

Abundances of P, S, and K in 58 bulge spheroid stars from APOGEE

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ABSTRACT

Context. We have previously studied several elements in 58 selected bulge spheroid stars, based on spectral lines in the H band. We now derive the abundances of the less studied elements phosphorus (P; $Z=15$), sulphur (S; $Z=16$), and potassium (K; $Z=19$).

Aims. The abundances of P, S, and K in 58 bulge spheroid stars are compared both with the results of a previous analysis of the data from the Apache Point Observatory Galactic Evolution Experiment (APOGEE), and with a few available studies of these elements.

Methods. We derived the individual abundances through spectral synthesis, using the stellar physical parameters available for our sample from the DR17 release of the APOGEE project. We provide recommendations for the best lines to be used for the studied elements among those in the H -band. We also compare the present results, together with literature data, with chemical-evolution models. Finally, the neutrino-process was taken into account for the suitable fit to the odd- Z elements P and K.

Results. We confirm that the H -band has useful lines for the derivation of the elements P, S, and K in moderately metal-poor stars. The abundances, plotted together with literature results from high-resolution spectroscopy, indicate that moderately enhanced phosphorus stars are found, reminiscent of results obtained for thick disc and halo stars of metallicity $[\text{Fe}/\text{H}]\approx-1.0$. Therefore, for the first time, we identify that this effect occurs in the old stars from the bulge spheroid. Sulphur is an α -element and behaves as such. Potassium and sulphur both exhibit some star-to-star scatter, but fit within the expectations of chemical evolution models.

Key words. stars: abundances – stars: atmospheres – Galaxy: bulge

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1. Introduction

Galaxy bulges and inner haloes form first (e.g. Gao et al. 2010); the oldest stars in the Milky Way were very likely formed in the early bulge. We have identified a sample of 58 stars that appear to represent the stellar population of an early bulge spheroid, based on chemical, kinematical, and dynamical criteria, as described in Razera et al. (2022). These old bulge stars were selected to have metallicities of $[\text{Fe}/\text{H}] < -0.8$, in order to exclude most of the bulge stars that are metal-rich and include stars from a small metallicity peak at $[\text{Fe}/\text{H}] \sim -1.0$ that is detected in both bulge globular clusters (Rossi et al. 2015; Bica et al. 2016; Pérez-Villegas et al. 2020; Bica et al. 2024) and field stars (Lucy et al. 2021). This relatively high metallicity for the oldest stars is due to rapid chemical enrichment in the central region of the Galaxy (Chiappini et al. 2011; Wise et al. 2012; Friaça & Barbuy 2017; Barbuy et al. 2018a; Matteucci 2021). There is some evidence that many of the stars located in the bulge with a metallicity of $[\text{Fe}/\text{H}] \lesssim -1.5$ should be assigned to a halo origin, as is discussed in Lucy et al. (2021), and also suggested by Fernández-Trincado et al. (2020) and Geisler et al. (2023) as concerns the bulge globular clusters.

The kinematical characteristics of the sample stars, together with analysed abundances of C, N, O, Mg, Si, Ca, and Ce, are reported in Razera et al. (2022). Barbuy et al. (2023) and Barbuy et al. (2024) derived Na and Al, and iron-peak elements V, Cr, Mn, Co, Ni, and Cu, respectively, and interpreted these abundances in terms of their chemical-evolution models. In Sales-Silva et al. (2024) the neutron-capture elements Nd and Ce, including some stars in common with the present sample, were analysed. In this paper, we analyse the less studied elements P, S, and K. These studies are based on data from the Apache Point Observatory Galactic Evolution Experiment (APOGEE; Majewski et al. 2017).

The elements P, S, and K are produced in type II supernovae, also called core-collapse supernovae (CCSNe) (e.g. Woosley & Weaver 1995, hereafter WW95). The chemical evolution models of Cescutti et al. (2012) succeeded in reproducing P in disc star data by using the massive star yields of Kobayashi et al. (2006) multiplied by a factor of 2.75. Based on data for halo stars, Jacobson et al. (2014) favoured the yields from hypernovae (HNe) by Kobayashi et al. (2006) to reproduce their data. These elements were little studied in the literature, primarily because they possess only a few measurable and unblended spectral lines in the optical. For phosphorus, the H -band shows one measurable line and a second fainter line. Sulphur has lines in the region 600–900 nm that are non-negligibly affected by NLTE, and one measurable line in the H band. Potassium has two lines in the H band, both only blended with rather weak CN features. In the present work we derive the abundances of these elements using spectrum synthesis.

This paper is organised as follows. In Section 2, we summarise the selection of sample stars and the available data. In Section 3, we describe the calculations for spectral synthesis and the lines studied. In Section 4, we present the results, along with the data from the literature on the elements studied. In Section 5, the chemical-evolution models are compared with the data. We summarise our conclusions in Section 6.

2. Data and selection sample

As is described in Razera et al. (2022), our selection is based on the reduced proper motion (RPM) stars from Queiroz et al. (2021) that served as a pool to select the sample. The RPM sample from Queiroz et al. (2021), in turn, is based on the stars

observed by APOGEE, combined with StarHorse distances (Santiago et al. 2016; Queiroz et al. 2018), and cross-matched with proper motions from the Gaia Early Data Release 3 (Gaia Collaboration 2021). The selection identified stars with a distance to the Galactic centre of $d_{\text{GC}} < 4$ kpc, a maximum height of $|Z|_{\text{max}} < 3$ kpc, an eccentricity > 0.7 , and orbits not supporting the bar (where the orbits were computed in Queiroz et al. 2021), and imposing a metallicity cut of $[\text{Fe}/\text{H}] < -0.8$. By applying kinematical criteria and imposing that the selected stars have spectra from the second generation of the APOGEE-2 – Majewski et al. 2017, a sample of 58 stars was built. A Kiel diagram of this sample was presented in Fig. 3 of Razera et al. (2022).

APOGEE is one of the surveys of the Sloan Digital Sky Survey IV (Blanton et al. 2017, SDSS-IV/V). The APOGEE spectra have a high resolution ($R \sim 22\,500$) and high signal-to-noise ratios in the H band (15 140–16 940 Å) (Wilson et al. 2019), and include about 7×10^5 stars. APOGEE-1 and APOGEE-2 used the 2.5 m Sloan Foundation Telescope at the Apache Point Observatory in New Mexico (Gunn et al. 2006), and the 2.5 m Irénée du Pont Telescope at the Las Campanas Observatory in Chile (Bowen & Vaughan 1973), respectively. Santana et al. (2021) and Beaton et al. (2021) describe the targeting of APOGEE for the southern and northern hemispheres, respectively. The detectors are H2RG (2048 × 2048) Near-Infrared HgCdTe Detectors with 18 micron pixels.

The analysis of H -band spectra in the APOGEE project is carried out through a Nelder-Mead algorithm (Nelder & Mead 1965), which simultaneously fits the stellar parameters effective temperature (T_{eff}), gravity ($\log g$), metallicity ($[\text{Fe}/\text{H}]$), and microturbulence velocity (v_t) – together with the abundances of carbon, nitrogen, and α -elements with the APOGEE Stellar Parameter and Chemical Abundances Pipeline (ASPCAP) (García Pérez et al. 2016), which is based on the FERRE code (Allende Prieto et al. 2006) and the APOGEE line list (Smith et al. 2021). In the present work, we use APOGEE Data Release 17 – DR17 (Abdurro'uf et al. 2022) for the stellar parameters, to derive the abundances with spectrum synthesis, in particular for phosphorus, that are not available from DR17; the P abundance is available for only one star in APOGEE DR17, to be discussed in Sect. 3.2.

