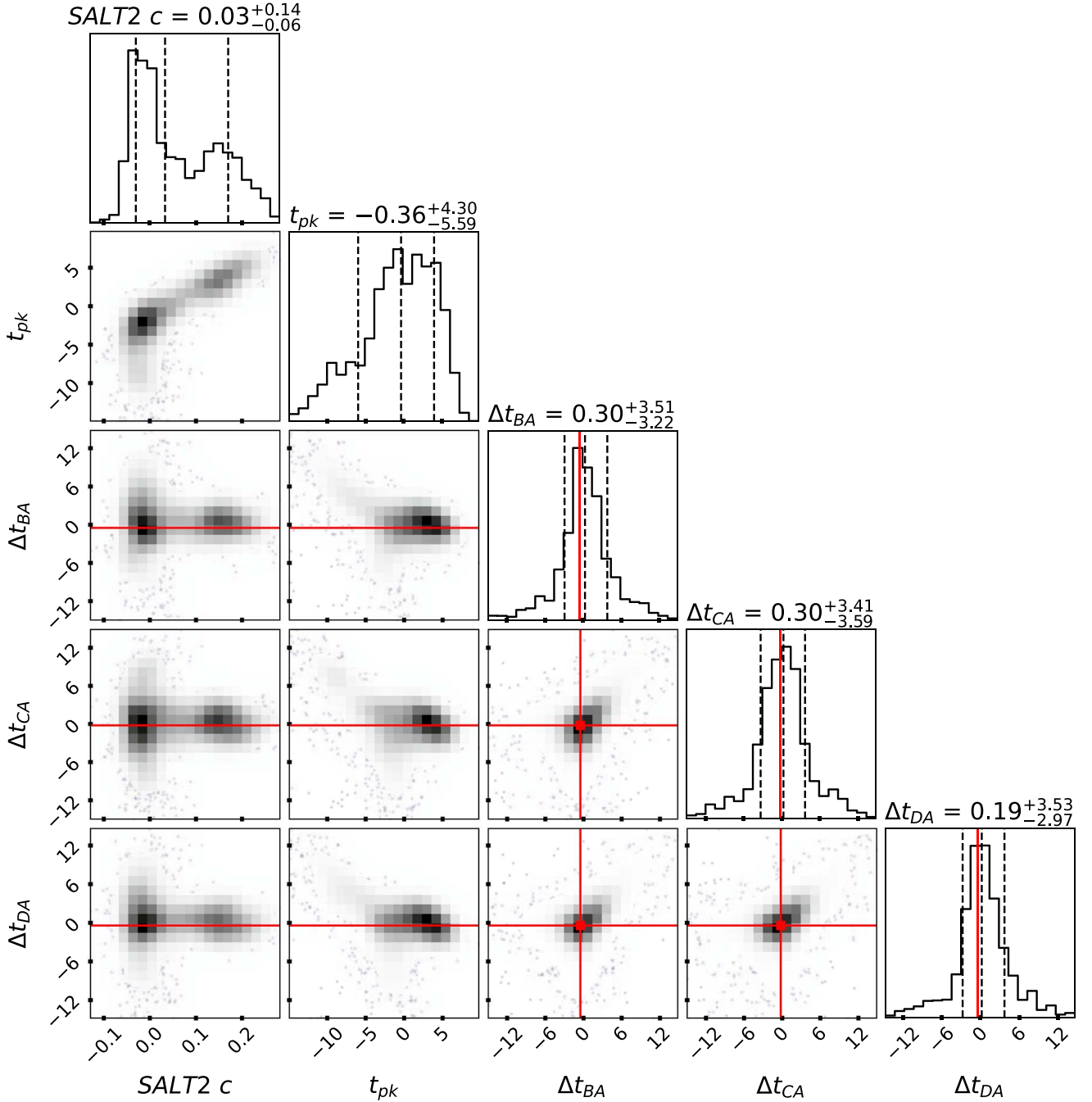


Figure 9. Posterior distributions of the SNTD Color method fitting of the HST photometry. The dashed vertical lines correspond to the distribution 16th, 50th, and 84th quantiles and the solid red lines show the final lens-model-predicted time delays and magnifications (those of A and C are off of the plot, see Section 3.1). The t_{pk} parameter is given relative to the G22 value of 59808.6.

5. Discussion

We have presented the first analysis of the LensWatch collaboration, which includes the only space-based observations of the gravitationally lensed and quadruply imaged SN Zwicky. The images are resolved with HST WFC3/UVIS (PSF FWHM $\sim 0''.07$) but not with WFC3/IR (PSF FWHM $\sim 0''.15$). We have measured photometry for each SN image in three optical HST filters, and we use the resulting colors (with an additional uncertainty due to chromatic microlensing) to infer time delays ($0.30^{+3.51}_{-3.22}$, $0.30^{+3.40}_{-3.59}$, and $0.19^{+3.53}_{-2.97}$ days, respectively, for images B–D relative to image A) and fluxes to infer

magnification ratios using the SNTD software package. Leveraging the fact that SN Zwicky is of Type Ia and therefore has a standardizable light curve, we apply a fiducial light-curve standardization and obtain absolute magnification estimates of $8.31^{+4.16}_{-1.43}$, $3.24^{+1.69}_{-0.57}$, $6.73^{+3.38}_{-1.16}$, and $3.39^{+1.65}_{-0.62}$ for images A–D, respectively.

We have also carried out an analysis of the lensing system using five distinct methodologies, of which we combine four primary methods to obtain our best constraints on the magnifications, time delays, and Einstein radius (θ_E) for the lensing system. We infer the smallest Einstein radius yet seen

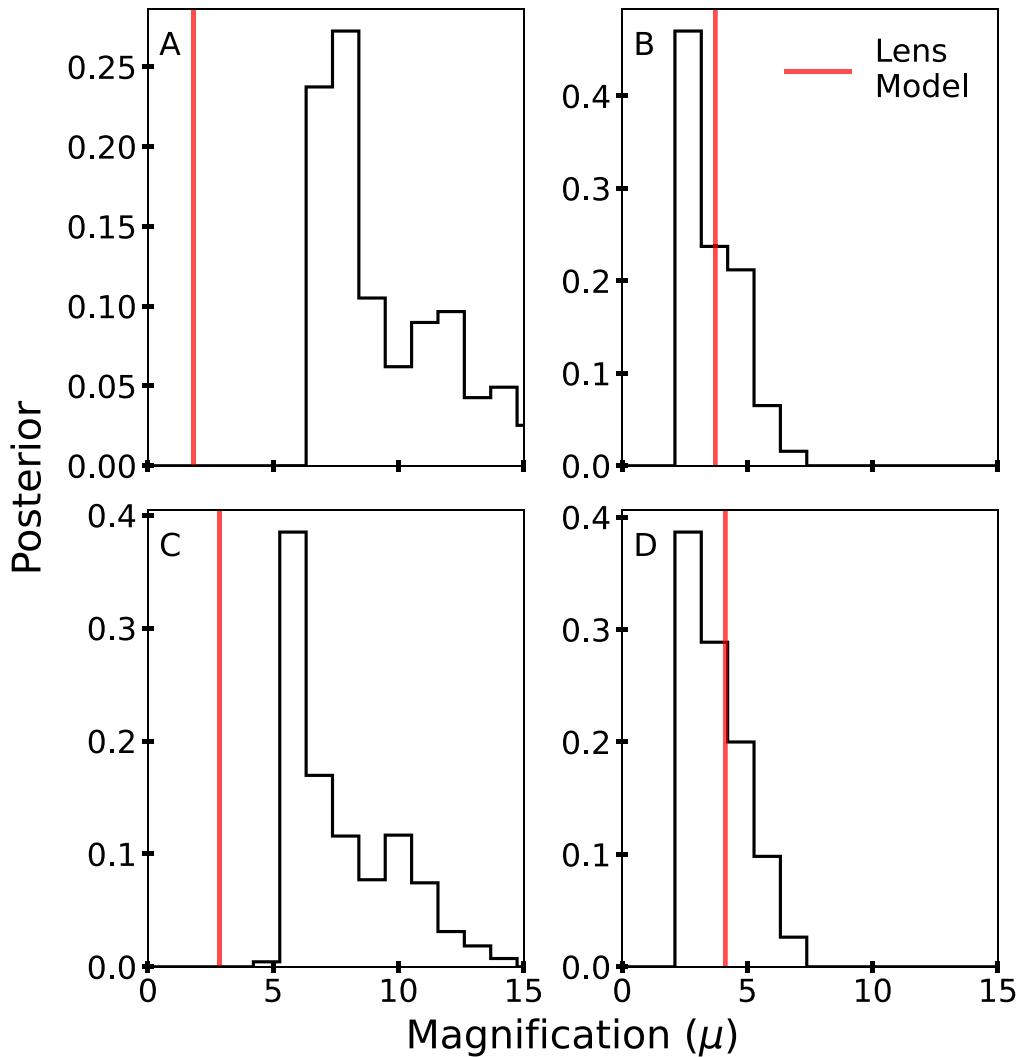


Figure 10. Posterior distributions of the magnification estimations (black), after fixing the light-curve and time-delay parameters from Figure 9. The solid red lines show the final lens-model-predicted time delays and magnifications.

in a lensed SN ($\theta_E = (0.168^{+0.009}_{-0.005})''$) and short time delays of $\lesssim 1$ day, consistent with our color curve fitting results and G22. For SN images B and D, we find consistent magnification predictions across all of our lens models that are in good agreement with the measured values. However, our lens models are unable to explain the observed fluxes for images A and C fully, and we see a significant discrepancy between the measured and predicted magnifications ($\sim 1.7^{+0.8}_{-0.6}$ mag and $\sim 0.9^{+0.8}_{-0.6}$ mag, respectively).

