



Low-efficiency long gamma-ray bursts: a case study with AT2020blt

N. Sarin¹,^{2,3,4,5} R. Hamburg,^{5,6} E. Burns,⁷ G. Ashton^{1,2,8}, P. D. Lasky^{1,2} and G. P. Lamb⁹

¹*School of Physics and Astronomy, Monash University, Vic 3800, Australia*

²*OzGrav: The A.R.C. Centre of Excellence for Gravitational Wave Discovery, Clayton VIC 3800, Australia*

³*Nordita, KTH Royal Institute of Technology and Stockholm University Roslagstullsbacken 23, SE-106 91 Stockholm, Sweden*

⁴*The Oskar Klein Centre, Department of Physics, Stockholm University, AlbaNova, SE-106 91 Stockholm, Sweden*

⁵*Department of Space Science, University of Alabama in Huntsville, Huntsville, AL 35899, USA*

⁶*Center for Space Plasma and Aeronomic Research, University of Alabama in Huntsville, Huntsville, AL 35899, USA*

⁷*Department of Physics and Astronomy, Louisiana State University, Baton Rouge, LA 70803, USA*

⁸*Royal Holloway University of London, Egham, Surrey, TW20 0EX, UK*

⁹*School of Physics and Astronomy, University of Leicester, University Road, LE1 7RH, UK*

Accepted 2022 March 2. Received 2022 February 24; in original form 2021 October 8

ABSTRACT

The Zwicky Transient Facility recently announced the detection of an optical transient AT2020blt at redshift $z = 2.9$, consistent with the afterglow of an on-axis gamma-ray burst. However, no prompt emission was observed. We analyse AT2020blt with detailed models, showing the data are best explained as the afterglow of an on-axis long gamma-ray burst, ruling out other hypotheses such as a cocoon and a low-Lorentz factor jet. We search *Fermi* data for prompt emission, setting deeper upper limits on the prompt emission than in the original detection paper. Together with *KONUS-Wind* observations, we show that the gamma-ray efficiency of AT2020blt is $\lesssim 0.3$ – 4.5 per cent. We speculate that AT2020blt and AT2021any belong to the low-efficiency tail of long gamma-ray burst distributions that are beginning to be readily observed due to the capabilities of new observatories like the Zwicky Transient Facility.

Key words: gamma-ray bursts.

1 INTRODUCTION

The interaction of an ultra-relativistic jet launched in a gamma-ray burst with the surrounding interstellar medium is known to produce broad-band synchrotron radiation referred to as an afterglow. Observed extensively in X-rays, optical, and radio, these phenomena are observationally confirmed to be linked to the collapse of massive stars and compact object mergers (e.g. Abbott et al. 2017; Cano et al. 2017).

Afterglows of gamma-ray bursts are predominantly observed following-up the prompt emission trigger. There are a few exceptional cases such as iPTF11gg (Cenko et al. 2012), iPTF14yb (Cenko et al. 2015), FIRSTJ141918 (Marcote et al. 2019), and AT2020blt (Ho et al. 2020). These transients were either independently detected from the prompt emission and later associated to a gamma-ray counterpart that were missed in low-latency (Cenko et al. 2015), initially believed to be observed off-axis, in which case the high-energy prompt emission was missed due to relativistic beaming (Rhoads 1997; Huang, Dai & Lu 2002; Nakar, Piran & Granot 2002), or later shown to not be afterglows at all (Lee et al. 2020).

The transient AT2020blt (Ho et al. 2020) is a fast-fading optical transient at $z = 2.9$ observed by the Zwicky Transient Facility without any high-energy trigger from gamma-ray satellites. This transient has characteristic features akin to afterglows from gamma-ray bursts and broad-band observations in X-rays, optical, and radio (Ho et al. 2020;

Singer et al. 2020). However, the non-detection of prompt gamma-ray emission is puzzling. Ho et al. (2020) suggest that this non-detection may be due to one of three reasons: (1) The prompt emission expected in gamma-rays was weak or missed due to occultation by the Earth. (2) The transient was the afterglow from an off-axis gamma-ray burst where the prompt gamma-rays were missed because of relativistic beaming. (3) AT2020blt was the afterglow of a dirty fireball, i.e. a gamma-ray burst with a low Lorentz factor ($\Gamma \lesssim 100$), such that the optically-thick environment absorbed the gamma-ray photons.

Prompt gamma-ray emission observed for both long and short gamma-ray bursts is likely produced by internal dissipation in an ultra-relativistic jet (e.g. Kumar 1999). To produce a gamma-ray burst, the jet must be ultra-relativistic to avoided the compactness problem (Ruderman 1975), i.e. that gamma-ray photons are above the pair production threshold and should only be observable if the jet is moving ultra-relativistically. This theoretical constraint has led to the placement of lower limits on the initial bulk Lorentz factor of $\Gamma_0 \gtrsim 100$ (e.g. Lithwick & Sari 2001). The existence of relativistic jets following gamma-ray bursts has also been shown observationally, for e.g. through multiwavelength observations of GRB 170817A (e.g. Kasliwal et al. 2017; Lamb, Mandel & Resmi 2018; Mooley et al. 2018; Ghirlanda et al. 2019; Matsumoto, Nakar & Piran 2019; Beniamini, Granot & Gill 2020).

As an ultra-relativistic jet passes through the stellar/ejecta envelope, it creates a cocoon of shocked material. If the jet stalls within this envelope then such a *choked* jet will dissipate energy into the surrounding bubble of matter, forming a quasi-spherical *cocoon* that will produce minimal prompt gamma-ray emission (e.g.

* E-mail: nsarin.astro@gmail.com

Gottlieb et al. 2018). The interaction of the jet and/or cocoon with the interstellar medium is ultimately responsible for producing the broad-band afterglow we see following almost all gamma-ray bursts.

The broad-band observations of GRB 170817A confirmed that a relativistic jet successfully broke out of the ejecta and that the jet was likely structured i.e. the energy and Lorentz factor of the jet had some angular dependence. Although the exact jet structure is unknown, various phenomenological models such as a Gaussian or power-law structure can successfully explain the broad-band afterglow observations of GRB 170817A (e.g. Troja et al. 2017; Lamb et al. 2018; Lazzati et al. 2018; Gill & Granot 2018; Kathirgamaraju, Barniol Duran & Giannios 2018; Ryan et al. 2020). We note that although a jet without any angular dependence is possible (e.g. Aloy, Janka & Müller 2005), it is unlikely as both the jet-launching mechanism and jet-breakout will likely produce some jet structure (e.g. Nakar 2020).

In this paper, we investigate why there was no detection of prompt gamma-rays from AT2020btl by analysing the multiwavelength data from Ho et al. (2020) with detailed afterglow models. In Section 2, we use the multiwavelength observations and physical arguments to estimate when the associated gamma-ray burst happened. Using this estimated time, we fit the data with a structured-jet model and cocoon model. We introduce our structured-jet and cocoon models in Section 3 and perform Bayesian model selection to identify the more likely scenario. We explore the dirty fireball hypothesis by estimating the Lorentz factor in Section 4. We perform a sub-threshold search in *Fermi* data and discuss the efficiency of the unobserved prompt gamma-ray emission in Section 5. We discuss the implications of our results and conclude in Section 6. Our analysis suggests that AT2020btl is likely an on-axis low efficiency long gamma-ray burst. The lack of gamma-ray radiation can be attributed to the low radiative efficiency of AT2020btl. We find that the non-detection in gamma rays with *KONUS-Wind* and *Fermi* implies AT2020btl has a radiative efficiency $\lesssim 0.3$ –4.5 per cent.

2 ESTIMATING THE BURST TIME

AT2020btl was first observed in the *r* band on 2020 January 28 by the Zwicky Transient Facility (Bellm et al. 2019). Follow-up observations provided detections in *g* and *i* bands along with detections in radio and X-rays (Ho et al. 2020; Singer et al. 2020). No coincident gamma-ray trigger was found (Ho et al. 2020; Ridnaia et al. 2020). The lack of a gamma-ray trigger implies that we do not know the burst time t_0 , which is critical for discerning physics and testing the various hypotheses in detail.

