

# A multiwavelength study of Swift GRB 060111B constraining the origin of its prompt optical emission<sup>\*</sup>

G. Stratta<sup>1</sup>, A. Pozanenko<sup>2</sup>, J.-L. Atteia<sup>3</sup>, A. Klotz<sup>4</sup>, S. Basa<sup>5</sup>, B. Gendre<sup>5</sup>, F. Verrecchia<sup>1</sup>, M. Boër<sup>6</sup>, S. Cutini<sup>1</sup>, M. Henze<sup>7</sup>, S. Holland<sup>8</sup>, M. Ibrahimov<sup>9</sup>, F. Ienna<sup>3</sup>, I. Khamitov<sup>10</sup>, S. Klose<sup>7</sup>, V. Rumyantsev<sup>11</sup>, V. Biryukov<sup>12</sup>, D. Sharapov<sup>9</sup>, F. Vachier<sup>13</sup>, S. Arnouts<sup>5,14</sup>, and D. A. Perley<sup>15</sup>

<sup>1</sup> ASDC, ASI Science Data Center, via Galileo Galilei, 00044 Frascati, Italy\*\*  
e-mail: [giulia.stratta@asdc.asi.it](mailto:giulia.stratta@asdc.asi.it)

<sup>2</sup> Space Research Institute IKI, Profsoyznaya 84/32, 117910 Moscow, Russia

<sup>3</sup> LATT, Université de Toulouse, CNRS, 14 Av. E. Belin, 31400 Toulouse, France

<sup>4</sup> CESR, Université de Toulouse, CNRS, BP 4346, 31028 Toulouse Cedex 04, France

<sup>5</sup> LAM, Observatoire Astronomique de Marseille, Pôle de l'Étoile, Site de Château-Gombert, 38 rue Frédéric Joliot-Curie, 13388 Marseille Cedex 13, France

<sup>6</sup> Observatoire de Haute Provence, 04870 Saint Michel l'Observatoire, France

<sup>7</sup> Thüringer Landessternwarte Tautenburg, Sternwarte 5, 07778 Tautenburg, Germany

<sup>8</sup> NASA, Goddard Space Flight Center, Greenbelt, MD 20771, USA

<sup>9</sup> Ulugh Beg Astronomical Institute, Tashkent 700052, Uzbekistan

<sup>10</sup> TÜBİTAK National Observatory, Akdeniz Üniversitesi, 07058 Antalya, Turkey

<sup>11</sup> SRI Crimean Astrophysical Observatory, Nauchny, Crimea 98409, Ukraine

<sup>12</sup> Crimean Laboratory of the Sternberg Astronomical Institute, Nauchny, Crimea 98409, Ukraine

<sup>13</sup> IMCCE, Observatoire de Paris, 77 Avenue Denfert Rochereau, 75014 Paris, France

<sup>14</sup> Canada France Hawaii Telescope corporation, 65-1238 Mamalahoa Hwy, Kamuela, HI 96743, USA

<sup>15</sup> Department of Astronomy, University of California, Berkeley, CA 94720-3411, USA

Received 3 March 2009 / Accepted 13 June 2009

## ABSTRACT

**Context.** The detection of bright optical emission measured with good temporal resolution during the prompt phase of GRB 060111B makes this GRB a rare event that is especially useful for constraining theories of the prompt emission.

**Aims.** For this reason an extended multi-wavelength campaign was performed to further constrain the physical interpretation of the observations.

**Methods.** In this work, we present the results obtained from our multi-wavelength campaign, as well as from the public Swift/BAT, XRT, and UVOT data.

**Results.** We identified the host galaxy at  $R \sim 25$  mag from deep  $R$ -band exposures taken 5 months after the trigger. Its featureless spectrum and brightness, as well as the non-detection of any associated supernova 16 days after the trigger, enabled us to constrain the distance scale of GRB 060111B within  $0.4 \leq z \leq 3$  in the most conservative case. The host galaxy spectral continuum is best fit with a redshift of  $z \sim 2$ , and other independent estimates converge to  $z \sim 1$ – $2$ . From the analysis of the early afterglow SED, we find that non-negligible host galaxy dust extinction, in addition to the Galactic one, affects the observed flux in the optical regime. The extinction-corrected optical-to-gamma-ray SED during the prompt emission shows a flux density ratio  $F_{\gamma}/F_{\text{opt}} = 10^{-2}$ – $10^{-4}$  with spectral index  $\beta_{\gamma, \text{opt}} > \beta_{\gamma}$ , strongly suggesting a separate origin of the optical and gamma-ray components. This result is supported by the lack of correlated behavior in the prompt emission light curves observed in the two energy domains. The temporal properties of the prompt optical emission observed during GRB 060111B and their similarities to other rapidly-observed events favor interpretation of this optical light as radiation from the reverse shock. Observations are in good agreement with theoretical expectations for a thick shell limit in slow cooling regime. The expected peak flux is consistent with the observed one corrected for the host extinction, likely indicating that the starting time of the TAROT observations is very near to or coincident with the peak time. The estimated fireball initial Lorentz factor is  $\Gamma \geq 260$ – $360$  at  $z = 1$ – $2$ , similar to the Lorentz factors obtained from other GRBs. GRB 060111B is a rare case of a GRB with both a bright, well-observed optical counterpart and a “canonical” early X-ray light curve, thus providing a good test case of the reverse shock emission mechanism in both energy ranges.

**Key words.** gamma rays: bursts – radiation mechanisms: non-thermal – methods: data analysis

\* Based on observations performed with: TAROT at the Calern observatory. The Nordic Optical Telescope, operated on the island of La Palma jointly by Denmark, Finland, Iceland, Norway, and Sweden, in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias. CFHT, which is operated by the National Research Council of Canada, the Institut National des Sciences de l'Univers of the CNRS of France, and the University of Hawaii. The

2.6 m Shajn telescope at the Crimean Astrophysical Observatory. Some of the data presented here were obtained at the W. M. Keck Observatory, which is operated as a scientific partnership among the California Institute of Technology, the University of California and the National Aeronautics and Space Administration. The Observatory was made possible by the generous financial support of the W. M. Keck Foundation.

\*\* INAF personnel resident at ASDC

## 1. Introduction

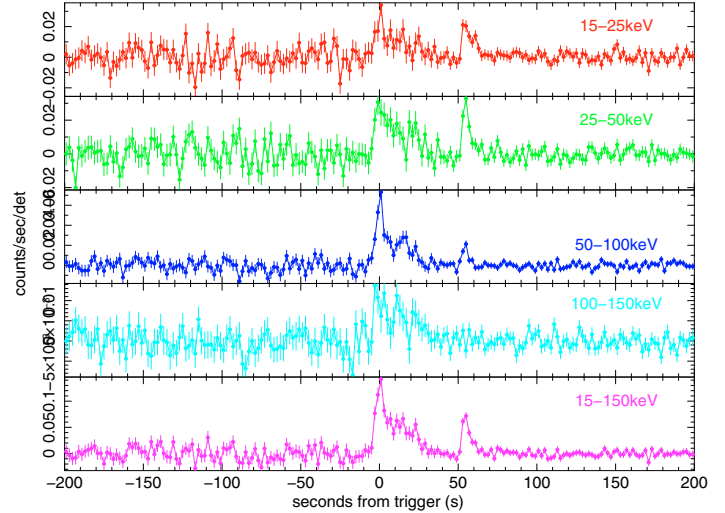
Measuring the multi-wavelength spectrum of the prompt emission of gamma-ray bursts (GRBs) is a challenge, requiring coordinated observations both from space and on the ground that are performed on very short notice (a few seconds). Such observations are nevertheless sorely needed to constrain the nature of the prompt emission. Rapid-response telescopes (e.g. Akerlof et al. 2003; Boër 2001; Pérez-Ramírez et al. 2004; Covino et al. 2004; Bloom et al. 2006) and developments in the technique of wide-field optical surveys (e.g. Vestrand et al. 2002; Pozanenko et al. 2004; Cwiok et al. 2005; Tamagawa et al. 2005) are fundamental to this purpose. So far, a few dozen GRBs have been detected at optical wavelengths when the GRB was still active, or soon after its end (e.g. Blake et al. 2005; Yuan et al. 2009; Gendre et al. 2008b). In several cases (e.g. GRB 990123, GRB 080319B), the optical flux after the peak emission shows a very steep decay that renders the detection of this component even more challenging (e.g. Melandri et al. 2008). The demanding rapidity of observing these bright optical flashes (about a dozen seconds after the trigger or less) is on average fulfilled by automatic ground-based telescopes but not by the onboard *UV-Telescope Swift*/UVOT, whose first observation times range from 40 to 200 s after the trigger (Romano et al. 2009). It is thus difficult to infer the true fraction of GRBs showing this emission.

The common interpretation of the observed emission from GRBs relies on the fireball internal-external shock model (see e.g. Zhang 2007, for a recent review). Whether the origin of the optical counterpart observed during the prompt emission is from internal shocks, like the high-energy emission, or from an external (reverse) shock, is still an open debate in the context of this model (Genet et al. 2007b). Answering this question will provide useful insight into the formation of GRB shocks. In general, the early optical emission might be a superposition of internal shocks, reverse shock and forward shock, with different contributions at different phases of the time history.

The first Swift-BAT position notice of GRB 060111B was distributed at 20:16:03 UT, 20 s after the trigger, allowing the TAROT and ROTSE III robotic telescopes to observe during prompt emission (Klotz et al. 2006; Yost et al. 2006; Rykoff et al. 2009). TAROT Calern (France) started observing at 20:16:11 UT (28 s after trigger) and detected a bright optical emission with  $R \sim 13.7$ , fading as  $t^{-2.4 \pm 0.1}$  until 80 s (Klotz et al. 2006). The position of the afterglow as measured by MITSuME from 45-min exposure images taken about 10 min after the trigger is  $RA = 19^h05^m42.47^s$ ,  $Dec = +70^\circ22'33.1''$  (J2000.0), with a precision of  $0.3''$  (Yanagisawa et al. 2006). This is fully compatible with the position measured by UVOT with an accuracy of  $0.5''$  ( $RA = 19^h05^m42.48^s$ ,  $Dec = +70^\circ22'33.6''$ , Perri et al. 2006a).

Because of the high declination of the source, ground-based optical follow-up has been challenging: the position cannot be observed from any major South Hemisphere observatory, and while it is circumpolar in the northern hemisphere, at late times the hour angle of the source is very large throughout the night and not accessible to equatorial telescope mounts. Despite these unfavorable conditions we set-up an observing campaign that provided many important results: the identification of the host galaxy, strong constraints on any supernova associated with the GRB and on the host-galaxy redshift. We present in Table 2 the list of the observations of the GRB error box that led to the results presented in this paper.

The first part of this paper deals with the observations of GRB 060111B (Sect. 2). Results from the multiwavelength



**Fig. 1.** Swift/BAT count rate of GRB 060111B in the ranges (from top to bottom): 15–25 keV, 25–50 keV, 50–100 keV, 100–150 keV and 15–150 keV. Count rate units are counts per seconds per illuminated-detector (note illuminated-detector =  $0.16 \text{ cm}^2$ ).

prompt and afterglow analysis are presented in Sect. 3. Results from the host galaxy identification, as well as the redshift constraints of the source, are presented in Sect. 4, while in Sect. 5 we discuss the nature of the prompt optical emission in light of our multiwavelength campaign results.