We resort to variations from microlensing and/or millilensing (Metcalf & Zhao 2002; Foxley-Marrable et al. 2018; Goldstein et al. 2018; Hsueh et al. 2018; Huber et al. 2019) to explain this inconsistency. We do see some evidence for differential dust extinction and/or chromatic microlensing across the four images of SN Zwicky with relative differences in the measured $F475W - F814W$ of up to $\sim 0.18 \pm 0.08$ mag, and based on the work of Goldstein et al. (2018) our HST epoch is of order three rest-frame weeks outside of the “achromatic phase” of SN Ia microlensing where the impact on optical colors is expected to be $\lesssim 0.1$ mag (95% confidence). We include the additional predicted uncertainty due to microlensing in our fitting. However, the F475W filter was

most discrepant with the SALT2 fitting, suggesting a possible systematic error in the photometry, though regardless this differential extinction is insufficient to explain the discrepancies we observe in images A and C relative to B (the largest color difference). We therefore estimate the additional (roughly achromatic) magnification to be ~ 1 mag in both images, which is significant but well within expectations for average galaxy-scale microlensing (Pierel et al. 2021) or millilensing (Metcalf & Zhao 2002). While it is suggested that saddle images are more likely to be demagnified by microlensing (e.g., Schechter & Wambsganss 2002), we expect the method of detection for SN Zwicky would be significantly biased toward microlensing events with high magnifications. A template epoch is already scheduled to occur after SN Zwicky will have faded, which will drastically simplify and improve the photometry and lens modeling processes and provide more stringent constraints.

Although the multiply imaged SN population is still small, this decade it is expected to grow by orders of magnitude with the Vera C. Rubin Observatory Legacy Survey of Space and Time (LSST; Ivezić et al. 2019) and Nancy Grace Roman Space Telescope (Pierel et al. 2021). Though both telescopes are expected to find a large number of spatially resolved lensing

systems, unresolved lensed SNe discovered with Rubin will still be common, and a dedicated follow-up campaign from space similar to LensWatch will be necessary to provide accurate photometry and lens modeling for such systems. SN Zwicky is the first lensed SN analysis presented by the LensWatch program, and is an excellent example of the coordination that will be required for upcoming lensed SN cosmology efforts. The unresolved, ground-based discovery with ZTF and subsequent follow-up with HST is a glimpse at likely future discoveries with Rubin and follow-up with Roman (or HST), a strategy that can be extremely fruitful for the field of gravitationally lensed SN cosmology.

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Appendix Individual Lens Modeling Results

Appendix A Modeling with GLEE

A.1. Lens Model Parameterization

We modeled SN Zwicky with an automated modeling pipeline that is based on the modeling software GLEE (Suyu & Halkola 2010; Suyu et al. 2012; Ertl et al. 2023), where we adopt the SPEMD (Barkana 1998) profile whose dimensionless surface mass density (or convergence) is given by

$$\kappa_{\text{SPEMD}}(x, y) = \frac{3 - \gamma}{2} \left[\frac{\theta_E}{\sqrt{q_m(x - x_m)^2 + \frac{(y - y_m)^2}{q_m}}} \right]^{\gamma-1}, \quad (\text{A1})$$

where (x_m, y_m) is the lens mass centroid, q_m is the axis ratio of the elliptical mass distribution, θ_E the Einstein radius, and γ is the power-law slope. The mass distribution is then rotated by the position angle ϕ_m , where an elliptical mass distribution with an angle of $\phi_m = 0$ corresponds to elongation along the y -axis (after converting to the conventional definition of the position angle). The external shear strength is described by $\gamma_{\text{ext}} = \sqrt{\gamma_{\text{ext},1}^2 + \gamma_{\text{ext},2}^2}$, with $\gamma_{\text{ext},1}$ and $\gamma_{\text{ext},2}$ the components of the shear. For a shear position angle $\phi_{\text{ext}} = 0$, the system is sheared along the x -direction.

A.2. Results Based on SN Image Positions

First, we model the light of the lens galaxy with two Sérsic profiles, and the light of multiple lensed SN images by fitting a PSF model constructed from multiple stars in the field of the drizzled data. Ertl et al. (2023) showed that for lensed quasars we can achieve astrometric accuracy of 2 mas from the surface brightness (SB) fit, by comparing the modeled image positions

Table 7
GLEE: Astrometry and Brightness of SN Images—Best-fit SN Image Positions and Amplitudes from SB Fitting

Image		F475	F625	F814
A	x''	1.688	1.691	1.686
	y''	1.781	1.783	1.783
	amplitude	9.96	29.11	62.60
B	x''	1.789	1.788	1.791
	y''	1.545	1.547	1.549
	amplitude	4.17	12.36	23.61
C	x''	1.575	1.577	1.576
	y''	1.504	1.504	1.500
	amplitude	9.08	27.39	51.27
D	x''	1.470	1.466	1.469
	y''	1.671	1.667	1.669
	amplitude	4.62	12.72	25.49

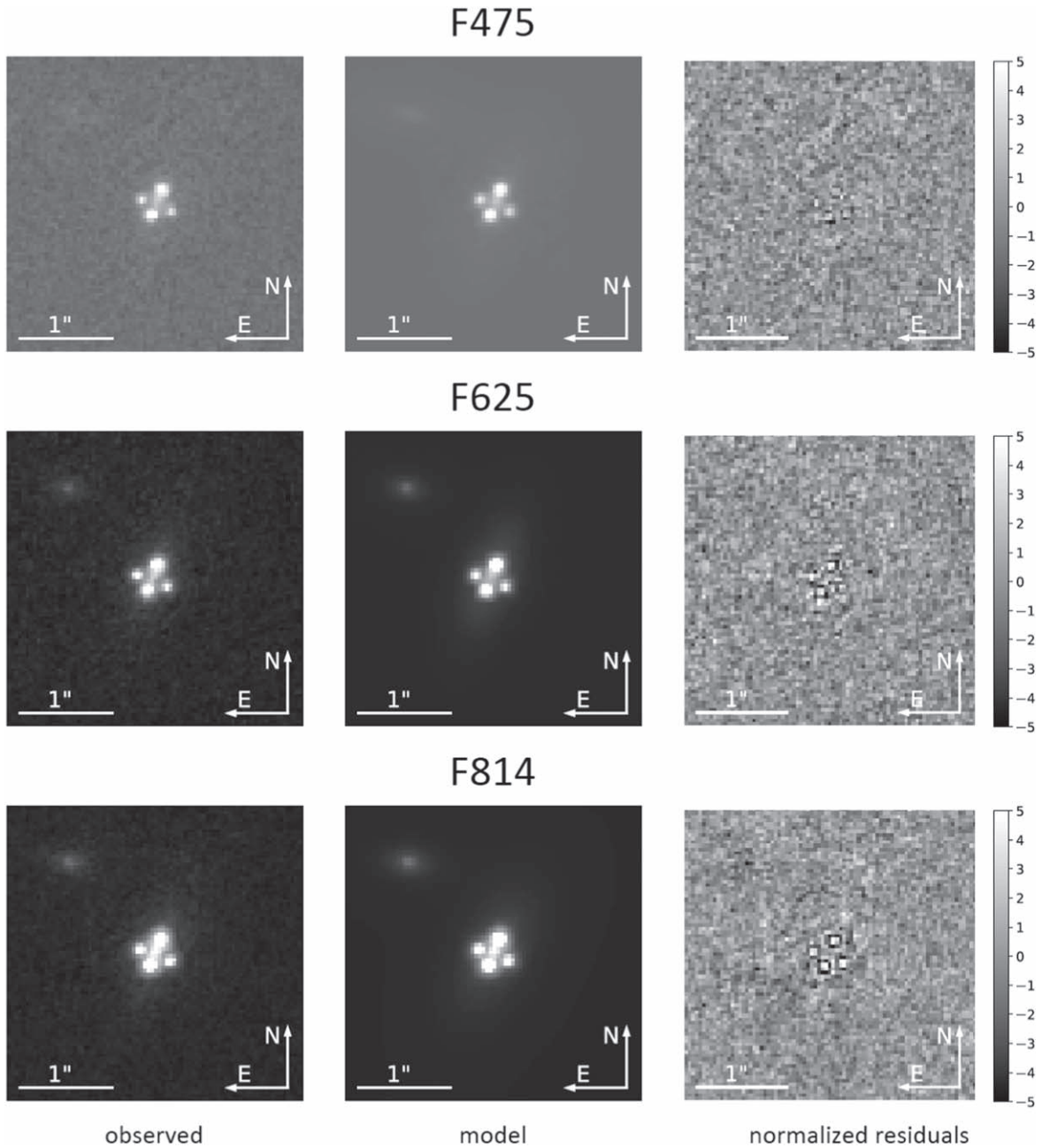


Figure 11. GLEE: SB fitting with GLEE in the three HST filters, as shown in the different rows. From left to right: observed image, model, and normalized residuals after modeling the light of the lens galaxy and the four SN images.

to those measured by the Gaia satellite. We use our SN image positions (from PSF fitting) to constrain the mass parameters, since we did not find (and do not expect) any substantial lensed arc light (from the SN host galaxy) in the modeling residuals of the three UVIS bands. We show the results of our SB fit in Figure 11. The measured astrometric positions of the four SN images in all three modeled bands are summarized in Table 7, and the light properties (based on the first and second brightness moments of the modeled lens light distribution that is a combination of the two Sérsic profiles) in the F814W filter is in Table 8. The positions in the F475W and F625W bands are aligned with the F814W coordinate frame.

