

Dwarf galaxies in the Coma cluster – II. Spectroscopic and photometric fundamental planes^{★†}

E. Kourkchi,^{1,2‡} H. G. Khosroshahi,¹ D. Carter³ and B. Mobasher⁴

¹*School of Astronomy, Institute for Research in Fundamental Sciences (IPM), PO Box 19395-5531, Tehran, Iran*

²*Department of Physics, Sharif University of Technology, PO Box 11155-9161, Tehran, Iran*

³*Astrophysics Research Institute, Liverpool John Moores University, Twelve Quays House, Egerton Wharf, Birkenhead CH41 1LD*

⁴*Department of Physics and Astronomy, University of California, Riverside, CA 92521, USA*

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ABSTRACT

We present a study of the Fundamental Plane (FP) for a sample of 71 dwarf galaxies in the core of the Coma cluster in the magnitude range $-21 < M_I < -15$. Taking advantage of the high-resolution DEIMOS spectrograph on Keck II for measuring the internal velocity dispersion of galaxies and high-resolution imaging of the *Hubble Space Telescope* (HST)/ACS, which allows an accurate surface brightness modelling, we extend the FP of galaxies to luminosities of ~ 1 mag fainter than all the previous studies of the FP in the Coma cluster. We find that the scatter about the FP depends on the faint-end luminosity cut-off, such that the scatter increases for fainter galaxies. The residual from the FP correlates with the galaxy colour, with bluer galaxies showing larger residuals from the FP.

We find $M/L \propto M^{-0.15 \pm 0.22}$ in the *F814W* band, indicating that in faint dwarf ellipticals, the M/L ratio is insensitive to the mass. We find that less massive dwarf ellipticals are bluer than their brighter counterparts, possibly indicating ongoing star formation activity. Although tidal encounters and harassment can play a part in removing stars and dark matter from the galaxy, we believe that the dominant effect will be the stellar wind associated with the star formation, which will remove material from the galaxy, resulting in larger M/L ratios. We attribute the deviation of a number of faint blue dwarfs from the FP of brighter ellipticals to this effect.

We also study other scaling relations involving galaxy photometric properties including the Photometric Plane. We show that compared to the FP, the scatter about the Photometric Plane is smaller at the faint end.

Key words: galaxies: clusters: individual: Coma – galaxies: dwarf – galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: fundamental parameters – galaxies: kinematics and dynamics.

1 INTRODUCTION

The photometric and dynamical properties of elliptical galaxies could be explained by multivariate analysis which reveals the fun-

damental scaling relations of these galaxies. One of the most well-known manifolds showing a tight correlation between R_e (i.e. the radius encompassing the half-light of the galaxy), I_e (i.e. the mean surface brightness within R_e in flux units) and σ (i.e. the galaxy internal velocity dispersion) is the Fundamental Plane (hereafter FP; Djorgovski & Davis 1987; Dressler et al. 1987; Faber et al. 1987; Bernardi et al. 2003).

In theory, the FP is derived from the virial theorem as $R_e \propto \sigma^2 I_e^{-1} (M/L)^{-1}$, where I_e is the effective surface brightness in flux units, calculated within the half-light radius, R_e , of the galaxy, σ is the galaxy internal velocity dispersion and M/L is its mass-to-light ratio. Assuming that M/L is expressed by a power-law function of σ and/or the effective surface brightness, I_e , the FP relation is simplified as

$$\log(R_e) = A \log(\sigma) + B \langle \mu \rangle_e + C, \quad (1)$$

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[‡]E-mail: ehsan@ipm.ir

where $\langle\mu\rangle_e$ is the mean surface brightness in mag arcsec^{-2} and is defined as $\langle\mu\rangle_e = -2.5 \log(I_e) + \text{constant}$. Although the shape of the FP and its coefficients differ for different gravitationally bound systems from globular clusters (Burstein et al. 1997) to galaxy clusters (Schaeffer et al. 1993; Fritsch & Buchert 1999; Zaritsky, Gonzalez & Zabludoff 2006a, hereafter ZGZ06), there is no doubt about its existence (Lucey, Bower & Ellis 1991, and references therein). In reality, the coefficients of the FP relation differ from the prediction of the virial theorem. The observed coefficients are $A = 1.24$ and $B = 0.33$ (Jørgensen, Franks & Kjærgaard 1996, hereafter JFK96) while the virial theorem predicts $A = 2.0$ and $B = 0.4$. This difference, often referred to as the ‘tilt’ of the FP, is mainly attributed to different formation histories and evolutionary processes.