## 2. Observations and data analysis

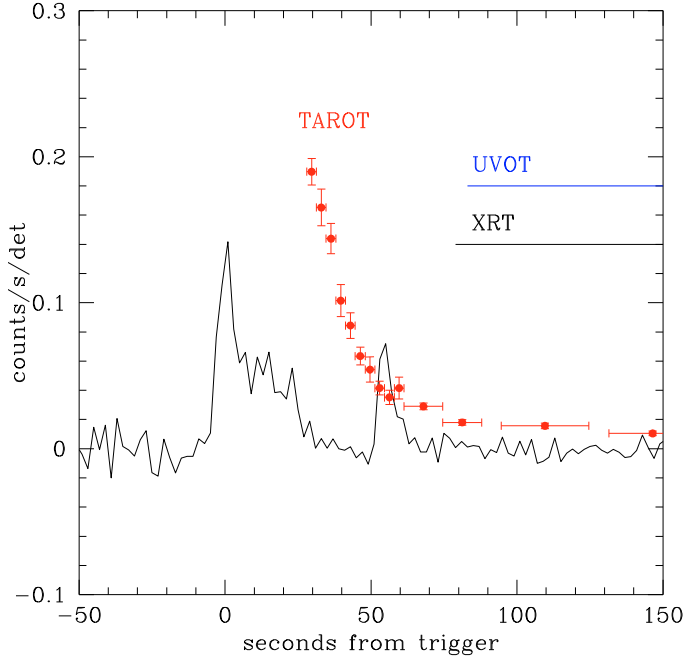
### 2.1. Swift/BAT and WAM/Suzaku

GRB 060111B was a moderately bright GRB detected on 2006, January 11 at 20:15:43.24 UT (hereafter  $t_{\text{trig}}$ ) by the BAT instrument on the Swift spacecraft (Perri et al. 2006a; Tueller et al. 2006). This burst exhibits two peaks reaching their maxima at  $t_{\text{trig}} + 2$  s and  $t_{\text{trig}} + 55$  s. While the first peak is clearly detected in the full BAT energy band, the second is significantly softer (Fig. 1). The duration of the burst is  $T_{90} = 59 \pm 5$  s in the range 15–150 keV. The time-averaged spectrum (from  $t_{\text{trig}} - 2$  to  $t_{\text{trig}} + 63$  s) is well-fit by a power-law model, with best-fit photon index  $\Gamma_\gamma = 0.9 \pm 0.1$ . The observed peak energy flux and fluence in the 15–150 keV band are  $2.6 \times 10^{-8} \text{ erg cm}^{-2} \text{ s}^{-1}$  and  $1.6 \times 10^{-6} \text{ erg cm}^{-2}$ , respectively.

The prompt high-energy emission was also observed by WAM on Suzaku, with a duration of  $T_{90} = 25.4$  s and a fluence of  $(5.6 \pm 0.8) \times 10^{-6} \text{ erg cm}^{-2}$  in the 100–1000 keV band (Sato et al. 2006). A joint spectral analysis of the prompt emission with Swift/BAT and Suzaku/WAM data assuming a Band model, provides a best-fit, observer-frame peak energy of  $E_{\text{peak}} = 462^{+290}_{-242} \text{ keV}$  (Fukazawa et al. 2007). This value is high enough to consider the spectral index measured by Swift to be the low-energy spectral index of the Band function describing the energy spectrum of the prompt emission.

### 2.2. Swift/XRT

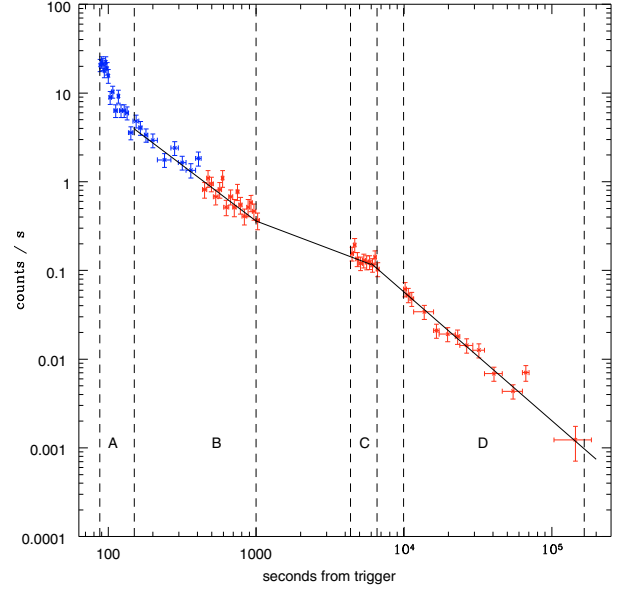
Swift slewed immediately to the burst, and the X-Ray Telescope (XRT, Burrows et al. 2005) began data acquisition at  $t_{\text{trig}} + 87$  s in windowed timing (WT) mode. At  $t_{\text{trig}} + 153$  s the count rate was at a critical level for XRT acquisition mode and an automatic



**Fig. 2.** Swift/BAT count rate (in units of counts per seconds per illuminated-detector) of GRB 060111B in the range 15–150 keV. The TAROT light curve superimposed on the BAT count rate is  $c_R = k10^{-0.4m_R}$  with  $k = 10^4$  in arbitrary units. The temporal ranges covered by UVOT and XRT data are also plotted.

flickering between the WT mode and the photon counting (PC) mode started. At that time, the afterglow was too bright for the PC mode, and the PC data were affected by moderate pile-up. At  $t_{\text{trig}} + 426$  s, the XRT switched permanently to the PC mode and continued observations until  $t_{\text{trig}} + 2.3$  days. In the following, we used WT data taken between  $t_{\text{trig}} + 87$  s and  $t_{\text{trig}} + 426$  s, and PC data taken later on. The XRT data were processed following standard procedures (Capalbi et al. 2005<sup>1</sup>), by using the version 0.11.6 of the dedicated pipeline xrtpipeline. Grade filtering was applied by selecting the 0–2 and 0–12 ranges for the WT mode and PC mode data, respectively. During the first orbit, the PC mode data show a count rate above 0.5 counts  $s^{-1}$  thus likely affected by moderate pile-up. From point-spread function fitting (Moretti et al. 2005) we excluded a circular region of 1 px radius from the extraction region of the source.

The X-ray (0.3–10 keV) light curve from 87 s to 2.3 days after the trigger shows a steep initial decay up to 150 s after the trigger that is poorly approximated by a power law. Starting from 150 s after the trigger we find that the data are best fitted by a double broken power law model, with decay indexes  $\alpha_{X,1} = 1.25 \pm 0.15$  up to 1 ks,  $\alpha_{X,2} = 0.6 \pm 0.2$  up to  $t_b = 6 \pm 1$  ks and  $\alpha_{X,3} = 1.4 \pm 0.1$  later on ( $\chi^2/\text{d.o.f.} = 40/43$ ). These values are quite typical for Swift/XRT afterglow light curves (e.g. Liang et al. 2007). We note that these results are marginally consistent with the results from an automatic analysis of 318 GRBs by Evans et al. (2008) where, for GRB 060111B,  $\alpha_{X,2} \sim 1$  from 175 s to  $10^4$  s after the trigger, corresponding to the average of our  $\alpha_{X,1}$  and  $\alpha_{X,2}$ . By fitting PC data only, a simple power law and a broken power law model are rejected with  $\geq 99.99\%$  and  $\sim 80\%$  confidence levels, respectively.



**Fig. 3.** The 0.3–10 keV light curve (blue data are taken in Windowed Timing mode and red data are taken in Photon Counting mode by XRT). The continuum solid line shows the best fit temporal model (see Sect. 2.2). For each plotted temporal interval (A, B, C and D) we extracted an energy spectrum (Table 4).

In order to investigate the spectral properties of each segment individuated in the light curve, we extracted the 0.3–10 keV spectrum of the X-ray afterglow at 4 different epochs, indicated with letters A, B, C, and D as plotted in Fig. 3. Assuming an absorbed power-law model we found no evidence of spectral variation during the evolution of the afterglow, with a photon index of  $\sim 2$  and an equivalent hydrogen column density of  $N_H \sim 2 \times 10^{21} \text{ cm}^{-2}$  in addition to the Galactic one ( $N_{H,\text{Gal}} = 7.65 \times 10^{20} \text{ cm}^{-2}$  from Kalberla et al. 2005). These results are summarized in Table 4. Our temporal and spectral results are in agreement with results found by Racusin et al. (2008b).

### 2.3. Swift/UVOT

The Swift follow-up with UVOT started 84 s after the trigger (Perri et al. 2006a,b; Holland 2006). We analyzed the UVOT images starting with the reduction of raw data using the standard Swift/UVOT software reduction tasks to create the sky images. We performed the analysis with the release 2.7.2 (HEADAS version 6.3.1) and the 20070711 version of the CALDB.

Our analysis involved the steps listed below. We performed the “standard” aperture photometry of the object in the sky images for each single exposure and, as comparison, in the images obtained as the sum of all the exposures in each filter. Photometry was performed using the UVOT-HEADAS tasks “uvotsource” and “uvotmaghist”. For this purpose we first selected the best source and background aperture regions from the first single exposure frame of the V filter. We applied the standard sizes for the source and background apertures, excluding the nearby contaminating objects from the background extraction apertures. The results of the photometry of each single exposure in each filter is reported in Table 1. We also report the dereddened flux obtained by correcting the observed magnitude using a mean Galactic interstellar extinction law (Fitzpatrick 1999) with a value of  $E(B - V) = 0.11 \text{ mag}$

<sup>1</sup> <http://swift.gsfc.nasa.gov/docs/swift/analysis/>

**Table 1.** Swift/UVOT single exposure frames analyzed following standard procedure and photometry results.

DATE UT	Filter	Exposure (s)	mag <sup>a</sup>	mag <sup>a</sup> (dereddened)	$F_{\nu}^a$ ( $10^{-28}$ erg cm <sup>-2</sup> s <sup>-1</sup> Hz <sup>-1</sup> )	Log frequency	Log ( $\nu F_{\nu}$ )
2006-01-11 20:17:06	V	196.6	17.59±0.11	17.24 ± 0.11	46.19 ± 4.67	14.74	-11.59 ± 0.10
2006-01-11 20:29:35	V	196.6	>18.86	>18.51	>14.34	14.74	-12.10 ± 0.00
2006-01-11 21:52:36	V	118.8	>18.86	>18.51	>14.34	14.74	-12.10 ± 0.00
2006-01-11 23:12:09	V	196.6	>19.29	>18.94	>9.65	14.74	-12.27 ± 0.00
2006-01-11 20:20:33	B	196.6	19.73 ± 0.32	19.27 ± 0.32	79.46 ± 2.34	14.84	-12.26 ± 0.29
2006-01-11 20:33:01	B	93.4	>19.24	>18.78	>12.48	14.84	-12.06 ± 0.00
2006-01-11 22:06:15	B	65.5	>18.98	>18.52	>15.85	14.84	-11.96 ± 0.00
2006-01-11 23:37:36	B	221.2	>19.84	>19.38	>71.80	14.84	-12.30 ± 0.00
2006-01-11 23:41:24	B	221.2	>19.67	>19.21	>83.97	14.84	-12.23 ± 0.00
2006-01-11 23:45:12	B	126.6	>19.43	>18.97	>10.47	14.84	-12.14 ± 0.00

[<sup>a</sup>] The upper limits are given at  $3\sigma$  confidence level.