For each band, we use the image positions reported in Table 7 and adopt an uncertainty on the image positions of 4 mas to constrain the lens mass parameters. The 4 mas

Table 8

GLEE: Centroid, Axis Ratio, and Position Angle of the Lens Light Computed from the Second Brightness Moments of the Two Sérsic Profiles in Our Best-fit Model of the F814W Filter

Parameter	Description	
x_s (")	x -centroid	1.641
y_s (")	y -centroid	1.650
q_s	axis ratio	0.52
ϕ_s (deg)	position angle	155

uncertainty is an estimate based on the astrometric accuracy of 2 mas (Ertl et al. 2023) and to account for substructure lensing, which can perturb the image positions at the few milliarcseconds level, as shown by Chen et al. (2007). We impose

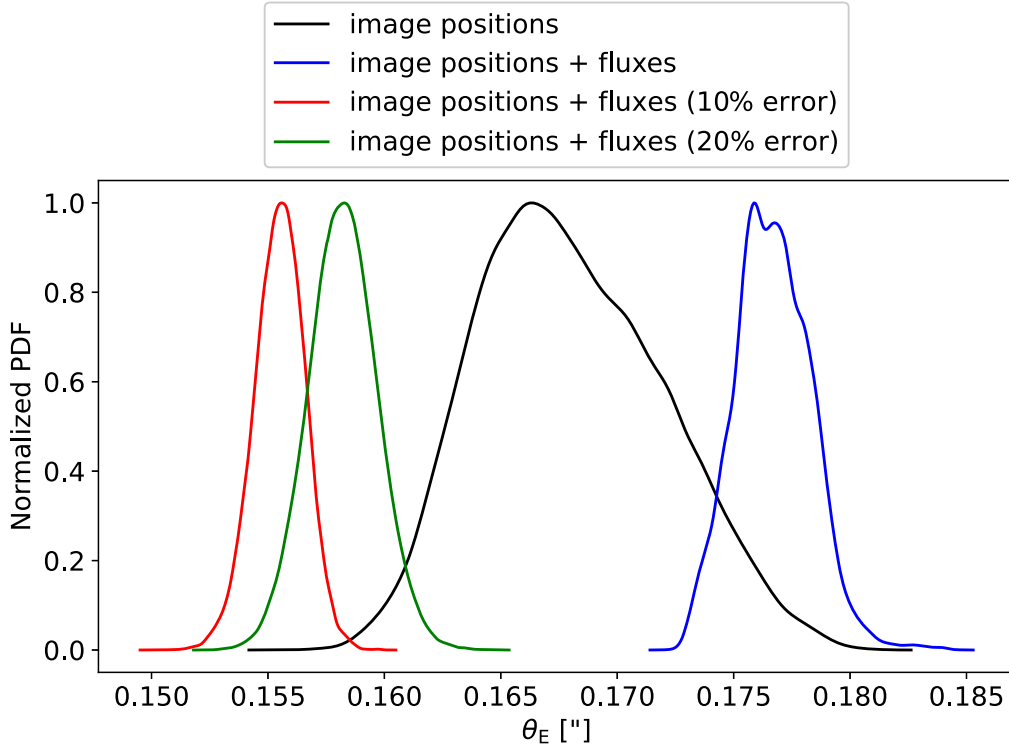


Figure 12. Exploration of the impact of flux constraints on the lens models with GLEE: distribution of Einstein radii for the four different model classes. The flux constraints are based on the measured SN image amplitudes in Table 7, with typical uncertainties of $\sim 1\%$, unless boosted to 10% or 20% as indicated in the legend.

Table 9

GLEE: Modeled Mass and Shear Parameters in the F814W Band

Parameter	Description	
$x_m(\prime\prime)$	x -centroid	$1.645^{+0.002}_{-0.002}$
$y_m(\prime\prime)$	y -centroid	$1.657^{+0.006}_{-0.004}$
q_m	axis ratio	$0.47^{+0.18}_{-0.13}$
$\phi_m(\text{deg})$	position angle	159^{+2}_{-2}
$\theta_E(\prime\prime)$	Einstein radius	$0.168^{+0.005}_{-0.004}$
γ	power-law index	$2.04^{+0.14}_{-0.22}$
γ_{ext}	shear strength	$0.02^{+0.03}_{-0.01}$
$\phi_{\text{ext}}(\text{deg})$	shear position angle	52^{+37}_{-57}

Note. We present the median and 1σ uncertainties. Position angles are reported as east of north.

uniform priors on all eight lens mass parameters that are tabulated in the leftmost column of Table 9. We report the mass model results from the F814W band because we achieved the lowest image-position χ^2 in this band. The results in this band are consistent with those of the other two bands, and also with the results of our light fit of the lens galaxy. We do not combine the constraints from all three bands in one single model, because positional uncertainties due to astrometric perturbations from, e.g., substructures in the mass distributions, are dominant, and the measured SN image positions of the individual bands are thus not completely independent. Our final lens mass and shear parameters are presented in Table 9. Comparing to the modeled lens light in Table 8, the lens galaxy mass profile agrees well with the light in terms of centroid, axis ratio, and position angle. In Table 10, we present the

Table 10

GLEE: Convergence κ and Total Shear Strength γ_{tot} at the (Modeled) Image Positions

Image	κ	γ_{tot}
A	$0.75^{+0.15}_{-0.15}$	$0.95^{+0.26}_{-0.24}$
B	$0.26^{+0.13}_{-0.08}$	$0.43^{+0.04}_{-0.06}$
C	$0.65^{+0.09}_{-0.12}$	$0.81^{+0.18}_{-0.18}$
D	$0.27^{+0.13}_{-0.08}$	$0.45^{+0.05}_{-0.07}$

convergence κ and total shear strength γ_{tot} at the (modeled) image positions.

A.3. Impact on the Results due to Flux Constraints

To investigate the bias to higher magnification for images A and C, we include fluxes, which were obtained from the light fit to the SN images (listed in Table 7), in our model. The flux amplitude had typical uncertainties of $\sim 1\%$. We find that models based on image positions and fluxes try to fit the fluxes at the expense of the poorer image-position recovery, so the model cannot fit well to both positions and fluxes.

The image positions are close to a critical curve, so small shifts lead to large changes in magnification. This is especially evident for the model where we use flux uncertainties from the SB fit. Imposing higher flux uncertainties (either 10% or 20% of the modeled flux value) leads to a lower image-position χ^2 and a higher magnification χ^2 and brings the modeling results closer to the models where we used only image positions.