3. Calculations

We computed the abundances of P, S, and K in the H band using the code TURBOSPECTRUM from Alvarez & Plez (1998) and Plez (2012). The model atmosphere were interpolated within the CN-mild MARCS grids from Gustafsson et al. (2008). The solar abundances of the elements studied – that is, $A(\text{P}) = 5.41$, $A(\text{S}) = 7.12$, and $A(\text{K}) = 5.07$ – are from Asplund et al. (2021). Note that, from O and B stars, recently Aschenbrenner et al. (2025) derived a present-day abundance for P, of $A(\text{P}) = 5.36 \pm 0.14$, and there are present-day abundances for S, derived by Daflon et al. (2009), of $A(\text{S}) = 7.15 \pm 0.05$. We also recomputed the C, N, O abundances, given that in Razera et al. (2022) these abundances were computed with a different code, PFANT (Barbuy et al. 2018b). Here we assumed the solar C, N, O abundances of Asplund et al. (2021), $A(\text{C}) = 8.46$, $A(\text{N}) = 7.83$, and $A(\text{O}) = 8.69$, which are also different solar values from the ones assumed in Razera et al. (2022).

Table 1 reports the lines in the H band that we used to measure the abundances of the elements P, S, and K in the spectra of the sample stars. Oscillator strengths were adopted from the line list of the APOGEE collaboration, which were initially

Table 1. Line list and oscillator strengths.

Ion	λ (Å)	χ_{ex} (eV)	log gf		
			VALD3	Kurucz	APOGEE
PI	15711.522	7.176	-0.510	-0.720	-0.404
	16482.932	7.213	-0.290	-0.400	-0.273
SI	15469.816	8.046	-0.050	-0.220	-0.199
	15475.616	8.047	-0.520	-0.700	-0.744
	15478.482	8.047	0.180	0.000	-0.040
KI	15163.067	2.670	0.524	0.640	0.630
	15168.377	2.670	0.347	0.480	0.481

Notes. Air wavelengths are from APOGEE collaboration linelist. Note that somewhat different wavelengths for S I 15469.826, 15475.624 and 15478.496 Å are given in National Institute of Standards and Technology (NIST) by Kramida, A., Ralchenko, Yu., Reader, J., and the NIST ASD Team (2024). The NIST Atomic Spectra Database used is (ver. 5.12), [Online]¹ [2025, April 16]. Oscillator strengths of the P, S, and K lines from VALD3, Kurucz & Bell (1995), and the APOGEE collaboration (adopted) are reported.

adopted from the most recent line list of Kurucz & Bell (1995), with log gf values updated with National Institute of Standards and Technology Atomic Spectra Database – NIST values, and adjusted within 2-sigma uncertainties based on the Sun and Arcturus spectra (Smith et al. 2021). For comparison purposes, we also show the log gf values from the Kurucz CD ROM 23, and from the Vienna Atomic Line Database (VALD3): see the line lists of Kurucz & Bell (1995) and Ryabchikova & Pakhomov (2015). We note that log gf values are not available in the NIST database for any of these studied lines. The full atomic line list employed is that from the APOGEE collaboration, together with the molecular lines described in Smith et al. (2021).

We adopted the uncalibrated stellar parameters effective temperature (T_{eff}), gravity (log g), metallicity ([Fe/H]), and microturbulence velocity (v_t) from the APOGEE DR17 results. These parameters are reported in Table A.1, together with recomputed C, N, O, and the results for P, S, and K abundances.

The use of uncalibrated instead of calibrated parameters could be a matter of debate because the difference in parameters can lead to changes in the abundances derived. While using calibrated parameters may seem to be a more sensible choice for most of the giants of the APOGEE survey, we argue here that this may not be as obvious for our current sample. According to Jönsson et al. (2020), ASPCAP calibrated temperatures have been calibrated using the colour–temperature relation of González Hernández & Bonifacio (2009), which has been tested down to 4000 K, and surface gravities have been calibrated using seismic data from Pinsonneault et al. (2018) and validated down to log $g = 1.2$. Because our sample, as is reported here in Table A.1, contains only very cool (3600–4300 K) metal-poor ([Fe/H] ~ -1.0) low-gravity (log $g \lesssim 1.2$) giants, it is not clear anymore whether the calibrated parameters perform better than the uncalibrated parameters. In addition, we also use a model atmosphere grid that is slightly different in C, N, and α -element composition than APOGEE-ASPCAP. For the cooler stars ($T_{\text{eff}} \lesssim 3700$ K), that will also have a significant impact on the atmosphere structure, and thus on parameters and abundances. However, neither the MARCS atmosphere grid that we use here nor the APOGEE-ASPCAP one strictly matches the abundances of our sample stars and it is not clear which of them

is best for such a sample. We still chose the uncalibrated parameters in order to maintain some consistency with the raw results of the APOGEE-ASPCAP pipeline, as was previously discussed in Razera et al. (2022). We also note that in da Silva et al. (2024) we verified the reliability of ASPCAP for deriving stellar parameters, and concluded that the use of molecular-line intensities is a powerful method, in particular for the derivation of effective temperatures. We remind the reader that for DR17 the parameters were obtained with new spectral grids constructed using the Synspec spectrum synthesis code (Hubeny & Lanz 2017; Hubeny et al. 2021).

3.1. Phosphorus, sulphur, and potassium lines

We analysed the lines of P, S, and K in the H band. The fits were all carried out visually, adopting convolutions with a full width at half maximum (FWHM) from 0.65 to 0.75 Å in the range 15 000 to 17 000 Å. These FWHM values are compatible with the ones based on a directly measured FWHM of ~ 0.7 Å, with 10 to 20 percent variations seen across the wavelength range by Ashok et al. (2021) and Nidever et al. (2015). Although for most of the lines the resulting abundance is similar to that reported in APOGEE-ASPCAP DR17, there are cases in which a visual inspection is needed because of noise or defects. The results are given in Table A.1.

Phosphorus: the two available lines of P I are weak. The P I 15711.522 Å is often too weak, while P I 16482.932 Å is measurable. In about one quarter of the stars, the P I 15711.522 Å line is measurable, and with an abundance compatible with that of the P I 16482.932 Å line. The upper limit of the 15711.522 Å line is compatible with the measurement of the 16482.932 Å line for the other two thirds of stars. We also stress that P measurement was not possible in the spectra where the 16482.932 Å line was too close to the continuum, or too noisy and with strong artefacts – or even falling out of range given the detector edge.

A very important issue is the mixing of the main P I 16482.932 Å line with the molecular lines of CO, which was noted by Hayes et al. (2022) and Brauner et al. (2023). Although we have already derived C, N, and O abundances for these stars in Razera et al. (2022), it is crucial to precisely adjust the strength of the related molecular features to our current methodology. For that we basically used the region 15 520–15 590 Å, which contains a clear CO band-head, clean OH lines, and CN lines spread in this region, as is explained in Barbuy et al. (2021) and Razera et al. (2022). Then, with CO-band lines well fitted, the resulting estimate of P becomes reliable. Figure 1 shows the P I 16482.932 Å in the eight stars with the largest P abundances. The blending with molecular lines is also shown.

Sulphur: there are three measurable lines, but only one is unblended. S I 15403.784 Å is strongly blended with molecular OH lines, plus other atomic lines, and this applies also to 15469.826 Å, but the latter is still useful, as is discussed in Hayes et al. (2022). S I 15475.624 Å is reasonably sensitive to the S abundance but is non-negligibly blended with CN. Also, a few other S I lines fall in the APOGEE gaps. The more reliable feature that is essentially clean from molecular and atomic blending is the line S I 15478.406 Å. Fits to this feature are shown in Figure 2 for stars b11 2M17351981-1948329 and c2 2M17285088-2855427, including the calculations with only molecular lines. Even if this line had more weight, the three features were fitted simultaneously, i.e. 15469.826, 15475.624, and 15478.496 Å, and in some of the stars these features were reasonably fitted.