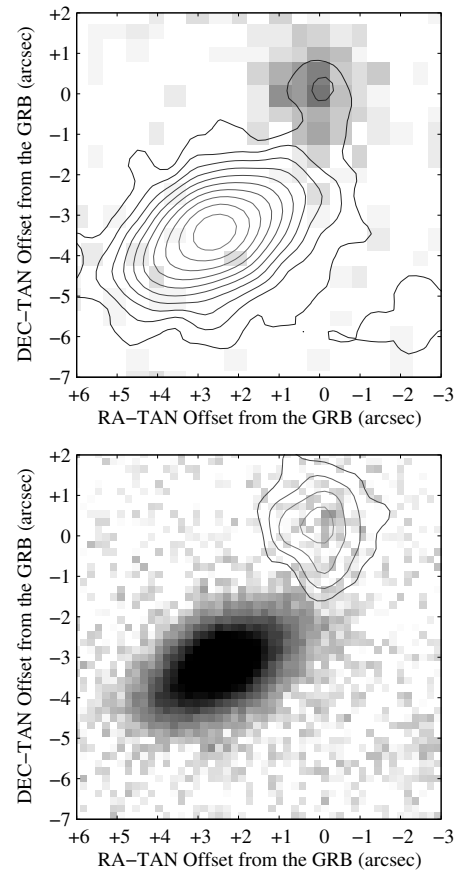
(Schlegel et al. 1998). The logarithm of the  $\nu F_{\nu}$  calculated from this flux, i.e. the Spectral Energy Distribution, is also reported.

We further refined the analysis of the first single exposure frame of the UVOT image of filter V by applying “non-standard” apertures for the source extraction and by applying the aperture correction procedure in order to enhance the signal to noise ratio. Various “non standard” extraction regions, such as with larger radii or different shapes (for example an elliptical region), have been tested, but generally give results similar to those obtained using the standard extraction radius. For the aperture correction procedure a first photometry on apertures of 1.9'' (the FWHM of the V filter) was performed. Then 6 objects not saturated in the field were selected and photometry was performed using FWHM apertures and standard apertures (5'' radius) for each of them. The difference of the magnitudes in the two different apertures was evaluated for the stars selected and finally the weighted mean among all the differences was calculated. This value was subtracted to the GRB magnitude obtained with FWHM aperture. The final value of the GRB magnitude thus obtained integrating over all the exposure time is  $V = 17.55 \pm 0.08$  mag that is consistent with our standard analysis. The upper panel of Fig. 4 shows the field of the GRB in filter V, during the first single exposure frame of sky image.

In order to increase the temporal resolution, the event data in all filters were also analyzed to extract the UVOT lightcurve of the object. Event data were first screened to filter out unwanted contaminations using standard selection criteria, as excluding the South-Atlantic Geomagnetic Anomaly and possible anomalous attitude variations, keeping the angle between the pointing and the Earth limb greater than 10 degrees and the same angle with “bright” Earth limb greater than 20 degrees. Moreover, only events classified as “good” were accepted. The lightcurve for each filter was extracted using the UVOT task “uvotevtcl”, with different binning for the V filter (30 s bin) as compared to the remaining (70 s bin except for the White for which 100 s bin has been used). The photometry is calibrated using the photometric zero points in the Swift/UVOT calibration data base. These zero point have a systematic error of 0.1 mag that is not included in the values quoted in Table 2.

#### 2.4. Ground based optical follow-up

Very early optical emission, measured with unprecedented temporal resolution of a few seconds by the TAROT robotic observatory, has been presented by Klotz et al. (2006). We quote here their results in Table 2. ROTSE-III observations started 33 s after the trigger and detected a bright optical flare which was initially at  $\sim 13$  mag and fading as  $\sim t^{-(2.35 \pm 0.10)}$ , in agreement with



**Fig. 4.** Top panel: R-band isophote curves of the host galaxy and the nearby bright galaxy from the CFHT image taken 159 days after the trigger. Superposition of the UVOT image taken 123 s after the trigger shows that the afterglow coincides with the host galaxy candidate from the CFHT image. Bottom panel: CFHT gray scale image is superposed to the UVOT isophotes of the afterglow.

TAROT observations (Yost et al. 2006; Rykoff et al. 2009). The field of GRB 060111B was also observed with the three-color MITSuME 50 cm telescope at Okayama, Japan starting at  $t_{\text{trig}} + 10.6$  min in the  $R_c$  and  $I_c$  band images (Yanagisawa et al. 2006).

Few hours after the trigger, GRB 060111B was observed with the Russian-Turkish 1.5-m telescope (RTT150, Bakırlıtepe, TÜBİTAK National Observatory, Turkey). In order to extract the maximum information on the OT, we carefully analyzed the sum and median stacks of 13 images of 300 s taken at RTT150 between 5.92 and 7.20 h after the trigger (see also



**Table 2.** Optical observations of GRB 060111B not corrected for Galactic extinction.

$T_{\text{mid}}^a$	$T_{\text{exp}}$ (s)	Filter	mag <sup>b</sup>	$\Delta\text{mag}^b$	Ref.
29.65 s	3.3	<i>R</i>	13.75	0.06	TAROT <sup>(1)</sup>
32.95 s	3.3	<i>R</i>	13.9	0.09	TAROT
36.3 s	3.4	<i>R</i>	14.05	0.09	TAROT
39.65 s	3.3	<i>R</i>	14.43	0.13	TAROT
42.95 s	3.3	<i>R</i>	14.63	0.13	TAROT
46.3 s	3.4	<i>R</i>	14.94	0.12	TAROT
49.65 s	3.3	<i>R</i>	15.11	0.2	TAROT
52.95 s	3.3	<i>R</i>	15.4	0.14	TAROT
56.3 s	3.4	<i>R</i>	15.58	0.17	TAROT
59.65 s	3.3	<i>R</i>	15.4	0.22	TAROT
67.95 s	13.3	<i>R</i>	15.79	0.10	TAROT
81.25 s	13.3	<i>R</i>	16.31	0.13	TAROT
109.6 s	30	<i>R</i>	16.46	0.15	TAROT
146.5 s	30	<i>R</i>	16.89	0.20	TAROT
183.3 s	30	<i>R</i>	17.53	0.30	TAROT
85.5 s	30	<i>V</i>	16.8	0.4	UVOT
115.5 s	30	<i>V</i>	17.2	0.5	UVOT
145.5 s	30	<i>V</i>	17.7	0.7	UVOT
175.5 s	30	<i>V</i>	17.6	0.7	UVOT
388 s	197	<i>B</i>	19.73	0.32	UVOT
792 s	300	<i>R</i>	18.9	0.4	MITSuME <sup>(2)</sup>
792 s	300	<i>I</i>	18.3	0.4	MITSuME <sup>(2)</sup>
0.28 days	3900	<i>R</i>	23.2	0.5	RTT-150
12.19 days	1440	<i>R</i>	>21.7		Maidanak-1.5
24.18 days	2700	<i>R</i>	>22.3		Maidanak-1.5
27.20 days	1800	<i>R</i>	>22.7		Maidanak-1.5
31.18 days	3420	<i>R</i>	>22.5		Maidanak-1.5
16.29 days	3360	<i>R</i>	25.1	0.5	ZTSh
108.0 days	1200	<i>V</i>	>21.6		ZTSh
26.3 days	3600	<i>R</i>	>21.0		OHP-T120
37.1 days	2700	<i>R</i>	>22.5		NOT
159 days	1840	<i>R</i>	24.8	0.4	CFHT
2.3 years	1800	<i>H</i>	> 23.0		CFHT
5.3 months	4800		3800–8500 Å	spectroscopy	Keck/LRIS

<sup>a</sup>  $T_{\text{mid}}$  is the mean epoch of the exposure of duration  $T_{\text{exp}}$ .

<sup>b</sup> Errors are given at  $1\sigma$ . The upper limits are given at  $3\sigma$  confidence level.

References. (1) all TAROT data are from Klotz et al. (2006); (2) Yanagisawa et al. (2006).

Khamitov et al. 2006). The afterglow is marginally detected with a signal to noise ratio of  $S/N = 2$ . A PSF fit of the theoretical location of the afterglow gives  $R = 23.2 \pm 0.5$ .

Further observations were performed with AZT-22 (Maidanak Astronomical Observatory, Uzbekistan) 2.6 m Shajn reflector (ZTSh, Crimean Astrophysical Observatory, Ukraine), OHP T120 (Observatoire de Haute-Provence, France) and NOT with only one detection at ZTSh 16.3 days after the trigger (see Table 2).

Late time observations of the host galaxy are discussed in Sect. 4.1.

### 3. Multiwavelength analysis

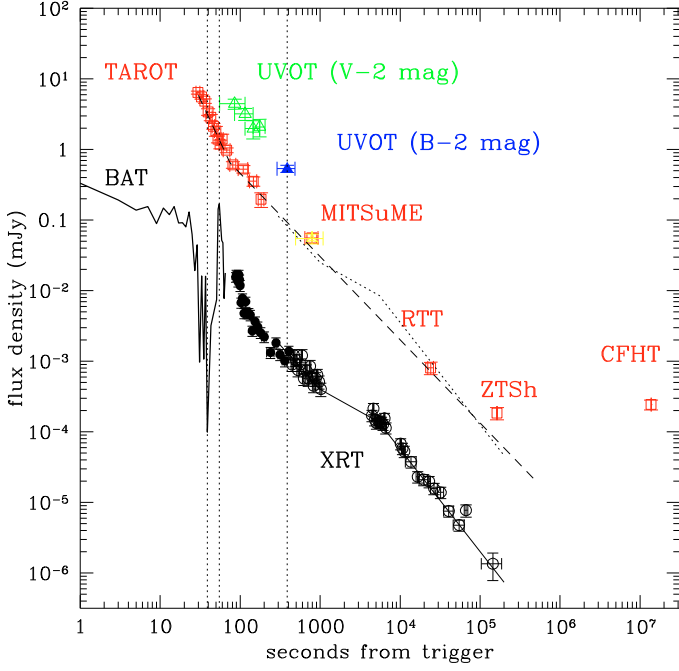
#### 3.1. Temporal evolution

The prompt emission of GRB 060111B was observed simultaneously in the optical band (*R*) and in the gamma-rays (15–150 keV), starting from 28 s up to 60 s ( $T_{90}$ ) after the trigger time. During this interval, the optical emission was found to decay steadily, in contrast with the gamma-ray emission which

exhibited flux variability of 2 orders of magnitude (Klotz et al. 2006).

The prompt optical emission from  $t_{\text{trig}} + 28$  s to  $t_{\text{trig}} + 200$  s does not track the X-ray (0.3–10 keV) emission either. The X-ray light curve, from  $\sim t_{\text{trig}} + 87$  s to  $\sim t_{\text{trig}} + 150$  s, shows a much steeper decay. However, starting from 200 s after the trigger the X-ray and optical decay indexes are consistent (see Table 4 and Fig. 5). Indeed, the TAROT data, the MITSuME *R*-band observations (taken 13 min after the trigger, Yanagisawa et al. 2006) and the RTT150 observations, show that the *R*-band light curve from 28 s to 7.2 h after the trigger can be modeled with a broken power law with best fit decay indexes  $\alpha_{R_1} = 2.4 \pm 0.2$  and  $\alpha_{R_2} = 1.18 \pm 0.05$  with a break at  $\sim 75$  s. The *V*-band decay index between  $\sim t_{\text{trig}} + 85$  and  $\sim t_{\text{trig}} + 200$  s is  $\alpha_V = 1.2 \pm 0.4$ , consistent with the *R*-band decay index. The ZTSh detection (taken 16.3 days after the trigger) slightly exceeds the power-law extrapolation, likely due to the superposition of the underlying host galaxy (see Sect. 4.1).

