



# The UNCOVER Survey: A First-look HST + JWST Catalog of 60,000 Galaxies near A2744 and beyond

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## Abstract

In 2022 November, the James Webb Space Telescope (JWST) returned deep near-infrared images of A2744—a powerful lensing cluster capable of magnifying distant, incipient galaxies beyond it. Together with existing Hubble Space Telescope (HST) imaging, this publicly available data set opens a fundamentally new discovery space to understand the remaining mysteries of the formation and evolution of galaxies across cosmic time. In this work, we detect and measure some 60,000 objects across the 49 arcmin<sup>2</sup> JWST footprint down to a  $5\sigma$  limiting magnitude of

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$\sim 30$  mag in  $0''.32$  apertures. Photometry is performed using circular apertures on images matched to the point-spread function (PSF) of the reddest NIRCcam broad band, F444W, and cleaned of bright cluster galaxies and the related intracluster light. To give an impression of the photometric performance, we measure photometric redshifts and achieve a  $\sigma_{\text{NMAD}} \approx 0.03$  based on known, but relatively small, spectroscopic samples. With this paper, we publicly release our HST and JWST PSF-matched photometric catalog with optimally assigned aperture sizes for easy use, along with single aperture catalogs, photometric redshifts, rest-frame colors, and individual magnification estimates. These catalogs will set the stage for efficient and deep spectroscopic follow up of some of the first JWST-selected samples in summer of 2023.

*Unified Astronomy Thesaurus concepts:* [Catalogs \(205\)](#); [Abell clusters \(9\)](#); [Photometry \(1234\)](#); [James Webb Space Telescope \(2291\)](#); [Hubble Space Telescope \(761\)](#); [Astronomical methods \(1043\)](#)

## 1. Introduction

The vast distance scales of our Universe relative to the human timescale implicitly means that there are very few astrophysical processes we can observe changing in real time. Dynamical processes in galaxies transpire over timescales of millions to billions of years. Thus to understand the formation and evolution of galaxies across cosmic time necessitates the study of statistically representative snapshots. Observational campaigns are forced to make decisions in survey design, generally prioritizing either depth (e.g., Williams et al. 1996; Giavalisco et al. 2004; Beckwith et al. 2006; Bouwens et al. 2011; Illingworth et al. 2013, 2016; Lotz et al. 2017), volume (e.g., Scoville et al. 2007; Jarvis et al. 2013; Abbott et al. 2018; Aihara et al. 2018), or a mix thereof (e.g., Grogin et al. 2011; Koekemoer et al. 2011) in order to assemble unbiased galaxy populations. All of these surveys share a common theme: they require a synthesis of panchromatic flux measurements of detected sources into a photometric catalog as the first step toward modeling these stellar populations. Such photometric catalogs serve as the foundation of any galaxy survey, necessary for identifying robust galaxy samples and for enabling the vast majority of subsequent science investigations.

The deepest surveys of our Universe to date come from single ultradeep pointings with the Hubble Space Telescope (HST), either of “blank fields” (i.e., relatively dark lines of sight through our own Milky Way galaxy; Williams et al. 1996; Beckwith et al. 2006; Bouwens et al. 2011; Illingworth et al. 2013) or by targeting known clusters of galaxies at intermediate redshifts (Lotz et al. 2017; Coe et al. 2019; Salmon et al. 2020; Sharon et al. 2020). One particular advantage of targeting galaxy clusters is the added boost from strong gravitational lensing; the richest clusters magnify background galaxies by factors of a few up to dozen typically, depending on the size and position of the background galaxy with respect to the lens (e.g., Coe et al. 2019). Strong lens clusters unveil some of the most distant (e.g., Zheng et al. 2012; Coe et al. 2013; Zitrin et al. 2014; Strait et al. 2021; Roberts-Borsani et al. 2022; Adams et al. 2023; Atek et al. 2023; Bradley et al. 2023; Furtak et al. 2023a; Hsiao et al. 2023; Williams et al. 2023) and lowest-mass (e.g., Bouwens et al. 2017; Livermore et al. 2017; Atek et al. 2018; Bhatawdekar et al. 2019; Kikuchihara et al. 2020; Furtak et al. 2021; Bouwens et al. 2022) galaxies known, even single candidate stars in some cases (Welch et al. 2022). However, this boost from cosmic telescopes comes with a cost in terms of contamination from cluster galaxies/intracluster light (ICL) and the complex and nonlinear distortions to galaxy morphologies (e.g., Shipley et al. 2018; Bhatawdekar et al. 2019; Pagul et al. 2021; Fox et al. 2022). In addition, the source-plane area that is being probed behind a lens is smaller. As such, there is a trade-off between detecting a higher number

of galaxies that are boosted in flux but for a smaller area probed relative to an unlensed field, an effect known as magnification bias (e.g., Broadhurst et al. 1995). However, since the luminosity function of high-redshift galaxies is steep enough (e.g., Finkelstein et al. 2015; Mason et al. 2015), the net effect is a gain in the number density of detections.

Despite these challenges, galaxy clusters afford our best opportunity to push to the most extreme depths and thus to the frontiers of galaxy formation. Campaigns such as the Director’s Discretionary Hubble Frontier Fields (HFF) program have amassed a rich archival data set of HST imaging (Lotz et al. 2017; Steinhardt et al. 2020) that has set the stage for JWST imaging and spectroscopic programs (e.g., Willott et al. 2017; Bezanson et al. 2022; Treu et al. 2022; Windhorst et al. 2023). One cluster is particularly compelling to study: in addition to a spectacular central core (Lotz et al. 2017; Shipley et al. 2018; Pagul et al. 2021; Kokorev et al. 2022), A2744 contains prominent lensing features within two additional massive cluster substructures in the north and northwest (Furtak et al. 2023b). A2744 thus contains an unusually large area of high magnification when combining the various structures. Several early JWST programs targeted A2744; here we combine publicly available HST and JWST photometry from the JWST-GO-2561, JWST-DD-ERS-1324, and JWST-DD-2756 programs.

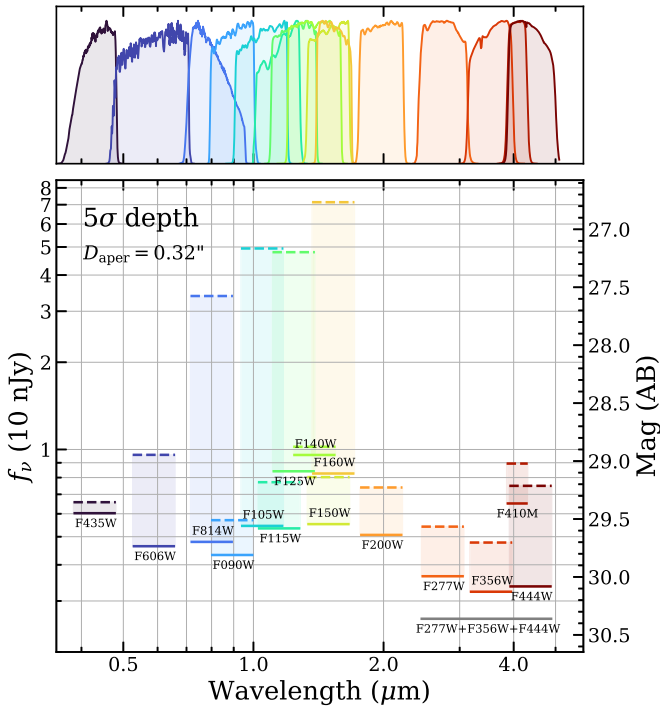
In this paper, we present the space-based photometric catalog for the UNCOVER survey (Bezanson et al. 2022) as part of Data Release 2 (DR2),<sup>41</sup> including derived photometric redshifts and magnification corrections from updated lensing models originally presented in Furtak et al. (2023b). Catalogs are available online.<sup>42</sup> The software used to produce these catalogs, *aperpy*, is generally applicable to any JWST/NIRCcam data and is freely available.<sup>43</sup>

In Section 2, we present an overview of the data processing, including the reduction and astrometric correction of images from HST and JWST. Section 3 describes two approaches for handling the added complexity of ICL in the A2744 cluster. Source detection and photometry are described in Section 4, including a description of the methodology adopted to homogenize the point-spread function (PSF), the measurement of total fluxes from aperture photometry, and a quantification of representative errors. We present the photometric catalog properties in Section 5, including depths, galaxy number counts, and comparisons to other surveys. Finally, we summarize our findings in Section 6.

<sup>41</sup> Note that an earlier set of mosaics, photometric catalog, and manuscript were provided as part of DR1; these too are available on our website (<https://jwst-uncover.github.io/DR1.html>) and Zenodo at doi:10.5281/zenodo.8199803.

<sup>42</sup> Catalogs are available on our website (<https://jwst-uncover.github.io/DR2.html>) and Zenodo at 10.5281/zenodo.8199802.

<sup>43</sup> *aperpy* is available through GitHub (<https://github.com/astrowhit/aperpy>) and Zenodo (Weaver & Cutler 2023).



**Figure 1.** Effective catalog depths over the A2744 JWST footprint for the 15 available photometric bands and their transmission curves. The effective depth of our LW detection image is also included. Depths are quoted in  $0''.32$  diameter apertures and correspond to the area-weighted 50th (median) and 10th percentiles (dashed and solid lines, respectively). Areas in HST imaging without JWST coverage are not considered. See the text for details.

An Appendix contains additional relevant information regarding the stability of the JWST PSF in time and across the detector.

All magnitudes in this paper are expressed in the AB system (Oke 1974), for which a flux  $f_\nu$  of 10 nJy ( $10^{-28}$  erg cm $^{-2}$  s $^{-1}$  Hz $^{-1}$ ) corresponds to  $AB_\nu = 28.9 - 2.5 \log_{10}(f_\nu/\mu\text{Jy})$ . When computing physical properties such as rest-frame fluxes, we adopt a standard  $\Lambda$ CDM cosmology with  $H_0 = 70$  km s $^{-1}$  Mpc $^{-1}$ ,  $\Omega_M = 0.3$ , and  $\Omega_\Lambda = 0.7$ .

## 2. Data

### 2.1. JWST

The photometric catalogs presented herein include all public JWST/NIRCam imaging of A2744 available to date: the Ultradeep NIRSPEC and NIRCam Observations before the Epoch of Reionization (UNCOVER) Treasury survey (PIs: Labbé & Bezanson, JWST-GO-2561; Bezanson et al. 2022), the Early Release Science (ERS) GLASS-JWST program (PI: Treu, JWST-DD-ERS-1324; Treu et al. 2022), and a Directors Discretionary Time (DDT) program (JWST-DD-2756, PI: Chen). As described in Bezanson et al. (2022), our data set combines the deep NIRCam imaging with 4–6 hr exposures in seven filters (F115W, F150W, F200W, F277W, F356W, F410M, and F444W) from UNCOVER, with the ultradeep imaging with 9–14 hr exposures from GLASS-ERS in seven filters (F090W, F115W, F150W, F200W, F277W, F356W, and F444W). The GLASS-ERS NIRCam pointing is taken in parallel and is thus offset to the cluster outskirts, thereby extending the combined science area. Additionally, the DDT program includes two epochs of NIRCam imaging in six filters (F115W, F150W, F200W, F277W, F356W, and F444W), totaling  $\sim 1$  hr per filter. All together, images in eight unique

**Table 1**

Effective Catalog Depths, Quoted in  $0''.32$  Diameter Apertures and that Correspond to the Area-weighted 10th, 50th (Median), and 90th Percentiles

Filter	Depth ( $5\sigma$ AB)			Area (arcmin $^2$ )			
	10th	50th	90th	10th	50th	90th	Total
LW DET	30.36	29.85	29.01	5.78	27.10	44.56	48.91
F435W	29.45	29.35	28.30	0.90	4.30	15.85	18.52
F606W	29.73	28.95	27.42	0.95	19.97	33.11	37.54
F814W	29.70	27.58	27.17	2.24	19.28	28.05	32.98
F090W	29.81	29.51	29.30	1.46	6.76	11.65	12.84
F105W	29.56	27.17	27.06	4.17	12.86	17.95	21.11
F115W	29.58	29.18	28.15	6.18	27.28	44.53	48.25
F125W	29.09	27.20	27.09	4.32	12.24	17.48	20.97
F140W	28.95	28.88	28.54	1.66	4.04	4.84	5.62
F150W	29.54	29.14	28.23	6.07	27.89	45.00	48.63
F160W	29.11	26.77	26.53	3.35	12.00	18.89	21.03
F200W	29.64	29.23	28.47	5.93	27.00	44.10	47.82
F277W	29.99	29.57	28.76	5.53	26.25	44.11	48.58
F356W	30.13	29.70	28.88	5.74	26.25	44.20	48.82
F410M	29.37	29.02	28.58	3.45	15.84	25.90	28.66
F444W	30.08	29.21	28.24	5.60	26.31	44.37	48.25

**Note.** Total areas reflect the union of the LW detection footprint with the coverage available for each band.

JWST filters from both the short-wavelength (SW) and long-wavelength (LW) channels are combined to extend the coverage of A2744 to include the nearby cluster substructures (Figure 1, see also Table 1). Although the UNCOVER NIRISS parallel imaging has been reduced and released in Bezanson et al. (2022), photometry and cataloging of this parallel imaging is planned for future work.

Next, we summarize the key steps of the image reduction, referring the reader to Section 3 of Bezanson et al. (2022) for further details. Imaging mosaics are produced from the flux-calibrated NIRCam exposures released in Stage 2b of the JWST calibration pipeline (v1.8.4) and combined with calibration set `jwst_1039.pmap`. The exposures are then processed, aligned, and coadded using the GRISM redshift and LINE analysis software for space-based spectroscopy (GRIZLI,<sup>44</sup> 1.8.16.dev12; Brammer 2019; Kokorev et al. 2022). The pipeline has been optimized to handle known JWST artifacts (Rigby et al. 2023). Our flat-field calibration image is custom made from on-sky commissioning data (COM-1063), updating the official calibration files to correct for smoothly varying large-scale structure in the flats and to further optimize pixel-to-pixel variations. The data reduction pipeline next subtracts a large-scale sky background, performs an astrometric alignment (see Section 2.4), identifies and removes hot pixels, and drizzles the images to a common pixel grid of  $0''.02$  pixel $^{-1}$  for SW bands and  $0''.04$  pixel $^{-1}$  for LW bands using ASTRODRIZZLE (Gonzaga et al. 2012).

### 2.2. HST

A wide range of imaging of the A2744 cluster and surrounding area exists within the public HST archive. Briefly, we summarize the programs relevant herein. Program HST-GO-11689 (PI: Dupke) and HST-GO-13386 (PI: Rodney) include deep HST/Advanced Camera for Surveys (ACS) imaging in the cluster center in three filters (F435W, F606W,

<sup>44</sup> <https://github.com/gbrammer/grizli>

and F814W), with program HST-DD-13495 (PI: Lotz; Lotz et al. 2017) acquiring complementary deep HST/Wide Field Camera 3 (WFC3) observations in four filters (F105W, F125W, F140W, and F160W). While each of the above programs are (deep) individual pointings limited by the ACS and WFC3 fields of view (FOVs), respectively, the data were later expanded by a factor of 4 with shallower imaging in two ACS filters (F606W and F814W) and three WFC3 filters (F105W, F125W, and F160W) by the BUFFALO survey (Program HST-GO-15117, PIs: Steinhardt & Jauzac; Steinhardt et al. 2020). Most recently, the deep optical coverage was further expanded by Program JWST-DD-17231 (PI: Treu). A summary of the instruments, filters, program IDs, and orbit depths can be found in Table 3 in Bezanson et al. (2022). Taken together, they contribute seven unique HST filters (Figure 1, see also Table 1). These images are reduced following the same procedure as described in Section 2.1 onto the same  $0''.04$  pixel grid as the JWST images.