¹ <https://physics.nist.gov/asd>

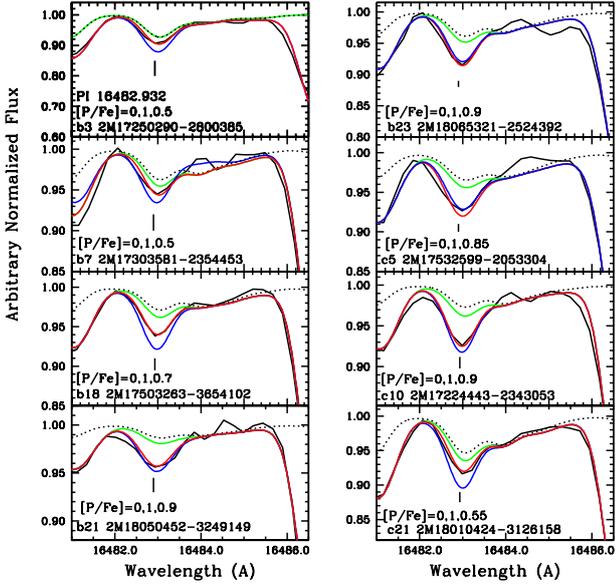


Fig. 1. P I 16482.932 Å line in eight sample stars, fitted with synthetic spectra computed with $[P/Fe] = 0.0$ (green), 1.0 (blue), and final values (red), if different from 0.0 or 1.0. The computations with molecular lines only are shown as dotted black lines.

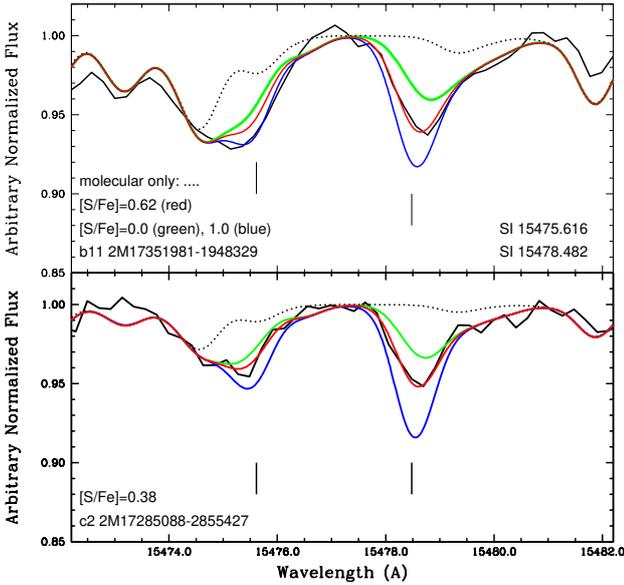


Fig. 2. S I 15475.624 and 15478.496 Å lines in stars b11 2M17351981-1948329 and c2 2M17285088-2855427. Observed spectra (black) are compared with synthetic spectra computed with $[S/Fe] = 0.0$ (green), 1.0 (blue), and final values (red). Dotted lines correspond to molecular lines only.

Potassium: the two available lines are clearly measurable; both have a small blend with CN lines, and they both yield results compatible with each other. Figure 3 illustrates the quality of these lines, as well as the blending CN features, for stars b17 2M17483633-2242483 and c12 2M17323787-2023013.

Table A.1 lists the APOGEE uncalibrated stellar parameters, and the present resulting abundances for the elements P, S, and K, as well as the APOGEE-ASPCAP abundances of S and K, for comparison purposes.

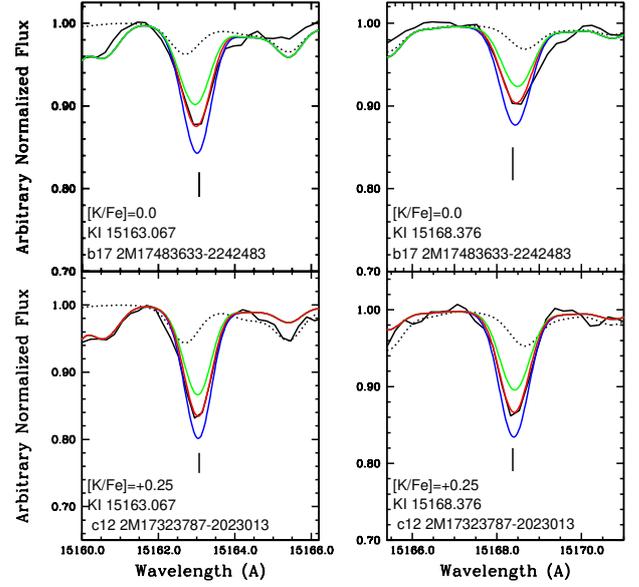


Fig. 3. K I 15163.067 and 15168.376 Å lines in the sample stars b17 2M17483633-2242483 and c12 2M17323787-2023013. The observed spectra (black) are fitted with synthetic spectra computed with $[K/Fe] = 0.0$ for b17 and $+0.25$ for c12, respectively (red); calculations for -0.2 and $+0.2$ and shown in green and blue, respectively. Dotted lines correspond to molecular lines only.

3.2. Uncertainties

There are no calculations of non-LTE deviations for the lines of P I analysed in this work; therefore, these uncertainties cannot be evaluated. Brauner et al. (2023) do not find any trend of P abundances versus temperature, suggesting that NLTE effects for the P lines in the *H* band should be negligible. For K I the synspec calculations from APOGEE DR17 take into account NLTE. Since there are also LTE calculations available, we inspected the differences between the two, and they are negligible, below 0.1 dex. For sulphur, the large spread in the abundance estimates may be improved with NLTE corrections, as it is in the case of lines studied by Korotin et al. (2017), lowering the LTE abundances. Korotin & Kiselev (2024) give grids of NLTE corrections for the lines studied here, indicating corrections: as examples, for $[Fe/H] \sim -1.0$, $T_{\text{eff}} = 4000$ K, and $[S/Fe] = +0.4$ the correction is of ~ -0.16 dex for a gravity $\log g = 0.0$, and ~ -0.06 dex for $\log g = 1.0$; for $[S/Fe] = +0.8$, the corrections are of ~ -0.19 and -0.07 , respectively. The grid is limited to $T_{\text{eff}} \geq 4000$ K, and the trend for lower temperatures is a decrease in the correction, whereas for higher effective temperatures the corrections increase. In summary, for the low temperatures of most of our sample stars, the corrections should be ≤ -0.06 dex.

Typical systematic uncertainties due to stellar parameters were computed by adopting uncertainties in the stellar parameters of $\Delta T_{\text{eff}} = 100$ K, $\Delta \log g = 0.2$ dex, and $\Delta v_t = 0.2$ km s $^{-1}$, shown in Table 2. These were applied to the cool star b14 2M17392719-2310311, and the warmer star c16 2M17310874-2956542. This table shows that the P abundance has a dependence on $\log g$, more so for the cooler star, and a low dependence on the effective temperature. Finally, the continuum placement is another source of uncertainty, in particular for the P I line; this adds another 0.1 dex of uncertainty for the noisier spectra. Finally, it is worth noting that this study of systematic errors offers a good representation of the impact on the abundances due to the use of either uncalibrated or calibrated parameters

Table 2. Phosphorus, sulphur, and potassium abundance uncertainties for stars b14 and c16, due to changes in stellar parameters of $\Delta T_{\text{eff}} = 100$ K, $\Delta \log g = 0.2$ dex, and $\Delta v_t = 0.2$ km s $^{-1}$. The corresponding total error is given in the last column.

Element	ΔT 100 K	$\Delta \log g$ 0.2 dex	Δv_t 0.2 km s $^{-1}$	$(\sum x^2)^{1/2}$
b14 = 2M17392719-2310311 – $T_{\text{eff}}=3643$ K, $\log g=0.67$				
P	0.01	0.10	0.0	0.10
S	0.09	0.10	0.0	0.14
K	0.05	0.01	0.0	0.05
c16 = 2M17310874-2956542 – $T_{\text{eff}}=4175$ K, $\log g=1.2$				
P	0.02	0.07	0.0	0.07
S	0.04	0.01	0.0	0.04
K	0.05	0.005	0.0	0.05

of APOGEE-ASPCAP DR17 as the difference between those parameters is on average similar to the ones used in Table 2.

Another way to estimate uncertainties on abundances is to compare with the literature. APOGEE-ASPCAP DR17 do not provide P abundances. Nevertheless, we have one star in common with Brauner et al. (2023), b2 2M17173693-2806495. Brauner et al. (2023) report a conservative upper limit of 1.24. The BACCHUS analysis of weak lines in APOGEE spectra (BAWLAS; Hayes et al. 2022, 2023) gives a P abundance for this same star of 0.83. The values are different from our own determination of 0.4, but we recall that the stellar parameters are different (calibrated vs uncalibrated stellar parameters from APOGEE).