GRB 060111B represents one of the few cases where color information during the early afterglow is available. The similarity of the *V* and *R* band decay indexes (within uncertainties) between 85 s and 200 s after the trigger is indicative of achromatic

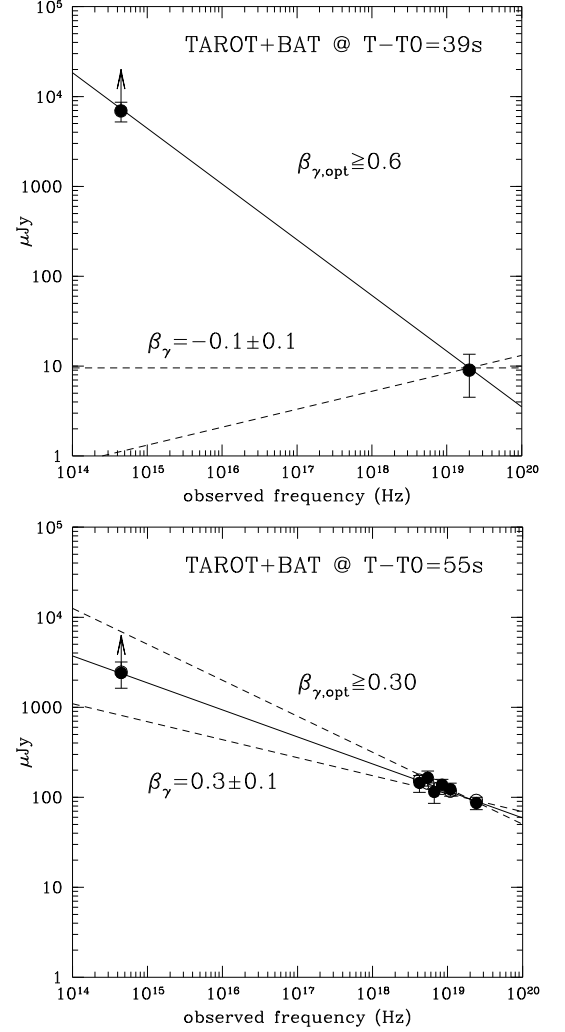


**Fig. 5.** Temporal evolution of the prompt and afterglow flux at several wavelengths. The *R*-band early data are clearly not correlated with the BAT prompt emission. The dashed curve is the *R*-band data best fit model while the solid curve is the X-rays best fit model. The dotted curve is the best fit X-ray model plotted over the *R*-band data. The three dotted vertical lines are the epochs at which a SED was extracted.

evolution of the early afterglow. However the large uncertainty affecting the *V*-band decay index  $\alpha_V$  and the proximity of *R* and *V* pass-bands prevent from more detailed investigation. The lack of a continuous coverage in the optical band after about  $t_{\text{trig}} + 1$  ks, makes possible a consistency of the *R*-band light curve at late times with the X-ray shallow-to-normal behavior with a temporal break at about 6 ks after the trigger.

### 3.2. Spectral energy distribution

In order to investigate on the origin of the prompt optical emission, we analyzed the spectral energy distribution (SED) from simultaneous BAT and TAROT observations. Two SEDs were extracted at 39 s after the trigger (from 28 s to 50 s, immediately before the second high energy prompt-emission peak) and at 55 s (from 50 s to 60 s, during the second peak). In the first interval, the available statistics from BAT data is too low to obtain spectral information and we could compute only the observed average 15–150 keV count rate, that is  $1.6 \times 10^{-3}$  counts  $\text{s}^{-1}$ . We converted this count rate into flux density assuming the  $T_{90}$ -averaged best fit spectral model (a simple power law with a photon index of  $\Gamma_\gamma = \beta_\gamma + 1 = 0.9 \pm 0.1$ , see Sect. 2.1). Extrapolating the power law model back to low energies, the TAROT flux density is about 3 orders of magnitude above the extrapolation of the gamma-ray flux density, with  $F_\gamma/F_{\text{opt}} \sim 10^{-3}$  and  $\beta_{\gamma,\text{opt}} \sim 0.6$  (Fig. 6). During the second peak of the burst (i.e. 55 s after the trigger), the available statistics enabled to extract a high energy spectrum best fitted by a simple power law with photon index of  $\Gamma_\gamma = \beta_\gamma + 1 = 1.3 \pm 0.1$ . Fixing this value within the uncertainties, we found that the *R*-band extrapolation is consistent with the observed flux (Fig. 6). The  $F_\gamma/F_{\text{opt}}$  is now  $\sim 0.1$  and  $\beta_{\gamma,\text{opt}} \sim \beta_\gamma = 0.3$ .



**Fig. 6.** Spectral Energy Distribution at  $t_{\text{trig}} + 39$  s and  $t_{\text{trig}} + 55$  s after the trigger where optical data have been corrected for Galactic extinction (Schlegel et al. 1998) but not for host galaxy extinction and should be considered as lower limits. The solid curve represents the optical-to-gamma-ray slope while the dashed lines individuate the best fit power law model from gamma-ray data only within  $1\sigma$  (at  $t_{\text{trig}} + 39$  s we have assumed the  $T_{90}$  integrated spectrum best fit model).

These results assume no additional extinction beyond that corrected for in our Galaxy. However, from X-ray spectral analysis there are hints of non-null absorption from the host galaxy. For example, assuming a Milky Way ISM, the rest frame  $A_V^{\text{host}}$  is expected to be  $N_H/1.8 \times 10^{21} \text{ mag cm}^{-2}$  (Predehl & Schmitt 1995), that for GRB 060111B is of about 1 mag or more since our  $N_H$  measure ( $N_H \sim 2 \times 10^{21} \text{ cm}^{-2}$  assuming  $z = 0$ , Sect. 2.2) should be considered as a lower limit given the unknown redshift. We thus attempted to estimate the host galaxy dust extinction affecting the TAROT data by exploiting our multiwavelength data set. We used Swift/XRT 0.3–10 keV and Swift/UVOT *B*-band and *V*-band data as well as the TAROT *R*-band and MITSuME *I*-band data. Since we have only one detection in the *B*-filter, the most sensitive wavelength to dust extinction, we decided to extract the SED at the time of the *B*-filter detection, that is at 388 s after the trigger. *I*-, *R*- and *V*-band data have been extrapolated at  $t_{\text{trig}} + 388$  s assuming the *R*-band decay slope at that time, that is  $1.18 \pm 0.05$ , and then corrected for Galactic extinction, that in each filter is  $A_B = 0.48$ ,  $A_V = 0.37$ ,

$A_R = 0.30$  and  $A_I = 0.22$  mag, using a standard galactic extinction curve.

The optical spectral index is unusually steep ( $\beta_{\text{opt}} \geq 2$ ). For comparison, other optical spectral indexes of GRBs with early optical steep-to-shallow light curve for which spectral data are available during the shallow phase are e.g.  $\beta_{\text{opt}} = 0.17$  for GRB 080319B (Woźniak et al. 2009) and  $\beta_{\text{opt}} = 0.95$  GRB 061126 (Perley et al. 2008). This lends additional support to our inference of significant extinction in the host galaxy.

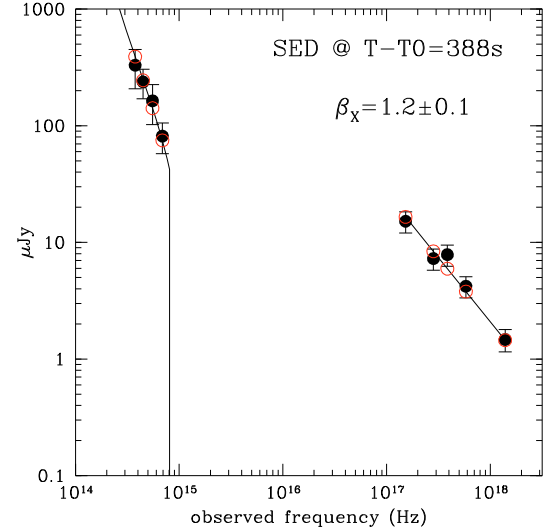
The consistency of the decay slopes in optical and X-rays at  $t_{\text{trig}} + 388$  s suggests a common cooling regime of the electron populations responsible for the observed emission in the two energy ranges (Sari et al. 1998). Specifically, the flux decay rate and the X-ray spectral index are consistent with the fireball expected closure relationship for an electron population in slow cooling regime and emitting at frequencies above the cooling frequency  $\nu > \nu_c$ . The expected spectral and decay indexes are  $p/2$  and  $(2 - 3p)/4$  respectively, where  $p$  is the electron distribution spectral index. We find a good agreement with the measured indexes for  $p \sim 2.3$ . The optical and X-ray SEDs, both extracted at  $t_{\text{trig}} + 388$  s, were therefore fitted simultaneously assuming a simple power law model with spectral index fixed (within  $1\sigma$ ) at the value obtained from X-ray data analysis only, that is  $\alpha_X = 1.2 \pm 0.1$  (e.g. interval B in Table 4). To account for dust extinction from the host galaxy, we tested three different extinction curves.

Assuming different redshifts, the best fit among Galactic (Cardelli et al. 1989), SMC (Pei 1992) and Starburst (Calzetti et al. 1994) extinction curves was obtained with the latter at  $z = 2$  with rest frame visual extinction  $A_V^r \sim 2$  mag (Fig. 7, Table 3). We have assumed a ratio of total to selective extinction of  $R_V = A_V/E(B - V) = 3.1$  for the Galactic and the Starburst extinction curves while  $R_V = 2.93$  for the SMC extinction curve (Pei 1992). At  $z < 2$  acceptable fits were obtained only with the Starburst extinction curve while at  $z \geq 2$  comparable good fits were obtained also assuming the Galactic extinction curve up to  $z = 2.5$ – $2.7$ , above which the Lyman  $\alpha$  cut-off due to intergalactic neutral hydrogen absorption would provide a  $B$ -band flux density more than  $2\sigma$  below the observed one.

At  $z = 2$ , the observed  $R$ -band corresponds to the UV band in the rest frame (central wavelength  $\sim 0.2 \mu\text{m}$ ), with best fit rest frame extinction of  $A_{\text{UV}} = 2.5 \pm 0.7$  mag. We note that at  $z = 1$  and  $z = 3$  the corresponding rest frame UV extinction is 2.7 and 1.7 mag, respectively. Indeed, if at  $z \geq 2$ , part of the  $B$ -band attenuation may be due to Lyman  $\alpha$  cut-off and not only to dust extinction. Despite the redshift uncertainties, these findings indicate that the observed optical flux is attenuated by host dust extinction. Correcting the observed  $R$ -band magnitude of  $\sim 2.5$  mag, that is the best fit value with a starburst extinction curve, the optical flux density increases by about a factor of 10. The above analysis of the gamma-to-optical SED would now provide  $F_\gamma/F_{\text{opt}} \sim 10^{-4}$  and  $\beta_{\gamma,\text{opt}} = 0.80 \pm 0.05$  at 39 s after the trigger and  $F_\gamma/F_{\text{opt}} \sim 0.01$  and  $\beta_{\gamma,\text{opt}} = 0.50 \pm 0.05$  at 55 s after the trigger.

These results make GRB 060111B consistent with the few GRBs with prompt optical detection for which the extrapolation of the prompt gamma-ray flux to the optical bands assuming a simple power law under-predicts the observed flux (cases “B” and “D” in Yost et al. 2007). Together with their different evolution in time, these findings suggest two separate origins for the prompt gamma-ray and the optical emission components.