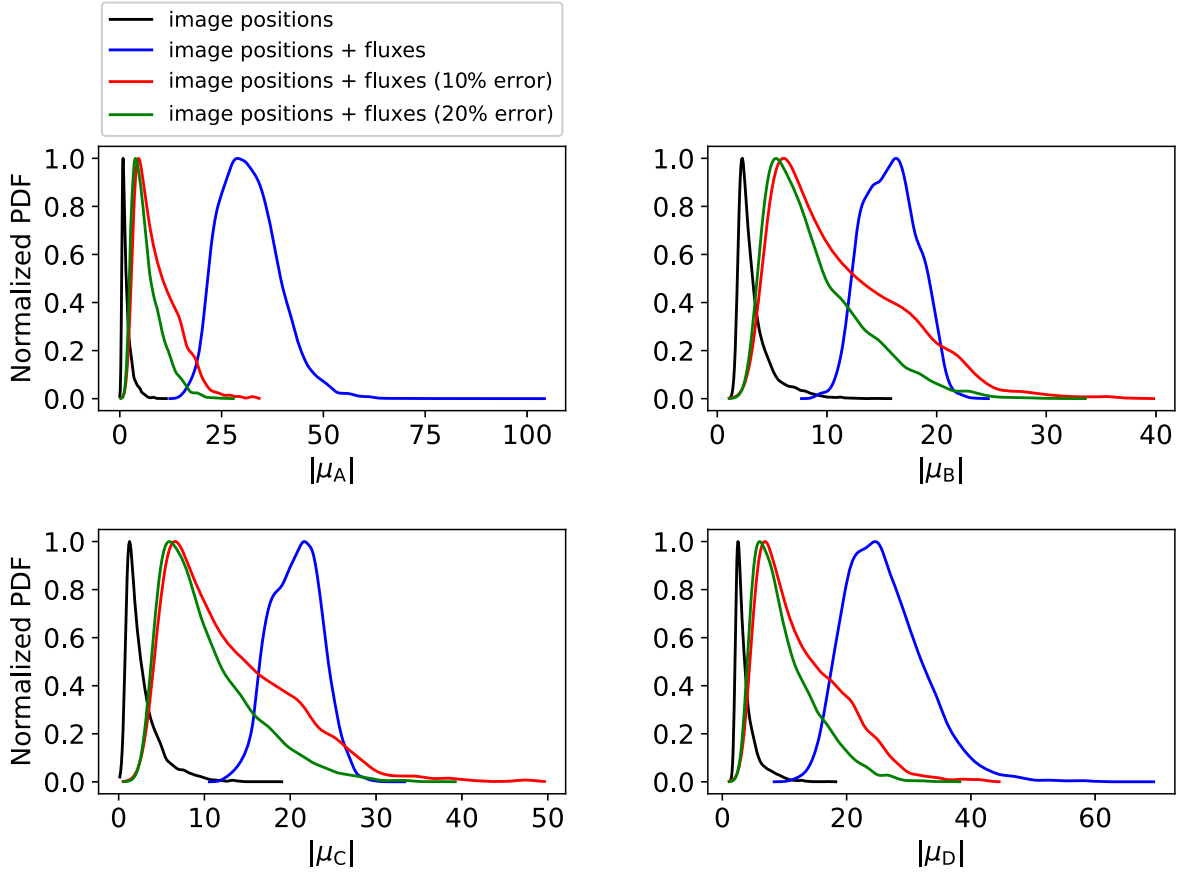


Figure 13. Similar to Figure 12 from GLEE but for the SN image magnifications from the four different model classes.

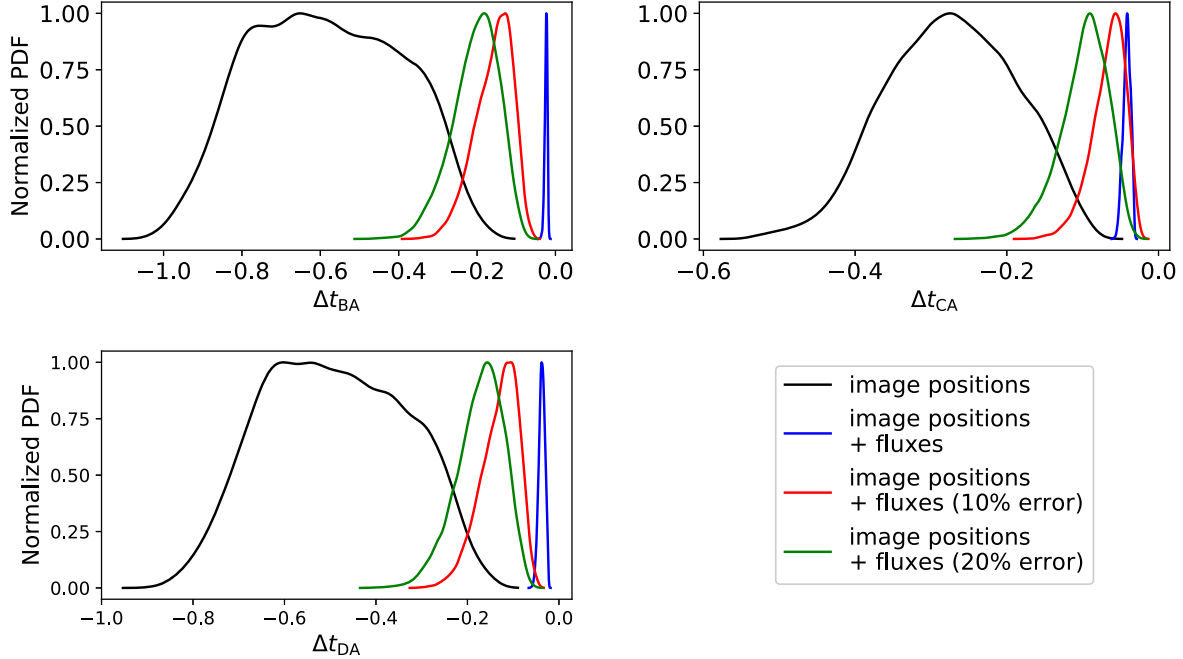


Figure 14. Similar to Figure 12 from GLEE but for the distribution of predicted time delays from the four different model classes.

We show the impact of including fluxes in our model by plotting the distribution of θ_E in Figure 12, magnifications in Figure 13, and predicted time delays in Figure 14 for the four different model classes. Since the models with flux constraints

do not fit well to both the image positions and fluxes (with $\chi^2 \gtrsim 10$), these models result in underestimated mass parameter uncertainties, as indicated by the narrower distributions for the blue, red and green models in Figure 12.

Appendix B

Modeling with `lfit_gui`

`lfit_gui` is a lens modeling software introduced by Shu et al. (2016), which has been applied to about 150 strong-lens systems (Shu et al. 2016, 2017, 2018; Marques-Chaves et al. 2017; Wang et al. 2017; Marques-Chaves et al. 2020). In order to maintain an independent analysis, in the `lfit_gui` approach, the positions and fluxes of the four SN images in the three optical filter bands are independently measured by fitting a photometric model consisting of two concentric Sérsic components, four PSF components, and a constant component to the drizzled data downloaded from the Barbara A. Mikulski Archive for Space Telescopes Portal.⁴⁰ The PSF models constructed in the GLEE approach are used. These photometric fitting results are shown in Figure 15. Overall speaking, the photometric model considered is able to reproduce the main structures in the data. Some residuals are seen at the lensed SN positions, which are primarily caused by PSF mismatches. The measured positions and fluxes of the four SN images are summarized in Table 11. The positional uncertainties are clearly correlated with the signal-to-noise ratios (S/Ns) of the lensed SN images. As a result, they are the smallest in F814W (≈ 1 mas) and the highest in F475W (≈ 4 mas). In general, the measured positions of the four lensed SN images agree well across the three bands. The largest differences are seen in the relative x positions of images C and D between F475W and F625W, which are about 1.8σ . The measured photometry for the four lensed SN images are found to be systematically brighter than the measurements in Table 4. The differences are typically within 0.15 mag in F625W and F814W and become 0.3–0.6 mag in F475W. We think this is likely related to the different treatments of the lens galaxy light. It affects photometry in the F475W the most because the brightness contrast between the lensed SN images and the lens galaxy is the smallest.

In terms of lens modeling, the `lfit_gui` approach considered an SIE lens model, the convergence of which follows the profile defined in Equation (A1) but with γ fixed to 2, and used the measured positions of the four SN images to constrain the five SIE parameters (as well as the source position) in the three bands separately. The sampling was done

using EnsembleSampler from the `emcee` package assuming uniform priors with sufficiently wide ranges for all the seven free parameters. The maximum a posteriori (MAP) estimation and marginalized posterior distribution for the three key SIE parameters, i.e., Einstein radius, axis ratio, and position angle, are reported in Table 12, and the posterior probability density distributions (PDFs) are provided in Figure 16. As shown in Figure 17, this lens model well reproduces the four lensed SN positions. The rms of the differences between the predicted and observed image positions is $0''.002$, $0''.0004$, and $0''.0001$ in F475W, F625W, and F814W, respectively. The tightest constraints on the lens model parameters are obtained in F814W, which have the smallest positional uncertainties, and the posterior PDFs are the broadest in F475W. Nevertheless, the lens model parameters are generally consistent within 1σ . We find a clear anticorrelation between the Einstein radius and axis ratio (Figure 16), which is also observed in other lens modeling methods (e.g., Figures 21 and 23).