Figure 4 shows the difference between the S and K abundances derived in this work and the ones from APOGEE-ASPCAP DR17. The mean difference between the present results and ASPCAP DR17 is found to be:

$$\begin{aligned} [\text{S}/\text{Fe}]_{\text{present}} - [\text{S}/\text{Fe}]_{\text{ASPCAP}} &= -0.08_{-0.17}^{+0.22} \\ [\text{K}/\text{Fe}]_{\text{present}} - [\text{K}/\text{Fe}]_{\text{ASPCAP}} &= -0.04_{-0.10}^{+0.07}. \end{aligned} \quad (1)$$

For S the differences are due to continuum placement, but more so due to measuring only the best line for S, because, as was explained above, the other lines are strongly blended with molecular lines. For K, we measure systematically lower abundances, and this could also be due to continuum placement. The fits are available on request.

4. Present results and literature data

The results in Table A.1 can be compared with literature data, all of which can essentially be described in this section, given that there have been few previous observational studies on P, S, and K.

Phosphorus: there are only a few measurements of P abundances in the literature. Caffau et al. (2011) used the CRyogenic high-resolution InfraRed Echelle Spectrograph (CRIRES) at the Very Large Telescope (VLT) spectra to measure the P I 10511.584, 10529.522, 10581.569, and 10596.900 Å lines in 20 disc dwarf stars. Caffau et al. (2016) used the GIANO spectrograph at the Telescopio Nazionale Galileo (TNG) to measure these same lines in another four disc dwarf stars. Roederer et al. (2014) identified six P I lines using the Space Telescope Imaging Spectrograph (STIS) on board the Hubble Space Telescope

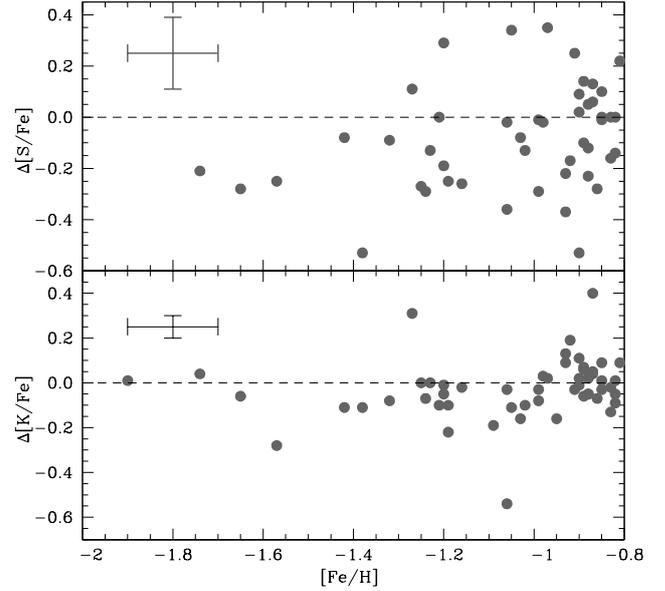


Fig. 4. [S, K/Fe] vs. [Fe/H], plotting the differences between the present results and the APOGEE DR17 values. Error bars correspond to the errors for the cool star b14 reported in Table 2.

(HST) in the near-UV, namely at 2154.080, 2154.121, 2533.976, 2535.603, 2553.262, and 2554.915 Å, relying in the end only on the P I 2136 Å line, for 14 halo stars. Roederer et al. (2016) used the Cosmic Origins spectrograph (COS), also on board the HST, to derive P and S abundances in carbon-rich metal-poor stars. Maas et al. (2019) used the Phoenix spectrograph at the Gemini-South telescope to measure the 10581.560 and 10596.960 Å lines in nine disc stars, and Maas et al. (2022) used the spectra obtained with the Hobby-Eberly Telescope at the McDonald Observatory to measure the P I 10529.52 Å line as the primary indicator of P abundances in 163 disc and outer halo stars. Sadakane & Nishimura (2022) used the ESPaDOnS spectrograph at the Canada–France–Hawaii telescope (CFHT) to observe the 9750.75 and 9796.83 Å lines in 45 main-sequence stars. Nandakumar et al. (2022) used the Immersion GRating INfrared Spectrometer (IGRINS) at the Gemini-South telescope to measure the P I 16482.92 Å line for 38 disc stars. Brauner et al. (2023) analysed 87 stars selected from the APOGEE DR17 survey, including the 16 P-rich stars previously found by Masseron et al. (2020a). Of this sample, 78 stars turned out to be enriched in P ([P/Fe] > +0.8).

Figure 5 shows the literature data compared with the present results. It is interesting that a significant part of the sample stars show a clearly enhanced phosphorus compared to the mean field value. P-rich stars have previously been found in other samples by Masseron et al. (2020a), Brauner et al. (2023), and Brauner et al. (2024). However, we note that Brauner et al. (2023) found more extreme cases of P enhancement than we observe here, which led them to consider a higher limit for the P abundance ([P/Fe] > +0.8) than we do here ([P/Fe] > +0.45) to consider the star as P-rich. Nevertheless, it is intriguing that the P-rich stars from Masseron et al. (2020a,b) and Brauner et al. (2023), discussed further in Brauner et al. (2024), have metallicities around [Fe/H] ~ -1.0 , coinciding with the same effect in our sample stars at the same metallicities. This might confirm the findings of our previous studies (Razera et al. 2022; Barbuy et al. 2023, 2024); namely, that there is an old population of stars that are the result of an early fast chemical enrichment in the Milky

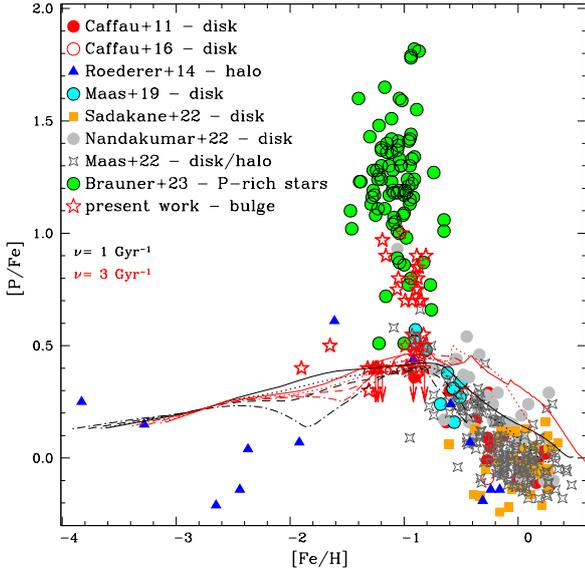


Fig. 5. $[P/Fe]$ vs. $[Fe/H]$ for the present results compared with literature data. Symbols – open red stars: present work, filled red circles: Caffau et al. (2011), open red circles: Caffau et al. (2016), filled blue triangles: Roederer et al. (2014), filled cyan circles and open black circles: Maas et al. (2019), filled light grey circles: Nandakumar et al. (2022), open four-pointed light grey stars: Maas et al. (2022), filled green circles and open black circles: Brauner et al. (2023). Different model lines correspond to the outputs of models computed for radii of $r < 0.5$, $0.5 < r < 1$, $1 < r < 2$, and $2 < r < 3$ kpc from the Galactic centre. Black lines correspond to specific star formation of $\nu = 1 \text{ Gyr}^{-1}$, and red lines to $\nu = 3 \text{ Gyr}^{-1}$.

Way towards the Galactic bulge. To confirm this conclusion, the orbital analysis of the Masseron–Brauner sample of P-rich stars led these authors to identify most of their sample as belonging to the thick disc and the inner Galactic halo, and thus also old populations. Moreover, the results obtained by Roederer et al. (2014) seem to indicate that no P-rich stars are present among the very metal-poor halo stars ($[Fe/H] \leq -2.0$), although this sample is small and thus statistically not very robust to draw firm conclusions.