The difference in FP coefficients of different spheroidal systems with different mass, size and luminosities can be explained by evolution of the M/L ratio as a function of stellar population [age, metallicity or initial mass function (IMF)] and/or dark matter content (Tortora et al. 2009; Graves & Faber 2010; Grillo & Gobat 2010). In addition, the absence of homology, i.e. the fact that the structure of spheroids is scalable regardless of their size, can be the source of the FP tilt (Trujillo, Burkert & Bell 2004; D’Onofrio et al. 2008). On the other hand, some authors studied the role of dissipation in explaining the nature of the FP (Hopkins, Cox & Hernquist 2008, hereafter HCH08; Ribeiro & Dantas 2010). HCH08 claimed that the non-homology or change in the dark matter distribution is not the main driver of the FP tilt. Studying the early-type galaxies in 59 nearby galaxy clusters, D’Onofrio et al. (2008) have found a strong correlation between the FP coefficients and the local cluster environment, and no strong correlations with internal galaxy properties (e.g. Sérsic index and galaxy colour). Moreover, FP coefficients are independent of global properties of clusters such as radius, X-ray luminosity and cluster velocity dispersion (D’Onofrio et al. 2008). On the other hand, while Reda, Forbes & Hau (2005) have shown that isolated early-type galaxies lie on the same FP as galaxies in high-density environments, cluster galaxies have also less intrinsic scatter in their properties compared to field galaxies (de Carvalho & Djorgovski 1992).

The study of dark matter in dwarf galaxies showed that below the critical virial velocity, which is estimated to be $\sim 100 \text{ km s}^{-1}$, interstellar gas removal via supernova explosions becomes important (Dekel & Silk 1986, hereafter DS86). This mechanism has also been invoked to explain the shape of the low-mass dwarf galaxies and their mass profiles (Governato et al. 2010; Sánchez-Janssen, Méndez-Abreu & Aguerri 2010). Differences between the FP of giant and dwarf galaxies have been known for some time (Nieto et al. 1990; Bender, Burstein & Faber 1992; Guzman, Lucey & Bower 1993). Peterson & Caldwell (1993), studying a sample of nucleated dwarfs, found a change in the M/L ratio with luminosity as predicted by the scaling relations of DS86. The study of 17 Virgo dwarfs ($-17.5 < M_V < -15.48$) by Geha, Guhathakurta & van der Marel (2002, 2003) placed dwarf ellipticals (dEs) on a plane parallel to, but offset from, that occupied by normal elliptical (E) galaxies. On the other hand, Graham & Guzmán (2003), based on a detailed surface photometry from the *Hubble Space Telescope* (HST) archival images for a sample of dEs in the Coma cluster, claimed that the E–dE dichotomy is a consequence of a continuously varying profile shape with galaxy luminosity. They attributed the E–dE dichotomy to the fact that the light profile of dwarf galaxies does not follow the de Vaucouleurs (1948) law and therefore concluded that giant and dEs were not formed by different mechanisms. However, as found by Geha et al. (2003), modelling the dEs with Sérsic functions (1968) does not explain why dEs lie on a dif-

ferent region of the FP. Although it is clear that dwarf ellipticals do not lie on the FP as defined by brighter ellipticals (Geha et al. 2003), ZGZ06 and Zaritsky, Gonzalez & Zabludoff (2006b) showed that the differences between the ‘fundamental planes’, defined by different families of spheroids from dEs to galaxy clusters, are the result of a non-linear relationship between $\log \sigma$ and $\log M/L$. Desroches et al. (2007) also found the sensitivity of the coefficient A of the FP to the luminosity faint-end cut-off. D’Onofrio et al. (2008) were among the first to question the universality of the FP and proposed the bending of the FP of early-type galaxies. On the other hand, Fraix-Burnet et al. (2010) studied a sample of galaxies with the redshift range 0.007–0.053, and found that the global FP is a mixture of several fundamental planes, with different thicknesses and orientations, and concluded that the FP is not a bent surface. They found seven groups of fundamental planes which are assigned to different galaxy assembly histories and formation scenarios.