### 2.3. Astrometry

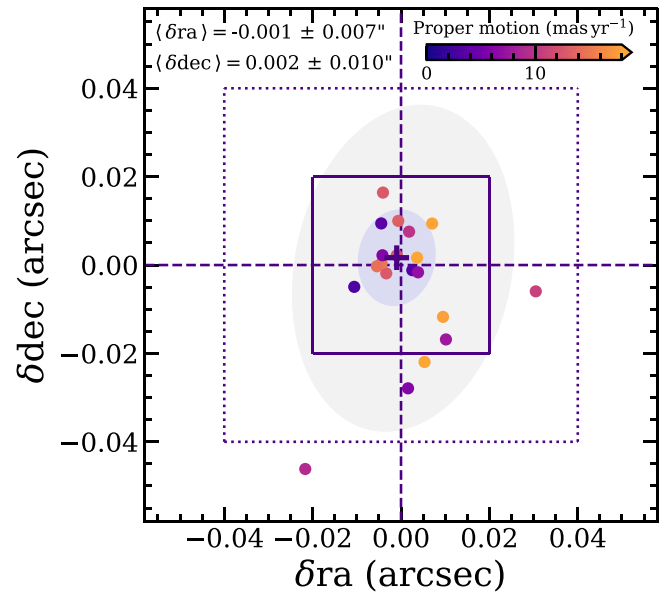
Astrometric registration of the images is performed by GRI<sub>Z</sub>LI using F444W. We adopt star positions from Gaia DR3 (Gaia Collaboration et al. 2016, 2023). Using their well-known proper motions, the positions of Gaia stars observed in 2015 are projected to 2022 November, the epoch during which the JWST imaging was acquired. The remaining images are then registered consistently to the F444W filter. In order to test the accuracy of the astrometry independently, we opt to compare to stars in the F160W filter instead of F444W, where saturation and central star clipping are less of an issue. We perform an additional correction to the proper motions to shift to the median epoch of the wider F160W BUFFALO HST imaging in 2019 July (Steinhardt et al. 2020). Figure 2 demonstrates our achieved precision of  $\approx 0''.008$  or one-fifth of a pixel, measured by the standard deviation of the median absolute deviation for the innermost 50% of sources (purple shaded region). The median bias, based on the same sources, is also small at  $\approx 0''.002$ , or 5% of a single pixel.

### 2.4. Spectroscopic Redshifts

Spectroscopic redshifts (hereafter  $z_{\text{spec}}$ ) over our survey footprint are taken from a compilation by Kokorev et al. (2022). We find 518 secure entries with  $z_{\text{spec}}$  values with confidence flags 3 or 4 within  $0''.3$  of our sources. Generally, a flag of 3 or 4 notes a secure  $z_{\text{spec}}$  from a single strong emission line, multiple emission lines, or easily distinguished continuum features; note that flagging definitions can vary somewhat between subsets of the compilation. We find 156 in the NASA/IPAC Extragalactic Database,<sup>45</sup> where 340 are cluster members with  $z_{\text{spec}}$  from Richard et al. (2021) with the Very Large Telescope/MUSE and 22 are grism redshifts from GLASS (Treu et al. 2015). The corresponding values are stored in the catalog in the  $z_{\text{spec}}$  column. Please note that we do not include any  $z_{\text{spec}}$  from the UNCOVER follow-up program (see Price et al. 2023).

## 3. Removal of Sky, ICL, and Bright Cluster Galaxies

In order to achieve the science objectives, we need to first mitigate the contamination of foreground light from the many



**Figure 2.** Astrometric performance of the imaging data set based on the positions of bright stars in HST/F160W compared to Gaia DR3. 1 and 4 pixel areas are shown by the solid and dashed squares, respectively. Filled purple and gray elliptical contours enclose roughly 50% and 90% of bright stars, respectively. The median deviation (purple cross) and the standard deviation of the absolute median deviation are quoted for each axis corresponding to the innermost 50% of stars.

bright cluster galaxies (bCGs) and the powerful ICL. Otherwise, the photometry of distant sources seen through this foreground light will be inaccurate, potentially mischaracterized, or missing altogether, the rare high- $z$  galaxies magnified by the strong gravitational lensing of these cluster members. We note that throughout this paper we adopt the term bCG, which is not synonymous with the traditional brightest cluster galaxy.

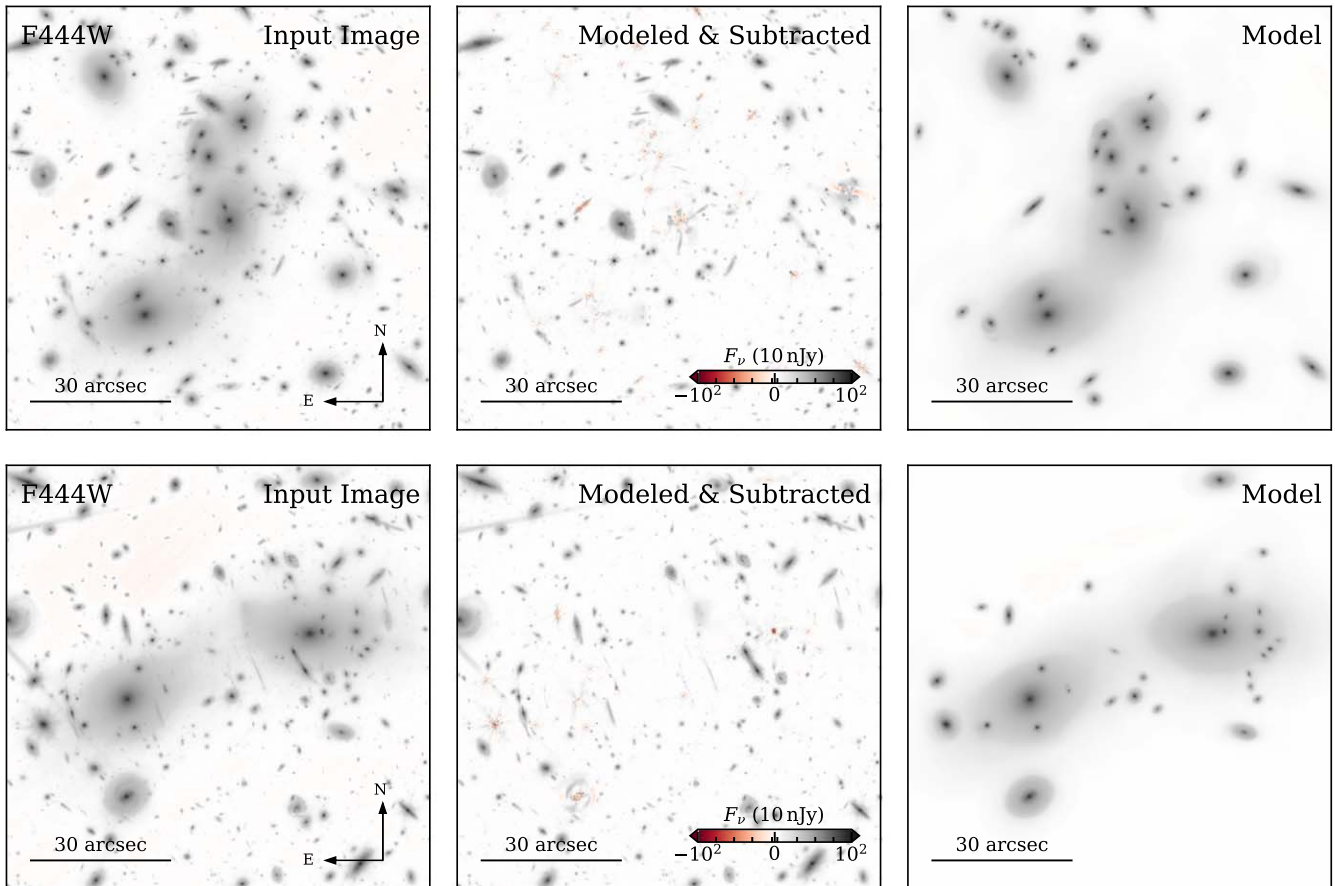
### 3.1. Subtracting Fitted Models

For robust and tested bCG and ICL subtraction, we adopt the method described in Ferrarese et al. (2006) and implemented by Shipley et al. (2018) in the HFF-DeepSpace (HFFDS) photometric catalogs of six lensing clusters, including A2744.

The bCGs that contribute significantly to the total cluster luminosity are first identified from the HFFDS catalogs; we refer the reader to Shipley et al. (2018) for a more detailed description of the selection process. We further expand our selection to accommodate the wider footprint of the present data set (see Figure 2 of Bezanson et al. 2022). We note that fewer bCGs are subtracted from the F410M mosaic, which has a smaller footprint. To minimize computation time, we generate “cropped mosaics” using the IRAF IMCOPY task. The boundaries of these mosaics are defined by the outermost isophotal cluster radii.

We use SOURCE EXTRACTOR to create a crude mask of all background sources (excluding cluster members). This is done by using the parameters `DETECT_THRESH = 1.2`, `DEBLEND_NTHRESH = 10`, and `DEBLEND_MINCONT = 0.01`, which identifies more isolated sources. We repeat this detection on the masked image, providing a more accurate and precise mask, especially near tightly clustered galaxies. These two masks are combined to isolate the cluster galaxies that are ultimately smoothed with a Gaussian kernel to account for nearby, poorly modeled light.

<sup>45</sup> <https://ned.ipac.caltech.edu/>



**Figure 3.** Zoom-in view around the two primary cluster cores of A2744 in F444W. The bCGs and ICL visible in the input image (left) are removed by subtracting fitted models (middle), with the models themselves on the right.

Using this mask to isolate each cluster member, we run ELLIPSE to measure and extract the isophotal parameters out to an arbitrary radius. This is then given to BMODEL to create the galaxy model. This galaxy model is subtracted from the cropped mosaic. This process is repeated for each cluster member, yielding an initial residual image.

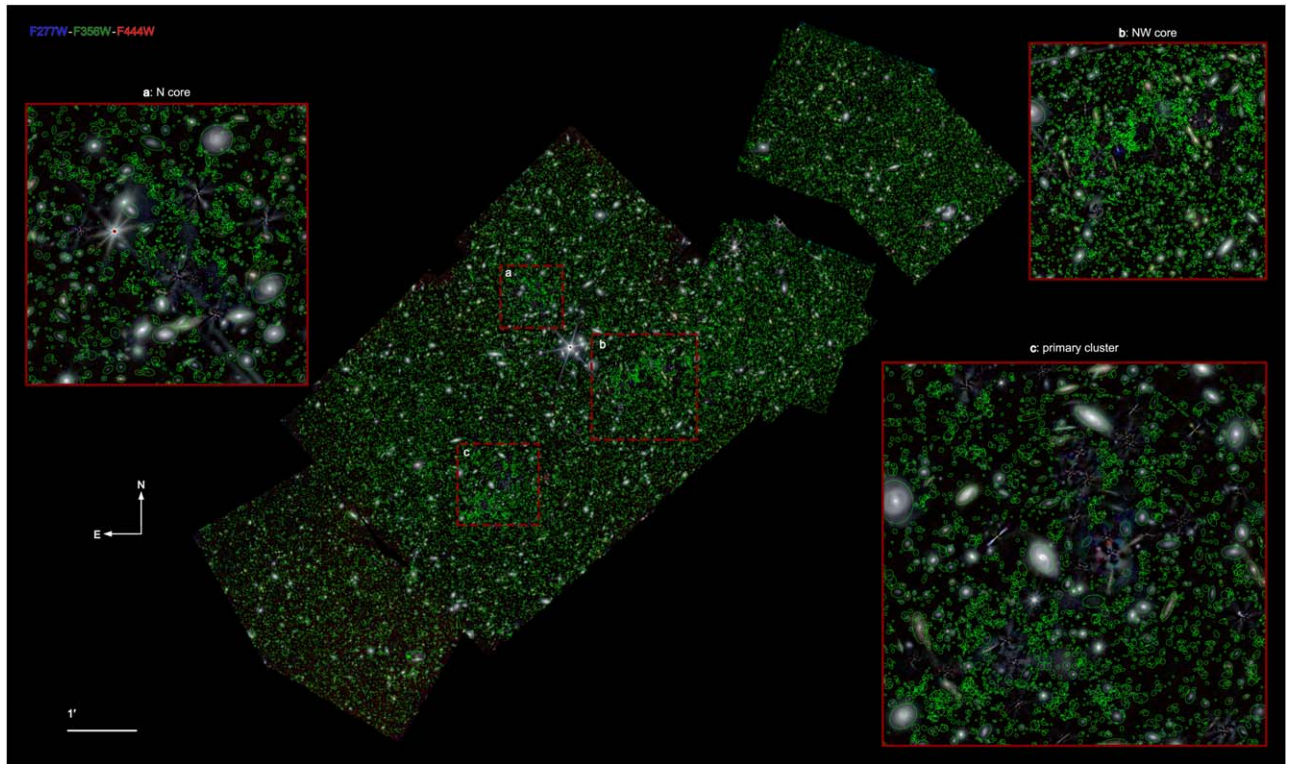
Although these initial models provide a good approximation of the bCGs and ICL, we adopt an iterative approach, repeating the modeling process on residual mosaics to improve the models and subsequent residuals. To do this, we construct intermediate images whereby all but one galaxy is subtracted. We model and subtract the best-fit model, and repeat the process for the rest of the sample. This allows us to model better the galaxy without any contamination from neighbors. We see a convergence after 10 rounds after which the models do not visibly change (see Shipley et al. 2018 for details). For crowded regions with multiple bCGs we allow for SOURCE EXTRACTOR to recreate and improve the original mask as a replacement for the primary mask. The improvement is most noticeable for galaxies whose masks are strongly affected by multiple nearby bCGs.

The final model images are produced by averaging over the 10 individual model images from each iteration. As found by Shipley et al. (2018), further noise reduction of the averaged model image can be achieved by rejecting the lowest four and highest two values on a pixel-by-pixel basis. We use the IRAF task IMCOMBINE with the parameters `combine = "average,"` `reject = "minmax,"` `nlow = "4,"` and `nhigh = "2"` (see

Section 3.1.3 of Shipley et al. 2018 for details). Finally, we subtract the average galaxy model from the cropped mosaic to produce the final residual mosaic. Figure 3 shows the effect of this careful bCG and ICL subtraction relative to the original mosaic near the primary cluster cores.

With this final subtracted mosaic, we revert our initial cropping by using the IRAF IMCOPY task. We copy our subtracted mosaic onto the original mosaic. We then perform a sky subtraction to remove excess light near the edges of the galaxy models using a Gaussian interpolation. The background is measured using the SOURCE EXTRACTOR AUTO setting with a mesh size of 192 for SW bands ( $0''.02$  scale) and 96 for LW bands ( $0''.04$  scale), limiting magnitude of 15, and a maximum threshold of 0.01. The background subtraction does not significantly change the residual mosaic, but near the borders where the differences are well defined due to our initial cropping this step smooths the previously defined edges.

As shown in Figure 3, there are residuals left near the bCG centers due to real structures not described by the smooth elliptical models. Given that the positive residuals are liable to be detected as sources and that photometry within their vicinity will be unreliable, we opt to mask objects detected near the known bCGs (see Section 4.6 for details). Furthermore, there is some stray ICL not immediately associated with known bCGs that is bright enough to not be removed by the final background subtraction. While photometry of objects detected near this stray residual ICL will be biased, the regions affected make up less than 1% of the total area.



**Figure 4.** Color composite image of the JWST footprint of A2744, with three cluster cores highlighted. bCGs and ICL have been subtracted. Apertures consistent with our super photometric catalog are shown in green for reliable objects (see Section 4.6).

**Table 2**

SEP Parameters Used for Source Detection in the Noise-equalized F277W +F356W+F444W Coadded Image

Name	Value
KERNEL	3.5 pixel FWHM Gaussian
MINAREA	3 pixels
THRESH	$1.2\sigma$
DEBLEND_NTHRESH	32
DEBLEND_CONT	0.0001
CLEAN	N

**Note.** Not supplying a weight map implicitly tells SEP to use `THRESH_TYPE = ABSOLUTE`, suitable for noise-equalized detection images.

#### 4. Source Detection and Photometry

In this section we discuss the construction of our catalogs. Briefly, the subsections include source detection, PSF homogenization, aperture photometry, corrections to photometry accounting for magnification via strong gravitational lensing, identification of star candidates, and a general recommended “use” flag. Catalogs corresponding to DR2 are available online (see footnote 42). The software used to produce these catalogs, `aperpy`, is generally applicable to any JWST/NIRCam data and is freely available (see footnote 43). We report flux densities and their uncertainties in  $F_\nu$  units of 10 nJy, corresponding to an AB magnitude zero-point of 28.9.