Ho et al. (2020) estimated the burst time by fitting a broken power-law simultaneously to the *r*, *g*, and *i* band data assuming constant colour offsets. This allowed them to estimate t_0 as January 28.18. However, by fitting for t_0 with a broken power-law, Ho et al. (2020) have estimated the peak time of an afterglow t_{peak} . For a typical on-axis relativistic jet, this peak time is likely not significantly different to t_0 . However, for a mildly relativistic jet viewed on-axis, t_{peak} can be up to a few days after t_0 (e.g. Sari & Piran 1999), while for an off-axis system, the peak time could be several months after t_0 (e.g. Granot et al. 2002).

To investigate the reason for the lack of observed prompt gamma-ray emission with detailed afterglow models, we must first estimate the burst time more robustly. For a relativistic outflow viewed on-axis the peak time-scale in optical is the deceleration time-scale (e.g. Sari & Piran 1999; Chevalier & Li 2000)

$$t_{\text{dec}} \approx 10 \text{ s } (1+z) \left(\frac{E_k}{10^{52} \text{ erg}} \frac{1}{A_*} \frac{1}{\Gamma_0^4} \right), \quad (1)$$

for a stellar wind medium, as expected for long gamma-ray bursts. Here, E_k is the kinetic energy of the outflow, z is the redshift, A_* is wind parameter typically of order unity (Chevalier & Li 2000), and Γ_0 is the initial Lorentz factor of the outflow. For an on-axis observer, the optical lightcurve is well modelled as a rising power-law till t_{dec} (e.g. Sari & Piran 1999).

The location of AT2020btl was observed prior to the first detection on January 27 with an *r* band upper limit of >21.36 mag. We use this non-detection combined with a rising power-law on the observed flux

$$F = A (t - t_0)^m \lesssim t_{\text{dec}}, \quad (2)$$

which describes the pre-deceleration physics, to estimate t_0 and t_{dec} . We use broad uninformative priors on t_0 , t_{dec} , and fix A to ensure the pre-deceleration power-law smoothly connects with the observations. If AT2020btl is observed on-axis, the pre-deceleration behaviour must explain the non-detection on January 27 and the afterglow flux at the start of the observations. The prior on m is informed by physics and depends on whether the forward or reverse shock dominates, or on the thickness of the shells launched in the burst. To minimize the effect of the prior, we choose a broad uniform prior from 0.5 to 20, which covers all possible scenarios ranging from a forward or a reverse shock and a combination (e.g. Zhang 2018) and the possibility of different stratification in the environment. We note that typical observed pre-deceleration behaviour follows $m \approx 3$ (e.g. Zhang 2018). We examine whether the pre-deceleration power law would produce a signal above the upper limit on January 27 and explain the peak flux at the time of the first detection, ruling out the parameter space that violates this constraint.

The above analysis implies that if AT2020btl is observed on-axis (i.e. the assumption of pre-deceleration is correct) then we measure $t_0 = 58875.1^{+0.6}_{-1.1}$ MJD (i.e. Jan 27.1 $^{+0.6}_{-1.1}$). This initial analysis can already offer some clues into the nature of AT2020btl. For shallow indices $m \lesssim 3$, as would be expected for a cocoon or a relativistic jet where the forward shock dominates the pre-deceleration physics, a start time before \sim January 27 can not explain the non-detection. While for steeper indices $m \gtrsim 3$ (i.e. where reverse shock emission dominates), the range of t_0 is significantly broader, potentially as early as \sim January 26 as the rapid rise can accommodate both the non-detection and the peak observations. We estimate the deceleration (peak) time to $t_{\text{dec}} = 58876.1^{+0.1}_{-0.2}$ MJD i.e. Jan 28.1 $^{+0.1}_{-0.2}$, ≈ 5 h before the first observation. Our posterior on t_0 and m is shown in Fig. 1.

The above deceleration analysis can also be extended to an off-axis observer. We emphasize that this latter physics is a consequence of relativistic beaming and not deceleration. A relativistic jet viewed off-axis will also rise with $m \gtrsim 3$. However, if the jet has an extended structure such that a significant fraction of the jet energy covers the observers' line of sight, the rise will be much shallower (e.g. Nakar 2020). This suggests that if AT2020btl was an off-axis afterglow, the observer must satisfy $\theta_{\text{obs}} > 1/\Gamma$ (e.g. Nakar 2020) at the time of the non-detection or for the gamma-ray burst to occur after the non-detection on January 27. These two conditions combined with the steep late-time decay of AT2020btl suggests an off-axis interpretation is unfeasible. We note that the above analysis may not be valid for complex jet structures where there are two peaks in the lightcurve due to deceleration of material along the line of sight or other effects due to jet structure that produce a negative temporal slope even with relativistic beaming (Beniamini et al. 2020). We explore the off-axis interpretation in more detail in Appendix A.

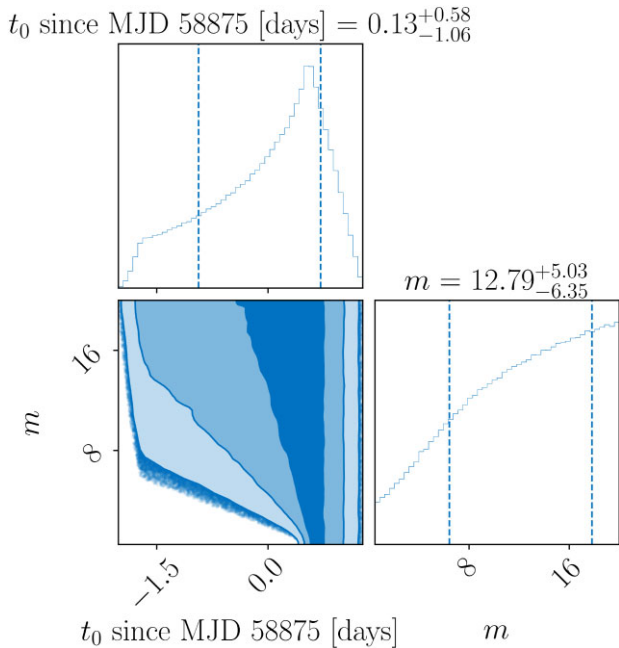


Figure 1. Posterior distribution for the burst time t_0 since MJD 58875 i.e. 2020 January 27 and m the pre-deceleration index. The uncertainties are the 68 per cent credible interval. The shaded contours indicate the 1σ – 3σ credible intervals. We note that MJD 58875 is arbitrarily chosen to show the posterior distribution more clearly.

3 AFTERGLOW CONSTRAINTS

If AT2020blt is the afterglow of a typical gamma-ray burst, the absence of observed gamma-ray emission immediately points towards two hypotheses: either the observer was off-axis, and therefore the prompt gamma-ray emission was missed due to relativistic beaming (e.g. Granot et al. 2002), or the jet did not successfully break through the ejecta and cocoon emission was responsible for producing the afterglow (e.g. Nakar & Piran 2017). We first test both these hypotheses by fitting the broad-band afterglow to a power-law structured jet and cocoon model. We note that it is also possible that the gamma-ray emission was missed, and we explore this hypothesis later in Section 5.

The power-law jet model is an angular structured jet with an energy distribution defined as

$$E(\theta_{\text{observer}}) = E_{\text{iso}} \left(1 + \frac{\theta_{\text{observer}}^2}{\beta \theta_{\text{core}}^2} \right)^{-\beta/2}. \quad (3)$$

Here, β is the exponent dictating the slope of the power-law jet structure, θ_{observer} is the observers viewing angle, θ_{core} is the half-width opening angle of the jet core. The Lorentz factor of the jet is proportional to $E(\theta_{\text{observer}})^{1/2}$ and E_{iso} is the on-axis isotropic equivalent energy. The cocoon model is a spherical outflow with velocity stratification with an energy distribution,

$$E(u) = E_0 \left(\frac{u}{u_{\text{max}}} \right)^{-k}. \quad (4)$$

Here, u is the dimensionless four-velocity, E_0 is the kinetic energy of the fastest material, and k is the power-law index. Both the cocoon and structured jet outflows interact with the surrounding interstellar medium accelerating a fraction of electrons, ξ_n , with a fraction of the total energy, ϵ_e , and the fraction of the energy in the magnetic field,

ϵ_b . The radiation produced by these electrons is responsible for the observed broad-band afterglow.