Another fundamental scaling relation, based on pure photometric observables, was introduced by Khosroshahi et al. (2000a) and Khosroshahi, Wadadekar & Kembhavi (2000b), called the Photometric Plane (hereafter PHP) of the galaxies. The Sérsic index in the PHP, which replaces the velocity dispersion in the FP, carries information on the light profile in galaxies and it is correlated with the dynamical properties of the galaxies such as the total mass and internal stellar velocity dispersion (Graham 2002, hereafter GR02). The steeper the light profile at the centre of the galaxy (higher Sérsic index), the more massive the galaxy and larger the central velocity dispersion are.

FP and PHP could potentially be used to derive distance to any single galaxy. GR02 has compared both photometric and fundamental planes for exactly the same sample of early-type galaxies in the Virgo and Fornax clusters, and concluded that the scatter about the PHP is ~ 15 – 30 per cent more than that about the FP. Since the PHP is purely based on photometric quantities, it can be replaced with the FP wherever the expensive kinematic data are not available and/or a high level of accuracy is not required.

In Paper I of this series (Kourkchi et al. 2012), we used the high spectral resolution of the DEep Imaging Multi-Object Spectrograph (DEIMOS) (Faber et al. 2003) on the Keck II telescope to measure precise velocity dispersions for a sample of 41 faint elliptical galaxies in the Coma cluster. Of this sample, 28 galaxies lie in the observed footprints of the *HST*/ACS Coma Treasury Survey (Carter et al. 2008). To this sample, we add a further 43 galaxies with velocity dispersions from Matković & Guzmán (2005, hereafter MG05) or Cody et al. (2009, hereafter Co09), which also lie in these footprints, giving us a combined sample of 71 objects with $-22 < M_R < -15$ with precise velocity dispersions measured from one of these sources. We noted that fainter dwarfs show a departure from brighter ellipticals on the Faber–Jackson relation (Faber & Jackson 1976), which indicates that they have higher velocity dispersion and therefore are more massive or less luminous than predicted by the L – σ linear trend.

Any study of the differences in the FP of dwarf and giant ellipticals helps us to test galaxy formation scenarios. In this paper, we present the structural parameters (effective radius, effective surface brightness, central surface brightness, Sérsic index and concentration parameter) for galaxies in our sample, based on the Treasury *HST*/ACS images in the *F814W* and *F475W* filters. The precise measured structural parameters from the *HST*/ACS images help to examine the form of the scaling relations at faint luminosity and low surface brightness, and to investigate how the Coma dwarfs fit on to the well-known FP. In addition, we study the scatter of faint dEs about the FP of giant ellipticals and examine whether it depends

upon photometric and structural parameters of galaxies, such as Sérsic index, luminosity and colour. Moreover, we examine the form of the PHP of our sample dEs.

This paper is arranged as follows. Section 2 describes the photometric data analysis and modelling of the surface brightness of galaxies. Section 3 is devoted to studying the scaling relations of our sample galaxies, such as the FP, PHP, size–surface brightness and size–luminosity relations, as well as the correlation between different photometric and kinematic parameters. The departure of faint dEs from the FP of bright ellipticals is discussed in Section 4. In Section 5, we study the variation of the M/L ratio of the galaxies in our sample across their total mass. The discussion and summary are presented in Sections 6 and 7, respectively.

Throughout this paper, we assume that the Coma cluster is located at a distance of 100 Mpc ($z = 0.0231$), which corresponds to a distance modulus of 35.00 mag and angular scale of $0.463 \text{ kpc arcsec}^{-1}$ for $h = 71$, $\Omega_m = 0.27$ and $\Omega_\Lambda = 0.73$ (Carter et al. 2008). In this paper, all magnitudes are in the AB system.