##### 4.1. Source Detection

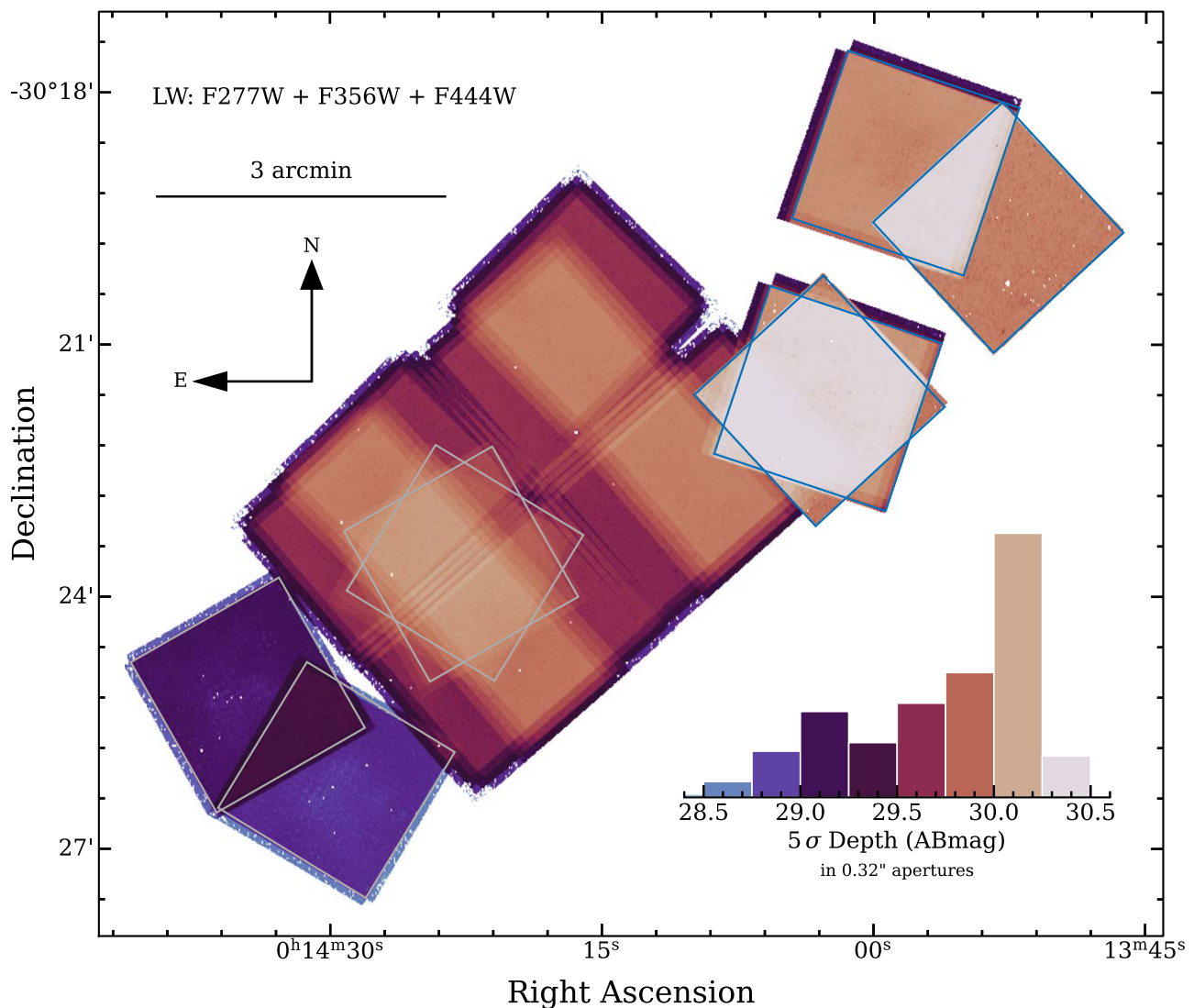
Sources are detected on a sky-subtracted noise-equalized (i.e., inverse variance weighted) coadded image based on our deepest JWST imaging in the three LW broadband filters

F277W, F356W, and F444W. Detection is performed with SEP (Barbary 2016), adopting the configuration listed in Table 2. While aperture photometry is performed on PSF-matched images (Section 4.2), we combine LW images at their native resolution to maximize sensitivity to identify faint sources. We detect 61,648 sources across the 49 arcmin<sup>2</sup> area. Figure 4 shows a RGB image of the bCG-subtracted detection coadded images with ellipses marking all unflagged (see Section 4.6) sources detected from the bCG-subtracted LW coadded image. The effective catalog depth in 0.32 arcsec apertures of the noise-equalized LW bCG-subtracted coadded detection image is shown in Figure 5 (see Section 5.1 for further details).

##### 4.2. PSF Matching

Before extracting aperture photometry, the PSF of each image is matched (or “homogenized”). This approach allows for consistent photometric measurements within the same aperture size across all bands, which leads to a better recovery of source colors,  $z_{\text{phot}}$ , and physical parameters. We adopt our longest-wavelength NIRCam band, F444W, as our target PSF. This choice is motivated by F444W being our broadest NIRCam PSF, meaning that the corresponding images are matched to the lowest resolution. Additionally, F444W will probe the reddest rest-frame light (e.g., 1–2  $\mu\text{m}$  stellar bulk) at  $z \gtrsim 1$ , making it an ideal band with which to derive total fluxes from our aperture photometry. Preserving the original F444W image properties will maximize consistency within the photometry.

Following methodology described in Skelton et al. (2014) and Whitaker et al. (2019), we generate empirical PSFs in all HST and JWST bands using stars identified within the FOV. Point sources are known to inhabit a locus within a size–magnitude plane where size is approximated here by the ratio



**Figure 5.** Schematic of the depth variation across the noise-equalized LW coadded detection image aggregated from the DDT (gray), GLASS (blue), and UNCOVER programs. Combining F277W, F356W, and F444W, the  $5\sigma$  depths in our LW detection band measured in  $0''.32$  apertures span 28.5–30.0 mag. Poisson contributions of the brightest objects feature prominently in our JWST weight maps. Although provided in our mosaic release, the UNCOVER NIRISS parallel imaging is not yet cataloged and so is not shown.

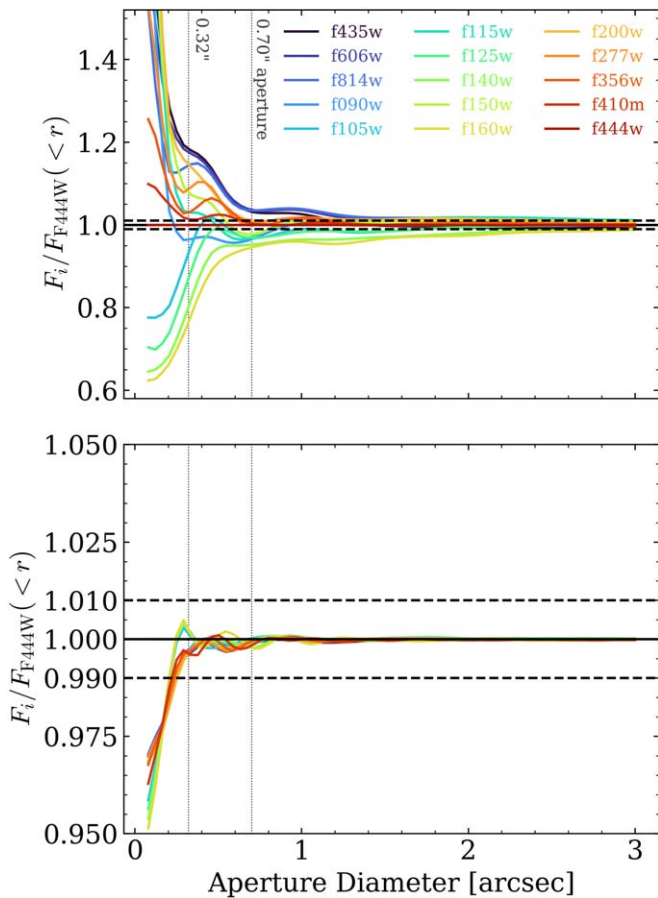
of flux in  $0''.16$ – $0''.32$  diameter apertures and magnitude within  $0''.32$  diameter apertures. Instead of a selection box, we fit the slope based on a first pass selection of stars. Sources with aperture magnitudes fainter than 24 AB or outside the locus by  $2.8\sigma$  are rejected, with the total number of candidates in each filter varying from 14 (F090W) to 120 (F814W). The candidates are extracted in stamps, recentered using cubic interpolation based on their center of mass evaluated in a window around the initial stamp center, and then normalized to unity within 16 pixels ( $0''.64$  diameter). Final PSFs for each filter are constructed by averaging the centered and normalized stamps, discarding any pixel whose flux is outside  $3\sigma$  of the flux distribution at that position. Lastly, we renormalize each empirical PSF such that its energy enclosed within a  $4''$  diameter aligns with typical calibration levels.<sup>46</sup>

<sup>46</sup> HST/ACS: <https://www.stsci.edu/hst/instrumentation/acs/data-analysis/aperture-corrections/>; HST/WFC3: <https://www.stsci.edu/hst/instrumentation/wfc3/data-analysis/photometric-calibration/ir-encircled-energy/>; JWST/NIRCam: <https://jwst-docs.stsci.edu/jwst-near-infrared-camera/nircam-performance/nircam-point-spread-functions>.

We choose to build empirical PSFs for JWST/NIRCam as opposed to using the simulated PSFs from WebbPSF (Perrin et al. 2012, 2014) as we find that all PSFs provided by WebbPSF are 1%–2% narrower than real stars in our mosaics even after accounting for broadening introduced during image reduction. Further details can be found in Appendix B.

We note that our relatively simple technique does not consider the inhomogeneous position angles (PAs) of the various observations in each mosaic. For example, due to a guide star failure one of the four visits of UNCOVER was observed at a slightly different PA whose effect is visible in the mosaics. This is in addition to the significantly different PAs of the other two programs in the field. Additionally, while the expected spatial and temporal variations of the PSF as measured in circular apertures appear negligible based on predictions from WebbPSF (see Appendix C), this has not yet been measured empirically. A more sophisticated treatment of the PSF including spatial and rotational variations will be explored in future work.

Kernels are produced using PYPHER (Boucaud et al. 2016). PYPHER generates PSF-matched kernels using an algorithm



**Figure 6.** PSF growth curves for each filter before (top) and after (bottom) matching to F444W. After matching, all filters have deviations below the 1% level at the smallest aperture diameter used ( $0''.32$ ). Growth curves are shown relative to the F444W growth curve; a value of 1 indicates perfect matching with F444W. Dashed lines indicate the  $\pm 1\%$  deviations from exact matching (solid black line). Dotted lines indicate the location of  $0''.32$  and  $0''.70$  aperture diameters.

based on Wiener filtering (Wiener 1949) with a tunable regularization parameter, which we set to  $3 \times 10^{-3}$ . Matching is done in  $3 \times$  oversampled space, after which the kernels are rescaled to the original pixel scale. The lower value of the regularization parameter helps avoid high-frequency noise in the kernels, while still maintaining  $< 1\%$  deviations at all aperture diameters of interest. All filters are matched to the reference filter (F444W). Figure 6 shows the PSF growth curves of every filter relative to the F444W growth curve, both before (top) and after (bottom) convolving with a matching kernel. Matched PSFs have almost identical growth curves to the reference filter, with deviations below the 1% level. The HST NIR filters (F105W, F125W, F140W, and F160W) have the most variation, but this is only significant for aperture diameters smaller than  $0''.32$ . This discrepancy is unavoidable given the systematic differences in the shapes of the PSF between HST and JWST, where the latter has significantly more substructure and “snowflake-like” patterns that complicate the PSF matching. This substructure should not significantly impact our photometric measurements, as it is azimuthally averaged out within the circular apertures.

Alternatively, we have similar success matching PSFs with PHOTUTILS (Bradley et al. 2022), which uses ratios of Fourier transforms to generate a matching kernel. PHOTUTILS also

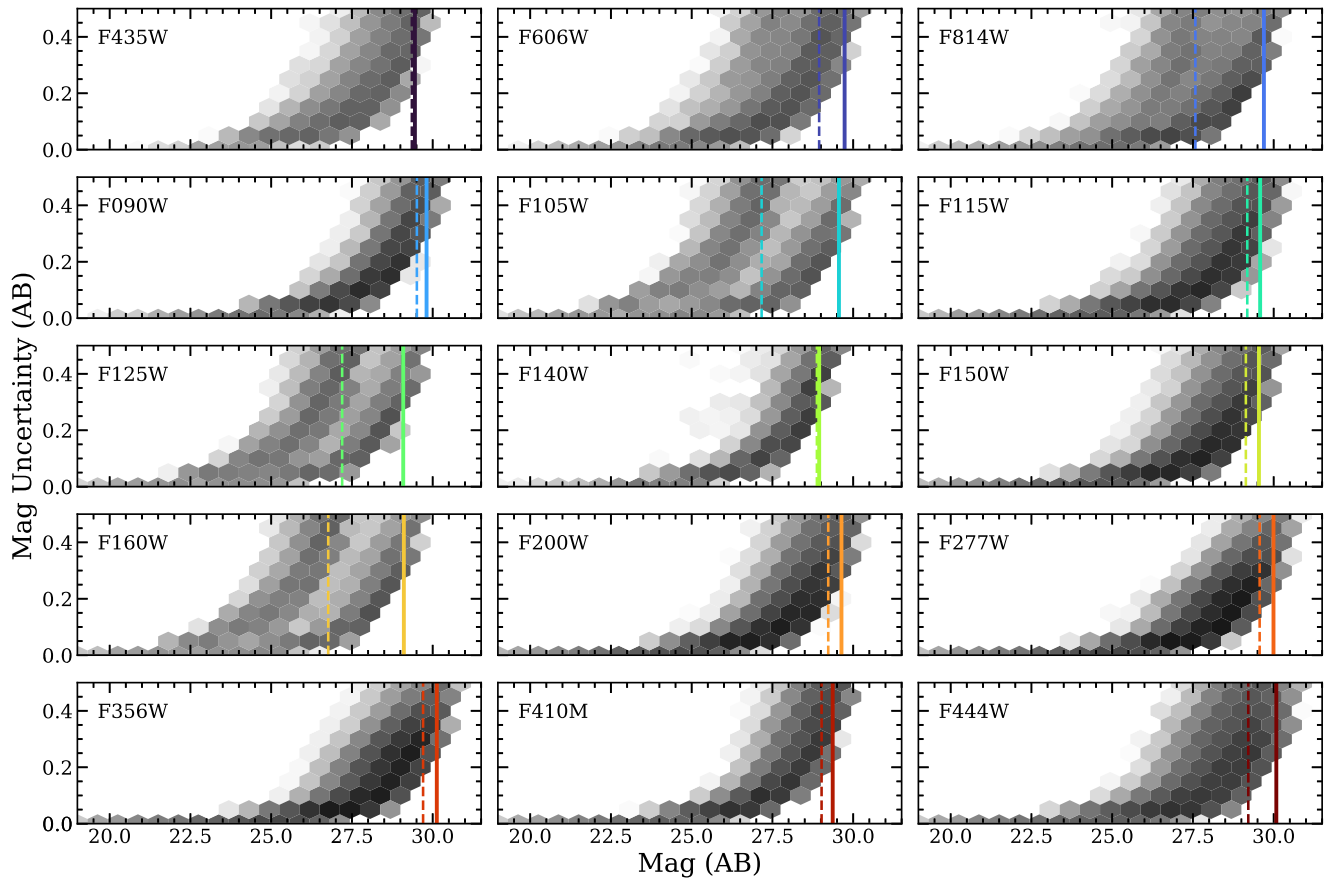
requires the selection of one of several window functions, used to filter high-frequency noise from the Fourier ratios. The best window function for a given reference PSF and the ideal values for the tuning parameters used to scale the window function are not straightforward and usually require testing different combinations. While PHOTUTILS has similar PSF-matching success to PYPHER, it is less convenient due to the additional parameters that need tuning. Finally, we find that PSF-matching methods that utilize linear combinations of shapelets (e.g., Gauss–Hermite or Gauss–Laguerre polynomials) to generate kernels (e.g., Skelton et al. 2014) should be avoided when matching to JWST filters. The intrinsic symmetry of these functions cannot effectively match the rotational asymmetry of the diffraction spikes and extended structure of JWST PSFs. However, shapelet-based algorithms perform the best for PSF matching to an HST reference filter (e.g., F160W).