We fit the multiwavelength flux density data of AT2020blt (including the upper-limits) using the power-law structured jet and cocoon model described above implemented in AFTERGLOWPY (Ryan et al. 2020) and a Gaussian likelihood. We set broad priors on all 15 parameters for each models. We include both synchrotron and inverse Compton emission and account for potential host galaxy extinction. We estimate the uncertainty on each data point as the quadrature sum of the measurement uncertainty reported by Ho et al. (2020) and an estimated systematic uncertainty which we model. Our prior on this systematic uncertainty is uniform from 10^{-4} to 10^{-3} mJy. The values being motivated by the flux data. We infer the parameters of the system using BILBY (Ashton et al. 2019) and the DYNesty sampler (Speagle 2020). Our fits to the broad-band data are shown in Fig. 2. The full list of parameters and their respective priors for the structured-jet model are shown in Table B1, while Fig. B1 shows the posterior distribution. We note that we do not fit the X-ray data as it is computationally expensive to simultaneously fit the flux density and integrated flux data with our model. However, after fitting the flux density data we ensure that our result is consistent with the X-ray observations. Excluding X-ray data from our fit may also be necessary as AFTERGLOWPY does not incorporate synchrotron self-Compton emission which may lead to biased estimation of parameters if included when fitting the multiwavelength data (e.g. Nakar, Ando & Sari 2009).

We find that the cocoon model does not explain the r -band data at late times well, while the structured-jet model can successfully explain all the observations. In particular, the last data point is a 2σ outlier from the posterior prediction for the cocoon model. Furthermore, given the pre-deceleration power-law exponent of a cocoon is $m \lesssim 3$, to explain the observations with a cocoon the associated gamma-ray burst must have occurred after the non-detection on January 27. This requires the cocoon emission to rise rapidly to explain the observations on January 28, which is difficult (e.g. Nakar & Piran 2017). The rapid decay of AT2020blt is also difficult to expect from a cocoon. Ho et al. (2020) measure the temporal decay of AT2020blt at late-times as ~ -2.6 , and such a rapid decay is difficult to expect with a cocoon for a broad range of parameters (Lamb et al. 2018; Troja et al. 2018).

More quantitatively, we perform Bayesian model selection between the two hypotheses. Assuming both models are equally likely *a priori*, the structured-jet model is ~ 4 times more likely than the cocoon interpretation, favouring the hypothesis that an ultra-relativistic structured jet broke out of the ejecta and later interacted with the surrounding environment to produce AT2020blt.

Although we focus here on the power-law structured jet model for simplicity, we find a similar preference for other jet models, including a Gaussian structured jet and a top-hat jet with Bayes factors of ~ 4 and ~ 6 in favour of the jet hypothesis, respectively. We emphasize that these Bayes factors are predicated on the assumption that AT2020blt was a gamma-ray burst i.e. we are comparing the cocoon and structured jet hypotheses given AT2020blt is a gamma-ray burst. In Section 6, we speculate that it is possible, although unlikely, that AT2020blt may have a different origin, in which case this analysis would not be valid.

As we are also modelling the noise, another way to compare the model fit is by the size of the estimated noise. We find $\sigma = 7.7 \times 10^{-4}$ and $\sigma = 8.2 \times 10^{-4}$ for the power-law structured jet and cocoon models, respectively, i.e. to fit the data with the cocoon model, we need the data to be noisier, suggesting that the structured-jet model is favoured over the cocoon.

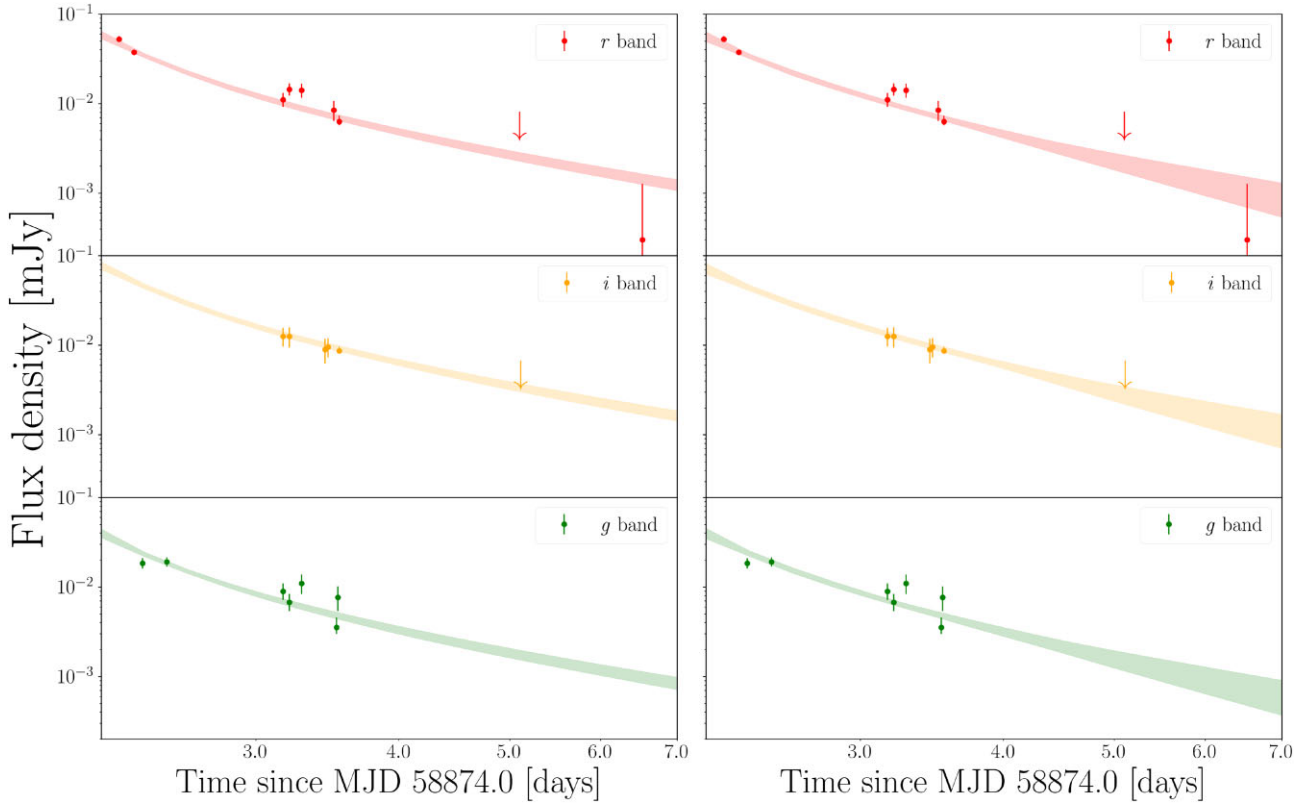


Figure 2. Data for AT2020blt in *r*, *g*, and *i* band, with arrows indicating upper limits. In the left-hand panel we show the 95 per cent credible interval predicted lightcurves from our posterior samples for the cocoon model. Similarly, the right-hand panels show the 95 per cent credible interval predicted lightcurves for the power-law jet model. The errors on the data are the measurement uncertainty reported by Ho et al. (2020) combined with our modelled uncertainty for each model. We note that MJD 58874 is arbitrarily chosen to show the posterior distribution more clearly. We also note that only *r* band data is shown here but the fits are performed on the full multiwavelength data.

The inability to explain the observations with a cocoon while successfully explaining the observations with a structured jet suggests AT2020blt had a successful jet breakout, as typical for most observed gamma-ray bursts. Therefore, the lack of observed prompt gamma-ray emission could be a consequence of relativistic beaming, i.e. that we observed AT2020blt outside the ultra-relativistic core. We explore this hypothesis in detail in Appendix A finding that we can not fit the data well, with a Bayes factor of $\sim 3 \times 10^{64}$ in favour of the on-axis hypothesis. This is likely due to the sharp rise required to explain the non-detection and first observation around January 28 while also explaining the subsequent steep decay thereafter, which is difficult to expect for an observer located off-axis (e.g. Granot et al. 2002; Nakar 2020).