2 THE PHOTOMETRIC DATA

The *HST* ACS Coma Treasury Survey is a deep two-passband imaging survey of one of the nearest rich clusters of galaxies. The complete survey covers a 274 arcmin^2 area of the sky in the core and infall region of the Coma cluster. 25 fields were imaged by the ACS Wide Field Camera (WFC) with the $F475W$ (g -band) and $F814W$ (i -band) filters. Of the 25 fields, 19 were located within 0.5 Mpc (0.3 deg) of the Coma cluster centre. For the purpose of this study, we used the images of Data Release 2 (DR2), which include several improvements to the initial release.

2.1 Photometry of sample galaxies

Among the galaxies with DEIMOS spectroscopic data, 32 have *HST*/ACS images. Reliable velocity dispersions were derived for 28 of these galaxies. Four remaining galaxies were therefore excluded from the analysis. Moreover, velocity dispersion measurements were available for 41 more galaxies (34 from MG05 and 9 from Co09) in the Coma cluster for which ACS images at the above bands were available. The final sample consists of 71 galaxies with ACS images and determined velocity dispersion covering a luminosity range from $M_R \approx -22$ to -15 .

For the purpose of this study, galaxies in both the $F814W$ and $F475W$ images were extracted using SExtractor (Bertin & Arnouts 1996) for photometry and measuring the initial model-independent shape parameters such as the effective radius, R_e , position angle and ellipticity. For each galaxy, the initial central position as well as the concentration parameter [i.e. $C = 5 \log_{10}(r_{80}/r_{20})$, where r_{80} (r_{20}) is the radius within which 80 per cent (20 per cent) of the total galaxy light is collected] and the Kron radius, i.e. a characteristic radius as weighted by the light profile originally defined by Kron (1980), were measured using SExtractor. For each galaxy, we used both the $F814W$ and $F475W$ bands to run SExtractor in dual-image mode where the $F814W$ band was used for object detection. The SExtractor input parameters are taken directly from table 1 of Hammer et al. (2010).

2.2 Galaxy surface brightness profile

The Sérsic function, defined as $\log I(r) \propto r^{1/n}$, describes the structure of most elliptical galaxies remarkably well (Kormendy et al. 2009). To find the best Sérsic fit to the light profile, we used GALFIT (version 3; Peng et al. 2010). For ACS images, the anisotropic

point spread function (PSF) shape depends on the location of each object on ACS CCD chips (WFC1 and WFC2) and was modelled by TINYTIM (Krist 1993). To run GALFIT, initial values of the R_e , μ_e , galaxy position and galaxy position angle were taken from the SExtractor initial run, and the initial value of the Sérsic index n was set to 3. Tests show that with the well-defined PSF of the ACS data, the final solution does not depend strongly upon the initial value of n (see also Hoyos et al. 2011). Any object in the vicinity of the target galaxies was masked out. It is important to leave enough sky background for a reliable estimation of the background level, as the estimated Sérsic index, n , is slightly sensitive to the masked regions. This sensitivity is higher for larger Sérsic indices. Comparing the estimated effective radius (R_e , i.e. the radius encompassing the half-light of the galaxy) and the effective surface brightness ($\langle \mu \rangle_e$, i.e. the mean surface brightness within R_e) from GALFIT and SExtractor, we found that the difference in results (the scatter around the line with the slope ‘one’) is minimized when the size of the fitted area (galaxy and background) is about 2.5 times the Kron radius. In each case, the modelled galaxy and the GALFIT residual images were inspected by eye to identify galaxies that are well described by the Sérsic model. The galaxies either with internal spiral structure or with a poor fit were excluded from the analysis.

To study the dependence of measured Sérsic indices on the observing wavelengths, surface brightness fitting was performed on both $F475W$ and $F814W$ images. The Sérsic indices in the $F814W$ band are about 10 per cent higher than those in the $F475W$ band (see Fig. 1). To study the wavelength dependence of the FP and PHP in Section 3.2, for each band, the corresponding Sérsic index and effective radius were used.