### 4.3. Aperture Photometry

Photometry is measured in  $0''.32$ ,  $0''.48$ ,  $0''.70$ ,  $1''.00$ , and  $1''.40$  diameter circular apertures with SEP. These “color” aperture measurements are then corrected to total fluxes by scaling them by the ratio of the flux estimated from elliptical Kron-like apertures (Kron 1980) to the flux measured in consistently sized color apertures. To ensure that all detected objects have robust Kron measurements, photometry is extracted on an F444W-matched inverse variance weighted F277W + F356W + F444W coadded image (consistent with the construction of our detection image) in elliptical apertures whose semimajor and semiminor axes are grown by a factor of  $2.5 \times$  Kron “radius” (a unit-less factor), with unity as the minimum allowed factor. Then, for sources whose circularized Kron radius is less than the circular aperture radius, the circular apertures are used instead of the Kron-scaled ellipse. The resulting measurements from this procedure are commonly referred to as “AUTO” flux densities. Note that unlike SOURCE EXTRACTOR, SEP does not mask neighboring objects by default and so may produce catastrophically large Kron radii in certain cases. As of v1.1, SEP now provides functionality to mask neighboring sources which in our case dramatically reduces the number of such failures. We leverage this functionality.

We additionally correct each measurement for light missed at large radii (especially important for JWST) by dividing the flux measurement of each source by the fraction of the total light from the F444W PSF curve of growth within each respective circularized Kron radius. We stress that this correction for the JWST bands in particular is on the order of 10%–20%, significantly larger than for HST due to the large fraction of light characteristically scattered to large radii in JWST PSFs.

Photometric uncertainties are derived by means of an independent estimate of the background noise. For each band we place 10,000 circular apertures in regions outside detected sources, within our detection footprint, and with good coverage in that band. We opt to use the segmentation maps to mark detected sources, making the placement of the apertures dependent not only on the footprint of the detection image, but also its union with the footprint for that particular filter. For a given filter, apertures are placed on the corresponding noise-equalized image to account for the variation in depth across the field (e.g., see Figure 5). Outlier measurements greater than  $5\sigma$  are removed, and the width of the flux distribution is estimated by the standard deviation, and is unit-less. To obtain physical noise estimates consistent with our photometry, the width is



**Figure 7.** Photometric uncertainty as a function of magnitude shown by  $\log_{10}$ -scaled 2D histograms for total fluxes derived in  $0''.32$  diameter apertures. Catalog depths for each filter measured in the same aperture size corresponding to the 50th and 10th percentile areas are indicated by the dashed and solid lines, respectively.

multiplied by the effective local noise around each source estimated by the inverse of the square root of the median weight computed within a  $9 \times 9$  pixel box. Total flux errors are then computed by multiplying the resulting per-object noise by the ratio of total to aperture flux measured on the F444W-matched F277W + F356W + F444W coadded image.

Figure 7 illustrates the growth of photometric uncertainty with magnitude. While the growth of photometric uncertainty in some HST bands shows dual loci corresponding to the shallower BUFFALO and deeper HFF observations, those of the JWST bands follow many loci which are blended together due to the multimodal depths produced by the overlapping DDT, UNCOVER, and GLASS programs (see also Section 5.1).

Given the extreme depths of our detection images and the crowded nature of galaxy clusters, 17,695 objects (29%) are flagged as potentially being blended (FLAG\_KRON). Estimates of Kron radii are known to fail in such cases, and so for these sources we do not correct to total flux using their likely corrupted Kron aperture fluxes. Instead we simply correct for the missing light beyond the color aperture using the F444W curve of growth, assuming point-source morphology. In most cases this will underestimate the total flux and some physical parameter estimates (e.g., stellar mass), while leaving their redshifts and colors robust. There also exist an additional 5561 sources (9%) small enough that their circularized Kron diameters are less than  $0''.32$  and so their total flux is robustly estimated by applying the same point-like correction to the total using the F444W curve of growth. Together these two subsamples make up the 23,256 objects (38%) flagged as USE\_CIRCLE.

While providing photometry in five different aperture sizes is useful for comparisons and cross-checks, it is liable to produce awkward workflows and confusion. To improve the accessibility of these early JWST catalogs, we additionally build a “super” catalog following Labbé et al. (2003) who use the isophotal areas for each object based on the detection image (i.e., the number of pixels in their segment). In short, unblended objects are assigned photometry corresponding to the smallest aperture that does not exceed the size of an aperture that would enclose the isophotal area. For blended objects, however, their true isophotal area is uncertain and so to be conservative we shrink the equivalent isophotal aperture diameter by 20%. This factor is dependent on the particular data set and is determined experimentally, see Section 5.2 of Labbé et al. (2003) for details. In rare cases where the most suitable aperture does not provide reliable photometry, such as a large aperture containing a masked pixel, the next largest usable aperture is chosen. Overall, 48,360 sources (78%) of the super catalog are assigned photometry based on  $0''.32$  apertures, with 5648 and 3267 sources being assigned photometry based on  $0''.48$  and  $0''.70$  apertures, respectively; 4373 sources are assigned either  $1''.0$  or  $1''.4$  apertures. We encourage users to default to this super catalog for most science applications.

Photometry is corrected for line-of-sight attenuation through the galaxy, adopting the dust maps of Schlafly & Finkbeiner (2011) and the attenuation law of Fitzpatrick & Massa (2007). Given the small footprint, we opt to apply the median  $E(B - V) = 0.01$ . The dust column density in the direction of

A2744 is favorably low, resulting in attenuation corrections in each band on the order of 1% or less. We report flux densities and their uncertainties in  $F_\nu$  units of 10 nJy, corresponding to an AB magnitude zero-point of 28.9.

#### 4.4. Source Magnification

A2744 is one of the most powerful lensing clusters known (e.g., Merten et al. 2011; Jauzac et al. 2015). According to the most recent estimates of all three cluster cores by Furtak et al. (2023b), A2744 magnifies most objects in our survey footprint by at least  $\mu = 2$ , with objects seen nearest to the cluster cores magnified by  $\mu \sim 10$ –100 (see Figure 5 in Furtak et al. 2023b). Magnifications and shear parameters for sources in each catalog are computed from the latest version (v1.1<sup>47</sup>) of the Furtak et al. (2023b) analytic strong-lensing mass model assuming  $z_{\text{phot}}$  derived from EAZ<sub>Y</sub> in Section 5.3 (and from  $z_{\text{spec}}$  where available). The v1.1 of the lens model includes an additional spectroscopic redshift from Bergamini et al. (2023) and two additional multiple image systems compared to v1.0 presented in Furtak et al. (2023b). The model is thus constrained by 141 multiple images belonging to 48 sources and achieves a lens plane rms of  $\Delta_{\text{RMS}} = 0''.51$ . Future work presented in Wang et al. (2023a) will provide updated magnification estimates using  $z_{\text{phot}}$  from PROSPECTOR- $\beta$  (Johnson et al. 2021; Wang et al. 2023b). Note that while the magnifications are provided in the catalogs, measurements (e.g., fluxes) are not corrected for magnification.

#### 4.5. Identifying Stars

Constructing reliable photometric catalogs of extragalactic sources requires identifying foreground stars and other spurious sources that would otherwise contaminate galaxy samples. While no one identifier is complete and strictly pure, stars in HST photometric surveys have been traditionally identified by the fact that they are typically bright and unresolved point sources, noting that not all point sources are stars (e.g., quasars). However, due to the staggering efficiency of JWST/NIRCam, stars of similar brightness often saturate the detector pixels making their identification surprisingly difficult.

To overcome this new obstacle, we first identify stars using traditional methods as described in Skelton et al. (2014), with fluxes measured on the original, non-PSF-matched mosaics. Figure 8 presents stars selected in F200W (blue) having a flux ratio in  $0''.70$  over  $0''.32$  apertures brighter than 25.0 AB. While this selection alone accounts for most of the stars, there are a handful that are better identified from HST F160W where their profiles are less affected by saturation. However, despite measuring stellar sizes on the native-resolution HST images, we found that the width of the stellar locus was surprisingly large. Further inspection revealed that a number of stars have sufficiently high proper motions that both of their JWST-derived centroids miss their HST-era positions. As a work-around, we have added an additional flag that identifies stars detected on the F160W image itself having a flux ratio in  $0''.70$  over  $0''.32$  apertures between 1.5 and 1.65, and are brighter than 23 AB in  $0''.7$  diameter apertures. For completeness, we also flag any object that contributed to a PSF model (see Section 4.2) that has a match in our catalog (green). A total of 113 (<1%) stars are flagged by at least one of these criteria.

Note that bright stars that are saturated in our LW detection image are not flagged as stars because they are not detected as single objects; instead each drives multiple spurious detections near their saturated cores that we classify as artifacts (see Section 4.6).

However, a point-like indicator is unsuitable for the faintest stars that are likely to intermix with the general galaxy population. Some literature studies have opted to identify stars by comparing their fit quality between galaxy and stellar templates, either alone or in combination with resolution criteria, e.g., Weaver et al. (2022). To enable these comparisons, we use EAZ<sub>Y</sub> to quantify the goodness-of-fit of our spectral energy distributions (SEDs) to theoretical PHOENIX BT-Settl stellar templates (Allard et al. 2012) and include the corresponding  $\chi^2$  estimates in our catalog. However, we do not use  $\chi^2$  to flag additional faint stars. Such a comparison is prone to rejecting incorrectly a nonzero number of potentially interesting and little known objects for which we have poor spectral templates, e.g., high- $z$  galaxies. Consequently, galaxy candidates fainter than  $\sim 25$  mag are liable to be contaminated by difficult to identify foreground stars and so require additional scrutiny.

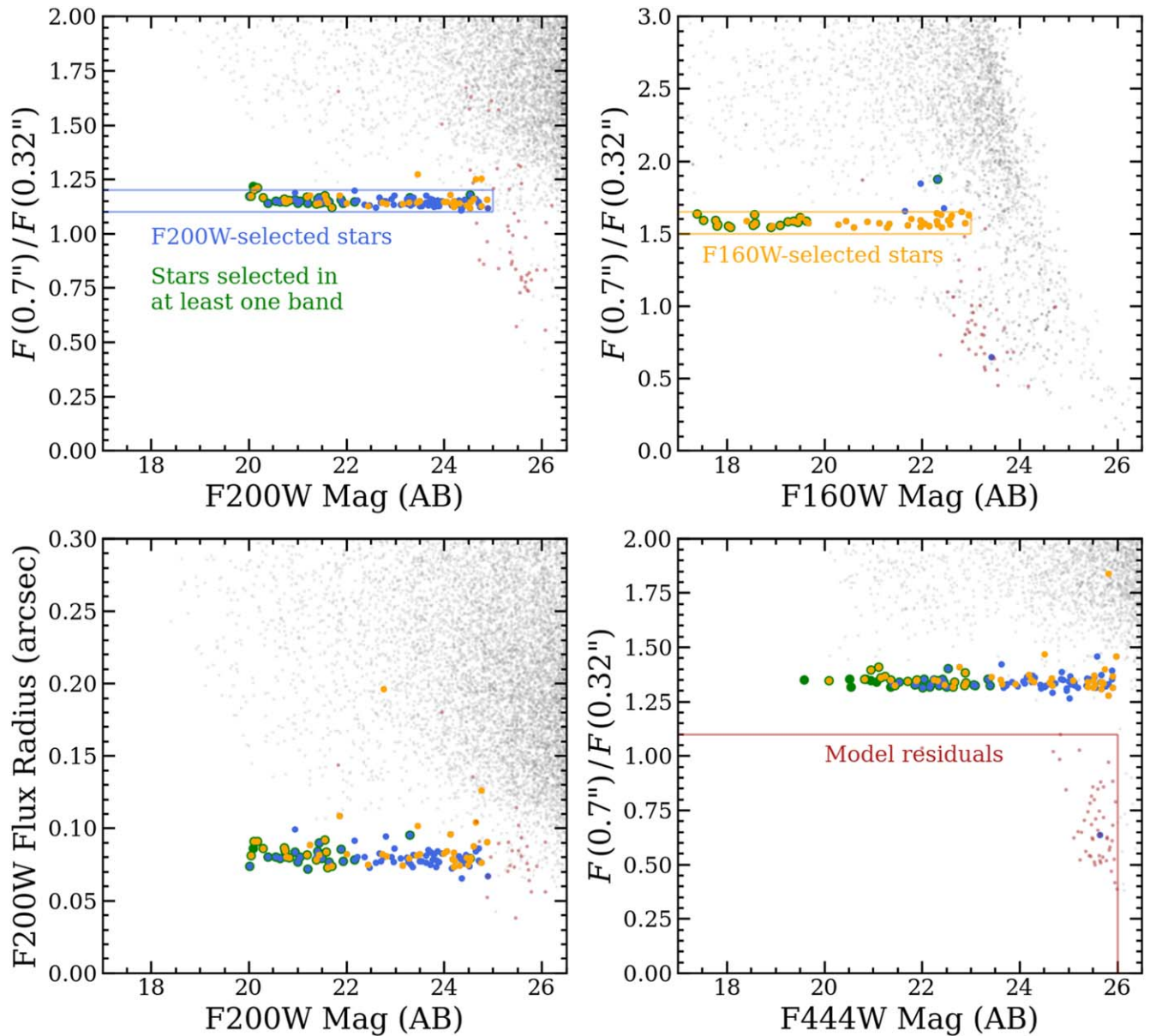
#### 4.6. Getting Started: The Use Flag

A “use” flag is particularly useful when familiarizing oneself with any given photometric catalog. Ideally, one simple selection yields a clean sample of galaxies across cosmic time for further analysis. Following Skelton et al. (2014), the USE\_PHOT flag in our photometric catalog requires stars to be removed (see Section 4.5) and sets a minimum signal-to-noise ratio (S/N) of 3 in a color aperture measured on the F277W + F356W + F444W coadded image to ensure robust Kron radii and aperture-to-total corrections. However, with the addition of the novel JWST photometry, there are further selections required to clean up the catalog sufficiently as described below. For this first generation, we opt to be more conservative at the expense of completeness to benefit more exotic parameter spaces (e.g., extreme high redshifts).

While our bCG modeling and subtraction enables searches for objects within the immediate vicinity of the bCGs, it also produces undesirable residual features in some cases which are detected in our catalogs. We find that these artifacts can be identified in F444W having a flux ratio less than 1.1 with magnitudes brighter than 26 AB (see the bottom right panel of Figure 8). Additionally, there are a number of artifacts found near the cores of bright stars that are saturated in JWST. We identify them as being within  $3''$  of the center of groups of saturated pixels at least 10 pixels in size. We also manually build a star spike mask and flag any object whose centroid is within the mask as an artifact. Several bright, saturated stars are flagged as artifacts as a bycatch; they are not flagged as stars even though they fall within the stellar locus of one or more bands in Figure 8. The catalog includes a column FLAG\_ARTIFACT to signal the 5112 objects (8%) satisfying any of these selections (and others, see below).

The areas immediately surrounding the subtracted bCGs are in all cases contaminated by residual structures. We further flag all detections within a conservative  $3''$  radius of a known bCG centroid computed as part of the modeling procedure. The catalog includes a column FLAG\_NEARBCG to signal the 3336 objects (5%) satisfying this selection. These objects are also flagged under FLAG\_ARTIFACT, making up 65% of the objects flagged as such. While we caution that some residual

<sup>47</sup> Publicly available at <https://jwst-uncover.github.io/DR1.html#LensingMaps>.



**Figure 8.** The selection of stars and spurious objects (bad pixels or model residuals) is determined from the flux ratio in large to small apertures above a limiting magnitude threshold. Stars are identified from the JWST/F200W image (blue) at positions corresponding to the catalog, and/or directly from the HST/F160W image (yellow) to account for high proper motion. We also include any star that contributed to PSF models in any band (green) that has a match to an object in the catalog. Some artifacts such as bCG model residuals can also be similarly flagged (red).