We use the Bayesian information criterion (BIC) to combine the results from the three bands, which is defined as

$$\text{BIC} = p \ln(n) - 2 \ln(\hat{L}), \quad (\text{B1})$$

where p is the number of free parameters (i.e., seven), n is the number of constraints (i.e., eight), and \hat{L} is the maximum likelihood of a model. The BIC values are 15.9757, 14.9495, and 14.6745 in F475W, F625W, and F814W, respectively. Weighting the results in F475W and F625W relative to F814W as

$$\text{weight}_{\text{F475W/F625W}} = \exp(-0.5 \times (\text{BIC}_{\text{F475W/F625W}} - \text{BIC}_{\text{F814W}})), \quad (\text{B2})$$

the combined results suggest that the Einstein radius is $0.1664^{+0.0010}_{-0.0019}$ arcsec, the axis ratio is $0.632^{+0.027}_{-0.019}$, and the position angle is $69.0^{+0.3}_{-0.3}$ degrees (i.e., the angle between the major axis of the lens surface mass density distribution and the x -axis, measured counterclockwise). The total lensing mass (within the ellipse that corresponds to $\kappa = 1$) is thus estimated to be $7.47^{+0.09}_{-0.17} \times 10^9 M_{\odot}$. The predicted magnifications, time delays, and convergence/shear values for the four lensed SN images are reported in Table 13 (and also in Table 3). We note

Table 11

Relative Positions (with Respect to Image A That Has the Highest S/N in All Three Bands and Thus the Smallest Positional Uncertainties) and AB Magnitudes Measured by the `lfit_gui` Approach in F475W, F625W, and F814W, Respectively

		A	B	C	D
F475W	Δx	0	0.1039 ± 0.0049	-0.1059 ± 0.0033	-0.2105 ± 0.0047
	Δy	0	-0.2351 ± 0.0049	-0.2785 ± 0.0031	-0.1123 ± 0.0044
	m_{AB}	22.9 ± 0.1	23.8 ± 0.1	22.9 ± 0.1	23.7 ± 0.1
F625W	Δx	0	0.1011 ± 0.0018	-0.1121 ± 0.0012	-0.2187 ± 0.0018
	Δy	0	-0.2360 ± 0.0017	-0.2817 ± 0.0012	-0.1149 ± 0.0017
	m_{AB}	21.6 ± 0.1	22.5 ± 0.1	21.7 ± 0.1	22.5 ± 0.1
F814W	Δx	0	0.1044 ± 0.0011	-0.1110 ± 0.0008	-0.2204 ± 0.0012
	Δy	0	-0.2345 ± 0.0012	-0.2816 ± 0.0009	-0.1137 ± 0.0012
	m_{AB}	20.5 ± 0.1	21.5 ± 0.1	20.7 ± 0.1	21.4 ± 0.1

Note. A constant 0.1 mag uncertainty is assumed.

⁴⁰ The .drc files from <https://mast.stsci.edu/portal/Mashup/Clients/Mast/Portal.html>.

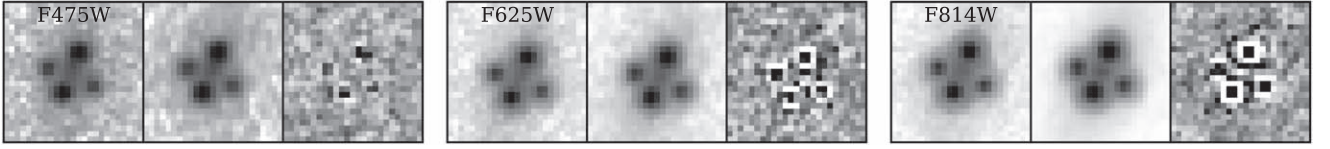


Figure 15. Photometric fitting results in F475W (left), F625W (middle), and F814W (right) in the `lfit_gui` approach (every cutout is approximately $1''$ by $1''$). In each subpanel, the data, best-fit model, and residuals are shown from left to right.

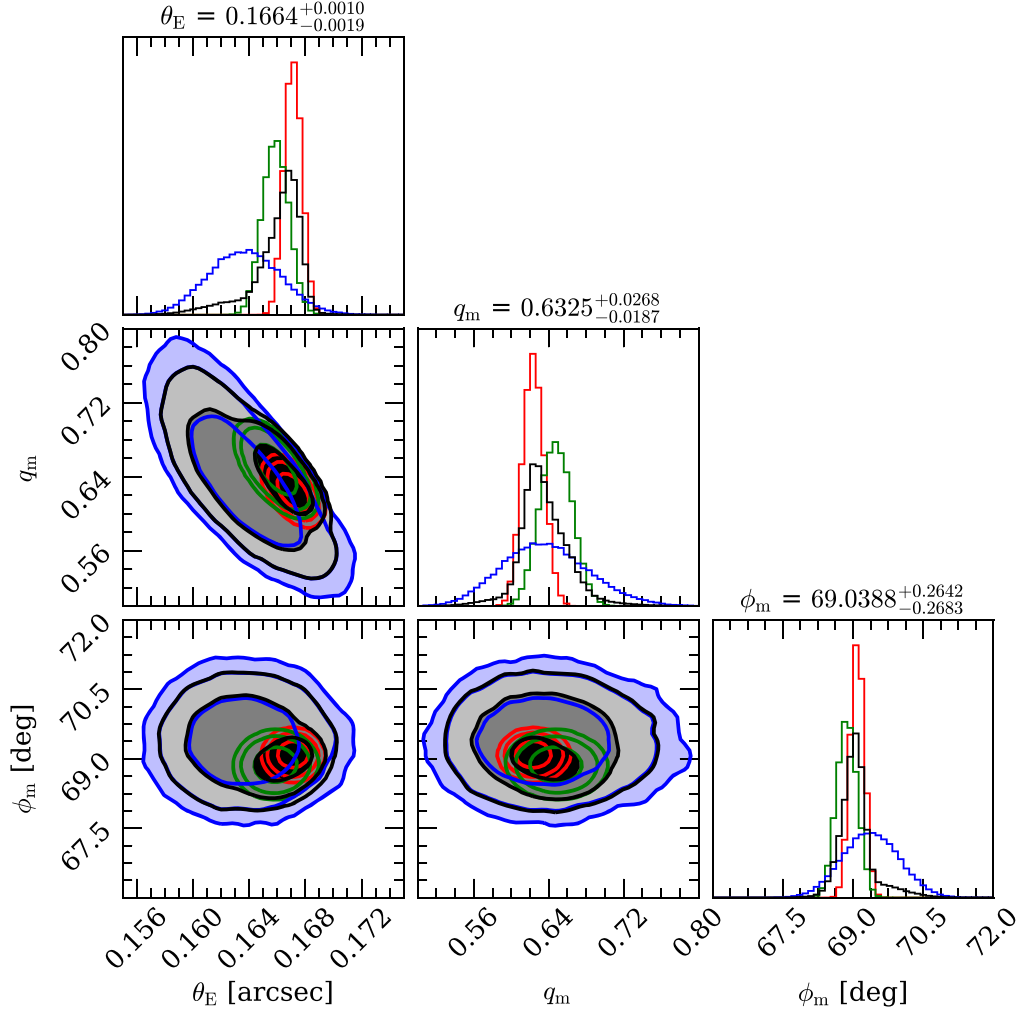


Figure 16. Posterior PDFs for the key SIE model parameter in the `lfit_gui` approach. The blue, green, and red contours and histograms correspond to results inferred from the F475W, F625W, and F814W data, respectively, and the black contours and histograms correspond to the combined results.

that, strictly speaking, the BIC values can only be compared and combined when different models are constrained by the same data set, which does not apply to our results from the

three different bands. Nevertheless, the adopted weighting scheme is equivalent to weighting by the likelihoods, which is still a sensible treatment.

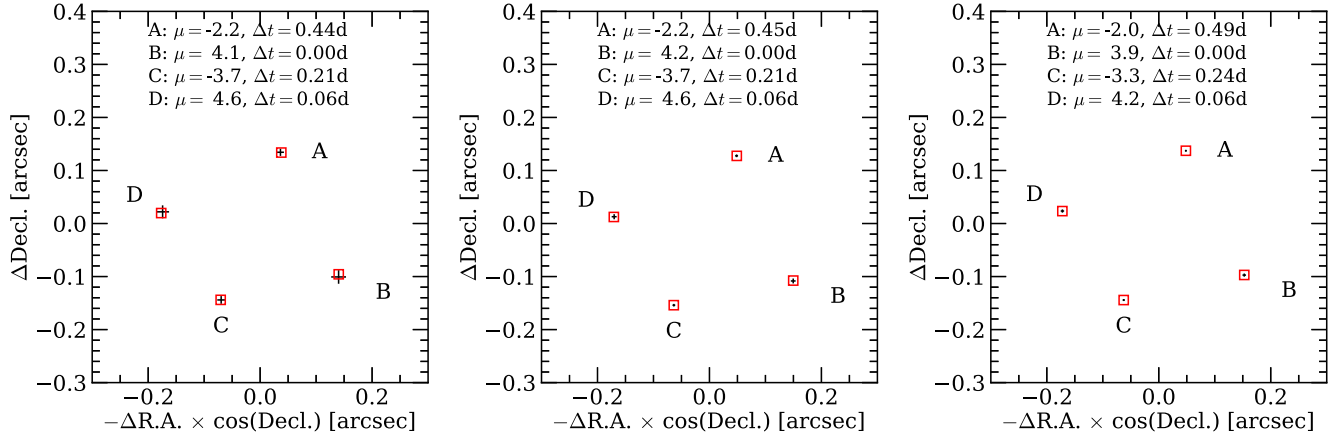


Figure 17. Modeling results using the MAP solution in F475W (left), F625W (middle), and F814W (right) in the `lfitt_gui` approach. The black symbols with error bars (3σ) indicate the measured SN positions, and the red squares correspond to the predicted image positions. The predicted magnifications and relative time delays are also shown.