Finally, we verified possible correlations between the P abundances and the other elements, C, N, O, Na, Al, Si, and Ca, and iron-peak elements previously studied (Razera et al. 2022; Barbuy et al. 2023, 2024). Brauner et al. (2023) found a strong correlation between P and Si in their P-rich star sample. However, this is not as clear in our study. While it is not possible to make a direct qualitative comparison between those two studies because of differences in stellar parameters and procedures, the difference in the correlation could also be due to the limited P abundance range of the current sample ($0.5 \leq [P/Fe] \leq 1.0$). Indeed, in Brauner et al. (2023), within such a limited range of P abundances, the P-Si correlation tends to vanish.

Sulphur: Figure 6 shows the present sulphur abundances compared with literature data, described below. Francois (1987) and Israelian & Rebolo (2001) analysed the S I 8693.958, 8694.641 Å lines. Takada-Hidai et al. (2002) gathered High Resolution Echelle Spectrometer (HIRES) observations with the samples observed by Francois (1987, 1988) and Clegg et al. (1981) for these same lines, and derived both LTE and non-LTE abundances. We adopted their NLTE $[Fe/H]$ and $[S/Fe]$ results for comparison purposes. Chen et al. (2002) used these same lines plus the lines S I 6046.03, 6052.67, and 6757.17 Å. Takada-Hidai et al. (2002) used the 8693.958 and 8694.641 Å

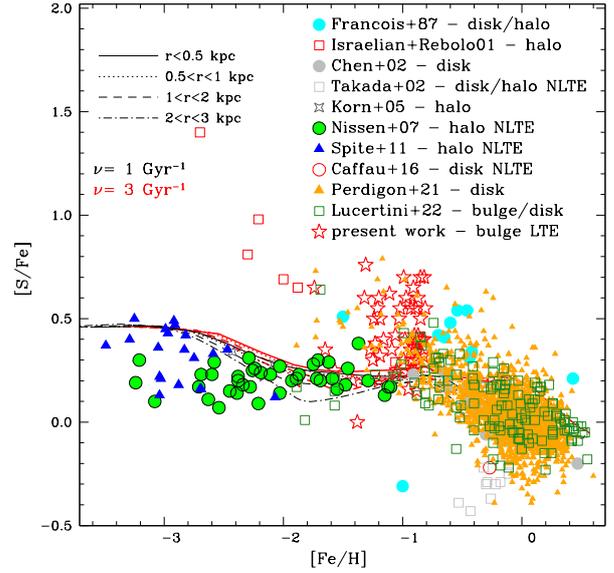


Fig. 6. $[S/Fe]$ vs. $[Fe/H]$ for the present results compared with literature data. Symbols – open red stars: present work, filled cyan circles: Francois (1987), open red squares: Israelian & Rebolo (2001), filled light grey circles: Chen et al. (2002), four-pointed light grey stars: Korn & Ryde (2005), filled green circles and open black circles: Nissen et al. (2007), filled blue triangles: Spite et al. (2011), open red circles: Caffau et al. (2016), filled orange triangles: Perdigon et al. (2021), open forest green squares: Lucertini et al. (2022). Same as in Fig. 5 for the model lines.

lines to derive the S abundances of giants and dwarfs from several samples, including their own, plus the ones from Francois (1987, 1988) and Clegg et al. (1981). Korn & Ryde (2005) used these lines and the S I 9212.9, 9228.1, and 9237.5 Å lines. Nissen et al. (2007) analysed the weaker S I 8694.6 and the stronger S I 9212.9 and 9237.5 Å lines, and presented abundances derived in LTE and NLTE. Spite et al. (2011) used several lines, namely S I 6756.851, 6757.007, 6757.171, 8693.931, 8694.626, 9212.863, 9228.093, 9237.538, 10455.449, 10456.757, 10459.406, and 10821.176 Å lines, and carried out NLTE calculations. Caffau et al. (2016) measured the 10455.449, 10456.757, and 10459.406 Å lines, and computed NLTE abundances. Perdigon et al. (2021) analysed the S I 6743.440, 6743.531, 6743.640, 6748.573, 6748.682, 6748.837, 6756.851, 6757.007, and 6757.171 Å lines, and computed LTE abundances, arguing that NLTE corrections would be smaller than 0.1 dex, following Takeda et al. (2016) and Korotin et al. (2017). Lucertini et al. (2022) measured the lines S I 6757.171, 8694.626, 9212.863, 9228.093, and 9237.538 Å.

Potassium: Figure 7 shows the K abundances for our sample of 58 stars compared with literature data. Zhang et al. (2006) derived K abundances from the K I 7664.92, and 7698.98 Å lines for moderately metal-poor disc stars, providing results corrected for NLTE effects. Andrievsky et al. (2010) computed NLTE K abundances for the sample halo stars from Cayrel et al. (2004), for the lines K I 5801.752, 7698.965, and 12432.274 Å. Reinhard & Laird (2024) derived LTE K abundances for 20 halo stars.

5. Chemical-evolution of P, S, K

5.1. Nucleosynthesis

Phosphorus: Brauner et al. (2024) considered the nucleosynthesis options to explain the P excess in stars of metallicity

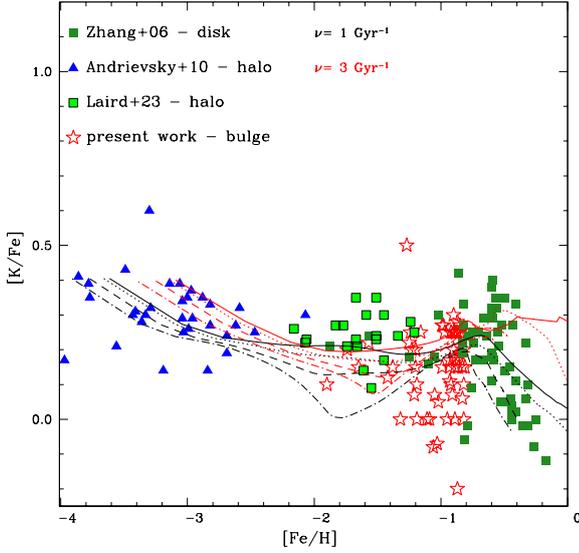


Fig. 7. $[K/Fe]$ vs. $[Fe/H]$ for the present results compared with literature data. Symbols – open red stars: present work, dark green squares: Zhang et al. (2006), filled blue triangles: Andrievsky et al. (2010), filled green squares and open black squares: Reinhard & Laird (2024). Same as in Fig. 5 for the model lines.

$[Fe/H] \approx -1.0$, and could not identify a process responsible for this effect. Normally, P is produced in CCSNe through neutron captures on the Si isotopes. ^{31}P , together with ^{28}Si and ^{30}Si , are produced in the oxygen- and neon-burning shells, and then ejected at the supernova core-collapse event (Woosley & Weaver 1995). During the explosion, according to Pignatari et al. (2016), there is some production of ^{31}P by explosive C burning and explosive He burning. One possible explanation for the large excess of P was given by Bekki & Tsujimoto (2024) through P production in oxygen-neon (ONe) novae. However, it is not clear why this would take place only in the environment where the old bulge stars of a moderate metallicity were formed.

There is also the possibility of enrichment by asymptotic giant branch (AGB) stars, where ^{31}P is the result of neutron-capture on ^{30}Si . However, this latter contribution is less likely as this is expected to occur in only very metal-poor AGBs (Karakas & Lugaro 2016).

Sulphur: sulphur is an α element; therefore, it is expected to behave like O, Mg, Si, and Ca, being produced in massive stars and ejected by SNe II.

Potassium: ^{39}K is the majority isotope, comprising 93.132%, followed by ^{41}K with 6.721% and ^{40}K with 0.147% (Asplund et al. 2009). The $^{39,41}\text{K}$ isotopes are dominantly produced in CCSNe (Pignatari et al. 2016).

5.2. Chemical-evolution models

The chemical-evolution model for the Galactic bulge derives from the chemical-evolution model for elliptical galaxies of Friaca & Terlevich (1998). This model consists of a multi-zone chemical evolution coupled with a hydrodynamical code. For the Galactic bulge, a classical spheroid with a baryonic mass of $2 \times 10^9 M_{\odot}$ and a dark halo mass of $1.3 \times 10^{10} M_{\odot}$ are assumed. Cosmological parameters from the Planck Collaboration VI (2020) are adopted; namely, $\Omega_m = 0.31$, $\Omega_{\Lambda} = 0.69$, Hubble constant $H_0 = 68 \text{ km s}^{-1} \text{ Mpc}^{-1}$, and an age of the Universe of $13.801 \pm 0.024 \text{ Gyr}$.