In Fig. 3, we show the one and two-dimensional posterior distributions for the observers viewing angle θ_{observer} , the half-width jet core θ_{core} , the isotropic equivalent energy E_{iso} , and the ambient interstellar medium density n_{ism} . Our measurement of the ultra-relativistic core of the jet is $\theta_{\text{core}} = 8^\circ \pm 2^\circ$, while $\theta_{\text{observer}} = 3^\circ_{-2^\circ}^{+3^\circ}$, implying that we observed AT2020blt on-axis much like most other gamma-ray bursts. The credible intervals are 68 per cent unless specified otherwise. We also measure the on-axis isotropic equivalent energy $\log_{10}(E_{\text{iso}}/\text{erg}) = 53.6^{+0.3}_{-0.4}$, and the ambient interstellar medium density $\log_{10}(n_{\text{ism}}/\text{cm}^{-3}) = 2^{+1}_{-2}$ consistent with the population of other gamma-ray bursts (e.g. Nava et al. 2014; Beniamini et al. 2015; Fong et al. 2015; Wang et al. 2015). Our constraint on the fraction of energy in electrons $\log_{10}\epsilon_e = -1.1 \pm 0.3$ is also consistent with

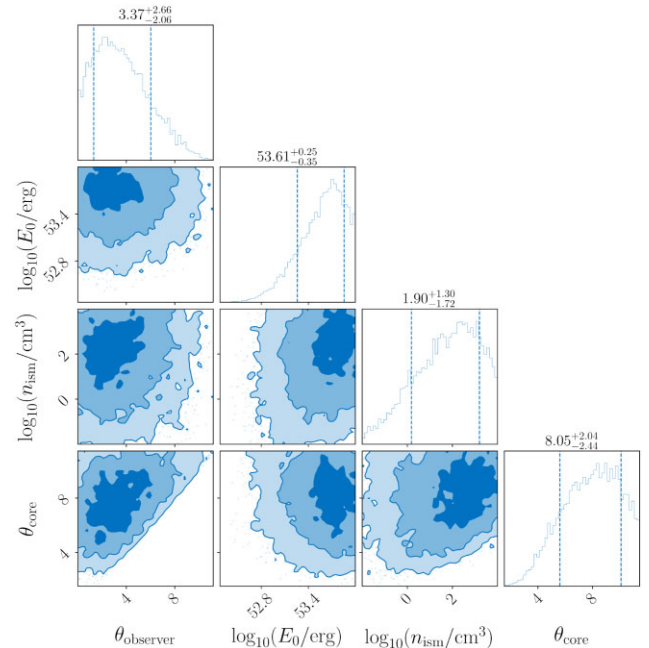


Figure 3. Posterior distributions for the structured jet model observers viewing angle θ_{observer} , the half-width jet core θ_{core} , the on-axis isotropic equivalent energy E_{iso} , and the ambient interstellar medium density n_{ism} for AT2020blt. The shaded contours indicate the 1σ – 3σ credible intervals.

measurements from radio afterglows of gamma-ray bursts (Nava et al. 2014; Beniamini & van der Horst 2017).

These measurements suggest that AT2020btl is consistent with being the afterglow of a relatively typical long gamma-ray burst where the observer is located within the ultra-relativistic core. This implies that the non-detection of prompt gamma-ray emission cannot be due to relativistic beaming i.e. we did not observe AT2020btl off-axis.

Our results confirm that assuming AT2020btl was related to the gamma-ray burst phenomena, the data of AT2020btl is best explained as the afterglow produced by an on-axis jet that successfully broke out of the ejecta. This implies that we can rule out the non-detection of prompt gamma-ray emission due to relativistic beaming or that the jet that produced AT2020btl did not successfully break out of the ejecta. We now explore whether the jet launched in AT2020btl could produce prompt gamma-ray emission.

4 LORENTZ FACTOR

As discussed above, gamma-ray production requires a relativistic jet to alleviate the compactness problem (Ruderman 1975). Naturally, this implies that if a jet breaks out of the ejecta and it is not sufficiently relativistic (for potential explanations of non-relativistic jets, see e.g. Huang et al. (2002)), the jet will not produce detectable prompt gamma-ray emission.

In Section 3, we showed that AT2020btl is on-axis and likely successfully launched a jet that broke out of the ejecta. Here, we explore if the jet was above the prompt emission threshold through back-of-the-envelope estimates and detailed fitting.

The threshold for producing prompt gamma-ray emission is typically assumed to be $\Gamma_0 \gtrsim 100$, with jets with Lorentz factors below this threshold referred to as ‘dirty’ fireballs or failed gamma-ray bursts (Huang et al. 2002; Rhoads 2003). However, indirect measurements of the Lorentz factor as low as $\Gamma_0 \sim 20$ have been made for jets following some successful gamma-ray bursts (e.g. Ghirlanda et al. 2018) with weak prompt emission. Accurately defining the gamma-ray prompt emission threshold requires knowing the radius at which prompt emission is produced, through what mechanism, and the fraction of photons above the pair production threshold in the co-moving frame. We do not know any of these constraints. However, all of these unknowns serve to lower the threshold Γ_0 value for producing prompt gamma-ray emission. As such, we take the conservative value of $\Gamma_0 \sim 100$ as the gamma-ray emission threshold, with $\Gamma_0 \sim 20$ serving as the absolute lower limit. We note that the latter threshold comes from assuming a wind-like interstellar medium.

We estimate the Lorentz factor by fitting the same power-law structured jet model described earlier with a finite deceleration radius allowing us to estimate the Lorentz factor. Our posterior on the Lorentz factor is $\Gamma_0 = 500 \pm 300$. The probability of having the Lorentz factor lower than the prompt emission threshold ($\Gamma_0 \lesssim 100$) is $\lesssim 7$ per cent and $\lesssim 1$ per cent for $\Gamma_0 \lesssim 20$. However, we note that because AT2020btl is missing early time data, our result is strongly dependent on the prior we assume for Γ_0 and implicitly on the full distribution of t_0 derived in Section 2. For the analysis above, we place a uniform prior on Γ_0 between 1 and 1000. To mitigate this dependency on the prior, we also estimate the Lorentz factor through back-of-the-envelope arguments.

The Lorentz factor can be estimated by measuring the afterglow onset time, also referred to as the deceleration time-scale t_{dec} . For an on-axis observer, the deceleration time-scale is the peak of the optical lightcurve, with the relativistic jet starting to decelerate on this time-scale. This allows us to place a lower limit on Γ_0 . The Lorentz factor is related to the deceleration time-scale and weakly to

the interstellar medium density and jet energy. We can approximate this relationship by (e.g. Sari & Piran 1999; Nakar 2007)

$$\Gamma_0 \approx 40 \left(\frac{E_{k,\text{iso},50}}{n_{\text{ism}}} \right)^{1/8} \left(\frac{100(1+z)}{t_{\text{peak}} - t_0} \right)^{3/8}. \quad (5)$$

Here, $E_{k,\text{iso},50}$ is the isotropic equivalent kinetic energy in units of 10^{50} erg, which we measure from the fitting of the afterglow, and z is the redshift. We do not observe this peak in AT2020btl but our analysis from Section 2 provides a conservative estimate of $t_{\text{peak}} = 58876.1^{+0.1}_{-0.2}$ MJD. With this estimate for t_{peak} we can use equation (5) and the derived values of $E_{k,\text{iso}}$ and n_{ism} from Section 3 to set a lower limit on the Lorentz factor. We emphasize that the derived estimates of E_{iso} and n_{ism} are robust to the choice of the prior. Moreover, given the weak dependence of Γ_0 to $E_{k,\text{iso}}$ and n_{ism} , the predominant source of uncertainty is from the estimate of t_0 and t_{peak} itself. Taking the median values of our estimated parameters suggests, conservatively, $\Gamma_0 \gtrsim 15$, indicating that AT2020btl may have successfully produced prompt gamma-ray emission. The full distribution of t_0 , t_{peak} etc implies a range on Γ_0 from 4 to 400.