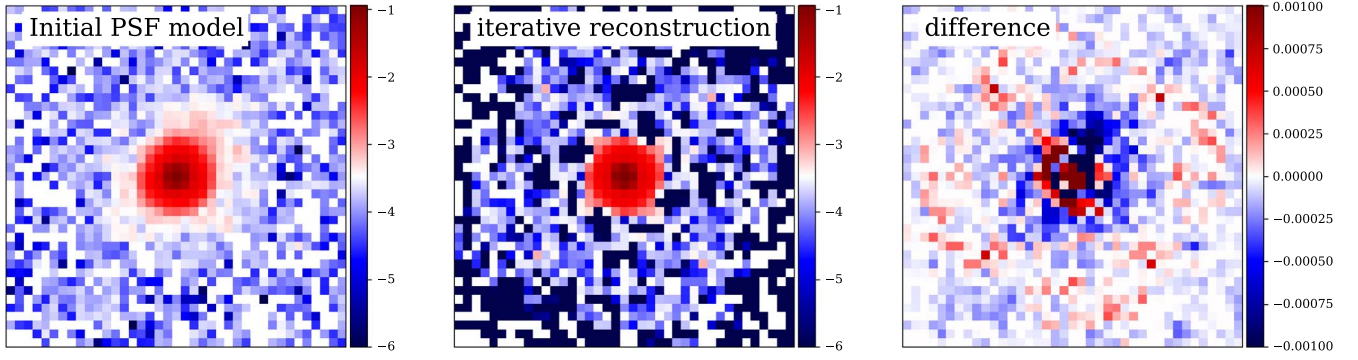


Figure 18. The initial PSF model for the F814W band (left panel), the reconstruction after iterative improvement by LENSTRONOMY (middle panel), and the difference between the two (right panel).

Table 12
Key SIE Model Parameters from MAP Estimation and Marginalized Posterior Distributions in the `lfitt_gui` Approach

Parameter	MAP			Marginalization			
	F475W	F625W	F814W	F475W	F625W	F814W	Combined
θ_E (arcsec)	0.1632	0.1658	0.1671	$0.164^{+0.002}_{-0.002}$	$0.1659^{+0.0009}_{-0.0009}$	$0.1671^{+0.0006}_{-0.0006}$	$0.1664^{+0.0010}_{-0.0019}$
q_m	0.647	0.649	0.624	$0.64^{+0.05}_{-0.04}$	$0.648^{+0.018}_{-0.017}$	$0.623^{+0.012}_{-0.011}$	$0.632^{+0.027}_{-0.019}$
ϕ_m (deg)	69.4	68.9	69.1	$69.4^{+0.6}_{-0.6}$	$68.9^{+0.2}_{-0.2}$	$69.1^{+0.2}_{-0.2}$	$69.0^{+0.3}_{-0.3}$

Appendix C

Modeling with LENSTRONOMY

LENSTRONOMY is a multipurpose, open-source, community-lead, ASTROPY-affiliated gravitational lensing and image modeling package (Birrer & Amara 2018; Birrer et al. 2021).⁴¹ LENSTRONOMY supports a large variety of lens models and SB profiles, as well as multiple numerical options to treat point sources. The modularity of LENSTRONOMY supports imaging modeling as well as catalog-based model fitting. LENSTRONOMY has been applied for time-delay cosmography of lensed quasars (Birrer et al. 2019; Shajib et al. 2020) and lensed SNe as well as a variety of other lens modeling and image analysis applications (Gilman et al. 2020; Shajib et al. 2021; Schmidt et al. 2022).

C.1. The LS1 Method

We described the mass profile of the lens galaxy by an SIE (Kormann et al. 1994), where the convergence is given by

$$\kappa(x, y) = \frac{\theta_E}{2\sqrt{q_m(x - x_m)^2 + \frac{(y - y_m)^2}{q_m}}}. \quad (C1)$$

Here, θ_E , q_m , and (x_m, y_m) are defined similarly as for Equation (A1). In order to fit the lens galaxy light profile, we stacked two Sérsic profiles (De Vaucouleurs 1948; Sérsic 1963)

$$I(R) = I_e \exp \left\{ -b_n \left[\left(\frac{R}{R_e} \right)^{1/n} - 1 \right] \right\}, \quad (C2)$$

where I_e is the intensity at the half-light radius R_e . The constant b_n is equal to $1.9992n - 0.3271$ (Birrer & Amara 2018), and

⁴¹ <https://github.com/lenstronomy/lenstronomy>

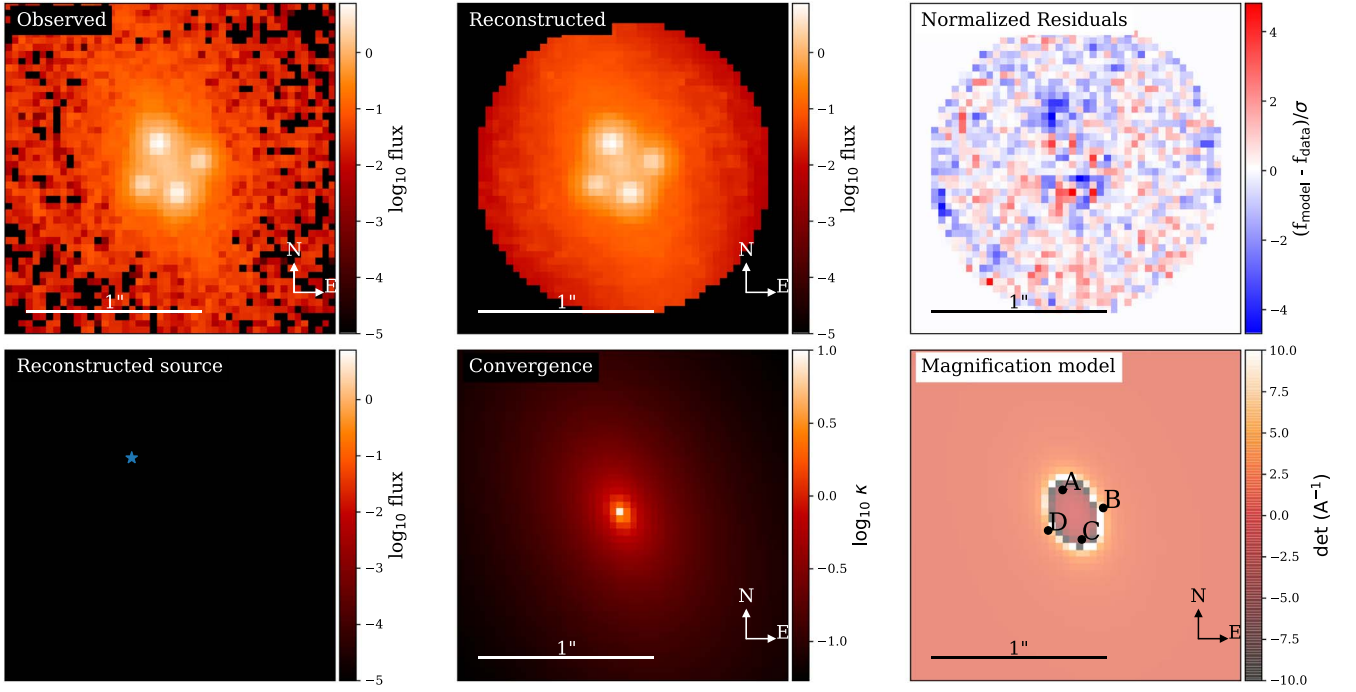


Figure 19. The observed HST F814W-band data compared to the reconstructed model. Upper panel from left to right: the observed image, the reconstructed light intensity, and the normalized residuals. The circular mask illustrates the image region used in computing the likelihood. Lower panel from left to right: the reconstructed source, convergence (projected surface mass density), and magnification model with the four SN images. In the reconstructed source panel, the blue star marks the unlensed position of the SN, and no light from the host galaxy is detected above the noise level.

Table 13

Predicted Magnifications, Relative Time Delays, and κ/γ Values for the Four Lensed SN Images by the `lfitt_gui` Approach (Using the MAP Estimation for Individual Bands and BIC-weighted Average for the Combined Result)

Image	Magnification				Time Delay (day)				κ/γ Combined
	F475W	F625W	F814W	Combined	F475W	F625W	F814W	Combined	
A	-2.2	-2.2	-2.0	$-2.0^{+0.2}_{-0.3}$	0.44	0.45	0.49	$0.48^{+0.03}_{-0.05}$	$0.74^{+0.02}_{-0.03}$
B	-4.1	-4.2	-3.9	$4.0^{+0.3}_{-0.2}$	0.00	0.00	0.00	0.00	$0.374^{+0.009}_{-0.006}$
C	-3.7	-3.7	-3.3	$-3.5^{+0.2}_{-0.4}$	0.21	0.21	0.24	$0.23^{+0.02}_{-0.02}$	$0.644^{+0.011}_{-0.016}$
D	-4.6	-4.6	-4.2	$4.4^{+0.4}_{-0.2}$	0.06	0.06	0.06	$0.06^{+0.01}_{-0.01}$	$0.385^{+0.009}_{-0.006}$

Note. For the SIE model, the κ and γ values are always the same.