In addition to GALFIT and SExtractor, the IRAF task ISOPHOT in the STSDAS package was used to find and compare the radial light profile of each galaxy and its corresponding estimated Sérsic model. This helps us to examine the reliability of the fitted profiles. Moreover, the central surface brightness of galaxies, μ_0 , was obtained from their estimated surface brightness profiles. Fig. 2 shows the example

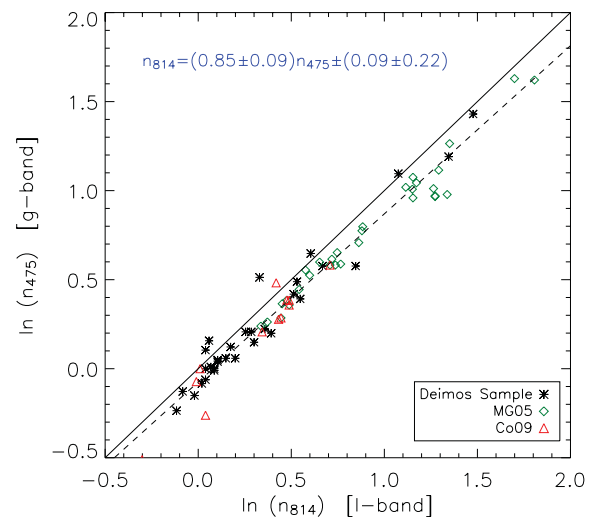


Figure 1. Comparing the Sérsic parameters in two different filters ($F814W$ and $F475W$). For both filters, the initial parameters to run GALFIT are chosen from the SExtractor catalogue of $F814W$ images. The solid black line is the locus of $n_{814} = n_{475}$. The blue dashed line is the best linear regression between the two parameters. n_{814} values are 10 per cent higher than n_{475} values, on average. Red triangles and green diamonds represent galaxies of the Co09 and MG05 samples, respectively. Black asterisks are dwarf galaxies of our DEIMOS sample.

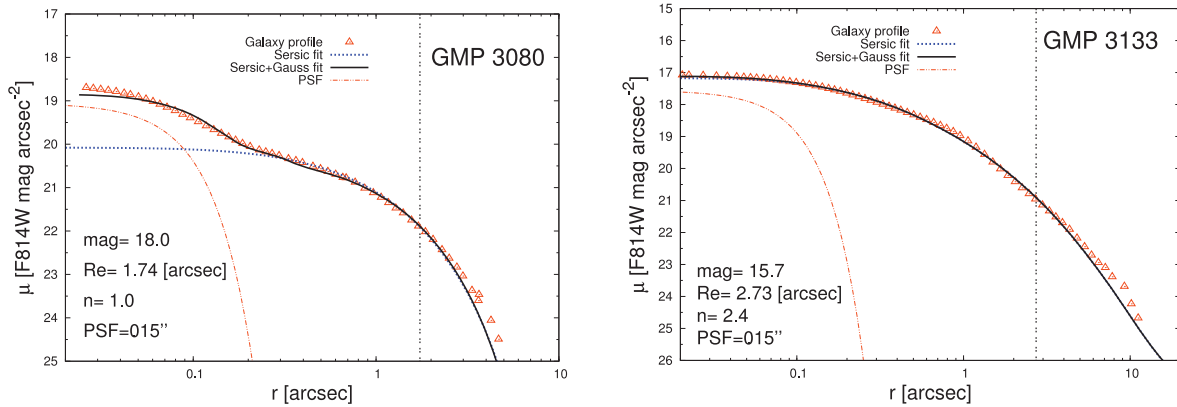


Figure 2. Two examples of the fitted Sérsic function to the galaxy light profile. Red open triangles show the real profile of the galaxies while the blue dotted and black solid curves represent the best fit from a single Sérsic and a Sérsic+Gaussian model, respectively. The vertical black double-dotted lines show the location of the effective radius, R_e , of these galaxies. The horizontal axes show the average radius of the ellipsoidal isophotes ($r = \sqrt{ab}$). The red dash-dotted curves represent the typical profile of the PSF. Left: the surface brightness profile of GMP3080. The central excess light (CEL) forced us to add an extra Gaussian component, which results in the best fit with the Sérsic index of $n = 1.0$ (the black solid curve). The FWHM of the added central Gaussian function is 0.13 arcsec. The Sérsic index corresponding to the blue dotted curve is 0.9. Right: the surface brightness profile of GMP3133, which is accurately described with a single Sérsic profile.

of two galaxies modelled by the Sérsic function. GMP3080 seems to have an extra component at its centre which is well modelled by an additional Gaussian function with full width at half-maximum (FWHM) ≈ 0.13 arcsec, in another iteration. The importance and physical meaning of this excess light is discussed in Section 4.2.