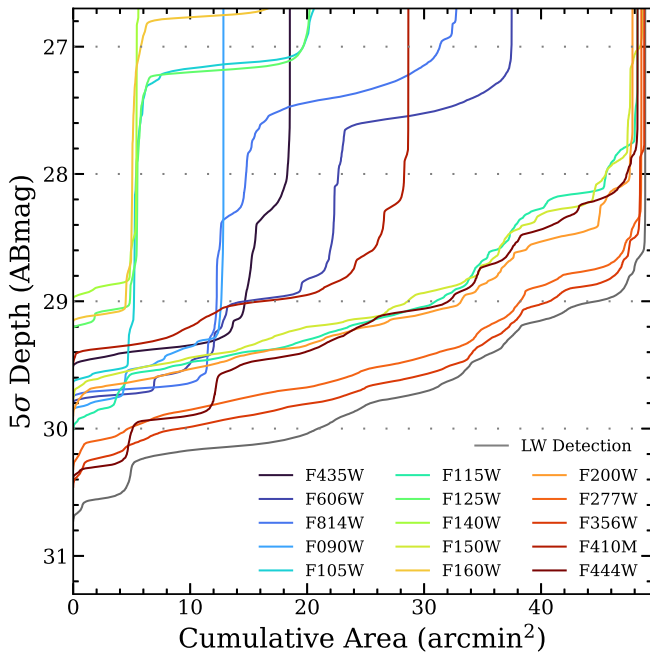
features outside this radius remain unflagged, as it is not possible to flag genuine residual features perfectly without also catching real sources of interest. A more sophisticated treatment will be explored in future work.

Together, we build a single `USE_PHOT` flag which removes undesirable objects, including known stars (`FLAG_STAR`), sources within  $3''$  of bCGs (`FLAG_NEARBCG`), as well as objects affected by bright star spikes, spurious detections near saturated star cores, and other model residual features (`FLAG_ARTIFACT`). While the super catalog contains zero sources with unreliable photometry (`FLAG_NOPHOT`), this number is nonzero in the single aperture catalogs, especially in the larger ones. To ensure the reliability of total photometry based on the F277W+F356W+F444W Kron radii, the “use”

flag also excludes 956 objects (2%) with an  $S/N < 3$  in the F277W+F356W+F444W color aperture, or was masked after detection (`FLAG_LOWSNR`). By selecting all objects where `USE_PHOT = 1`, this reduces our total sample to 55,613 galaxy candidates with reliable photometry.

## 5. Catalog Properties

With photometry in hand, we summarize the key properties of our catalog including effective catalog depths, galaxy number counts, photometric redshifts, and a brief comparison of measured JWST photometry to that expected from the best-fit SED template solutions. Although we provide catalogs based on five individual aperture sizes, in this section we will refer to the more broadly applicable “super” catalog.



**Figure 9.** Depth as a function of cumulative area for each filter mosaic as well as the LW detection image. Measurements are taken from 10,000 empty apertures of  $0''.32$  diameter, per filter. Gray dotted lines mark depths at 27, 28, 29, and 30 AB mag.

### 5.1. Photometric Depths

As introduced in Section 1, the A2744 imaging consists of three overlapping JWST and several overlapping HST programs. Consequently, the depth of any of these mosaics cannot be fully described by a single value. While the photometric uncertainties are computed to account for this variation, Figure 9 explicitly illustrates the effective catalog depth of each filter as a function of cumulative area corresponding to  $0''.32$  apertures. The depths of some of the HST mosaics are roughly bimodal as a result of the deep HFF observations contrasted with the wider but shallower BUFFALO coverage. Meanwhile for JWST, the contribution of the shallow DDT observation from NIRC*am*'s Module A is readily visible and the depths produced by including UNCOVER and GLASS create many small regions of varying depths elsewhere.

### 5.2. Galaxy Number Counts

We compute the galaxy number counts for our catalog in three JWST/NIRC*am* bands (F150W, F356W, and F444W) and one HST/WFC3 band (F160W), shown in Figure 10. The variation in image depth as well as source magnification make constructing number counts nontrivial, and so instead we estimate counts without correcting for magnification. We assume the nominal total science areas listed in Table 1. Areas for literature counts are taken from their respective papers. We remove any unreliable source with `USE_PHOT = 0`. We stress that these counts should not be used to quantify completeness nor survey depth precisely, but are merely an accessible means by which to validate our catalogs against literature measurements. Efforts to quantify completeness are in progress (R. Pan et al. 2023, in preparation).

While these lensed counts cannot be directly compared to those from unlensed field surveys, they can be readily compared to other lensed counts also from A2744 (Shipley

et al. 2018; Kokorev et al. 2022) where F150W/F160W, F356W/Ch1, and F444W/Ch2 are sufficiently similar such that deviations above 5% are significant (see Appendix A). To maximize consistency, we also remove any bCGs present in the literature counts that are removed as part of this work. We find good agreement with the literature counts at bright magnitudes ( $\lesssim 22$  AB) for all bands.

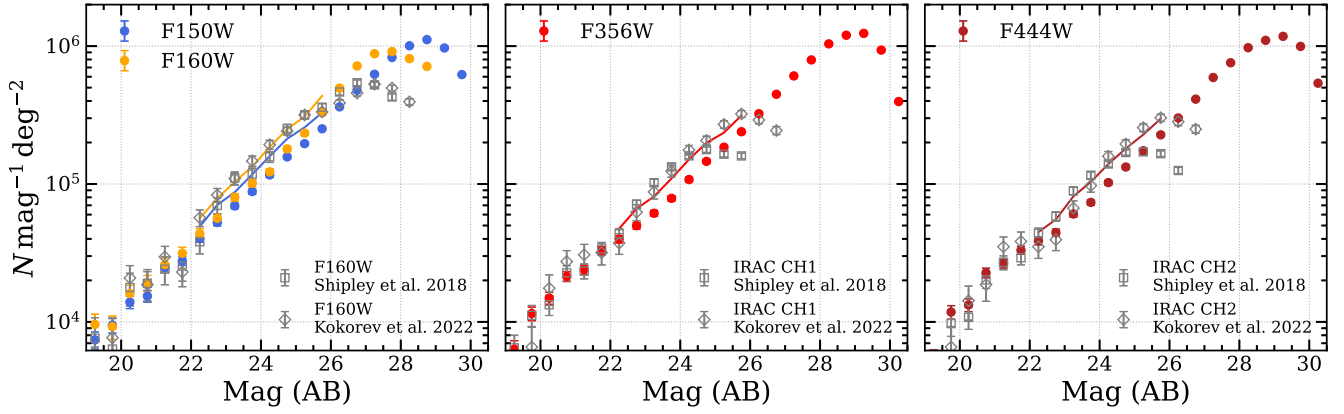
At intermediate magnitudes, our counts are lower due to real photometric differences. We hypothesize that the higher literature counts are caused by a combination of factors including blending from the shallower and lower-resolution F160W imaging, the relatively large Kron apertures used by the literature catalogs, and the conservative circular apertures adopted for blended objects in our catalog (which are not included in the comparisons in Appendix A). To study this further, we allowed all objects to have their aperture fluxes corrected using the Kron ellipses despite obvious cases of blending confirmed by visual inspection. Doing so brings our number counts in agreement with the literature (solid curves in Figure 10), suggesting that the literature counts are driven high by blending and hence justifying our conservative approach. In order to quantify the extent to which our fluxes are underestimated we compute aperture photometry on up a series of images containing PSF-convolved circularly symmetric galaxy light profiles with effective radii from van der Wel et al. (2014) corresponding to a  $10^9 M_\odot$  system at  $z = 1, 2, 3$ , and 4. Variation in the light missed due to the assumed Sérsic index and/or redshift is  $\lesssim 10\%$ . After applying aperture corrections assuming a point-like object, we find that our smaller  $0''.32$ – $0''.70$  apertures miss approximately half of the total light, with the larger apertures missing  $< 10\%$ . We note, however, that this is insufficient to elevate our number counts to match those in the literature and so magnification notwithstanding, the true counts likely lie somewhere in between.

At faint magnitudes, our deep LW-selected catalog provides significantly more objects near the depth limit of the images out to 27 AB in F160W relative to the literature. F150W counts go further still to  $\approx 29$  AB, although they do not include the very bluest objects missed by our redder LW selection function. Furthermore, our F356W and F444W counts extend to significantly fainter populations compared to existing Spitzer/IRAC data, making them some of the deepest ever obtained at these wavelengths. Although elementary, these observed galaxy number counts serve to demonstrate that we may confidently springboard from well-studied HST surveys to probe orders of magnitude deeper with JWST.

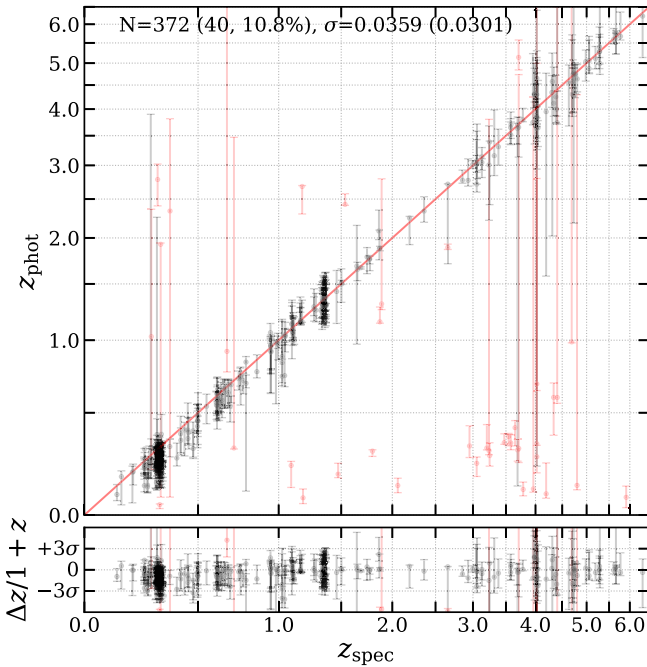
### 5.3. Photometric Redshifts

In order to provide an impression as to photometric performance, we compute  $z_{\text{phot}}$  using EAZY (Brammer et al. 2008). We use all HST and JWST bands available for each object and set a minimum error floor of 5%, an increase from the default 1% to reflect the calibration uncertainties of JWST/NIRC*am* more realistically. Given the considerable uncertainty as to the real high- $z$  galaxy SEDs, we forgo the usual preprocessing step of iteratively tuning zero-points to avoid biasing our colors to those of the SED templates. Furthermore, any disagreements between our photometry and that predicted from EAZY's models will be more readily visible (see Section 5.4). We also turn off both magnitude and  $\beta$ -slope priors for similar reasons.

We compute  $z_{\text{phot}}$  for all five sets of photometry independently based on our bCG-subtracted imaging which are then combined to complement the super catalog. We additionally



**Figure 10.** Galaxy number counts for sources identified from our LW detection image and measured in total magnitudes from F150W, F160W, F356W, and F444W. Stars and objects with unreliable photometry are removed. Effects of blending at intermediate brightness are shown by the colored curves where we allow Kron aperture corrections for blended sources. Literature F160W-selected counts in A2744 by Shipley et al. (2018) and Kokorev et al. (2022) are shown by the unfilled gray boxes and diamonds, respectively; note that bCGs have been excluded. Error bars denote  $1\sigma$  Poisson uncertainties. Counts are not corrected for magnification.



**Figure 11.** Performance of  $z_{\text{phot}}$  assessed by comparison with known  $z_{\text{spec}}$  described in Section 2.4. Outlier sources with  $z_{\text{phot}}$  wrong by more than  $0.15\Delta z/(1+z_{\text{spec}})$  are colored red. The lower panel shows residuals in normalized units of  $\sigma$  away from  $z_{\text{spec}}$ .  $N=372$  sources are compared finding 40 outliers (10.8%), an overall tightness  $\sigma_{\text{NMAD}}=0.0359$  and an outlier-removed  $\sigma_{\text{NMAD}}=0.0301$ .  $z_{\text{phot}}$  are derived from super catalog SEDs fit with EA $z$ Y using the SFHZ\_CORR template set.

compute separate  $z_{\text{phot}}$  from two SED template sets: the default FSPS\_FULLL set used frequently in the literature, and the newer SFHZ\_CORR set which features  $z$ -dependent priors on allowable star formation histories and an observed  $z=8.5$  emission line galaxy spectrum from Carnall et al. (2023) to provide realistic line ratios—two considerations especially important for identifying robust high- $z$  galaxy candidates. We set  $z=20$  as an upper limit.

One popular way of judging photometric accuracy is to compare  $z_{\text{phot}}$  and  $z_{\text{spec}}$ , where available. Of the  $N=518$  spectroscopic sources (see Section 2.4), we remove  $N=146$  of them including 10 failed EA $z$ Y fits, and 136 unreliable objects include bright bCGs subtracted in our mosaics, stars, and

artifacts (`USE_PHOT=0`).  $z_{\text{phot}}$  performance is then assessed using the normalized median absolute deviation (NMAD; Hoaglin et al. 1983), defined following Brammer et al. (2008),

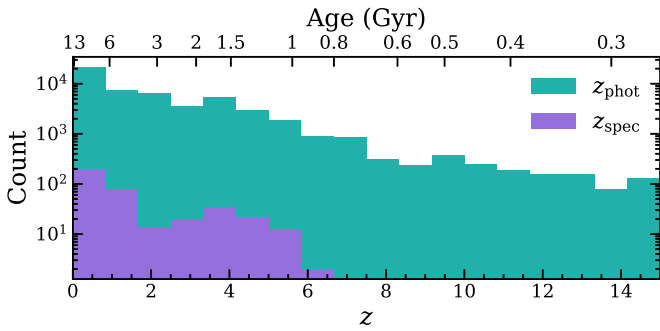
$$\sigma_{\text{NMAD}} = 1.48 \times \text{median} \left( \frac{|\Delta z - \text{median}(\Delta z)|}{1 + z_{\text{spec}}} \right), \quad (1)$$

as it is less sensitive to outliers compared to other definitions (e.g., Ilbert et al. 2006). We additionally quantify the outlier fraction as the fraction of objects with  $|z_{\text{phot}} - z_{\text{spec}}| \geq 0.15(1 + z_{\text{spec}})$ .

In general, the  $z_{\text{phot}}$  performance in comparison to  $z_{\text{spec}}$  is equally good and essentially agnostic to the photometric color aperture size. This, however, is expected as objects in our spectroscopic sample are notably brighter than many of the sources found in this new, deep imaging. This has two consequences: first, the assessable  $z_{\text{phot}}$  performance does not depend strongly on source magnitude and second, the  $z_{\text{phot}}$  performance reflects that of bright and easy-to-measure sources.

Unsurprisingly, we find significantly better  $z_{\text{phot}}$  performance using the SFHZ\_CORR SED templates for higher-redshift sources  $z \gtrsim 1$  compared to compared to the default FSPS\_FULLL that produces about twice as many catastrophic  $z_{\text{phot}}$  underestimates. Therefore the addition of  $z$ -dependent star formation histories and realistic line emission serves to enhance the ability of EA $z$ Y to recover the  $z_{\text{phot}}$  of known, distant spectroscopic sources correctly. However, we find an unexpected  $z_{\text{phot}}$  bias for  $z \lesssim 1$  sources when using SFHZ\_CORR that is not present when using FSPS\_FULLL. The sources of this bias is not yet understood, but given the success with FSPS\_FULLL, it seems unlikely to be driven by issues in photometry. Further exploration of this effect will be left to future work when the field will be complemented with a slew of grism redshifts and medium-band data as part of several Cycle 2 programs, and so here we limit our use of EA $z$ Y to basic photometric validation in the following section. For simplicity in this early catalog aimed at studying the high- $z$  Universe, we report only  $z_{\text{phot}}$  computed with SFHZ\_CORR templates and caution that the redshifts of  $z \lesssim 1$  objects may be underestimated.