We note that if AT2020btl was a ‘dirty’ fireball or failed gamma-ray burst, the threshold for producing gamma-rays would be higher. This is due to gamma-ray photons being produced when the jet is still optically thick and subsequently reabsorbed into the outflow, raising the kinetic energy of the outflow (e.g. Lamb & Kobayashi 2016). A scaling for this threshold is $\Gamma_0 \sim 16(E_{k,50})^{0.15}$, which is weakly dependent on the efficiency in turning the gamma-ray energy into kinetic. For our estimated kinetic energy from Section 3, this implies $\Gamma_0 \sim 56$ for producing gamma-ray emission, above our conservative estimate. However, our estimated Lorentz factor above is conservative and if AT2020btl is not a dirty fireball then this analysis does not hold. We therefore work with the observationally supported threshold where a jet with $\Gamma_0 \sim 20$ produced a successful gamma-ray burst.

Our analysis indicates that AT2020btl is likely not the afterglow from a ‘dirty’ fireball and that the jet that broke out of the ejecta was likely above the threshold for producing prompt gamma-ray emission.

The lack of observed gamma-rays is therefore puzzling. As we show in Section 3, we observed AT2020btl on-axis, and therefore relativistic beaming cannot explain the non-detection of prompt gamma-rays. Furthermore, our analysis indicates that AT2020btl can not be from a cocoon. Having ruled both these hypotheses out, we now turn to look at the prompt emission efficiency itself.

5 PROMPT EMISSION EFFICIENCY

Despite over three decades of observations, we still do not understand how prompt gamma-ray emission is produced. Given this uncertainty, we do not have a robust generative model to predict the energetics of the prompt gamma-ray emission. Given the lack of a model, a common approach in the field is to compare the energetics of the prompt and afterglow phases and compute a radiative efficiency n_γ ,

$$n_\gamma = \frac{E_{\gamma,\text{iso}}}{E_{k,\text{iso}} + E_{\gamma,\text{iso}}}, \quad (6)$$

where $E_{\gamma,\text{iso}}$ is the observed isotropic energy in gamma-rays. This efficiency has been calculated for a large catalogue of long and short gamma-ray bursts (Fong et al. 2015; Wang et al. 2015) using various techniques that have their own associated problems, most notably, fixing the energy in the magnetic-field ϵ_b , rather than marginalising over the uncertainty in this parameter. The efficiencies have a broad distribution ranging from ~ 1 to ~ 90 per cent. In

principle, the radiative efficiency can offer a clue into the prompt emission mechanism. However, this is fraught with uncertainties due to detector selection effects, uncertain physics (e.g. Lloyd-Ronning & Zhang 2004; Zhang et al. 2007) and likely incorrect modelling assumptions such as ignoring synchrotron energy losses due to synchrotron self-Compton emission, which is inconsistent with recent constraints (Beniamini, Nava & Piran 2016).

Ho et al. (2020) searched for sub-threshold triggers in *Fermi* (Meehan et al. 2009) and *KONUS-Wind* (Aptekar et al. 1995) finding no potential counterpart. Based on IPN observations, Ho et al. (2020) placed an upper limit on the gamma-ray energy of $E_\gamma \lesssim 7 \times 10^{52}$ erg by taking a nominal fluence threshold of 10^{-6} erg cm $^{-2}$. Given our estimate for the kinetic energy through the afterglow fitting (see Section 3) this implies $n_\gamma \lesssim 4.5$ per cent. However, this threshold is potentially conservative. A search for coincident gamma-rays from *KONUS-Wind* has instead set a deeper upper limit on the fluence of 6.1×10^{-7} erg cm $^{-2}$ (Ridnaia et al. 2020). With our estimate of the kinetic energy, this implies $n_\gamma \lesssim 2.8$ per cent for gamma-rays not to be observed by *KONUS-Wind* or a gamma-ray energy $E_\gamma \lesssim 1.2 \times 10^{52}$ erg. We note that the latter upper limit is set for a typical gamma-ray burst spectrum lasting 2.9 s and may not apply to AT2020btl.

To augment this upper limit, we perform a sub-threshold search of *Fermi* data from 1 min before the first non-detection (i.e. on January 27.54) to 2 min after the first detection (i.e. on January 28.28). Unlike *KONUS-Wind* which did not have any interruptions to observations, *Fermi* observations were periodically interrupted due to occultation by Earth and *Fermi* passing through the South Atlantic Anomaly. In our given search window, *Fermi* was observing the location of AT2020btl ~ 64 per cent of the total time making it plausible that a gamma-ray burst weaker than *KONUS-Wind* was missed. However, we note that around our most probable start time $t_0 = 58875.53$ MJD, *Fermi* did not have any interruptions. The constraints on the gamma-ray prompt emission from *Fermi* are more sensitive than *KONUS-Wind*, with an upper limit on $E_{\gamma, \text{iso}}$ from 1 to 6×10^{51} erg depending on the spectrum of the source and spectral template used in the search. This corresponds to an upper limit on $n_\gamma \lesssim 1.3$ per cent and $\lesssim 0.3$ per cent, respectively. For details of the sub-threshold search, source spectrum and templates, we refer the reader to Blackburn et al. (2015) and Goldstein et al. (2016, 2019). Taking the lower value implies that for AT2020btl to not produce detectable prompt emission, the radiative efficiency must have been lower than 0.3 per cent, challenging several prompt emission models such as photospheric emission (e.g. Lazzati et al. 2013) and magnetic field dissipation (e.g. Zhang & Yan 2011), but plausible for internal shock models (e.g. Kobayashi, Piran & Sari 1997; Daigne & Mochkovitch 1998; Kumar 1999). Such a low efficiency is only measured in one other gamma-ray burst, GRB190829A which had an efficiency of ~ 0.1 per cent (e.g. Salafia et al. 2021).

In Fig. 4, we show these upper limits on the isotropic energy in gamma-rays for AT2020btl. The green, black, blue, and purple curves are the cumulative distributions of gamma-ray burst energies for AT2020btl calculated using the distribution of efficiencies derived from the Nava et al. (2014), Wang et al. (2015), Fong et al. (2015), and Beniamini et al. (2015) analyses, respectively. The latter three analyses all calculate efficiencies for a sample of gamma-ray bursts with various assumptions. In particular, the distribution of efficiencies from Beniamini et al. (2015) is calculated using the GeV flux rather than the X-ray flux as the former is likely a better proxy for the true kinetic energy of the afterglow. The green curve, however, is calculated assuming the average efficiency of gamma-ray bursts is 0.2, consistent with other estimates (e.g. Beniamini et al. 2016) and using the constraint that the distribution of efficiencies is narrow

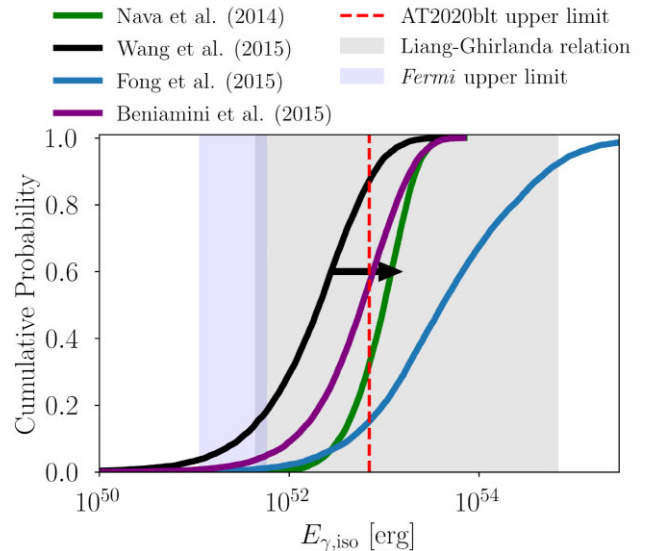


Figure 4. Cumulative probability distributions of the isotropic equivalent energy in gamma-rays $E_{\gamma, \text{iso}}$ of AT2020btl. The green, black, blue, and purple curves are the cumulative distributions of gamma-ray burst energies for AT2020btl calculated using the distribution of efficiencies derived from the Nava et al. (2014), Wang et al. (2015), Fong et al. (2015), and Beniamini et al. (2015) analyses, respectively. The black arrow indicates that the distribution derived in (Wang et al. 2015) is likely a lower limit. The red line indicates the conservative upper limit for AT2020btl based on IPN observations set by Ho et al. (2020), while the blue shaded region shows the upper limit from a sub-threshold search we perform with *Fermi*. The grey band is the prediction for the gamma-ray energy from Liang–Ghirlanda relation.

as determined by Nava et al. (2014). The true average efficiency of gamma-ray bursts is unknown which would shift the green curve left or right for lower and higher average efficiencies, respectively. We note that as Wang et al. (2015) fix $\epsilon_b = 10^{-5}$ for all gamma-ray bursts, the efficiency derived in their sample are systematically biased to low values and serve as a lower limit on the observed population. A similar bias due to fixing of parameters is present in other analyses (see for e.g. Racusin et al. 2011; Fong et al. 2015) and therefore comparisons to populations should be done cautiously.