$R \equiv \sqrt{q_s x^2 + y^2/q_s}$ with q_s being the axis ratio of the Sérsic profile. The SN images were fitted as point sources on the image planes with a PSF model. We initiated the model fitting with the PSF model constructed by the GLEE team and then further improved the PSF model using a built-in feature in LENSTRONOMY’s that minimizes the residuals between the observed and reconstructed image around the SN positions (Shajib et al. 2019). The comparison between the initial PSF model in the F814W band and the final reconstructed one is illustrated in Figure 18. Additionally, we adopted a circular region around the lensing system for likelihood computation to avoid the boundary effect of the PSF convolution in the evaluated likelihood function.

We fitted the pixel-level data from the three optical HST bands in a joint likelihood. The uncertainties on the model parameters were obtained from a Markov Chain Monte Carlo (MCMC) sampling. The flux ratios of the SN images were not included in our lens model, because they failed to provide a good fit to the data and increased the reduced χ^2 from 1.17 to 5.59 (for the F814W filter). The reconstructed image model,

source, convergence, and magnification model using the best-fit parameters from the converged MCMC chain are shown in Figure 19. Our measured κ and γ parameters at each image location are given in Table 14.

C.2. The LS2 Method

The “LS2” team used the catalog-data modeling functionality of the LENSTRONOMY software package, using the positions and positional uncertainties and redshifts reported in this paper. We adopted an elliptical power-law mass profile (Tessore & Metcalf 2015) plus external shear, and allowed all parameters to vary. Models were computed for each of the three WFC3/UVIS filters, F475W, F625W, and F814W. The MCMC parameter sampling for F475W is shown in Figure 20, the posterior computed model parameters for the same filter are shown in Figure 21, and the κ , γ results are given in Table 15.

C.3. The SALT+LS Method

As a proof of concept, in lens modeling we use the expected SN Ia brightness as prior with a broad standard deviation of

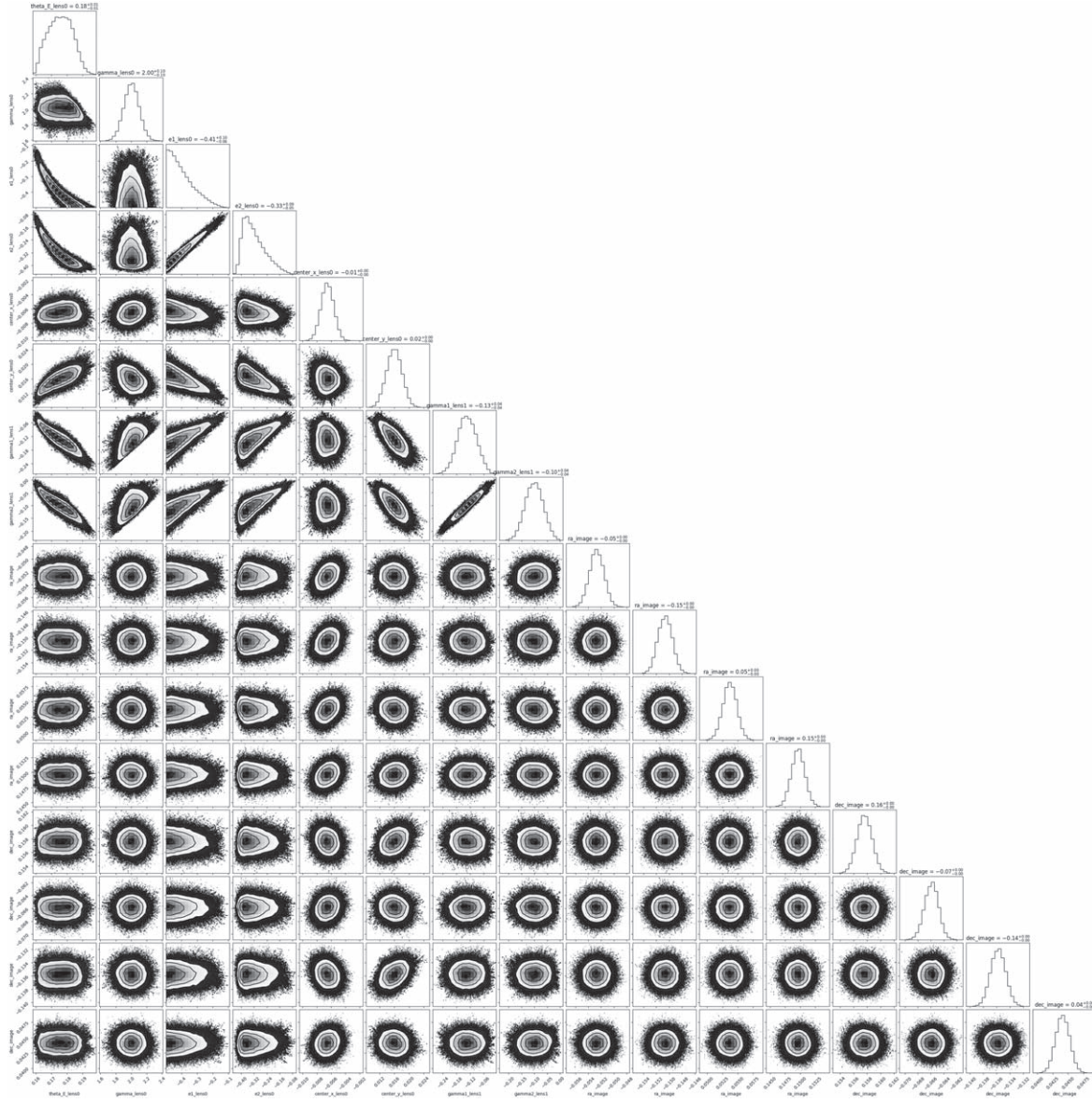


Figure 20. Parameter sampling for the LS2 LENSTRONOMY method, for the F475W filter position measurements.

0.3 mag. First, using SNCosmo (Barbary et al. 2016), we fit SALT2.4 for the publicly available ZTF photometric data in the *g* and *r* bands. We obtain $x_1 = 1.11 \pm 0.43$, $c = -0.071 \pm 0.029$, with maximum light at $\text{MJD} = 59808.54 \pm 0.43$ (Figure 22). These values are in good agreement with the best-fit SALT parameters from G22. Note that G22 also used data from the Liverpool Telescope, which provided additional observations in the *griz* bands. To standardize the SN Ia brightness, we adopt a peak *B* magnitude of $M_B = -19.05$, $\alpha = -0.141$, and $\beta = 3.101$ (Betoule et al. 2014).

We combine the three HST optical bands by averaging the best-fit PSF centroids and adding the fluxes. We assume a flat Λ CDM universe with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_M = 0.3$. Our model consists of an elliptical power law (EPL) main lens and external shear. Our loss function combines the summed squares of the difference in the delensed image positions and summed squares of the difference between the observed and the model-predicted fluxes. The relative weight between the flux and position terms is adjusted to achieve the best overall fit.

The results are summarized in Figure 23. We find a somewhat steep mass profile slope for the lensing galaxy: $\gamma_{\text{EPL}} = 2.50$. As G22 pointed out, with such a small θ_E , this system is in a regime of lensing galaxies that have seldom been studied before. Such systems can be used to probe the density profile at subkiloparsec scales within the lensing galaxy core. The image positions and model predictions agree to better than $0''.023$ (Figure 24). The magnifications from this model for images B and D (Table 3) are in fairly good agreement with the corresponding expectations in Table 6, whereas for images A and C, the model predictions are lower by $<2\sigma$ (see also Figure 24). It appears that without taking into account microlensing and/or differential dust extinction, this is the best compromise the model can achieve. The total predicted magnification is 17.73. Compared with the expectation of 24.3 ± 2.7 from G22, it is smaller by 2.5σ .

We find that without using the SN Ia brightness as prior, it is possible to find models with acceptable predictions for both image positions and flux values, but they tend to have a much shallower slope ($\gamma_{\text{EPL}} \lesssim 1.75$). In contrast, using the SN Ia



Figure 21. A full corner diagram plot for the computed posterior estimates for the LS2 lens modeling fit, for the F475W filter position measurements.

Table 14
Best-fit κ/γ Values for the LS1 Lens Model

Image	κ/γ^a
A	$0.703^{+0.011}_{-0.009}$
B	$0.387^{+0.003}_{-0.004}$
C	$0.625^{+0.005}_{-0.005}$
D	$0.397^{+0.003}_{-0.003}$

Note.

^a Note that $\kappa = \gamma$ for an SIE model.

brightness prior, we find γ_{EPL} to be consistently $\gtrsim 2.5$, whether we use single-band data or combine the different bands. We also note that if the host galaxy identification in G22 is correct, this system is possibly in a unique situation in that the core of the host galaxy is not multiply imaged. This makes the modeling of this system especially challenging: we cannot

separately perform lens modeling using the lensed host galaxy in contrast to the other small- θ_E lensed SNe Ia (Dhawan et al. 2020).