In Klotz et al. (2006) a value of  $A_V^r \sim 4$  mag at  $z \sim 1$  was reported from our XRT and TAROT data at  $\sim t_{\text{trig}} + 140$  (assuming a simple underlying power-law spectral continuum and a Galactic



**Fig. 7.** Spectral Energy Distribution at  $t_{\text{trig}} + 388$  s where optical data have been corrected for Galactic extinction (Schlegel et al. 1998). The solid curve is the optical-to-X-ray best fit model obtained at  $z = 2$  assuming a Starburst extinction curve (see Sect 3.2) and the open red circles are the expected flux density values in each energy band.

**Table 3.** Best fit parameters obtained for the Spectral Energy Distribution extracted at  $t_{\text{trig}} + 388$  s assuming a simple power law model.

Redshift	Extinction Model	$A_V^r$ mag	$\chi^2/\nu$
0	Galactic	$5.2 \pm 0.2$	19.8/7
	SMC	$5.5 \pm 0.2$	20.7/7
	Starburst	$4.1 \pm 0.3$	5.2/7
1	Galactic	$2.2 \pm 0.2$	14.0/7
	SMC	$2.1 \pm 0.2$	15.2/7
	Starburst	$2.8 \pm 0.2$	3.5/7
2	Galactic	$1.5 \pm 0.1$	5.1/7
	SMC	$2.1 \pm 0.1$	15.2/7
	Starburst	$2.3 \pm 0.2$	2.8/7

extinction curve). This is significantly higher than our estimate at  $t_{\text{trig}} + 388$  s, possibly indicating an active dust destruction mechanism or the presence of a spectral break between the two bands at that time. In the latter case, the decaying flux behavior and the similarity of the decay rates observed at optical wavelengths and in X-rays from  $\sim 200$  s after the trigger suggest that the synchrotron cooling frequency crossed the rest frame UV band between  $t_{\text{trig}} + 140$  s and  $t_{\text{trig}} + 200$  s after the trigger.

## 4. The host galaxy of GRB 060111B

### 4.1. Imaging

Two deep exposures have been taken with the Canada-France-Hawaii Telescope (CFHT<sup>2</sup>) 5.3 months and 2.3 years after the burst with the  $R$  and  $H$  band filters, respectively. We detect a persistent source in the  $R$ -band, consistent with the afterglow position of GRB 060111B, with magnitude  $R = 24.8 \pm 0.4$ . The host

<sup>2</sup> Based on observations from director discretionary time obtained at the Canada-France-Hawaii Telescope (CFHT) which is operated by the National Research Council of Canada, the Institut National des Sciences de l’Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawaii.

**Table 4.** Spectral and temporal parameters at different epochs and energy ranges.

Epochs $T - T_{\text{trig}}$ (s)	Instrument	Energy range	Decay Index	Spectral Index
28–50	BAT	15–150 keV	–	–
50–60	BAT	15–150 keV	–	$0.3 \pm 0.2$
28–75	TAROT	<i>R</i> -band	$2.4 \pm 0.2$	–
75–26 000	TAROT+MITSUME+RTT	<i>R</i> -band	$1.18 \pm 0.05$	–
85–200	UVOT	<i>V</i> -band	$1.2 \pm 0.4$	–
A(87–150 s)	XRT/WT	0.3–10 keV	$3.8 \pm 0.5$	$1.0 \pm 0.2$
B(150–1000 s)	XRT/WT+PC	0.3–10 keV	$1.25 \pm 0.15$	$1.2 \pm 0.2$
C(4350–6800 s)	XRT/PC	0.3–10 keV	$0.6 \pm 0.2$	$1.3 \pm 0.2$
D(9930–185 650 s)	XRT/PC	0.3–10 keV	$1.4 \pm 0.1$	$1.1 \pm 0.1$

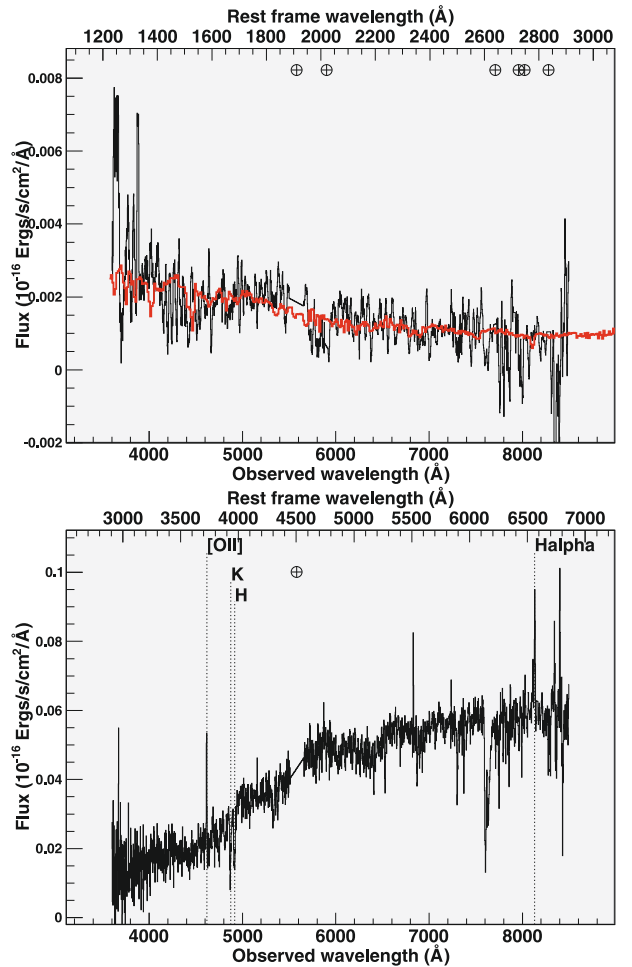
is 6.3'' off-axis from a nearby, bright spiral galaxy. From the positions derived from the relative astrometry performed on both the UVOT and CFHT images, we found that the afterglow and the host galaxy have an offset of less than 0.2 arcsec (Fig. 4). The GRB 060111B field was also observed in the *H*-band for an exposure time of 30 min with the WIRCAM wide-field NIR camera at the CFHT on May 18, 2008 (2.3 years after the burst). Data were preprocessed and calibrated by “I’iwi” pipeline at CFHT, and stacked with the SWARP software. We could not detect any source at the location of the burst down to the limiting magnitude of 23.0 (22.0) for a signal to noise ratio of 1.5 (3.0) (*AB* magnitude system).

From the CFHT image and considering that at the epoch of its exposure the afterglow has definitively faded away, it is clear that GRB 060111B occurred within a faint galaxy.

#### 4.2. Spectroscopy

A spectroscopic observation of the host galaxy has been performed by the LRIS instrument on the Keck I telescope on 2006 July 26 (UT). The 600/8500 grism and 600/4000 grating with the 560 dichroic were used in order to allow the covering of a wide wavelength domain (between 3800 and 8500 Å). The slit width was fixed to 1.5 arcsec and it was oriented in order to cover both the GRB host galaxy and the bright nearby galaxy (see Sect. 4.3). The total integration time was  $2 \times 2400$  s. Due to the position of the object in the sky, the airmass was relatively large during the observation, about 1.6. Data were reduced using the LRIS routines within the IRAF package. The two arms, Blue and Red, were reduced separately. In particular, the over-scan region was removed and the bias was subtracted by using the *lrsbias* task, which manages the dual amplifier (two-amp) readout mode. Since the two LRIS arms do not have exactly the same spectral resolution, the two spectra were therefore rebinned to the same wavelength resolution of 1 Å and finally combined together to produce the final spectrum. The junction region between the two arms, around 5800 Å, presents anomalous features that were cleaned “manually” (Fig. 8).

No clear emission or absorption lines are visible in the spectrum. No feature is compatible with a galaxy located close to, or at the same redshift than, the nearby large galaxy (see below). By applying a simple boxcar filter over 15 pixels, a weak continuum is observed, which is compatible with the one from a starburst spiral with  $0.39 < E(B-V) < 0.50$  (Kinney et al. 1996) at a redshift  $z \sim 1.9$ .



**Fig. 8.** Upper panel: optical spectrum of the host. The red curve is the template model. Bottom panel: optical spectrum of the nearby galaxy taken with the Keck Telescope.

#### 4.3. The bright nearby galaxy

A much brighter galaxy is present in our imaging as well, centered at an angular offset of only 4.5 arcsec from the GRB host galaxy. Astrometry based on CrAO *R*-band observations taken 16 days after burst and USNO-A2.0 reduction located the bright galaxy at coordinates (J2000) RA =  $19^{\text{h}}05^{\text{m}}43.06^{\text{s}}$ , Dec =  $+70^{\circ}22'29.7''$  with uncertainty of 0.2'' (Fig. 4). We performed first photometric and then spectroscopic observations of this galaxy in order to check whether any connection with the



**Table 5.** Photometry of the nearby galaxy (not corrected for Galactic extinction).

Filter	Magnitude	$\Delta\text{mag}$	Exposure (s)
<i>U</i>	23.34	0.3	1500
<i>B</i>	22.54	0.06	810
<i>V</i>	21.02	0.03	360
<i>R</i>	20.24	0.02	360
<i>I</i> <sub>int</sub> <sup>a</sup>	19.70	0.03	180
<i>J</i>	17.63	0.05	2880

<sup>a</sup> Filter *I*<sub>int</sub> is centered at  $\lambda = 7970 \text{ \AA}$  ( $FWHM = 157 \text{ \AA}$ ) #12 interference *I* filter more details in <http://www.not.iac.es/instruments/filters/filterlist.html>

host (e.g. a satellite galaxy) and/or the burst itself (e.g. a halo burst as if it originated from an old binary system) could exist. Photometric observations have been performed at the NOT telescope about four and eight months after the burst. *U*, *B*, *V*, *R*, *I* and *J* magnitude are quoted in Table 5. The best photometric redshift obtained with HyperZ code (Bolzonella et al. 2000) is  $z = 0.22^{+0.06}_{-0.07}$ . Its type is compatible with an early spiral galaxy, Sb. From *J* band images taken at NOT with WFCam, we find marginal evidence of two spiral arms. The angular scale at the measured redshift is  $3.76 \text{ kpc/arcsec}$ , making the size of this galaxy close to  $5.5'' \times 3.76 = 20.7 \text{ kpc}$ . Hereafter, we assumed a flat cosmology with  $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}$ ,  $\Omega_\Lambda = 0.73$ ,  $\Omega_m = 0.27$  (Wright 2006). From spectroscopic analysis performed of the Keck spectrum (see previous section), the galaxy displays several emission and absorption lines (Fig. 8), allowing redshift measurement of  $z = 0.239$  (1% accuracy).

#### 4.4. Limits on the SN component

No detection of any associated supernova (SN) down to  $R = 25.1$  was observed in the ZTSh telescope observations taken 16 days after the GRB (Table 2). Assuming as GRB-SNe templates the cases of SN1994I (GRB 021211, Della Valle et al. 2004) and SN1998bw (GRB 980425, Galama et al. 1998) and considering a host galaxy dust extinction in the *V* band of about  $\sim 1 \text{ mag}$  the lack of any SN detection in our images implies a redshift lower limit of about 0.4 for a SN1994I-like event and of about 0.7 for a SN1998bw-like one.

#### 4.5. Redshift estimate

In this section we attempt to estimate the distance of GRB 060111B following some considerations based on our multi-band observations of both the host and the afterglow.

The host galaxy spectrum shows a weak continuum that is compatible with a starburst spiral with  $0.39 < E(B - V) < 0.50$  (Kinney et al. 1996) at a redshift  $\sim 1.9$ . Unfortunately, no emission or absorption lines are present (Fig. 8). The two most common lines observed in long GRB host galaxies are the [OII] line and the  $H\alpha$ . Although the lack of these lines cannot give an absolute restriction to the redshift range, the lack of [OII] ( $\lambda 3727 \text{ \AA}$ ) emission line may imply  $z > 1.5$  while the lack of  $H\alpha$  ( $\lambda 6563 \text{ \AA}$ ) emission line may imply  $z > 0.4$ .