As shown in Figure 11, we achieve a  $\sigma_{\text{NMAD}}=0.0301$  after removing the 10.8% outlying sources, some of which are likely due to wrongly identified emission lines in the spectra. Figure 12 shows a comparison of the distribution of the total



**Figure 12.** Distribution of  $z_{\text{phot}}$  (teal) and  $z_{\text{spec}}$  (purple) shown up to  $z = 15$ . The  $z_{\text{phot}}$  values are derived from super catalog SEDs fit with EA $z$ Y using the SFH $z_{\text{CORR}}$  template set.

$z_{\text{phot}}$  sample to the redshifts of the generally lower- $z$ , bright spectroscopic sample. It is also worth noting that applying iterative zero-point corrections does not significantly improve the overall performance as assessed by our limited spectroscopic sample. This also simplifies interpreting differences between our photometry and the predicted photometry from EA $z$ Y as discussed below in Section 5.4.

For transparency, we provide the  $z_{\text{phot}}$  and several common rest-frame fluxes (e.g.,  $UVJ$  of Williams et al. 2009;  $ugi_s$  of Antwi-Danso et al. 2023) for all objects in the catalog. However, more sophisticated methods utilizing extensive physically based priors and advanced sampling techniques will enable even more robust  $z_{\text{phot}}$  and physical parameters (e.g., stellar mass). Wang et al. (2023a) details how we have applied Prospector- $\beta$  (Johnson et al. 2021; Wang et al. 2023b) to this end.

#### 5.4. Validation of JWST Photometry

Given that JWST is a relatively new facility, much is still to be learned about its performance. Consequently, photometric extractions of JWST imaging in the literature are only just becoming available and so the traditional demonstration of comparing our on-cluster JWST photometry to the published literature is not possible (see Merlin et al. 2022; Appendix D).

Thankfully, computing  $z_{\text{phot}}$  for every object provides access to the observed-frame fluxes predicted by the corresponding best-fit SED template.<sup>48</sup> While this is arguably a circular exercise, the SED template combinations allowed by EA $z$ Y still only span a finite volume in color–color space. By attempting to minimize the  $\chi^2$  of the fit, EA $z$ Y is effectively maximizing its agreement with the error-weighted colors that we provide it; yet large-scale biases in individual filters should stand out given our well-sampled 15 filter SEDs.

Figure 13 shows the photometric agreement  $\Delta\text{Mag}$  between our measured fluxes and those predicted by the best-fit model from EA $z$ Y using the best-fit template set. While there are no significant offsets, we note that in a handful of filters the difference trends upwards at the 50th percentile depth but flattens toward the 10th percentile depth. This behavior correlates with the variation in depth and so can be readily explained by differences in the photometric performance of similarly bright objects with a range (or bimodality) of S/Ns. We also investigate  $\Delta\text{Mag}$  as a function of  $z$ , finding similar levels of agreement even at  $z \gtrsim 6$ .

<sup>48</sup> Note that since EA $z$ Y uses linear combination of basis templates, the resulting best-fit model is not limited to physically allowed solutions.

Although this is a relatively basic comparison, it nonetheless demonstrates that our photometry appears reasonable and that the SEDs of most objects are describable by EA $z$ Y. Comparisons to literature HST and recent JWST photometric catalogs are included in Appendix A.

## 6. Summary

We present a first generation of high-angular-resolution space-based photometric catalogs of the strong lens cluster A2744, including a combination of archival HST imaging in seven bands with public JWST imaging in eight bands from three programs (UNCOVER, GLASS/ERS, and a DDT program). With an ultradeep noise-equalized F277W+F356W+F444W detection image at the native JWST resolution (Figure 5), we present 0.4–4.4  $\mu\text{m}$  panchromatic coverage of 61,648 sources within the extended cluster region, including two newly detailed structures to the north and northwest of the cluster heart. After removing stars, artifacts, and other spurious sources, we present reliable photometry for 55,613 galaxy candidates.

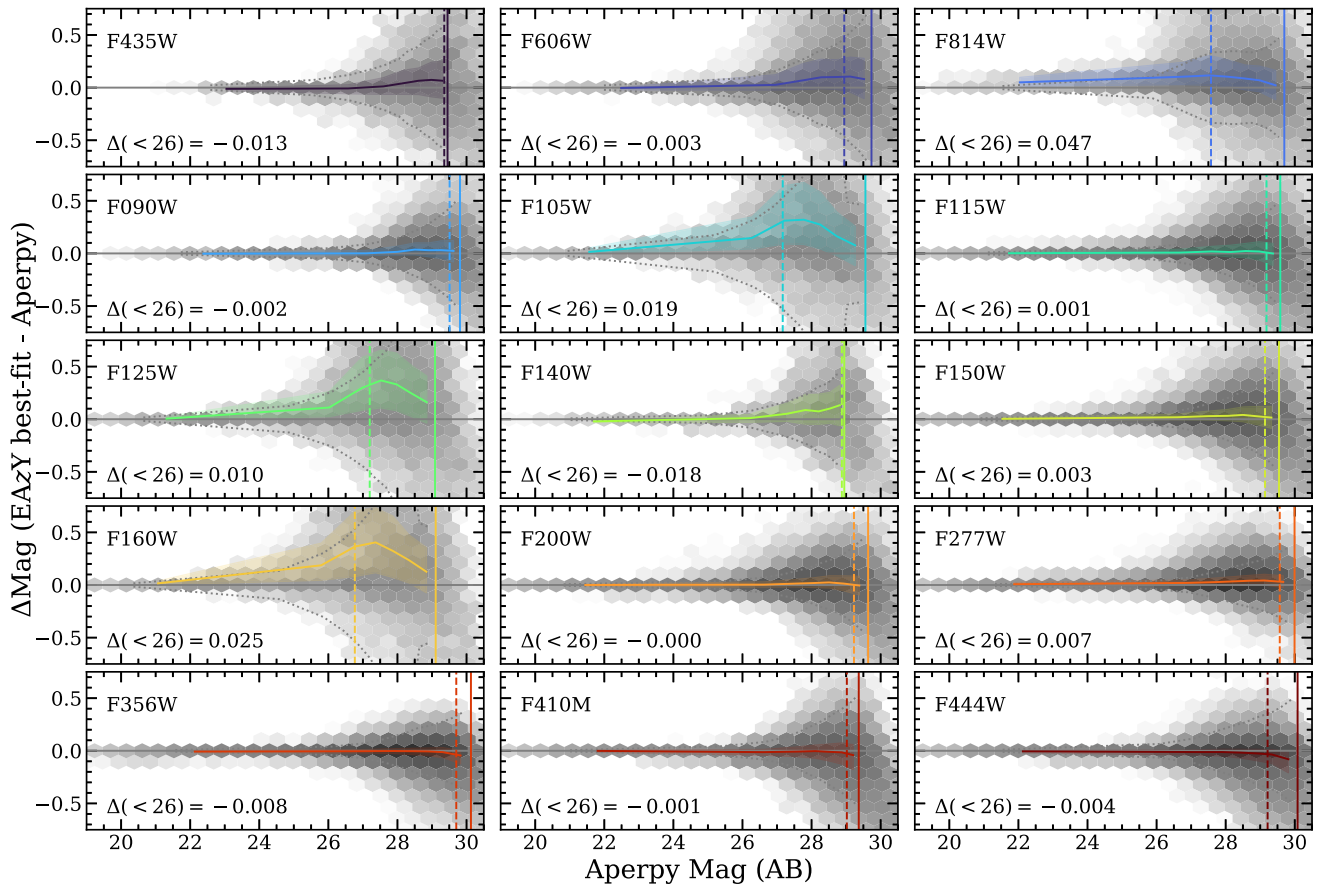
With this paper, we release the UNCOVER photometric catalogs derived from small ( $0''.32$  diameter) to large ( $1''.40$  diameter) circular apertures, including a straightforward aperture-combined “super” catalog for rapid, go-to use. Aperture photometry is measured on images PSF matched to JWST/F444W resolution that have been cleaned of contaminating light from BCGs and ICL. The photometry is corrected to total based on the ratio of flux within a Kron-like aperture relative to a circular aperture, plus an additional correction of order 5%–20% for missing light beyond the Kron radius as determined from the PSF curve of growth in the F444W filter. Details including how to access the catalogs, column descriptions, and general use recommendations can be found in Appendix D. Catalogs are available online (see footnote 42). The software used to produce these catalogs, *aperpy*, is generally applicable to any JWST/NIRCam data and is freely available (see footnote 43).

The UNCOVER photometric catalogs are among the deepest catalogs publicly available, reaching effective  $5\sigma$  depths greater than 29 AB in all 15 bands for the  $0''.32$  diameter apertures. These depths do not account for extra magnification factors from strong gravitational lensing for those background galaxies in optimal configurations. When combining the survey depths with strong lensing, UNCOVER is the deepest view into our Universe to date.

## Acknowledgments

The authors would like to acknowledge the generosity of Vasily Kokorev, as well as Diego Paris, Adriano Fontana, Emiliano Merlin, and Tommaso Treu for making their photometry available for comparison and for many useful discussions that improved our catalogs. We would also like to thank Olivier Ilbert for comparisons between Spitzer/IRAC and JWST/NIRCam bands. We are also grateful for the many helpful and constructive comments from the anonymous referee.

This work is based in part on observations made with the NASA/ESA/CSA James Webb Space Telescope. The data were obtained from the Mikulski Archive for Space Telescopes at the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc.,



**Figure 13.** Photometry measured in the super catalog of this work compared against predicted fluxes from EAzY in all 15 available bands, based on SFHz\_CORR templates. In each panel, the difference in AB magnitude ( $\Delta\text{Mag}$ ) as a function of observed magnitude is summarized by the  $\log_{10}$ -scaled 2D gray density histogram. Colored curves indicate binned medians with two-sided envelopes enclosing 68% of sources per bin. The overall median offset is labeled on each panel computed on all magnitudes up to the magnitude limit of the deepest 10% of the corresponding image indicated by the colored solid vertical line; that of the median depth is indicated by the colored dashed vertical line. The typical photometric error is indicated by the dotted gray curves.

under NASA contract NAS 5-03127 for JWST. These observations are associated with JWST-GO-2561, JWST-ERS-1324, and JWST-DD-2756. Support for program JWST-GO-2561 was provided by NASA through a grant from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Incorporated, under NASA contract NAS 5-03127. Financial support for this program is gratefully acknowledged. This research is also based on observations made with the NASA/ESA Hubble Space Telescope obtained from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555. These observations are associated with programs HST-GO-11689, HST-GO-13386, HST-GO/DD-13495, HST-GO-13389, HST-GO-15117, and HST-GO/DD-17231. All of the data presented in this paper were obtained from the Mikulski Archive for Space Telescopes (MAST) at the Space Telescope Science Institute. The specific observations used to produce these catalogs can be accessed via doi:[10.17909/nftp-e621](https://doi.org/10.17909/nftp-e621).

This work has made use of data from the European Space Agency (ESA) mission Gaia (<https://www.cosmos.esa.int/gaia>), processed by the Gaia Data Processing and Analysis Consortium (DPAC; <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the Gaia Multilateral Agreement.

Cloud-based data processing and file storage for this work is provided by the AWS Cloud Credits for Research program. This research has made use of the NASA/IPAC Extragalactic Database (NED), which is funded by the National Aeronautics and Space Administration and operated by the California Institute of Technology.

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de Investigación, and from Proyecto ASFAE/2022/025 del Ministerio de Ciencia y Innovación en el marco del Plan de Recuperación, Transformación y Resiliencia del Gobierno de España.

*Facilities:* JWST (NIRCam, NIRSpec, and NIRISS), HST (ACS and WFC3), Gaia

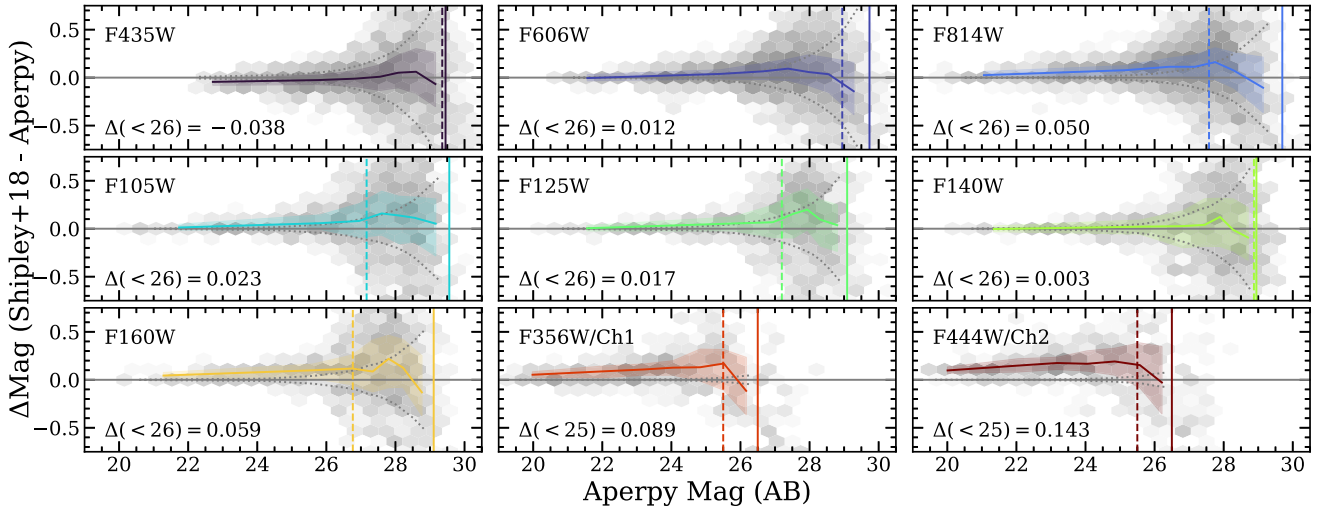
*Software:* ASTROPY (Astropy Collaboration et al. 2013, 2018, 2022), SOURCEEXTRACTOR (Bertin & Arnouts 1996), SEP (Barbary 2016), EXTINCTION (Barbary 2016), SFDMAP (Schlegel et al. 1998; Schlafly & Finkbeiner 2011; [github.com/kbarbary/sfdmap](https://github.com/kbarbary/sfdmap)), WEBBPSF (Perrin et al. 2012, 2014), EAzY (Brammer et al. 2008), PYPHER (Boucaud et al. 2016), PHOTUTILS (Bradley et al. 2022), ASTRODRIZZLE (Gonzaga et al. 2012), GRIZLI ([github.com/gbrammer/grizli](https://github.com/gbrammer/grizli)), NUMPY (van der Walt et al. 2011), and MATPLOTLIB (Hunter 2007).

### Appendix A Photometric Comparisons

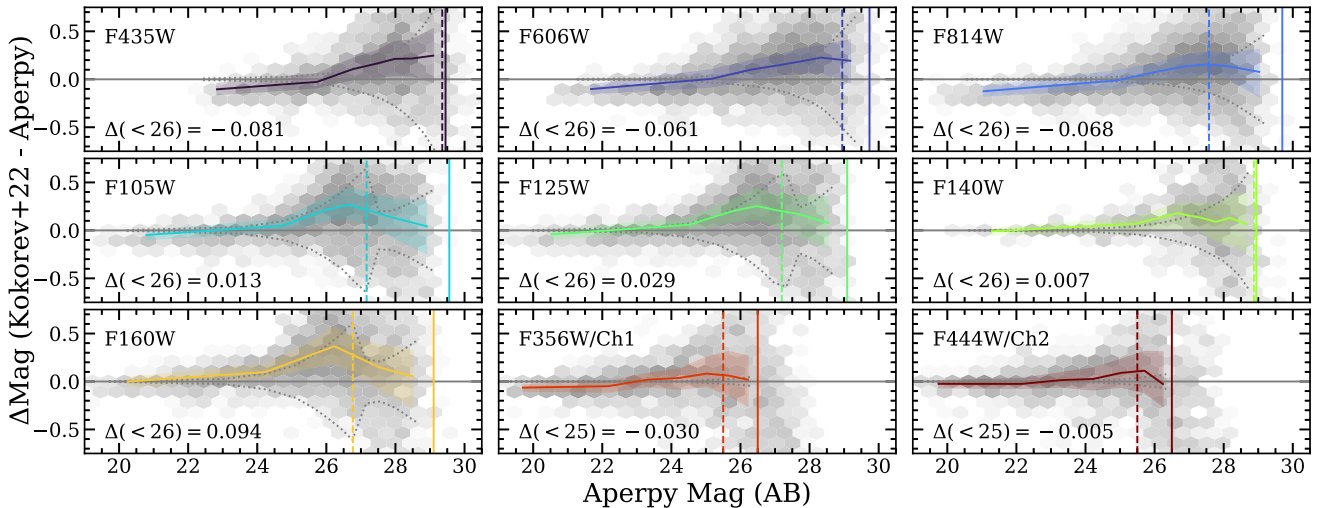
Although comparisons with predicted model fluxes from EAzY are useful (see Figure 13), we can also leverage the existing HST

catalogs to make comparisons for ACS- and WFC3-derived photometry. Comparisons for all seven common bands for Shipley et al. (2018) and Kokorev et al. (2022) are shown in Figures 14 and 15, respectively. Objects that are known blends, stars, or artifacts are not shown. We choose to show the total fluxes derived from 0".7 diameter apertures to be consistent with their choice of aperture size. We also compare to Spitzer/IRAC as the filter profiles of NIRCam F356W and F444W are sufficiently similar to those of IRAC Channels 1 and 2, noting that there likely exist systematics on the 5% level owing to the exact filter profiles and the method of photometric extraction from the literature IRAC measurements. For the comparison of IRAC to NIRCam we adopt the shallower 25.5–26.5 AB depths of the IRAC data upon which the photometry from Shipley et al. (2018) and Kokorev et al. (2022) are based.