For the nucleosynthesis yields, we adopted: (i) for massive stars, the metallicity-dependent yields from CCSNe/SNe II from

Woosley & Weaver (1995), and for low metallicities ($Z < 0.01 Z_{\odot}$, or $[Fe/H] < -2.5$), the yields from high-explosion-energy HNe from Nomoto et al. (2013); (ii) type Ia supernovae (SNIa) yields from Iwamoto et al. (1999) – their models W7 (progenitor star of initial metallicity $Z = Z_{\odot}$) and W70 (zero initial metallicity); and (iii) for intermediate-mass stars ($0.8\text{--}8 M_{\odot}$) with initial $Z = 0.001, 0.004, 0.008, 0.02$, and 0.4 , yields from van den Hoek & Groenewegen (1997) with variable η (AGB case).

Our nucleosynthesis prescriptions also consider neutrino-interaction processes, which could be important at very low metallicities. Core-collapse supernovae release large amounts of energy as neutrinos ($> 10^{53} \text{ erg}$) during the formation of the neutron star or black hole. In this case, the interaction of these neutrinos with matter could significantly increase the yields of odd-Z elements (Yoshida et al. 2008). Therefore, for the supernovae with very low-metallicity progenitors, our models include the enhancements by neutrino processes of the production of odd-Z elements such as phosphorus and potassium. In this paper, we adopt the results of the calculations of Yoshida et al. (2008) for the case with the total neutrino energy of the supernova explosion, $E_{\nu} = 3 \times 10^{53} \text{ erg}$, and a neutrino temperature of $T_{\nu} = 8 \text{ Mev}$.

In our calculations, we multiplied the yields from WW95 for sub-solar metallicities by a factor of 2 for P and a factor of 3 for K. The factor of 2 for P is the same factor adopted for cobalt, as suggested by Timmes et al. (1995) and adopted by Erandes et al. (2020). The value of this correction comes from the comparison of $[P/Fe]$ and $[Fe/H]$ given by Figure 23 of Timmes et al. (1995) and the observations of Caffau et al. (2011). In the case of K, the Galactic chemical evolution model of Timmes et al. (1995), using yields of WW95, shows a sharp drop of $[K/Fe]$ below $[Fe/H] \approx 0.6$, in serious disagreement with the observations (see their Figure 24). The factor of 2 for P is similar, but smaller than the factor of 2.75 adopted by Cescutti et al. (2012). These corrections reflect the perception that the WW95 yields for $Z < Z_{\odot}$ tend to underestimate the production of some odd-Z elements. On the other hand, for $[Fe/H] < -2.5$, our models, which include HNe and neutrino processes in supernovae, account for the observations and do not need further corrections.

Models were computed for radii of $r < 0.5$, $0.5 < r < 1$, $1 < r < 2$, and $2 < r < 3 \text{ kpc}$ from the Galactic centre, and for specific star formation rate values of $\nu = 1$ and 3 Gyr^{-1} .

Figure 5 overplots the chemical-evolution models for specific star formation rates of $\nu = 1$ and 3 Gyr^{-1} for the phosphorus data. It is interesting to note that, at metallicities of $[Fe/H] \sim -1.0$, there is a maximum $[P/Fe] = +0.423$ at $[Fe/H] = -0.825$ for the model with $\nu = 1$, and $[P/Fe] = +0.479$ at $[Fe/H] = -0.827$ for the model with $\nu = 3$. Moreover, the maximum $[P/Fe]$ in the models approximately coincides with the metallicities of the P-rich stars, although the very high levels of P are not reached. This maximum arises from the yields of the Woosley & Weaver (1995) models. About one third of the stars show higher values relative to the models. In summary, the nucleosynthesis scenario for P excess has not been identified yet, and different possibilities remain open.

Figure 6 overplots the chemical-evolution models for specific star formation rates of $\nu = 1$ and 3 Gyr^{-1} , for our results and literature data for sulphur described above. Our data seems to indicate that sulphur is enhanced in the bulge. Although a first thought is that NLTE effects could be artificially enhancing the abundance and the spread, the grid of NLTE corrections by Korotin & Kiselev (2024) indicates only rather low values for the corrections. On the other hand, the amplitude of such an effect on sulphur can be notably illustrated by comparing the

LTE studies (e.g. [Israelian & Rebolo 2001](#)) – which also found enhanced sulphur in halo stars – to other studies of a similar population but in NLTE ([Nissen et al. 2007](#), e.g.). We also argue that sulphur as an α element is not expected to behave distinctively from other α elements such as O, Mg, and Ca. Previous studies of bulge stars ([Gonzalez et al. 2011](#)) have demonstrated that Mg, Ca, or Ti abundances in metal-poor bulge stars are indistinguishable from the ones of the thick disc and halo stars at the same metallicities. Therefore, we conclude that our sulphur abundances are very likely overestimated, and that S abundances are expected to be compatible with the ones of the disc and halo at similar metallicities such as the ones found in other stars of the bulge by [Lucertini et al. \(e.g. 2022\)](#). Regarding the comparison of the data with the models, we confirm that the latter satisfactorily reproduce the behaviour of [S/Fe] against [Fe/H] over the full metallicity range.

Figure 7 overplots the results for K, with the models showing very good agreement with the data.

6. Conclusions

This is the fourth paper in a series dealing with the chemical-abundance analysis of 58 stars selected to have $[\text{Fe}/\text{H}] < -0.8$ and orbits compatible with being members of a spheroidal component of the Galactic bulge. In the present work, we used suitable lines to compute the abundances of the elements P, S, and K. The abundances of P were not available from the APOGEE-ASPCAP DR17 results.

A striking result is the relative overabundance of P for about one third of the stars in the sample, a phenomenon that may be related to the P-rich star phenomenon described by [Masseron et al. \(2020a,b\)](#), [Brauner et al. \(2023\)](#), and [Brauner et al. \(2024\)](#), although more moderate. In the present work, the over-abundance of P is seen in some of these old bulge stars, for the first time, and from literature data. For the other two thirds of the sample, the P abundances are compatible with the chemical evolution models. Further samples should be analysed in order to have more solid statistics on the number of P-rich versus P-normal stars in the Galactic bulge. It is remarkable that no P-rich stars are present in the most metal-poor bulge stars as well as in the very metal-poor halo stars. It appears that these over-abundances are not observed in other stellar populations, and this might indicate the contribution of a particular kind of early supernova operating in the central part of the Galaxy. The nucleosynthesis process involved is, however, not yet clearly identified.

Sulphur is an α element, and behaves as such, as is seen in the comparison with chemical-evolution models. We stress though that the present results rely on one line, and the abundances have not been corrected for non-LTE departures, given that the [Korotin & Kiselev \(2024\)](#) grid is limited to higher temperatures relative to our sample. Their corrections indicate a negligible correction, but this problem should be further investigated in order to explain the discrepancies between our results and the models.

Potassium is another little-studied element. The chemical-evolution models fit the behaviour of [K/Fe] with regard to metallicity remarkably well; this is partly due to the inclusion of the ν -process yields in the models.

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Appendix A: Additional table

Table A.1. Internal and 2MASS star identification, the ASPCAP uncalibrated stellar parameters, recomputed C, N, O abundances, the abundances of P, S, and K derived in this work, and the APOGEE-ASPCAP abundances of S and K.