The best-fit redshift from the afterglow SED was  $z = 2$ , consistent with the result obtained from fitting the host spectral continuum. We also used the X-ray light curve to estimate the

**Table 6.** Host galaxy and burst properties at different redshifts.

$z$	$d^a$ kpc	$M_B^b$ mag	$E_{\text{iso}}$ $10^{52} \text{ erg}$	$E_{\text{peak},i}$ MeV	$\theta_j^c$ deg
0.24	3.76	-15.4	$0.07 \pm 0.01$	$0.6 \pm 0.3$	4.5
1	8.0	-18.4	$1.5 \pm 0.2$	$0.9 \pm 0.5$	2.6
2	8.5	-20.3	$5.8 \pm 0.8$	$1.4 \pm 0.8$	1.9
3	7.8	-21.5	$12 \pm 2$	$1.8 \pm 1.0$	1.5

<sup>a</sup> Host size. <sup>b</sup> Absolute magnitude of the host. <sup>c</sup> Jet opening angle assuming that the temporal break observed in the X-ray light curve at a.  $\sim 6 \text{ ks}$  has a jet origin.

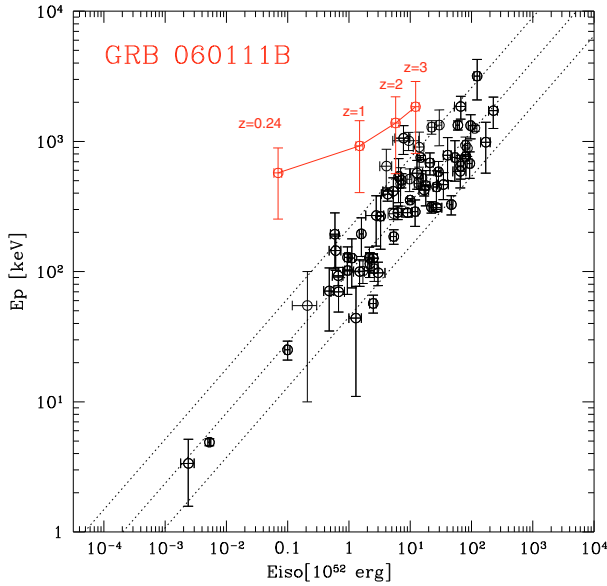
distance to this burst: following Gendre et al. (2008a) we estimate  $z = 2.2 \pm 1.0$ , consistent with the above results. Finally, comparing the host galaxy apparent magnitude of  $R \sim 25$  with the typical magnitudes observed for a sample of long GRBs hosts, we find that it is likely at  $z \sim 1-2$  (cf. for instance Fig. 9 of Berger et al. 2007), again consistent with previous findings.

We further investigate on the distance scale of this burst by comparing the energetics and the host galaxy properties obtained assuming different redshifts, with the typical values observed for long GRBs. In Table 6 we have computed the size of the host (assuming an angular size of  $1''$ , corresponding to the seeing of CFHT observation) and its absolute magnitude and the burst energetics at several different redshifts. The burst energetics was estimated with the isotropic equivalent energy  $E_{\text{iso}}$  that we have evaluated through the observed 100–1000 keV fluence  $S = (5.6 \pm 0.8) \times 10^{-6} \text{ erg cm}^{-2}$  (see Sect. 2.1) as  $S \times 4\pi D_L^2 / (1+z)$  where  $D_L$  is the luminosity distance. The absolute magnitudes have been computed from the observed *R* magnitude, considering that for  $z \geq 0.5$  it samples the rest frame *B* filter magnitude while for  $z \leq 1$  we assumed  $U - B \sim 0.75 \text{ mag}$  as typical  $U - B$  color of starburst galaxies (e.g. Coil et al. 2007; ?).

Long GRB host luminosities cover a wide range of values, most of them between  $M_B = -17$  and  $M_B = -22 \text{ mag}$ , with a peak at about  $M_B \sim -18 \text{ mag}$  (Berger et al. 2007) while the typical host sizes is  $\sim 1.7 \text{ kpc}$  (Wainwright et al. 2007). Considering that the host angular sizes quoted in Table 6 should be considered as an upper limit due to the seeing at CFHT at the epoch of the observations, the host size of GRB 060111B is consistent with typical long GRBs host sizes at all the tested redshifts while its absolute magnitudes would be too faint at the redshift of the nearby galaxy. From burst energetics considerations, Table 6 shows that at  $z = 0.24$ , GRB 060111B would be an outlier in the Amati  $E_{\text{peak},i}$  versus  $E_{\text{iso}}$  correlation (Fig. 9). We note however that even at higher redshift GRB 060111B seems to deviate from the correlation<sup>3</sup>. Other long GRBs detected by BATSE are inconsistent with the Amati correlation (Nakar & Piran 2005; Band & Preece 2005) and GRB 060111B may fall into this category (Fig. 9). However, the large uncertainties affecting the peak energy measure prevented us to reach any firm conclusion.

All of these considerations argue against any relation with the nearby, bright galaxy. Consistent with this conclusion is the lack of any supernova detection that imposes a redshift larger than  $z \sim 0.4-0.7$ . For the latter, we assumed an observed magnitude of  $R \geq 25.1 \text{ mag}$  for the SN; that is, the same brightness as the source detected in the ZTSh telescope images taken 16 days after the burst (see Sect. 4.4).

<sup>3</sup> This may be the reason why the pseudo-redshift (Atteia 2004) computations provides the uncomfortably high value of  $pz \sim 8.4$ .



**Fig. 9.** GRB 060111B is plotted (in red) in the  $E_{\text{peak}}$  vs.  $E_{\text{iso}}$  plot at four different redshifts:  $z = 0.24, 1, 2$  and  $3$ , the latter being the maximum redshift compatible with our observations (see Sect. 5.2). Other GRBs data and the best fit correlation within  $2\sigma$  are taken from Amati et al. (2008).

An upper limit on the redshift can be obtained from the lack of evidence of Lyman alpha cut-off due to the absorption from neutral hydrogen in the intergalactic medium in the spectral continuum of the host galaxy and by the  $B$ -band detection, providing  $z < 3$ .

Summarizing the above considerations, there is evidence against a possible relation of GRB 060111B with the nearby, bright galaxy and independent estimates converge to  $z \sim 1-2$ . Having GRB 060111B at this redshift would result in having a standard long GRB host galaxy.

#### 4.6. Outflow collimation

Using our redshift estimates above, we can compute the possible collimation of the outflow  $\theta_j$  assuming a jet origin of the temporal break  $t_b$  observed in the XRT light curve at about 6 ks. The angle estimations computed assuming a constant ISM density of  $n \sim 5 \text{ cm}^{-3}$  and radiative efficiency  $\sim 0.2$  are presented in Table 5. It occurs that the jet opening angle is within the distribution of angles of sample GRBs with well defined jet break in X-ray light curve (Racusin et al. 2008a). Alternatively, if the jet break epoch is beyond the end of XRT observations, that is  $t_b > 200 \text{ ks}$ , then  $\theta_j > 6 \text{ deg}$  ( $z \leq 2$ ) which is marginally consistent with the distribution. Thus we can support the break at  $t_b \sim 6 \text{ ks}$  as a real jet break. This result is in line with the anticorrelation found between the X-ray afterglow shallow phase durations and the redshift (Stratta et al. 2009), that, if interpreted as a selection effect, may reflects an anticorrelation of the X-ray shallow phase duration with the burst energy. Indeed, the latter is expected if the end of the shallow phase originates from a collimated outflow. The burst energy corrected for collimation at  $z = 1-2$  would be  $E_\gamma \sim 1 - 2 \times 10^{49} \text{ erg}$ .

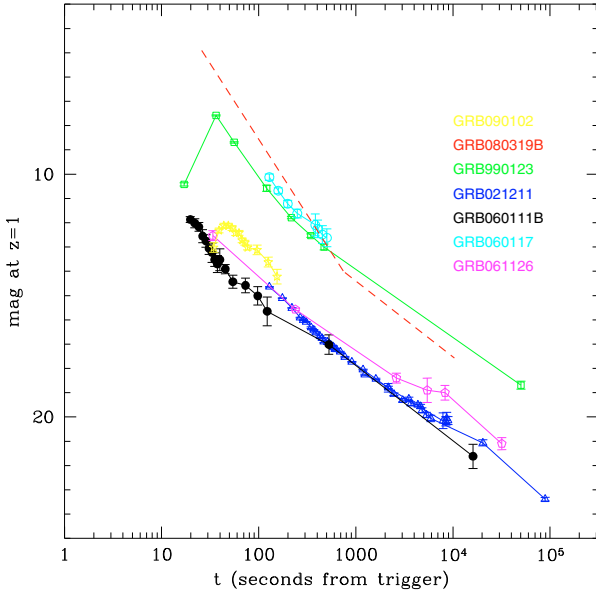
### 5. The nature of the prompt optical emission

The nature of the prompt optical emission of GRB 060111B has been already investigated in past works (e.g. Klotz et al. 2006;

Wei 2007). A reverse shock (RS, e.g. Meszaros & Rees 1994, 1997) emission has been invoked for this burst by Klotz et al. (2006), primarily based on the TAROT observations of very steep early-time optical decay ( $\alpha = -(2.4 \pm 0.1)$  in  $R$ -band, see Sect. 3.1). Further evidence includes the observed lack of temporal correlation between the simultaneous BAT and TAROT emissions and the similarities of its optical light curve with GRB 990123 and GRB 021211 (for which a reverse shock origin has been proposed to explain their optical light curves). A decay index of  $\sim -2$  has generally been considered as the evidence of a RS origin. Some concerns however have been addressed to the unknown RS zero-time that, if different than the trigger time, would make the real decay rate of the emission different from those measured, leading to different conclusions (Panaitescu 2009). A simultaneous spectral and temporal analysis would be the best way to investigate the nature of the observed optical radiation. Unfortunately, for GRB 060111B we do not have any spectral information in the optical regime during the steep decay. In this work we further investigate on the nature of the prompt optical emission by exploring its optical-to-gamma ray spectral energy distribution. We also compare GRB 060111B with other GRBs showing early optical emission with similar temporal properties. We then attempt to constrain the physical parameters of this burst using our combined early-through-late-time observations, including our constraints on the redshift and on the host-galaxy dust extinction.

Thanks to our multi-band analysis we were able to confirm the presence of non-negligible extinction from dust in the burst environment. By correcting the TAROT data for the estimated amount of host galaxy dust extinction, we could evaluate the intrinsic optical-to-gamma-ray flux density ratio during the prompt emission. The obtained values are orders of magnitude above unity, quite unlike most GRBs with simultaneous optical counterpart which have been reported to date (see Yost et al. 2007). Most importantly, the optical flux density is well above the extrapolation obtained from the best fit spectral model in the gamma-ray regime. This result, as well as the smooth, fast decay of the optical light curve simultaneous to the strong gamma-ray flux variability, enabled us to firmly exclude a common emission mechanism at high and low energies. To compare GRB 060111B with other GRBs with “steep-to-shallow” decay observed during early optical emission, we plotted the optical magnitudes as a function of time as if all bursts were at  $z = 1$  (Fig. 10). To compute the expected magnitudes at  $z = 1$  we simply apply the distance modulus definition without taking into account any  $k$ -correction and the host galaxy extinction correction since for almost all the plotted bursts we have no color information during the early optical emission and very uncertain or no host galaxy extinction measurements. For GRB 080319B, given the huge amount of optical data available, we decided to plot only the well constrained best fit light curve model (Racusin et al. 2008a) in the period of interest for our study (that is, when the flux features the steep-to-shallow decay). Inspection of the figure demonstrates the strong similarities in the initial steep decay indexes (all within  $\sim 2$  and  $\sim 2.5$ , see Table 7). This evidence possibly indicates the presence of a common starting time of the observed prompt optical emissions that, if coincident with the trigger time, provides good consistency with the RS predicted decay slopes.