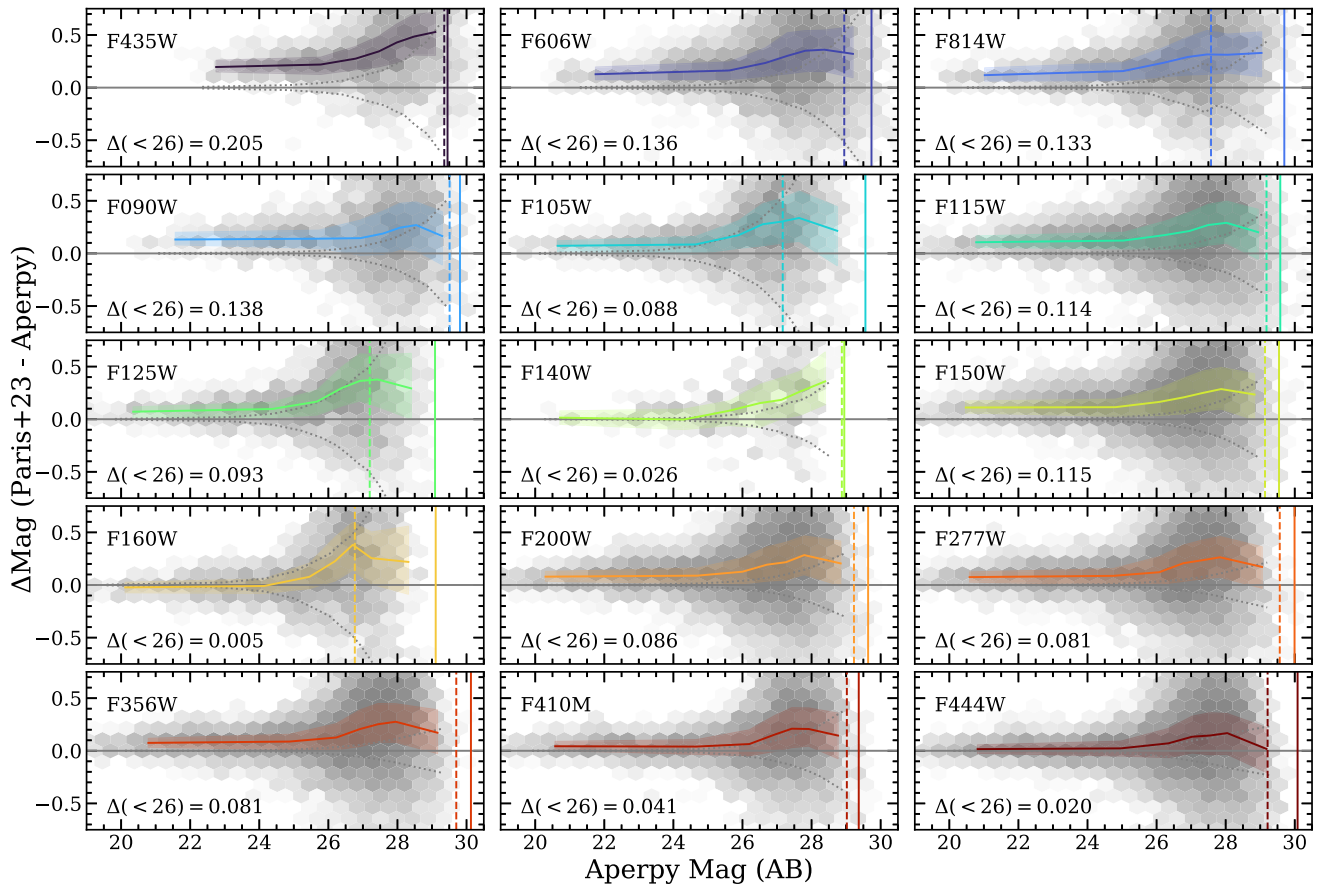
The agreement with literature photometry is generally excellent. Compared to Shipley et al. (2018), we generally find good agreement below 6% for all HST bands with somewhat larger offsets compared to Spitzer/IRAC. Compared to Kokorev et al. (2022), we again generally find good agreement except for HST/ACS bands where in our photometry bright objects are



**Figure 14.** Photometry measured in this work based on corrected 0".7 apertures compared to that of Shipley et al. (2018) from the apertures of the same diameter in the seven common HST bands and two IRAC bands similar to those in NIRCam. Format follows that of Figure 13.



**Figure 15.** Photometry measured in this work based on corrected 0".7 apertures compared to that of Kokorev et al. (2022) from the apertures of the same diameter in the seven common HST bands and two IRAC bands similar to those in NIRCam. Format follows that of Figure 13.



**Figure 16.** Photometry measured in this work based on corrected  $0''.48$  apertures compared to that of Paris et al. (2023) from  $0''.42$  apertures in the seven common HST bands and eight common JWST bands. Format follows that of Figure 13.

$\sim 10\%$  fainter and faint sources  $0.2\text{--}0.3$  mag brighter. However, as detailed in Kokorev et al. (2022), the authors elected to forgo PSF homogenization which makes this comparison particularly hazardous for faint, point-like sources. We achieve a surprisingly good agreement with the IRAC photometry from Kokorev et al. (2022) despite the order-of-magnitude difference in resolution. As seen before, the bimodality in depth for some of the HST bands drives up  $\Delta\text{Mag}$  for sources in the shallow areas.

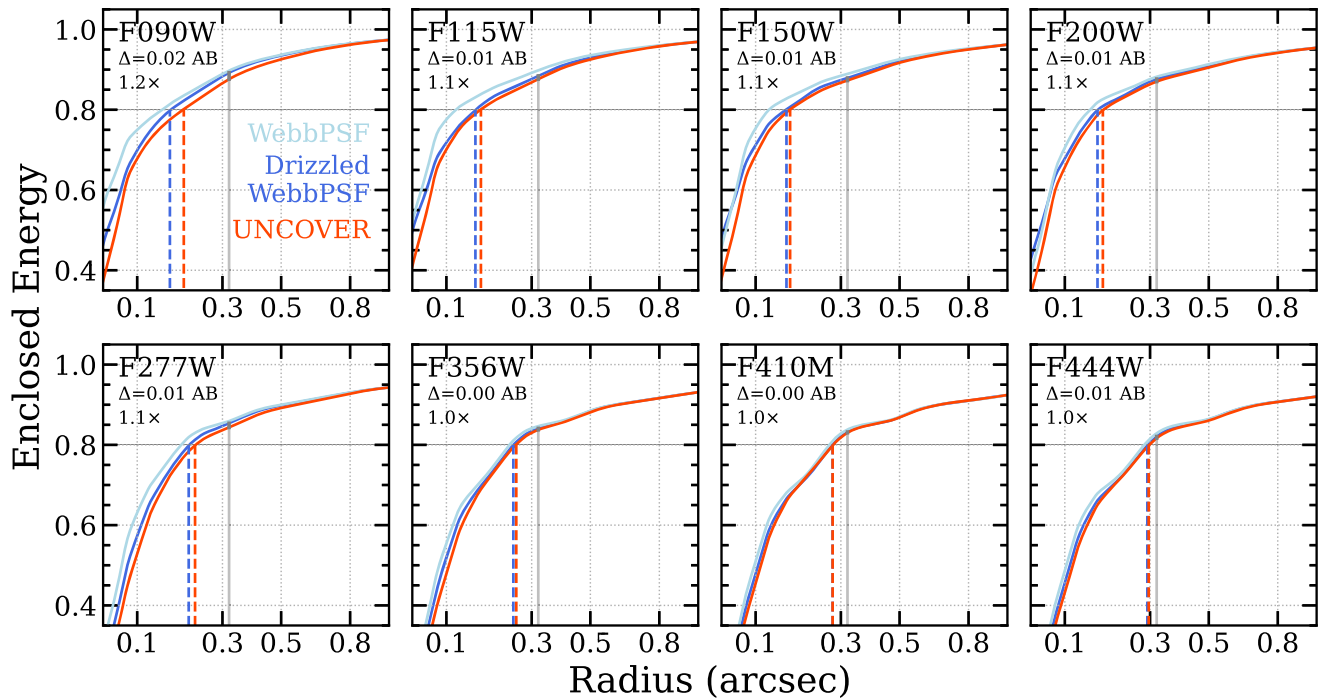
At the time of this writing there are no fully peer-reviewed catalogs with JWST photometry over A2744 produced with in-flight calibrations (see Merlin et al. 2022). Thankfully, we are fortunate in that we were able to compare with an updated photometric catalog developed by the GLASS team (Paris et al. 2023). Both catalogs are based on roughly the same public data sets from HST and JWST, though Paris et al. (2023) perform an independent data reduction and analysis to ours. Several notable differences between the two catalogs include (1) independent image reduction pipelines, (2) source detection (the GLASS catalog is F444W detected, whereas we use a noise-equalized LW detection), (3) the treatment of bCG/ICL (modeled out herein, and not in GLASS), and (4) different software and approaches to PSF modeling, PSF homogenization, aperture photometry, and total flux corrections. Figure 16 shows the photometric comparison in all 15 common filters for HST and JWST, adopting the  $0''.42$  corrected photometry from GLASS compared to our  $0''.48$  corrected photometry. The GLASS photometry is not already corrected for Galactic attenuation, and so we apply a correction consistent with our catalog (see Section 4.3). We note that while the depths indicated correspond

to those in Table 1 measured in this work, many of the faintest sources in our catalog do not have matches to those in GLASS due to their shallower F444W-only detection (see Figure 9).

Given the considerable differences in the construction of the two catalogs, the photometric agreement is reasonable. The agreement is worst in the HST/ACS where the photometry from Paris et al. (2023) is fainter by  $10\%\text{--}20\%$  relative to ours and that of Shipley et al. (2018) and Kokorev et al. (2022). However, while photometry from GLASS is  $\sim 10\%$  fainter for bluer HST/WFC3 bands, the HST/WFC3 F140W and F160W photometry is in excellent agreement. We suspect that the observed differences in HST photometry are driven by both the choice of PSF and the PSF-matching techniques. The JWST/NIRCam photometry from GLASS is similarly fainter by  $\sim 10\%$  in the SW bands, and  $<8\%$  in the LW bands. The result of these differences is that galaxies in the GLASS catalog are systematically redder than measured in this work.

## Appendix B JWST/NIRCam PSF Profiles

JWST is a new facility and so it is worthwhile investigating the PSFs predicted by WebbPSF to the profiles of bona fide point sources within the UNCOVER field. The generation of observed PSFs from point sources within the images follow the description in Section 4.2, noting again that they are normalized at  $4''$  diameter to match the tabulated enclosed energy (EE) provided by STScI. We adopt single JWST PSFs generated by WEBBPSF corresponding to 2022 November 5 near the



**Figure 17.** EE of the PSFs reported by WebbPSF (blue), with drizzling (light blue), and those determined in this work by stacking stars (orange). Yet even after accounting for drizzling, our observed PSFs are still wider at fixed EE, e.g., 80% (colored dashed lines) and, equivalently, they have less EE at fixed aperture diameter, e.g.,  $0/3$  (vertical gray lines). The corresponding photometric difference, however, is relatively small: 0.01–0.02 AB for SW and  $\lesssim 0.01$  for LW.

expected midpoint of the planned UNCOVER visits<sup>49</sup> at the nominal PA of UNCOVER ( $31^\circ.4$ ). Importantly, we set `normalization=exit_pupil` so that the  $4''$  FOV stamps are normalized such that the energy at large radii is accounted for correctly. Comparisons of our WebbPSF PSFs to simulated PSFs and tabulated EE measurements provided by STScI show subpercent agreement at all radii. We also measured the EE of observed bright stars from single exposures taken by the Absolute Flux Calibration Program of Gordon et al. (2022) and hosted on MAST, also finding subpercent agreement with WebbPSF.

Figure 17 shows the curves of growth predicted by WebbPSF as the energy enclosed as a function of diameter from the center (blue). These are single-exposure predictions, and so to account for PSF broadening introduced by imaging reduction we consistently drizzle each PSF and extract a representative stamp that is spatially averaged across the FOV (light blue). While the drizzled PSFs are only slightly broader in most cases, the EEs of our observed PSFs from real stars are still broader. These shape differences translate into photometric offsets of 0.01–0.02 AB for SW and  $\lesssim 0.01$  for LW. These differences may be driven by missing elements in the computation of the WebbPSF or minor inconsistencies in the drizzling procedure. As such, we recommend that, where possible, observed PSFs should be used for measurements on NIRCcam imaging.

### Appendix C JWST/NIRCcam PSF Stability

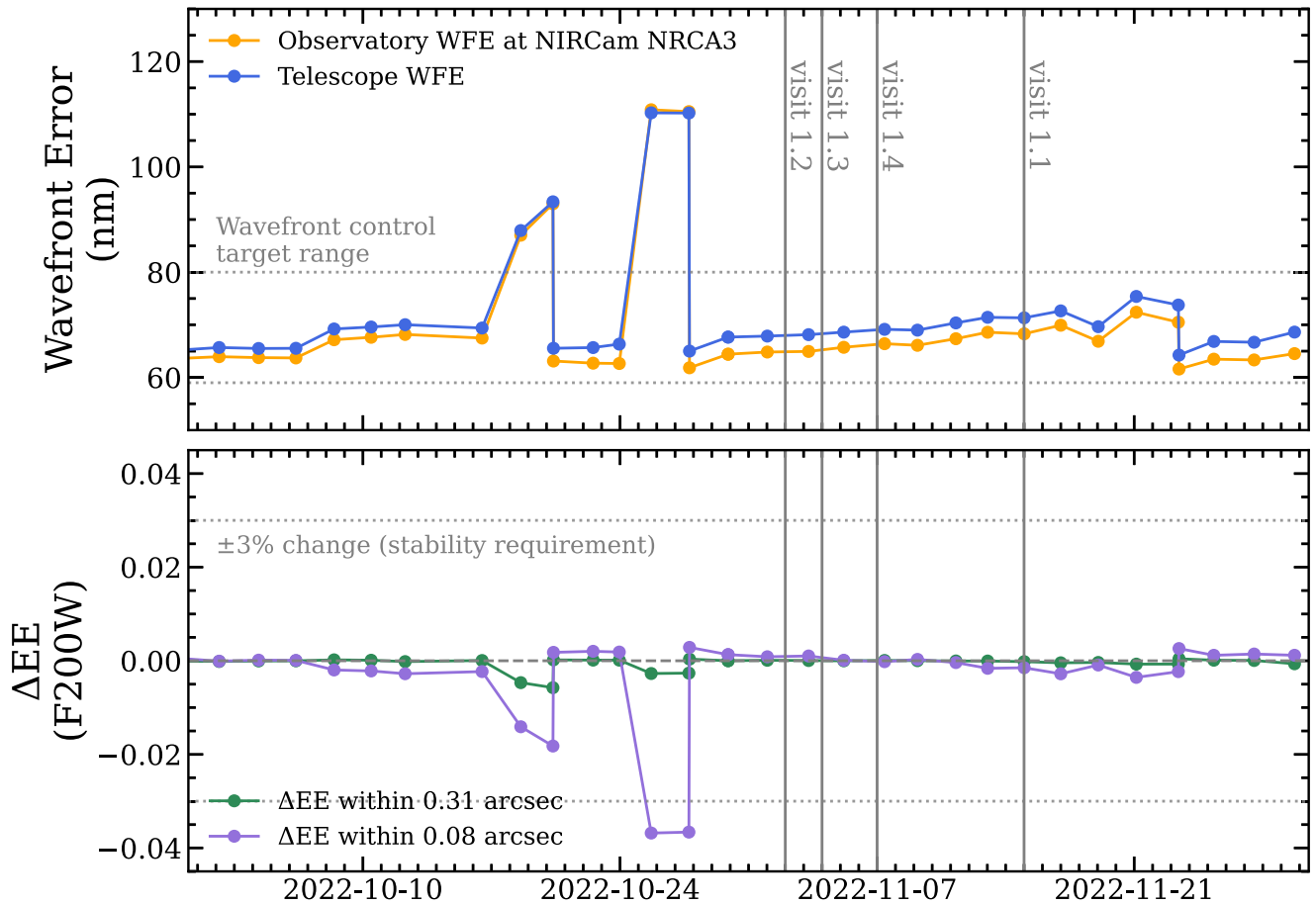
While JWST is already proving to be a revolutionary facility to advance nearly all areas of astrophysics, the typical behavior of the telescope resolution, both in time and across the

detectors, remains a concern. Unlike the generally stable PSF enjoyed by HST, the moving hexagonal mirrors of JWST means that the PSF can be highly variable. Thankfully, the wave front sensor (WFS) samples the observatory mirrors on an approximate 2 day cadence with sufficient detail to reconstruct the PSF at any detector location in any band. This is most easily accessed by WEBBPSF, which includes near-live reports from the WSS to enable PSF reconstruction on the fly, as well as tools to visualize the typical PSF behavior with time.