ID (internal)	ID (2MASS)	T_{eff} (K)	log g	[Fe/H]	v_r (km/s)	[C/Fe]	[N/Fe]	[O/Fe]	[P/Fe]	[S/Fe]	[S/Fe] _{APO}	[K/Fe]	[K/Fe] _{APO}
b1	2M17153858-2759467	3922.7±10	0.34±0.05	-1.65±0.02	2.62	b1 -0.60±0.03	+0.25±0.03	+0.20±0.03	+0.50±0.10	+0.35±0.05	+0.63/ 0.41 ±0.08	+0.15±0.03	+0.21±0.09
b2	2M17173693-2806495	3908.9±8	0.95±0.04	-0.97±0.01	2.20	b2 +0.15±0.03	+0.10±0.03	+0.38±0.03	+0.40±0.05	+0.60±0.05	+0.25±0.09	+0.25±0.03	+0.23±0.06
b3	2M17250290-2800385	3796.6±6	0.91±0.03	-0.82±0.01	2.39	b3 +0.15±0.03	+0.10±0.03	+0.40±0.03	+0.50±0.05	+0.40±0.05	+0.54/ 0.52 ±0.06	+0.00±0.03	+0.09±0.05
b4	2M17265563-2813558	4096.2±14	1.00±0.06	-1.32±0.02	1.89	b4 -0.30±0.03	+0.30±0.08	+0.20±0.03	+0.40±0.10	+0.60±0.05	+0.69±0.19	+0.00±0.08	+0.08±0.10
b5	2M17281191-2831393	4029.1±7	0.96±0.04	-1.19±0.01	1.73	b5 -0.30±0.03	+0.30±0.03	+0.40±0.03	<+0.40	+0.50±0.05	+0.75/ 0.58 ±0.01	+0.00±0.03	+0.22±0.06
b6	2M17295481-2051262	4205.9±13	1.50±0.05	-0.85±0.02	1.71	b6 +0.15±0.10	+0.20±0.10	+0.40±0.05	—	+0.40±0.05	+0.30±0.13	+0.19±0.03	+0.10±0.08
b7	2M17303581-2354453	3863.0±9	0.77±0.04	-0.98±0.02	2.13	b7 +0.07±0.03	+0.05±0.08	+0.50±0.03	+0.40±0.10	+0.40±0.05	+0.42±0.12	+0.27±0.03	+0.24±0.07
b8	2M17324257-2301417	3668.2±7	0.79±0.04	-0.82±0.02	2.30	b8 -0.10±0.03	+0.15±0.05	+0.35±0.05	<+0.40	+0.70±0.05	0.61 ±0.07	+0.20±0.03	+0.19±0.06
b9	2M17330695-2302120	3566.6±6	0.35±0.03	-0.93±0.01	2.42	b9 +0.30±0.03	+0.00±0.03	+0.55±0.03	+0.80±0.10	+0.55±0.05	—	+0.15±0.03	—
b10	2M17344841-4540171	3869.2±6	0.85±0.03	-0.88±0.01	2.16	b10 -0.30±0.03	+0.20±0.03	+0.35±0.03	+0.80±0.10	+0.15±0.10	+0.27±0.07	+0.18±0.03	+0.23±0.05
b11	2M17351981-1948329	3553.5±5	0.44±0.03	-1.11±0.01	3.06	b11 -0.10±0.03	+0.20±0.03	+0.30±0.05	—	+0.62±0.05	—	+0.00±0.03	—
b12	2M17354093-1716200	3895.5±7	1.01±0.03	-0.87±0.01	2.02	b12 +0.05±0.03	+0.30±0.03	+0.50±0.03	+0.70±0.10	+0.42±0.05	+0.29±0.08	+0.28±0.03	+0.24±0.05
b13	2M17390801-2331379	3740.4±6	0.83±0.03	-0.81±0.01	2.35	b13 +0.05±0.03	+0.20±0.03	+0.40±0.03	+0.90±0.10*	+0.55±0.05	+0.33/ 0.66 ±0.09	+0.15±0.03	+0.06±0.05
b14	2M17392719-2303131	3643.3±6	0.67±0.04	-0.87±0.01	2.55	b14 +0.20±0.03	+0.40±0.03	+0.60±0.03	+0.40±0.05	+0.38±0.05	—	+0.25±0.03	+0.20±0.06
b15	2M17473299-3224283	4018.3±9	0.47±0.05	-1.74±0.01	2.12	b15 -0.70±0.03	+0.80±0.03	+0.35±0.03	—	+0.65±0.05	+0.86/ 0.63 ±0.03	+0.20±0.10	+0.16±0.08
b16	2M17482995-2305299	4213.6±9	1.24±0.04	-1.03±0.01	2.10	b16 -0.30±0.10	+0.30±0.10	+0.30±0.08	—	+0.50±0.05	+0.58±0.09	-0.07±0.03	+0.09±0.06
b17	2M17483633-2242483	3651.5±6	0.44±0.03	-1.09±0.01	2.58	b17 -0.23±0.03	+0.10±0.03	+0.35±0.03	+0.40±0.05	+0.35±0.05	—	+0.00±0.03	+0.19±0.05
b18	2M17503263-3654102	3893.5±6	0.64±0.03	-0.99±0.01	2.19	b18 -0.10±0.03	+0.30±0.03	+0.25±0.05	+0.70±0.05	+0.30±0.05	+0.59±0.07	+0.27±0.03	+0.30±0.05
b19	2M17552744-3228019	4018.9±8	1.00±0.04	-1.06±0.01	2.00	b19 -0.40±0.03	+0.10±0.05	+0.10±0.08	—	+0.23±0.05	+0.59/ 0.46 ±0.08	-0.33±0.03	+0.21±0.06
b20	2M18020063-1814495	3988.8±10	0.80±0.05	-1.38±0.02	2.04	b20 -0.30±0.05	-0.10±0.05	+0.00±0.05	—	+0.00±0.10	+0.53±0.15	+0.15±0.03	+0.26±0.08
b21	2M18050452-3249149	3940.7±7	0.77±0.04	-1.16±0.01	2.08	b21 -0.40±0.03	+0.20±0.03	+0.40±0.03	+0.90±0.10	+0.40±0.05	+0.66±0.08	+0.25±0.03	+0.27±0.06
b22	2M18050663-3005419	3439.9±5	0.23±0.03	-0.92±0.01	2.52	b22 +0.12±0.03	+0.20±0.03	+0.35±0.03	<0.40	+0.30±0.05	—	+0.10±0.03	—
b23	2M18065321-2524392	3893.1±7	0.95±0.04	-0.89±0.01	2.02	b23 +0.00±0.03	+0.20±0.03	+0.38±0.03	+0.90±0.05	+0.38±0.05	+0.24±0.07	+0.00±0.03	+0.06±0.05
b24	2M18104496-2719514	4153.1±11	1.33±0.04	-0.82±0.02	2.05	b24 -0.10±0.10	-0.10±0.10	+0.10±0.10	—	+0.35±0.10	+0.35±0.11	+0.10±0.03	+0.15±0.07
b25	2M18125718-2732215	3617.2±6	0.44±0.03	-1.31±0.01	2.64	b25 +0.00±0.03	+0.40±0.03	+0.45±0.03	+0.30±0.05	+0.76±0.05	0.88 ±0.05	—	-0.85±0.06
b26	2M18200365-3224168	3976.6±6	0.95±0.04	-0.86±0.01	1.94	b26 -0.20±0.08	+0.20±0.08	+0.32±0.03	+0.50±0.05	+0.30±0.05	+0.58/ 0.41 ±0.05	+0.15±0.03	+0.22±0.05
b27	2M18500307-1427921	4077.0±9	1.23±0.04	-0.95±0.01	1.73	b27 -0.10±0.03	+0.20±0.03	+0.30±0.08	—	+0.16±0.05	+0.80±0.09	+0.00±0.03	+0.16±0.06
c1	2M17173248-2518529	3977.0±9	1.00±0.04	-0.91±0.02	1.81	c1 -0.10±0.03	+0.20±0.03	+0.33±0.03	+0.50±0.10	+0.65±0.05	+0.40±0.11	+0.07±0.03	+0.10±0.07
c2	2M17285088-2855427	3838.0±0	0.63±0.04	-1.23±0.02	2.18	c2 -0.30±0.03	+0.20±0.03	+0.50±0.03	<+0.40	+0.38±0.05	+0.51±0.13	+0.20±0.03	+0.20±0.07
c3	2M17301495-2337002	3814.0±6	0.69±0.03	-1.06±0.01	2.22	c3 -0.35±0.03	+0.00±0.03	+0.25±0.03	+0.75±0.05	+0.42±0.05	+0.44±0.08	-0.08±0.03	-0.05±0.05