In the context of the internal-external shock model (e.g. Sari & Piran 1997), the RS emission mechanism is a consequence of the external shock that the outflow produces with the surrounding interstellar matter, along with a forward shock (FS) that generates the afterglow emission. As for the FS, the RS



**Fig. 10.** Early optical emission light curves as if all bursts were at  $z = 1$  of GRB 060111B (filled circles), GRB 990123 (open squares, Akerlof et al. 1999), GRB 021211 (open triangles, Fox et al. 2003; Li et al. 2003), GRB 060117 (open circles, Jelínek et al. 2006), GRB 061126 (open pentagons, Perley et al. 2008), GRB 080319B (dashed line, best fit from Racusin et al. 2008b) and GRB 090102 (cyan pentagons, Klotz et al. 2009). For these GRBs the Reverse Shock model has been claimed to explain the prompt optical afterglow (see Table 7).

accelerates electrons in the outflow into a power law energy distribution with index  $p$  (e.g. Sari et al. 1998; Sari & Piran 1999). The light curve at a given frequency depends on the temporal evolution of the injection frequency  $\nu_m$ , the cooling frequency  $\nu_c$  and the synchrotron peak flux  $F_{\nu, \max}$  (e.g. Sari et al. 1998). Kobayashi (2000) derived a simple analytic solution of the RS light curve in two limits: the thick shell limit, when the RS becomes relativistic crossing the outflow and the thin shell limit, when the RS remains Newtonian. From an observational point of view the thick shell limit predicts that the RS peaks during the high-energy burst while a RS peaking after the end of the burst is predicted from the thin shell limit case. The combinations of the RS and FS light curves have been classified by Zhang et al. (2003) as of type I, II and III depending on RS and FS relative peak times and peak flux intensities. Klotz et al. (2006) pointed out that the optical prompt emission of GRB 060111B can be classified as type II (Zhang et al. 2003) where the RS emission is dominating over the peak of the FS component after the FS peak time.

In order to compare observations with theoretical predictions, we first attempt to constrain the fireball initial Lorentz factor  $\Gamma$  from the time at which we observe the prompt optical decay (Sari 1997; Zhang et al. 2003). Since the optical light curve is already fading at  $t = 28$  s after the trigger, the rest frame deceleration time is at  $t_d \leq t/(1+z)$ ; that is, the fireball has given the ISM an amount of energy comparable to its initial energy at  $t_d = [3E/32\pi n m_p c^5 \Gamma^8]^{1/3} \leq t/(1+z)$ . Assuming an interstellar matter density of  $n \sim 5 \text{ cm}^{-3}$  and a redshift of  $z = 1-2$  and estimating the energy of the shell  $E$  with  $E_{\text{iso}}$ , we find  $\Gamma \geq 260-360$ . This result is consistent with past estimates of the fireball Lorentz factor of about a few hundreds (e.g. Molinari et al. 2007; Gomboc et al. 2008; Ferrero et al. 2008).

As already mentioned by Klotz et al. (2006) GRB 060111B falls in the thick shell limit since the optical prompt emission is already fading before the end of the prompt, high-energy burst. In this limit, the deceleration time should coincide with the time at which the RS crosses the shell (e.g. Kobayashi 2000; Zhang et al. 2003) that is estimated with the high-energy burst duration. We use  $T_{90} \sim 25$  s observed in the Suzaku/WAM energy range (100–1000 keV) to estimate the burst duration. The injection frequency, the cooling frequency and the peak power at the shock crossing time for the thick shell limit are (Sari & Piran 1999; Kobayashi 2000):

$$\nu_m \sim 7.3 \times 10^{14} \epsilon_{e,0.6}^2 \epsilon_{B,0.01}^{1/2} n_5^{1/2} \Gamma_{300}^2 \text{ Hz} \quad (1)$$

$$\nu_c \sim 4.5 \times 10^{15} \epsilon_{B,0.01}^{-3/2} E_{52}^{-1/2} n_5^{-1} (T_{8.3})^{-1/2} \text{ Hz} \quad (2)$$

$$F_{\nu, \max} \sim 130 D_{L,28}^{-2} \epsilon_{B,0.01}^{1/2} E_{52}^{5/4} n_5^{1/4} \Gamma_{300}^{-1} (T_{8.3})^{-3/4} \text{ mJy} \quad (3)$$

where  $D_{L,28} = D_L/5 \times 10^{28} \text{ cm}$  and  $E_{52} = E_{\text{iso}}/5.8 \times 10^{52} \text{ erg}$  are the luminosity distance and the outflow energy computed at  $z = 2$  (Table 6),  $\epsilon_{B,0.01} = \epsilon_B/0.01$  and  $\epsilon_{e,0.6} = \epsilon_e/0.6$  are the fractions of the burst total energy that goes into the electrons and magnetic field respectively,  $n_5 = n/5 \text{ cm}^{-3}$  is the interstellar matter density,  $T_{8.3} = (T_{90}/(1+z))/8.3 \text{ s}$ , and  $\Gamma_{300} = \Gamma/300$  is the outflow Lorentz factor. Assuming the specified values, from Eqs. (1) and (2) we find that at the RS crossing time the observed  $R$ -band frequency ( $\nu \sim 1.2 \times 10^{15} \text{ Hz}$  in the rest frame), is between  $\nu_m$  ( $7.3 \times 10^{14} \text{ Hz}$ ) and  $\nu_c$  ( $4.5 \times 10^{15} \text{ Hz}$ ). Equation (3) shows that the RS peak flux density is expected at  $\sim 130 \text{ mJy}$  for  $z = 2$ . From Fig. 5 we see that at the beginning of the prompt optical decay, the flux density is  $\sim 10 \text{ mJy}$ . This value however is not corrected for the dust extinction in the host. The latter is highly uncertain but we find evidence that is not negligible (see Sect. 3.2). From our optical-to-X-ray best-fit obtained assuming a starburst extinction curve, we measure a rest frame UV extinction of about 2.5 mag that makes the corrected measured flux density consistent with the expected peak flux. This result indicates that the RS peak time should not be too far from the TAROT start observing time and that the outflow Lorentz factor is close to the estimated lower limit value.

From (1) and (2) we obtain that  $\nu_m < \nu_c$  at the RS crossing time. Therefore, for the TAROT observations we can assume a slow cooling regime. For  $\nu_m < \nu < \nu_c$  as observed,  $F_\nu = F_{\nu_m} (\nu/\nu_m)^{-(p-1)/2} \propto t^{-\alpha}$ . Since  $\nu_m \propto t^{-73/48} \sim t^{-3/2}$  and  $F_{\nu_m} \propto t^{-47/48} \sim t^{-1}$  (Kobayashi 2000), we expect  $\alpha = (3p+1)/4$  that is consistent with our measured decay index if  $p = 2.9 \pm 0.3$ . We note that this value is somewhat higher than what derived from the FS component (see Sect. 3.2). Different values of  $p$  for the FS and the RS have been invoked as a possible scenario also for GRB 061126 by Perley et al. (2008).

It has been argued that the RS emission mechanism cannot be the dominating mechanism to explain the bright optical early emission due to the lack of this component in many other GRBs despite the expected high flux (Romíng et al. 2005). We note that the prompt optical emission in GRB 060111B ends very early compared to the other events (Fig. 10), opening the possibility that this component may remain undetected in a significant fraction of GRBs. In addition, the magnetization of GRBs outflows plays a crucial role in suppressing the RS optical emission (e.g. Fan 2008).

In a scenario alternative to the RS, the prompt optical emission may be produced by internal shocks different from those responsible for the gamma-ray emission, which can be formed by collisions of different shells: these shells may be easily disrupted



**Table 7.** Best-fit decay indexes of the early optical emission for all GRBs with well monitored optical counterpart that shows a steep-to-shallow behavior (Fig. 10).

GRB	Filter	$\alpha_1$	$\alpha_2$	$t_{\text{break}}$ (s)	$z$
990123 <sup>(1)</sup>	V	$\sim 2.5$	$\sim 1.5$	700	1.60
021211 <sup>(2)</sup>	R	$1.97 \pm 0.05$	$1.01 \pm 0.02$	300	1.0
060117 <sup>(3)</sup>	R	$2.49 \pm 0.05$	$1.47 \pm 0.03$	300	1 <sup>a</sup>
061126 <sup>(4)</sup>	R	$2.0 \pm 0.3$	$0.76 \pm 0.06$	780	1.1588
060111B	R	$2.4 \pm 0.2$	$1.18 \pm 0.05$	75	2
080319B <sup>(5)</sup>	R	$2.49 \pm 0.09$	$1.25 \pm 0.02$	900	0.937
090102 <sup>(6)</sup>	R	$1.83 \pm 0.13$	$1.1 \pm 0.3$	400	1.547

<sup>a</sup> For GRB 060117 since the redshift is not known, we imposed  $z = 1$ .

References. (1) Akerlof et al. (1999); (2) Holland et al. (2004); (3) Jelínek et al. (2006); (4) Perley et al. (2008); (5) Racusin et al. (2008a); (6) Klotz et al. (2009).

by other shells, thus explaining why bright optical emission is not common in GRBs (Wei 2007). For GRB 060111B, Wei (2007) computed that the rest frame cooling frequency should be at energies  $\nu_c \sim 0.3\text{--}0.7$  keV, that is between the X-ray and the optical energy domains, at  $z = 1\text{--}2$ , 28 s after the trigger and assuming a Lorentz factor of the merged shell of 800. The  $\nu_c$  expected value at the shell crossing time is consistent with our observations assuming that  $\nu_c$  remains at energies higher than the optical up to  $t_{\text{trig}} + 200$  s after which the decay indexes in the two energy domains are consistent among themselves (see Sect. 3.1). However, the  $\nu_c$  theoretical estimation depends on the highly uncertain fireball Lorentz factor to the power of eight. If for example we assume  $\Gamma = 300$ , the cooling frequency would had been well below the R-band at  $t_{\text{trig}} + 28$  s. For this reason and for the numerous pieces of evidence supporting the RS emission exposed above, we do not favor the internal shock interpretation.

Finally we mention that the RS emission mechanism has been proposed to solve the difficulties of the “standard” forward shock model to explain the canonical (steep-flat-normal) evolution of the X-ray emission (Uhm & Beloborodov 2007; Genet et al. 2007a). GRB 060111B represents one of the few cases where both an optical flash and a “canonical” X-ray light curve (Nousek et al. 2006) are observed, enabling to test the RS in both the energy ranges. For example, to reproduce the observed X-ray light curves a long-lived RS emission is required, which is difficult to reconcile with bright optical early emission (Genet et al. 2007a).

## 6. Summary and conclusions

The presence of a bright optical emission during the prompt phase makes GRB 060111B an extremely interesting burst and for this reason an extended multi-wavelength campaign has been performed in order to provide further constraints to the physical origin of the observed optical emission. Hereafter we summarize our main results and conclusions.

Our late optical observations enabled us to identify the host galaxy of GRB 060111B. The featureless host spectrum, the non detection of any associated supernova 16 days after the burst as well as the inferred host luminosities at different redshift, allowed us to put some constraints on the distance scale of GRB 060111B. Independent estimates converge to  $z \sim 1\text{--}2$ .