The non-detection by *KONUS-Wind* and *Fermi* implies that if AT2020btl produced prompt emission, the isotropic energy released in gamma-rays must have been lower than $\sim 1 \times 10^{51}$ erg, an efficiency weaker than ~ 99.9 per cent of the observed distribution of long gamma-ray bursts (Fong et al. 2015) or 97.5 per cent if considering the distribution derived in Wang et al. (2015). In the grey shaded region, we show the predicted isotropic gamma-ray energy through the Liang–Ghirlanda relation, $E_\gamma \sim 10^{52} (\Gamma_0/182)^4$ (e.g. Liang et al. 2010; Ghirlanda et al. 2018). Here, we use our posterior, $\Gamma_0 = 500 \pm 300$ obtained by fitting the afterglow (see Section 3). Given the non-detection by *Fermi* and *KONUS-Wind*, AT2020btl likely did not follow the Liang–Ghirlanda relation, which may suggest a different origin for this transient. However, it is worth noting that there are notable exceptions to the Liang–Ghirlanda relation, such as GRB090510 (Ghirlanda, Ghisellini & Nava 2010).

Our results suggest that if AT2020btl is a typical gamma-ray burst, given the afterglow modelling indicates that this event was on-axis, the prompt emission must have been low-luminosity to not be detected by *Fermi* and *KONUS-Wind* or missed by *Fermi* and with a radiative efficiency of $n_\gamma \lesssim 2.8 - 4.5$ per cent. Or equivalently, weaker than up to 98.2 per cent of the population (Fong et al. 2015).

This low efficiency is not necessarily a problem for prompt emission models (e.g. Kobayashi et al. 1997; Daigne & Mochkovitch 1998; Kumar 1999) and might even support certain varieties of internal shock models. However, it is inconsistent with a large fraction of observed gamma-ray bursts and previously derived efficiency distributions (e.g. Racusin et al. 2011; Fong et al. 2015; Wang et al. 2015). It is also plausible that AT2020btl is a failed gamma-ray burst, but to confirm this hypothesis, we need to accurately determine the threshold for generating gamma-ray emission. If the threshold is $\Gamma_0 \sim 20$ (motivated by observations of successful gamma-ray bursts) then our most conservative analysis suggests AT2020btl is not a failed gamma-ray burst. However, if the threshold is higher, then there is some part of the parameter space where the jet launched in AT2020btl could fail to produce gamma-ray emission.

6 IMPLICATIONS AND CONCLUSION

Gamma-ray bursts have been observed extensively for over three decades, with their lower energy broad-band afterglows almost always observed in the follow-up to a gamma-ray trigger. The Zwicky Transient Facility is quickly changing this dynamic, with three afterglow-like transients already detected without a gamma-ray trigger (Ho et al. 2020, 2021), with a further four independent detections of an optical counterpart to a gamma-ray burst. Here, we have investigated the observations of AT2020btl in detail to ultimately determine why no prompt gamma-ray emission was observed?

In Section 3, we showed that AT2020btl is best interpreted as the afterglow produced by a successful on-axis structured jet. We ruled out the hypotheses that AT2020btl was the afterglow of a cocoon or an off-axis gamma-ray burst, with a cocoon and off-axis jet model unsuccessful in explaining the observations. In Section 4, we showed that the jet launched in AT2020btl was likely above the prompt emission generation threshold, i.e. the jet was relativistic enough to alleviate the compactness problem, potentially ruling out the hypothesis that AT2020btl was the afterglow of a dirty fireball. Moreover, given the rate of afterglow-like transients and gamma-ray bursts are roughly consistent, such phenomena can potentially already be ruled out (e.g. Cenko et al. 2015; Ho et al. 2020). In Section 5, we found that the non-detections of prompt gamma-ray emission in *Fermi* and *KONUS-Wind* imply that the prompt gamma-ray emission in AT2020btl is weaker than up to 99.5 per cent of the observed gamma-ray burst population.

We also predicted the prompt emission energy through the Liang–Ghirlanda relation. If AT2020btl was consistent with this relation, the prompt emission should have been detected by both *Fermi* and *KONUS-Wind*, implying the gamma-ray energy generated in AT2020btl must be lower than the predictions by the Liang–Ghirlanda relation. However, we note that there are other gamma-ray burst exceptions to this relation, particularly GRB 090510 (Ghirlanda et al. 2010), which like AT2020btl, had weaker gamma-ray emission than expected.

Given our inadequate knowledge of prompt emission generation, the only meaningful study about the prompt emission we can do is investigate the prompt emission efficiency. Depending on the upper limits used, the radiative efficiency of AT2020btl is $n_\gamma \lesssim 0.3$ –4.5 per cent, which is lower than between 86.3 and 99.5 per cent of the observed long gamma-ray burst population (Racusin et al. 2011). This low efficiency strongly favours internal shocks as the likely prompt emission generation mechanism (e.g. Kumar 1999).

In addition to AT2020btl, the Zwicky Transient Facility also detected AT2021any, another potential afterglow-like transient at

$z = 2.514$ (Ho & Zwicky Transient Facility Collaboration 2021; Ho et al. 2021) which was observed to be rapidly fading after ~ 22 mins and detected without an identified gamma-ray counterpart despite coverage from multiple gamma-ray telescopes (Ho et al. 2021). Given that AT2021any was rapidly fading after ~ 22 mins, it was also likely observed on-axis and the non-detection of gamma-ray emission could imply it also has a low prompt emission efficiency that is inconsistent with the population or probing the margins of the efficiency distribution of gamma-ray bursts (Nava et al. 2014). We note that the data for AT2021any is not yet available and therefore we can not perform the detailed analyses for AT2021any we presented here.

The radiative efficiencies of AT2020btl and AT2021any are potentially inconsistent with (or at least in the tails of) the observed population of gamma-ray bursts. The capabilities of the Zwicky Transient Facility, in particular, the high cadence and large survey volume (e.g. Bellm et al. 2019) provide an opportunity to detect afterglow transients without a gamma-ray counterpart. This implies that the detection of transients such as AT2020btl and AT2021any are independent of gamma-ray observatories’ observational biases. Therefore, it is conceivable that AT2020btl and AT2021any are the afterglows produced by the low-luminosity gamma-ray bursts that have so far been missed entirely (or less frequently observed) due to a Malmquist bias associated with gamma-ray observatories. Future all-sky gamma-ray detectors can avoid this bias and test whether these transients are genuinely part of the cosmological gamma-ray burst population or a distinct new class. Furthermore, new optical telescopes like the Vera Rubin observatory (e.g. Ivezić et al. 2019) will be capable of finding more transients like AT2020btl to deeper magnitudes and provide stringent constraints on various parameters to test the nature of these transients.

We also note that the potential inconsistency of the radiative efficiencies of AT2020btl and AT2021any with the observed population may be a product of inadequate modelling and systematic biases in previous studies into the radiative efficiencies of gamma-ray bursts. Inferring prompt emission efficiencies requires robustly determining the kinetic energy in the afterglow, which requires a good understanding of the jet structure (at least for off-axis sources). Furthermore, in previous analyses, the estimations have often been done by fixing the energy fraction in the magnetic field, ϵ_b , and the participation fraction, ξ_n , which significantly underestimates uncertainties at best or leads to biases at worst. This motivates the need for detailed afterglow modelling on individual gamma-ray bursts and entire populations with a detailed treatment of observational selection effects (e.g. Mandel, Farr & Gair 2019).