c4	2M17453659-2309130	4133.1±9	1.27±0.04	-1.20±0.01	1.08	c4 -0.30±0.05	+0.10±0.08	+0.25±0.05	—	+0.20±0.05	+0.39±0.11	+0.15±0.03	+0.20±0.07
c5	2M17532599-2053304	3896.9±6	0.91±0.03	-0.87±0.01	2.10	c5 -0.05±0.03	+0.25±0.03	+0.30±0.03	+0.85±0.05	+0.60±0.05	+0.54/ 0.60 ±0.03	-0.20±0.03	-0.60±0.05
c6	2M18044663-3132174	3832.6±6	0.92±0.03	-0.90±0.01	2.22	c6 +0.10±0.03	+0.35±0.05	+0.35±0.05	+0.45±0.10	+0.42±0.05	+0.33±0.07	+0.30±0.03	+0.28±0.05
c7	2M18080306-3125381	4310.0±12	1.57±0.04	-0.90±0.02	1.48	c7 -0.07±0.03	+0.25±0.03	+0.35±0.03	+0.50±0.10	+0.25±0.10	+0.23±0.11	+0.20±0.03	+0.09±0.07
c8	2M18195859-1912513	4102.0±10	1.05±0.04	-1.24±0.01	1.78	c8 -0.20±0.03	+0.40±0.03	+0.25±0.03	+0.40±0.10	+0.50±0.10	+0.79±0.11	+0.25±0.03	+0.32±0.07
c9	2M17190320-2857321	4139.6±12	1.19±0.05	-1.20±0.01	1.83	c9 -0.30±0.05	+0.20±0.08	+0.50±0.05	—	+0.60±0.05	+0.31±0.16	+0.20±0.10	+0.21±0.09
c10	2M17224443-2343053	4058.3±7	1.02±0.03	-0.88±0.01	1.97	c10 -0.10±0.03	+0.20±0.03	+0.37±0.03	+0.75±0.05	+0.30±0.05	+0.53/ 0.42 ±0.02	+0.25±0.03	+0.23±0.05
c11	2M17292082-2126433	3983.4±7	0.78±0.04	-1.27±0.01	2.59	c11 -0.20±0.03	+1.10±0.05	+0.75±0.05	+0.40±0.05	+0.35±0.10	+0.24±0.09	+0.50±0.03	+0.19±0.06
c12	2M17323787-2023013	3865.7±6	1.03±0.03	-0.85±0.01	1.94	c12 +0.07±0.03	+0.20±0.03	+0.40±0.03	+0.70±0.05	+0.70±0.10	+0.70/ 0.49 ±0.018	+0.25±0.03	+0.24±0.05
c13	2M17330730-2407378	4042.5±9	0.25±0.05	-1.90±0.01	1.88	c13 -0.60±0.03	+0.50±0.03	+0.30±0.03	+0.40±0.10	+0.20±0.10	-0.30±0.15	+0.10±0.10	+0.09±0.08
c14	2M18023156-2834451	3617.4±6	0.42±0.04	-1.19±0.01	3.02	c14 -0.05±0.03	+0.20±0.03	+0.30±0.03	+0.97±0.05	+0.32±0.10	—	+0.10±0.03	+0.20±0.06
c15	2M17291778-2602468	3844.3±7	0.71±0.04	-0.99±0.01	2.10	c15 +0.05±0.03	+0.30±0.03	+0.38±0.03	+0.50±0.10	+0.70±0.05	+0.71±0.08	+0.15±0.03	+0.23±0.06
c16	2M17310874-2956542	4175.7±10	1.20±0.05	-0.93±0.01	2.07	c16 -0.10±0.03	+0.20±0.03	+0.30±0.03	+0.55±0.10	+0.35±0.05	+0.57±0.10	+0.17±0.03	+0.08±0.07
c17	2M17382504-2424163	3880.4±6	0.99±0.04	-1.05±0.01	1.55	c17 -0.20±0.03	+0.30±0.03	+0.40±0.03	+0.80±0.05	+0.25±0.10	-0.09±0.08	+0.07±0.03	+0.18±0.05
c18	2M17511568-3249403	3921.2±9	0.98±0.05	-0.90±0.02	2.04	c18 +0.00±0.03	+0.00±0.03	+0.38±0.03	+0.75±0.10	+0.36±0.05	+0.89±0.10	+0.25±0.03	+0.26±0.07
c19	2M17552681-3334272	4051.0±10	1.08±0.05	-0.89±0.02	1.98	c19 -0.10±0.03	+0.00±0.03	+0.40±0.03	+0.85±0.05	—	+0.27±0.11	+0.25±0.03	+0.18±0.07
c20	2M18005152-2916576	4158.9±12	1.04±0.06	-1.02±0.02	2.21	c20 -0.35±0.03	+0.20±0.03	+0.40±0.03	+1.00±0.05	+0.55±0.05	+0.68±0.14	+0.05±0.10	+0.15±0.08
c21	2M18010424-3126158	3773.1±6	0.68±0.03	-0.83±0.01	2.20	c21 -0.05±0.03	+0.00±0.03	+0.38±0.03	+0.55±0.05	+0.32±0.05	+0.32/ 0.38 ±0.03	+0.15±0.03	+0.17±0.05
c22	2M18042687-2928348	4164.7±11	0.88±0.05	-1.21±0.02	2.14	c22 -0.30±0.03	+0.40±0.08	+0.30±0.05	—	+0.55±0.05	+0.55±0.13	+0.07±0.10	+0.17±0.08
c23	2M18052388-2953056	4252.9±14	0.92±0.06	-1.57±0.02	1.92	c23 -0.35±0.08	+0.40±0.10	+0.40±0.05	—	+0.25±0.10	+0.50±0.19	+0.22±0.20	+0.50±0.10
c24	2M18142265-0904155	3920.5±11	1.12±0.05	-0.85±0.02	2.13	c24 +0.00±0.03	+0.20±0.03	+0.33±0.03	+0.45±0.10	+0.50±0.10	+0.51±0.13	+0.27±0.03	+0.30±0.08
c25	2M17293482-2741164	4143.5±8	1.03±0.04	-1.25±0.01	1.85	c25 -0.20±0.03	+0.30±0.03	+0.40±0.03	<+0.40	+0.30±0.05	+0.57/ 0.33 ±0.10	+0.22±0.03	+0.22±0.06
c26	2M17341796-3905103	4163.5±13	1.42±0.05	-0.89±0.02	1.84	c26 +0.00±0.03	+0.25±0.03	+0.40±0.03	+0.60±0.10	+0.55±0.10	+0.65±0.13	+0.15±0.15	+0.09±0.08
c27	2M17342067-3902066	4380.4±16	1.40±0.05	-0.93±0.02	1.99	c27 +0.15±0.06	+0.20±0.06	+0.40±0.06	+0.70±0.10	+0.25±0.10	+0.62±0.15	+0.11±0.10	-0.02±0.09
c28	2M17503065-2313234	3821.4±6	0.98±0.03	-0.88±0.01	2.10	c28 +0.10±0.03	+0.10±0.03	+0.35±0.05	—	+0.40±0.05	+0.35±0.07	+0.17±0.03	+0.22±0.05
c29	2M18143710-2650147	4244.5±11	1.30±0.04	-0.92±0.01	1.97	c29 -0.10±0.08	+0.10±0.10	+0.40±0.08	—	+0.25±0.05	+0.42±0.10	+0.25±0.10	+0.06±0.07
c30	2M18150516-2708486	3833.4±9	1.00±0.04	-0.83±0.02	2.14	c30 +0.10±0.03	+0.00±0.05	+0.30±0.05	—	+0.40±0.05	+0.56±0.10	+0.06±0.03	+0.19±0.07
c31	2M18344461-2415140	4297.6±11	1.09±0.05	-1.42±0.01	1.83	c31 -0.45±0.10	+0.40±0.10	+0.25±0.10	—	+0.20±0.10	+0.28/ 0.12 ±0.22	+0.12±0.08	+0.23±0.08

Notes: [S/Fe] values in bold face are from the Value Added Catalogue (VAC) data derived with BAWLAS; * star b13 could be uncertain, because the PI line is at the border of the spectrum just before the wavelength gap.