From the afterglow SED we find that a non-negligible host galaxy dust extinction affects the observed flux in the optical

regime. After correcting for this effect, the optical-to-gamma-ray flux ratio during the prompt strongly suggests a separate origin of the optical component from that one at high energies. This was already pointed out by (Klotz et al. 2006) from the observed high energy prompt light curve not tracked by the simultaneous optical one. The similarities of the measured optical decay index of GRB 060111B with other observed prompt optical emissions suggest a common zero-time origin of the emission mechanism that, if coincident with the burst trigger time, makes the measured decay indexes consistent with reverse shock scenario.

From the time at which we observe the TAROT light curve starting to decay, we estimated the fireball initial Lorentz factor (Sari 1997) as  $\Gamma \geq 260\text{--}360$  at redshift  $z \sim 1\text{--}2$ , in agreement with past results obtained for other bursts. Assuming  $\Gamma \sim 300$ , the observations are in good agreement with theoretical predictions from the reverse shock emission in the thick shell limit. In particular, we find evidence that the prompt optical emission is in the slow cooling regime with  $\nu_m < \nu < \nu_c$  and that the peak flux is very near or coincident with the observed one corrected for host galaxy extinction.

The very short duration of optical RS emission in GRB 060111B (about 30 s in the restframe) raises the question of the true fraction of GRBs displaying this type of emission. GRB 060111B is among the few cases of a burst with bright early optical counterpart observed with good temporal resolution as well as a “canonical” Swift/XRT light curve in X-rays and provides an excellent test case for the reverse-shock emission mechanism in both the energy ranges.

*Acknowledgements.* The authors thank the anonymous referee for his/her useful comments and acknowledge Taka Sakamoto and Scott D. Barthelmy for making BAT data available to the community, and J. S. Bloom for obtaining the Keck data. B.G. is funded by a french CNES post-doctoral grant.

## References

- Akerlof, C., Balsano, R., Barthelmy, S., et al. 1999, *Nature*, 398, 400
- Akerlof, C. W., Kehoe, R. L., McKay, T. A., et al. 2003, *PASP*, 115, 132
- Amati, L., Guidorzi, C., Frontera, F., et al. 2008, *MNRAS*, 391, 577
- Band, D. L., & Preece, R. D. 2005, *ApJ*, 627, 319
- Berger, E. 2009, *ApJ*, 690, 231
- Berger, E., Fox, D. B., Price, P. A., et al. 2007, *ApJ*, 664, 1000
- Blake, C. H., Bloom, J. S., Starr, D. L., et al. 2005, *Nature*, 435, 181
- Bloom, J. S., Starr, D. L., Blake, C. H., Skrutskie, M. F., & Falco, E. E. 2006, in *Astronomical Data Analysis Software and Systems XV*, ed. C. Gabriel, C. Arviset, D. Ponz, & S. Enrique, *ASP Conf. Ser.*, 351, 751
- Boër, M. 2001, *Astron. Nachr.*, 322, 343
- Bolzonella, M., Miralles, J.-M., & Pelló, R. 2000, *A&A*, 363, 476
- Burrows, D. N., Hill, J. E., Nousek, J. A., et al. 2005, *Space Sci. Rev.*, 120, 165
- Calzetti, D., Kinney, A. L., & Storchi-Bergmann, T. 1994, *ApJ*, 429, 582
- Capalbi, M., Perri, M., Saija, B., Tamburelli, F., & Angelini, L. 2005, *The SWIFT XRT Data Reduction Guide*, <http://swift.gsfc.nasa.gov/docs/swift/analysis/>
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, 345, 245
- Coil, A. L., Hennawi, J. F., Newman, J. A., Cooper, M. C., & Davis, M. 2007, *ApJ*, 654, 115
- Covino, S., Stefanon, M., Sciuto, G., et al. 2004, in *SPIE Conf. Ser.* 5492, ed. A. F. M. Moorwood, & M. Iye, 1613
- Cwiok, M., Dominik, W., Kasprzowicz, G., et al. 2005, in *Astrophysical Sources of High Energy Particles and Radiation*, ed. T. Bulik, B. Rudak, & G. Madejski, *AIP Conf. Ser.*, 801, 143
- Della Valle, M., Malesani, D., Benetti, S., et al. 2004, in *AIP Conf. Ser.* 727, of *Discovery*, ed. E. Fenimore, & M. Galassi, 403
- Evans, P. A., Beardmore, A. P., Page, K. L., et al. 2008, *ArXiv e-prints*
- Fan, Y.-Z. 2008, in *AIP Conf. Ser.*, 1065, 159
- Ferrero, P., Klose, S., Kann, D. A., et al. 2008, *ArXiv e-prints*
- Fitzpatrick, E. L. 1999, *PASP*, 111, 63
- Fox, D. W., Price, P. A., Soderberg, A. M., et al. 2003, *ApJ*, 586, L5
- Fukazawa, Y., Ohno, M., Takahashi, T., et al. 2007, in *AIP Conf. Ser.* 921, ed. S. Ritz, P. Michelson, & C. A. Meegan, 106



- Galama, T. J., Vreeswijk, P. M., van Paradijs, J., et al. 1998, *Nature*, 395, 670
- Gendre, B., Galli, A., & Boër, M. 2008a, *ApJ*, 683, 620
- Gendre, B., Klotz, A., Stratta, G., et al. 2008b, in SF2A-2008: Proceedings of the Annual meeting of the French Society of Astronomy and Astrophysics, available online at <http://proc.sf2a.asso.fr>, ed. C. Charbonnel, F. Combes, & R. Samadi, 223
- Genet, F., Daigne, F., & Mochkovitch, R. 2007a, *MNRAS*, 381, 732
- Genet, F., Daigne, F., & Mochkovitch, R. 2007b, *A&A*, 471, 1
- Gomboc, A., Kobayashi, S., Guidorzi, C., et al. 2008, *ApJ*, 687, 443
- Holland, S. T. 2006, *GRB Coordinates Network*, 5820, 1
- Holland, S. T., Bersier, D., Bloom, J. S., et al. 2004, *AJ*, 128, 1955
- Jelínek, M., Prouza, M., Kubánek, P., et al. 2006, *A&A*, 454, L119
- Kalberla, P. M. W., Burton, W. B., Hartmann, D., et al. 2005, *VizieR Online Data Catalog*, 8076, 0
- Khamitov, I., Uluc, K., Aslan, Z., et al. 2006, *GRB Coordinates Network*, 4494, 1
- Kinney, A. L., Calzetti, D., Bohlin, R. C., et al. 1996, *ApJ*, 467, 38
- Klotz, A., Gendre, B., Boer, M., & Atteia, J. L. 2009, *GRB Coordinates Network*, 8764, 1
- Klotz, A., Gendre, B., Stratta, G., et al. 2006, *A&A*, 451, L39
- Kobayashi, S. 2000, *ApJ*, 545, 807
- Li, W., Filippenko, A. V., Chornock, R., & Jha, S. 2003, *ApJ*, 586, L9
- Liang, E.-W., Zhang, B.-B., & Zhang, B. 2007, *ApJ*, 670, 565
- Melandri, A., Mundell, C. G., Kobayashi, S., et al. 2008, *ApJ*, 686, 1209
- Molinari, E., Vergani, S. D., Malesani, D., et al. 2007, *A&A*, 469, L13
- Moretti, A., Campana, S., Mineo, T., et al. 2005, in *SPIE Conf. Ser.* 5898, ed. O. H. W. Siegmund, 348
- Nakar, E., & Piran, T. 2005, *MNRAS*, 360, L73
- Nousek, J. A., Kouveliotou, C., Grupe, D., et al. 2006, *ApJ*, 642, 389
- Panaitescu, A. 2009, in *AIP Conf. Ser.* 1133, ed. C. Meegan, C. Kouveliotou, & N. Gehrels, 127
- Pei, Y. C. 1992, *ApJ*, 395, 130
- Pérez-Ramírez, D., Park, H. S., Williams, G. G., on behalf of the Superlotus Team 2004, *Astron. Nachr.*, 325, 667
- Perley, D. A., Bloom, J. S., Butler, N. R., et al. 2008, *ApJ*, 672, 449
- Perri, M., Barthelmy, S., Boyd, P., et al. 2006a, *GRB Coordinates Network*, 4487, 1
- Perri, M., Moretti, A., Burrows, D. N., Fink, R., & Ajello, M. 2006b, *GRB Coordinates Network*, 4498, 1
- Pozanenko, A., Beskin, G., Bondar, S., et al. 2004, in *Gamma-Ray Bursts: 30 Years of Discovery*, ed. E. Fenimore, & M. Galassi, American Institute of Physics Conf. Ser., 727, 757
- Predehl, P., & Schmitt, J. H. M. M. 1995, *A&A*, 293, 889
- Racusin, J. L., Karpov, S. V., Sokolowski, M., et al. 2008a, *Nature*, 455, 183
- Racusin, J. L., Liang, E. W., Burrows, D. N., et al. 2008b, *ArXiv e-prints*
- Roming, P. W. A., Kennedy, T. E., Mason, K. O., et al. 2005, *Space Sci. Rev.*, 120, 95
- Roming, P. W. A., Koch, T. S., Oates, S. R., et al. 2009, *ApJ*, 690, 163
- Rykoff, E. S., Aharonian, F., Akerlof, C. W., et al. 2009, *ApJ*, submitted, [arXiv:0904.0261]
- Sari, R. 1997, *ApJ*, 489, L37
- Sari, R., & Piran, T. 1997, *MNRAS*, 287, 110
- Sari, R., & Piran, T. 1999, *ApJ*, 517, L109
- Sari, R., Piran, T., & Narayan, R. 1998, *ApJ*, 497, L17
- Sato, G., Nakazawa, K., Takahashi, T., et al. 2006, *GRB Coordinates Network*, 4571, 1
- Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525
- Stratta, G., Guetta, D., D'Elia, V., et al. 2009, *A&A*, 494, L9
- Tamagawa, T., Usui, F., Urata, Y., et al. 2005, *Nuovo Cimento C Geophysics Space Physics C*, 28, 771
- Tueller, J., Barbier, L., Barthelmy, S., et al. 2006, *GRB Coordinates Network*, 4492, 1
- Uhm, Z. L., & Beloborodov, A. M. 2007, *ApJ*, 665, L93
- Vestrand, W. T., Borozdin, K. N., Brumby, S. P., et al. 2002, in *SPIE Conf. Ser.* 4845, ed. R. I. Kibrick, 126
- Wainwright, C., Berger, E., & Penprase, B. E. 2007, *ApJ*, 657, 367
- Wei, D. M. 2007, *MNRAS*, 374, 525
- Woźniak, P. R., Vestrand, W. T., Panaitescu, A. D., et al. 2009, *ApJ*, 691, 495
- Yanagisawa, K., Toda, H., & Kawai, N. 2006, *GRB Coordinates Network*, 4496, 1
- Yost, S. A., Yuan, F., Swan, H., & Akerlof, C. 2006, *GRB Coordinates Network*, 4488, 1
- Yost, S. A., Aharonian, F., Akerlof, C. W., et al. 2007, *ApJ*, 669, 1107
- Yuan, F., Rykoff, E. S., Akerlof, C. W., et al. 2009, in *American Astronomical Society Meeting Abstracts*, 213, 476.06
- Zhang, B. 2007, *Chin. Astron. Astrophys.*, 7, 1
- Zhang, B., Kobayashi, S., & Mészáros, P. 2003, *ApJ*, 595, 950