It is also worth considering – albeit cautiously – whether AT2020btl and AT2021any are afterglows of typical gamma-ray burst progenitors at all. Instead, they may be the product of a different phenomenon or a subclass of typical gamma-ray burst progenitors that can launch a mildly relativistic, ultra-low efficiency jet that interacts with the interstellar medium to produce an afterglow-like transient but does not necessarily produce prompt gamma-ray emission. If the Zwicky Transient Facility and later the Vera Rubin observatory continues to find afterglow-like transients without prompt emission, it would be intriguing to consider the host galaxy properties and the population properties of such transients. Although, as we discussed previously, it is more likely that AT2020btl and AT2021any are simply afterglows of gamma-ray bursts we previously missed due to observational biases. By virtue of opening a new window into these phenomena, it is likely that the Zwicky Transient Facility and in the future, Vera Rubin observatory will continue to find afterglow-like transients without high energy counterparts.

ACKNOWLEDGEMENTS

We are grateful to Kendall Ackley, Moritz Hübner, Paul Easter, and the anonymous referee for their thoughtful comments on the manuscript. NS is supported by an Australian Government Research Training Program (RTP) Scholarship. PDL is supported through Australian Research Council Future Fellowship FT160100112 and CE170100004. PDL and GA are supported by ARC. Discovery Project DP180103155.

DATA AVAILABILITY

All the data used in this paper is available in Ho et al. (2020).

REFERENCES

- Abbott B. P., et al., 2017, *Astrophys. J.*, 848, L13
 Aloy M. A., Janka H. T., Müller E., 2005, *A&A*, 436, 273
 Aptekar R. L. et al., 1995, *Space Sci. Rev.*, 71, 265
 Ashton G. et al., 2019, *ApJS*, 241, 27
 Bellm E. C. et al., 2019, *PASP*, 131, 018002
 Beniamini P., van der Horst A. J., 2017, *MNRAS*, 472, 3161
 Beniamini P., Nava L., Duran R. B., Piran T., 2015, *MNRAS*, 454, 1073
 Beniamini P., Nava L., Piran T., 2016, *MNRAS*, 461, 51
 Beniamini P., Granot J., Gill R., 2020, *MNRAS*, 493, 3521
 Blackburn L., Briggs M. S., Camp J., Christensen N., Connaughton V., Jenke P., Remillard R. A., Veitch J., 2015, *ApJS*, 217, 8
 Cano Z., Wang S.-Q., Dai Z.-G., Wu X.-F., 2017, *Adv. Astron.*, 2017, 8929054
 Cenko S. B. et al., 2012, *ApJ*, 753, 77
 Cenko S. B. et al., 2015, *ApJ*, 803, L24
 Chevalier R. A., Li Z.-Y., 2000, *ApJ*, 536, 195
 Daigne F., Mochkovitch R., 1998, *MNRAS*, 296, 275
 Fong W., Berger E., Margutti R., Zauderer B. A., 2015, *ApJ*, 815, 102
 Ghirlanda G., Ghisellini G., Nava L., 2010, *A&A*, 510, L7
 Ghirlanda G. et al., 2018, *A&A*, 609, A112
 Ghirlanda G., et al., 2019, *Science*, 363, 968
 Gill R., Granot J., 2018, *MNRAS*, 478, 4128
 Goldstein A., Burns E., Hamburg R., Connaughton V., Veres P., Briggs M. S., Hui C. M., The GBM-LIGO Collaboration, 2016, preprint (arXiv:1612.02395)
 Goldstein A. et al., 2019, preprint (arXiv:1903.12597)
 Gottlieb O., Nakar E., Piran T., Hotokezaka K., 2018, *MNRAS*, 479, 588
 Granot J., Panaitescu A., Kumar P., Woosley S. E., 2002, *ApJ*, 570, L61
 Güver T., Özel F., 2009, *MNRAS*, 400, 2050
 Ho A. Y. Q., Zwicky Transient Facility Collaboration, 2021, GRB Coordinates Netw., 29313, 1
 Ho A. Y. Q. et al., 2020, *ApJ*, 905, 98
 Ho A. Y. Q., Perley D. A., Yao Y., Andreoni I., 2021, *Transient Name Serv. AstroNote*, 20, 1
 Huang Y. F., Dai Z. G., Lu T., 2002, *MNRAS*, 332, 735
 Ivezić Ž., et al., 2019, *ApJ*, 873, 111
 Kasliwal M. M., et al., 2017, *Science*, 358, 1559
 Kathirgamaraju A., Barniol Duran R., Giannios D., 2018, *MNRAS*, 473, L121
 Kobayashi S., Piran T., Sari R., 1997, *ApJ*, 490, 92
 Kumar P., 1999, *ApJ*, 523, L113
 Lamb G. P., Kobayashi S., 2016, *ApJ*, 829, 112
 Lamb G. P., Mandel I., Resmi L., 2018, *MNRAS*, 481, 2581
 Lazzati D., Morsony B. J., Margutti R., Begelman M. C., 2013, *ApJ*, 765, 103
 Lazzati D., Perna R., Morsony B. J., Lopez-Camara D., Cantiello M., Ciolfi R., Giacomazzo B., Workman J. C., 2018, *Phys. Rev. Lett.*, 120, 241103
 Lee K. H., Bartos I., Privon G. C., Rose J. C., Torrey P., 2020, *ApJ*, 902, L23
 Liang E.-W., Yi S.-X., Zhang J., Lü H.-J., Zhang B.-B., Zhang B., 2010, *ApJ*, 725, 2209
 Lithwick Y., Sari R., 2001, *ApJ*, 555, 540
 Lloyd-Ronning N. M., Zhang B., 2004, *ApJ*, 613, 477
 Mandel I., Farr W. M., Gair J. R., 2019, *MNRAS*, 486, 1086

- Marcote B., Nimmo K., Salafia O. S., Paragi Z., Hessels J. W. T., Petroff E., Karuppusamy R., 2019, *ApJ*, 876, L14
 Matsumoto T., Nakar E., Piran T., 2019, *MNRAS*, 483, 1247
 Meegan C. et al., 2009, *ApJ*, 702, 791
 Mooley K. P. et al., 2018, *Nature*, 561, 355
 Nakar E., 2007, *Phys. Rep.*, 442, 166
 Nakar E., 2020, *Phys. Rep.*, 886, 1–84
 Nakar E., Piran T., 2017, *ApJ*, 834, 28
 Nakar E., Piran T., Granot J., 2002, *ApJ*, 579, 699
 Nakar E., Ando S., Sari R., 2009, *ApJ*, 703, 675
 Nava L. et al., 2014, *MNRAS*, 443, 3578
 Racusin J. L., et al., 2011, *ApJ*, 738, 138
 Rhoads J. E., 1997, *ApJ*, 487, L1
 Rhoads J. E., 2003, *ApJ*, 591, 1097
 Ridnaia A. et al., 2020, GRB Coordinates Netw., 27039, 1
 Ruderman M., 1975, *Ann. New York Acad. Sci.*, 262, 164
 Ryan G., van Eerten H., Piro L., Troja E., 2020, *ApJ*, 896, 166
 Salafia O. S. et al., 2021, preprint (arXiv:2106.07169)
 Sari R., Piran T., 1999, *ApJ*, 520, 641
 Singer L. P., Ahumada T., Ho A. Y. Q., Zwicky Transient Facility (ZTF), Global Relay Of Observatories Watching Transients Happen (Growth) Collaboration, 2020, GRB Coordinates Netw., 26968, 1
 Speagle J. S., 2020, *MNRAS*, 493, 3132
 Troja E. et al., 2017, *Nature*, 551, 71
 Troja E. et al., 2018, *MNRAS*, 478, L18
 Wang X.-G. et al., 2015, *ApJS*, 219, 9
 Zhang B., 2018, *The Physics of Gamma-Ray Bursts*. Cambridge University Press
 Zhang B., Yan H., 2011, *ApJ*, 726, 90
 Zhang B. et al., 2007, *ApJ*, 655, 989