Nardiello et al. (2022) recently presented a detailed analysis of the PSF variability of NIRCcam, finding  $\sim 9\%$  variation across their FOV and 3%–4% over multiepoch exposures based on peak-to-peak variations measured from PSF residuals. However, their results are concerned with the shape of the PSF (especially its core) which is directly relevant to their PSF-fitting photometry of dense star clusters (see also Zhuang & Shen 2023). The power of aperture photometry, crucially, is that it is sensitive only to the energy enclosed by a given aperture and not about precisely where that energy is located within the aperture. By now studying the variation of the EE we only better understand the aperture photometric impact of these variations, but do so in the context of our particular observations from UNCOVER.

Our photometry is derived from images PSF matched to F444W and so their accuracy strongly depends on the F444W PSF characteristics. We adopt PSFs from WebbPSF in the same fashion as described above in Appendix B. In the following analysis we generate two regular grids of 144 PSF samples for each NIRCcam LW detector: NRCA5 and NRCB5. Here we focus on the UNCOVER data set only leaving the characterization of the PSF behavior during the GLASS and DDT visits to future work. Although shown below only for F444W, we have repeated the analysis for all NIRCcam bands and find similar results.

<sup>49</sup> Due to a guide star acquisition failure on October 31, visit 1.1 was repeated successfully on November 15.



**Figure 18.** Evolution of the WFE for JWST generated from WEBBPSF from 2022 October 1 to November 30. Top: the measured WFE for NIRCcam Module A3 (NRCA3, yellow) is shown along with that of the general telescope facility (blue). Bottom: the EE of the F200W PSF in  $0''.31$  and  $0''.08$  apertures (green and purple, respectively) relative to the median. Relevant control and stability criteria are indicated. behavior during the four UNCOVER visits is effectively unchanged, with  $\Delta EE \lesssim 0.1\%$ .

Figure 18 quantifies the wave front error (WFE) from 2022 October 1 to November 30, during which the UNCOVER program was executed. While there were significant anomalies in the WFE at the end of October, all four UNCOVER visits were executed with nominal PSF behavior with a WFE  $\approx 65\text{--}75$  nm, well within the target control range.<sup>50</sup> Of even greater significance to this work, the relative EE,  $\Delta EE$ , during that same window was effectively static, meaning that any uncertainty due to the time evolution of the F444W PSF when measuring photometry can be safely ignored.

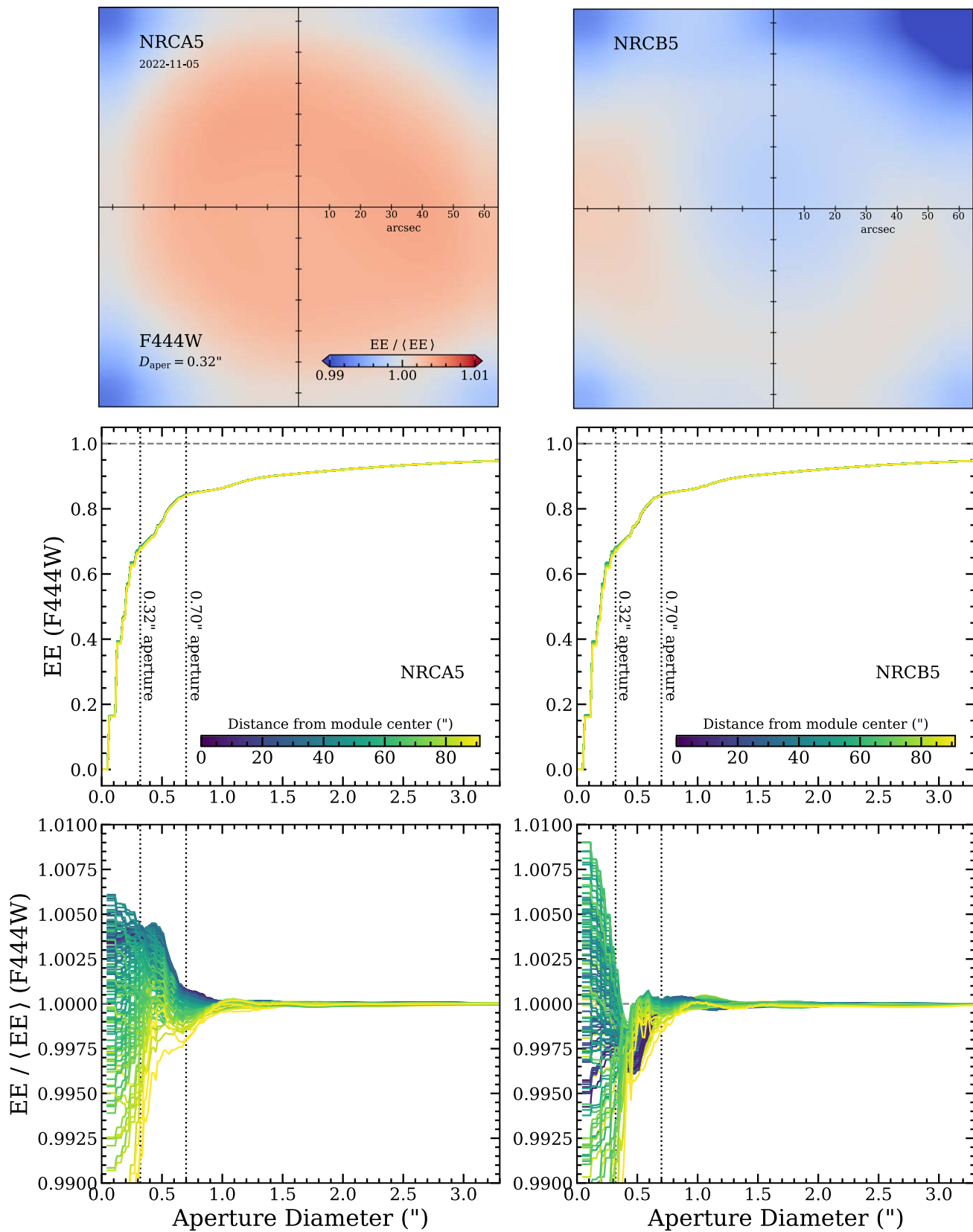
Figure 19 provides insight into the spatial variability of the PSF across NRCA5 and NRCB5 on November 5. Instead of examining the FWHM as provided by WEBBPSF, we measure the EE as it is more pertinent to assessing uncertainties in aperture photometry. While the EE relative to that of the

average PSF  $\langle EE \rangle$  differs between NRCA5 and NRCB5, we find that the spatial variation is generally smooth and increases with decreasing aperture size such that the variability in EE is  $\ll 0.1\%$  in  $0''.70$  apertures and  $\sim 1\%$  at  $0''.32$ .

We caution, however, that these results do not necessarily apply to model-based photometric techniques (e.g., GALSIM; Rowe et al. 2015) because although the summed energy at fixed radius is relatively constant, the exact shape of the PSF itself may vary enough that the spatial dependence may need to be taken into account, as stressed by Nardiello et al. (2022).

Together, these two pieces of evidence point to a favorably small contribution to the error budget from the JWST PSF behavior for UNCOVER—both in time evolution and spatial variation—such that we can confidently neglect these uncertainties in our photometric catalogs.

<sup>50</sup> <https://jwst-docs.stsci.edu/jwst-observatory-hardware/jwst-wavefront-sensing-and-control>



**Figure 19.** Behavior of JWST/NIRCam as measured by the WFS on 2022 November 5, from which our JWST PSFs are derived, based on a grid of 144 PSF realizations generated with WEBBPSF. Top: the spatial dependence of the energy enclosed in  $0''.32$  apertures for NIRCam Module A5 (NRCA5) and B5 (NRCB5) in F444W relative to that of the average PSF. Middle: EE as a function of aperture size colored by distance from the respective module center. Bottom: EE as a function of aperture size relative to that of the average PSF for each module, colored as before. PSF variability across NIRCam for F444W is minimal with  $\Delta EE \lesssim 0.1\%$  in even small  $0''.32$  apertures.

## Appendix D Catalog Column Descriptions

This paper is accompanied by a “super” photometric catalog and five others computed from individual aperture sizes from

0''32 to 1''40. Each is derived from the same bCG-subtracted images which have been PSF matched to that of F444W, and share the same LW-selected objects. Fluxes are reported in total as described in Section 4.3, accounting for any additional light

**Table 3**  
Catalog Columns

Column name	Description
id	Unique identifier
x/y	X/Y centroid in image coordinates
ra	R.A. J2000 (degrees)
dec	Decl. J2000 (degrees)
ebv_mw	Line-of-sight $E(B-V)$ , already applied, computed from Schlafly & Finkbeiner (2011)
faper_f277w+f356w+f444w	Flux computed on the F444W-matched LW stack in circular apertures ( $F_{\nu}$ , zero-point = 28.9)
eaper_f277w+f356w+f444w	1 $\sigma$ flux error computed from empty apertures on the F444W-matched LW stack ( $F_{\nu}$ , zero-point = 28.9)
fauto_f277w+f356w+f444w	Flux computed in Kron ellipse on the F444W-matched LW stack ( $F_{\nu}$ , zero-point = 28.9)
f_X	Total flux for each filter X ( $F_{\nu}$ , zero-point = 28.9)
e_X	1 $\sigma$ flux error for each filter X ( $F_{\nu}$ , zero-point = 28.9)
w_X	Weight relative to the maximum within image X (see text)
tot_cor	Aperture-to-total correction (including both aperture-to-Kron and Kron-to-total)
kron_radius	Kron radius factor (SOURCE EXTRACTOR-like, unit-less)
kron_radius_circ	Circularized Kron radius (arcseconds)
use_circle	1 for objects without Kron correction (faint and/or blended)
flag_kron	1 for blended objects without Kron correction
iso_area	Isophotal area based on the source segment, in square arcseconds
a_image	Semimajor axis (pixels)
b_image	Seminor axis (pixels)
theta_J2000	PA of the major axis (counterclockwise, measured from east)
flux_radius	Radius containing 50% of the flux with neighbors masked (arcseconds)
use_phot	1 for reliable sources, 0 if any of the following <code>flag_X = 1</code>
flag_nophot	1 if the object has no viable photometry in any band at this aperture size
flag_lowsnr	1 if source has aperture S/N < 3 in the LW stack
flag_star	1 if source is identified as a star (see Section 4.5)
flag_artifact	1 if source is identified as an artifact (see Section 4.6)
flag_nearbcg	1 if within 3'' of a known bCG (see Section 4.6)
z_spec	Spectroscopic redshift, where available (pre-JWST only)
id_DR1	ID of the source in the DR1 catalog release, if available
match_radius_DR1	Distance to DR1 catalog release match, if available (any within 0''08)
id_msa	Corresponding object ID from the UNCOVER MSA Spectral Catalog of Price et al. (2023), if available
match_radius_msa	Distance to object in the UNCOVER MSA catalog, if available (any within 0''24)
id_alma	ID of the source in the ALMA Catalog of S. Fujimoto et al. (2023, in preparation), if available.
f_alma	Atacama Large Millimeter/submillimeter Array (ALMA) flux (NaN for nondetections; $F_{\nu}$ , zero-point = 28.9)
e_alma	ALMA flux uncertainty (1 $\sigma$ upper limit for nondetections; $F_{\nu}$ , zero-point = 28.9)
match_radius_alma	Distance to object in the ALMA catalog (nearest within 1'')
use_aper	Aperture diameter used to measure fluxes, for super catalog only (arcseconds)
nusefilt	Number of filters used in the fit
z_phot	Best-fit photometric redshift via minimum $\chi^2$ using SFHz_CORR templates
z_phot_chi2	Total $\chi^2$ at <code>z_phot</code>
z025/160/500/840/975	Redshift posterior percentiles, e.g., <code>z025</code> $\rightarrow$ 2.5%
restU/V/J	Rest-frame UVJ-band flux ( $F_{\nu}$ , zero-point = 28.9)
restU/V/J_err	Rest-frame UVJ-band flux uncertainty ( $F_{\nu}$ , zero-point = 28.9)
restus/gs/is	Rest-frame synthetic <i>ugi</i> -band flux ( $F_{\nu}$ , zero-point = 28.9)
restus/gs/is_err	Rest-frame synthetic <i>ugi</i> -band flux uncertainty ( $F_{\nu}$ , zero-point = 28.9)
star_min_chi2	$\chi^2$ of best-fit stellar template
mu	Gravitational magnification $\mu$ (best fit; =1 for foreground objects)
mu025/160/840/975	Total magnification $\mu = \mu_r \mu_t$ posterior percentiles
mu_r	Radial magnification (best fit; =1 for foreground objects)
mu_r025/160/840/975	Radial magnification $\mu_r$ posterior percentiles
mu_t	Tangential magnification (best fit; =1 for foreground objects)
mu_t025/160/840/975	Tangential magnification $\mu_t$ posterior percentiles
shearone	Shear $\gamma_1$ (best fit; =1 for foreground objects)
shearone025/160/840/975	Shear $\gamma_1$ posterior percentiles
sheartwo	Shear $\gamma_2$ (best fit; =1 for foreground objects)
sheartwo025/160/840/975	Shear $\gamma_2$ posterior percentiles



**Note.** X = filter name, as defined in Section 2. Synthetic rest-frame *ugi*, filters are detailed in Antwi-Danso et al. (2023).

outside the Kron aperture. While the columns included in the catalogs are described in Table 3, we also include a dedicated README file with the catalogs in the case of future changes.

We strongly recommend using photometry from the super catalog, which adopts 0"32 diameter apertures for small objects (e.g., high- $z$  galaxies) that make up the majority of the catalog and larger 0"48–1"40 apertures for the relatively fewer bright, extended objects at lower- $z$ . The photometry is reported in  $F_{\nu}$  in units of 10 nJy, corresponding to a zero-point of 28.9 AB.  $z_{\text{phot}}$ , rest-frame fluxes, and magnification estimates are reported for all objects. No quantities are corrected for magnification. Clean samples of galaxies can be readily identified with `USE_PHOT = 1` (see Section 4.6).

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