We now briefly compare the SALT+LS model with the other four models presented in this paper that only use image positions. With regard to θ_E , when fluxes are taken into consideration by the GLEE team, they have found acceptable models with θ_E in agreement with the SALT+LS best-fit value. We further note that the time-delay predictions from the SALT+LS model are in good agreement with those from the GLEE model that has a similar θ_E . With regard to magnification: (1) the total magnification from these four models ranges from 9.2 to 15.7 (Table 3). The SALT+LS model predicts a magnification of 17.43, higher than these four models; and (2) whereas the SALT+LS model predicts the magnifications for the brighter two images, A and C, to be higher than those of B and D, the other four models predict the opposite. And yet, as mentioned before, even the total magnification from the SALT+LS model is $\sim 30\%$ lower than the expected magnification

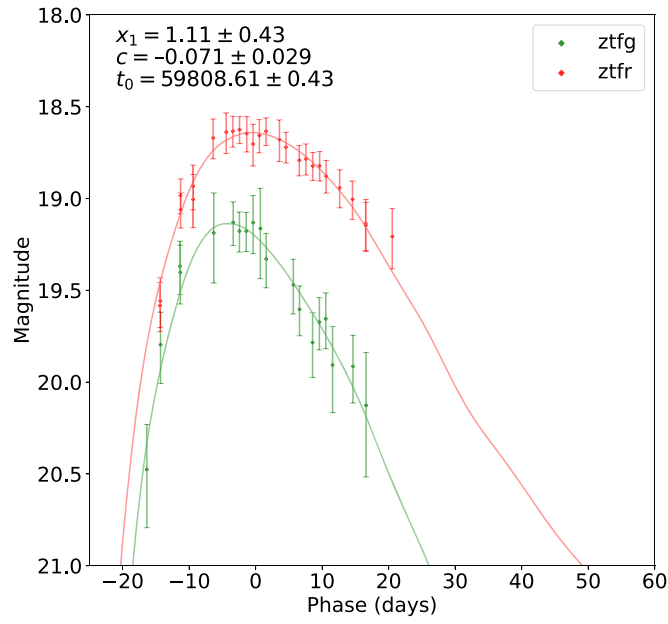


Figure 22. ZTF *g* and *r* band light curves and best-fit SALT2.4 parameters.

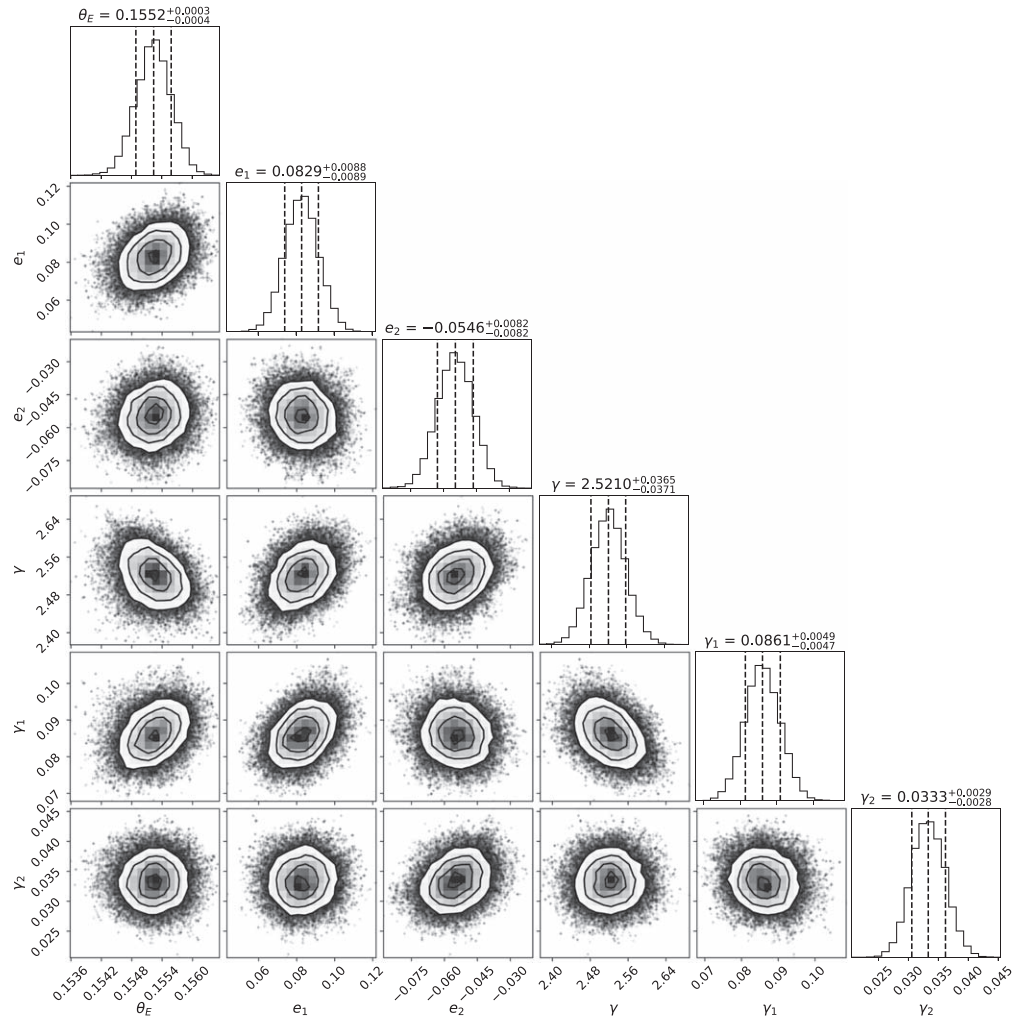


Figure 23. A corner plot of the posterior samples for the parameters of the SALT+LS lens model.

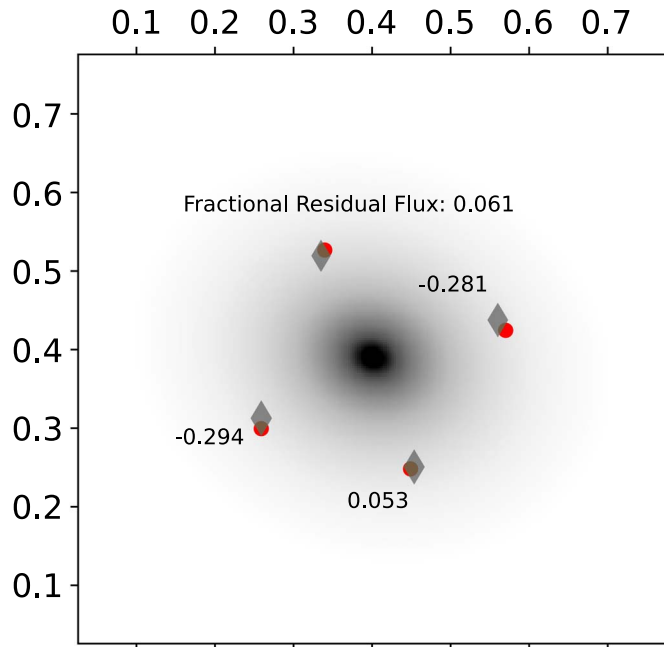


Figure 24. The SALT+LS lens model prediction for image positions and fluxes. The red dots are the observed positions of the images and the gray diamonds are the model-predicted positions. The number next to each image position is the fractional error of the model-predicted flux. The tick marks are in units of arcseconds. The gray-scale background shows the convergence (κ) map. This image is shown in a slightly different orientation relative to most of the figures in this paper, but is the same as Figure 19. The brightest image, A, is still at the very top, and the other images are in the same alphabetical order clockwise from this reference point.

Table 15
Best-fit κ/γ Values for the LS2 Lens Model

Image	κ	γ
A	$0.95^{+0.24}_{-0.20}$	$0.88^{+0.21}_{-0.16}$
B	$0.31^{+0.07}_{-0.06}$	$0.43^{+0.04}_{-0.04}$
C	$0.82^{+0.17}_{-0.15}$	$0.74^{+0.14}_{-0.11}$
D	$0.32^{+0.07}_{-0.06}$	$0.43^{+0.04}_{-0.04}$

from G22 or from Table 6 based on SN Ia brightness. The (κ , γ) values at the locations of the images are, in order of A–D: (0.25, 0.86), (0.22, 0.56), (0.26, 0.88), and (0.21, 0.60). Given that this appears to be a fairly normal SN Ia, it is possible that microlensing and/or differential dust extinction (likely to be small, given the small color differences for the four images shown in Table 4) have played a significant role. If so, for the SALT+LS model, the optimization can be distorted in a way to compensate for these effects. Thus, for the uncertainties for κ and γ , we report the largest uncertainties for the four images, which are 0.024 and 0.023, respectively (Table 16). Once follow-up HST observations are completed after the SN has faded and improved photometry has been obtained, we will revisit this model.

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Table 16
Best-fit κ/γ Values for the SALT+LS Lens Model

Image	κ	γ
A	0.25 ± 0.024	0.86 ± 0.023
B	0.22 ± 0.024	0.56 ± 0.023
C	0.26 ± 0.024	0.88 ± 0.023
D	0.21 ± 0.024	0.60 ± 0.023

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