APPENDIX A: OFF-AXIS AFTERGLOW

The lack of observed gamma-ray emission in AT2020btl could be a consequence of relativistic beaming, i.e. that we observed AT2020btl from outside the ultra-relativistic core. We test this hypothesis by fitting the structured jet models above with a broad uniform prior on t_0 from 2019 January 1 to 2020 January 27.54 (i.e. the first non-detection) and enforcing that the observer is located off-axis i.e. $\theta_{\text{observer}} > \theta_{\text{core}}$. Our fit to the multiwavelength data with the off-axis hypothesis is shown in Fig. A1. As we discuss in Section 3, an off-axis observer is unable to simultaneously explain the non-detection

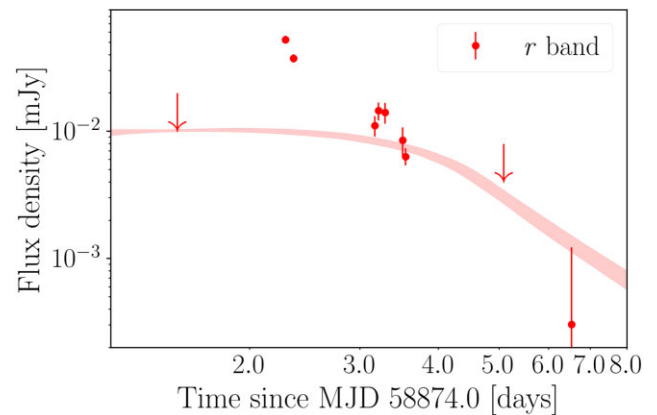


Figure A1. Data for AT2020btl in r band, with arrows indicating upper limits. The red shaded curve is the 95 per cent credible interval predicted lightcurves for an off-axis structured jet model. The errors on the data are the measurement uncertainty reported by Ho et al. (2020) combined with our modelled uncertainty for each model. We note that MJD 58874 is arbitrarily chosen to show the lightcurve more clearly. We also note that only r band data is shown here but the fits are performed on the full multiwavelength data.

on January 27, the rapid rise and subsequent rapid decay for all values of β from 0.5–10. We note that it is possible to explain the rapid rise and peak flux (but not the subsequent decay) for an observer located very close to the jet edge and with a jet to have $E_{k,iso} \gtrsim 10^{54}$ erg which is uncomfortably large. More quantitatively, we calculate a Bayes factor between a fit where the observer is forced off-axis and one where they are not. We calculate a Bayes factor of $\sim 3 \times 10^4$ in favour of the on-axis hypothesis. The fit and the Bayes factor clearly indicate that the off-axis hypothesis is not able to explain the observations.

APPENDIX B: SUPPLEMENTARY INFORMATION

In Table B1 we list the priors used in our analysis for the most preferred model i.e. the power-law structured jet along with a brief description of the parameters. In Fig. B1 we show the one and two-dimensional posterior distribution from our fit to the multiwavelength data using the power-law structured jet model. We note that the parameters not shown here are previously shown in Figs 1 and 3.

Table B1. Parameters associated with the smooth power-law structured jet model along with a brief description and the prior used in our analysis. We note that the last two parameters f_a and $\log_{10}(n_H/\text{cm}^{-2})$ are used to account for the host galaxy extinction and are motivated by the analysis from Güver & Özel (2009).

Parameter [unit]	Description	Prior
t_0 [MJD]	Burst time	Uniform[58872, 58876.28]
t_p [MJD]	Peak time i.e. the deceleration time	Uniform[58875.54, 58876.28]
m	Gradient of deceleration power law	Uniform[0.5, 7]
Γ_0	Initial Lorentz factor	Uniform[1, 1000]
θ_{observer} [rad]	Observers viewing angle	Cosine[0, 0.7]
$\log_{10}(E_{\text{iso}}/\text{erg})$	Isotropic-equivalent energy	Uniform[46, 54]
θ_{core} [rad]	Half-width of jet core	Uniform[0.02, 0.2]
θ_{wing}	Wing truncation angle of the jet	Uniform[1, 5] $\times \theta_{\text{core}}$
β	Power for power-law structure	Uniform[0.5, 10]
$\log_{10}(n_{\text{ism}}/\text{cm}^{-3})$	Number density of ISM	Uniform[−3, 4]
p	Electron distribution power-law index	Uniform[2, 3]
$\log_{10}\epsilon_e$	Thermal energy fraction in electrons	Uniform[−5, 0]
$\log_{10}\epsilon_b$	Thermal energy fraction in magnetic field	Uniform[−5, 0]
ξ_N	Fraction of accelerated electrons	Uniform[0, 1]
σ	Additional noise on measurements	Uniform[10^{-4} , 10^{-3}]
$\log_{10}(n_H/\text{cm}^{-2})$	Hydrogen column density of the host galaxy	Uniform[19, 24]
f_a	Extinction factor from Güver & Özel (2009)	Gaussian(2.21, 0.09)

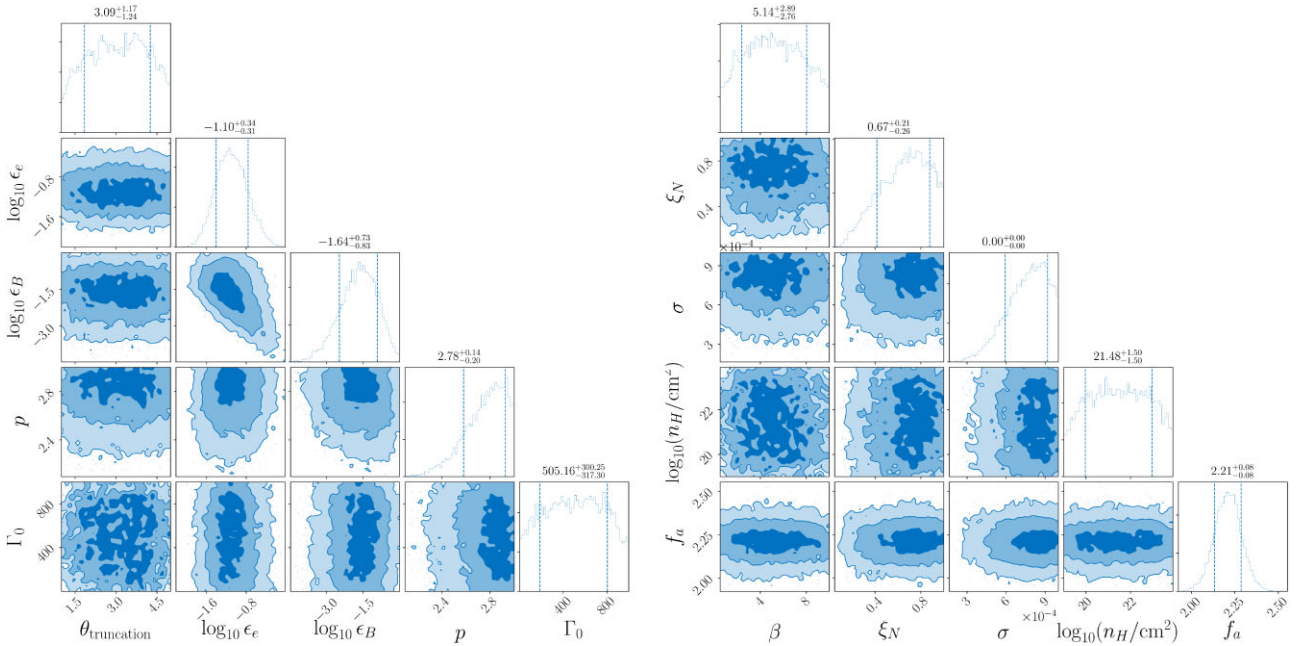


Figure B1. One- and two-dimensional posterior distribution for the structured jet model. The shaded contours indicate the 1σ – 3σ credible intervals. Note that the parameters not shown here are shown in Figs 1 and 3.

This paper has been typeset from a \LaTeX file prepared by the author.