The estimated photometric and kinematic parameters of our 71 sample galaxies are presented in Table 1. Hoyos et al. (2011) present a detailed comparison of effective radius and surface brightness with the ground-based data of Gutiérrez et al. (2004) and Aguerri et al. (2005). They find a good agreement, with a few outliers where complex structure is not well resolved in the ground-based data. Our derived values are, in general, in good agreement with those of Hoyos et al. (2011) although our derived values of the Sérsic index are somewhat lower. This difference can be attributed to the fact that we allow a separate nuclear component in some of the fits.

3 THE SCALING RELATIONS

Taking all essential kinematic and photometric parameters, we have investigated the most well-known scaling relations for our sample galaxies. Our sample consists of 71 dwarf galaxies which are fainter and less massive than the previously studied galaxies in the Coma cluster. The distribution of magnitude, velocity dispersion and Sérsic index of our galaxies is presented in Fig. 3. In this section, we also compare the scaling relations of our sample of dEs with other studies.

3.1 The Fundamental Plane

In order to fit the FP to our sample galaxies, we performed an orthogonal distance regression, in which the sum of the orthogonal residuals about the FP is minimized. Compared with the ordinary least-squares fitting method, in which the residuals in $\langle \mu \rangle_e$ or R_e are minimized, the orthogonal distance regression is less sensitive to the outliers and is more robust (JFK96; La Barbera et al. 2010). During the fitting process, the square inverse of the error bars of the measured quantities was used as weight numbers, and therefore the final fitting results are less affected by the values with large error bars. To perform the weighted orthogonal distance regression, we used the ODRPACK software (Boggs et al. 1989, 1992).

The conventional projection of the FP is shown in Fig. 4. The fitting was performed on bright galaxies (i.e. $\langle \mu \rangle_e^{F814W} < 21$). In Figs 4(c) and (e), we have excluded the galaxies with poor Sérsic fits. There is no significant difference in the fitted FPs in Figs 4(a) and (c) within the 1σ uncertainty.

We probe the dependence of the FP on the wavelength by performing the Sérsic fit to galaxies in the $F475W$ filter. In Section 2.1, we showed that the measured Sérsic parameter in the red filter ($F814W$) is 10 per cent higher than that in the blue filter ($F475W$). Fig. 4(e) shows the FP for the same galaxies in the $F475W$ ACS (g band) filter. The resulting FPs for both red and blue filters (Figs 4c and e) are compatible with each other within the error bars, which implies no colour dependence of the FP of our dEs.

For the 12 brightest galaxies ($M_{814} < -20$ mag) in our sample, the best orthogonal fit for the FP is derived as

$$\log(R_e) = (1.33 \pm 0.02) \log(\sigma) + (0.32 \pm 0.00) \langle \mu \rangle_e + (-8.74 \pm 0.11), \quad (2)$$

with an rms scatter of 0.031 dex in $\log(R_e)$. R_e and σ are in kpc and km s^{-1} , respectively. Even though our study is based on the precise measurements of the structural and photometric parameters from the *HST*/ACS imaging data, the coefficients of the FP in this analysis for bright galaxies are consistent with those derived from optical (V band) and near-infrared (K band) data for 48 giant ellipticals in the Coma cluster (Mobasher et al. 1999, hereafter Mo99). Moreover, our resulting FP for bright ellipticals agrees well with other studies of the FP. For instance, JFK96 found the coefficients of $\log(\sigma)$ and $\langle \mu \rangle_e$ as 1.3 and 0.32, respectively. In addition, La Barbera et al. (2010) have investigated the FP for a sample of local early-type galaxies ($z < 0.1$) from the Sloan Digital Sky Survey (SDSS) DR6, with $M_r \lesssim -20$, in the *grizYJHK* wavebands. In agreement with our results, they have shown that the coefficient of $\langle \mu \rangle_e$ is independent of the waveband and equals 0.32 by minimizing the orthogonal residuals about the FP. The rms scatter about the FP in a previous study of bright ellipticals in the Coma cluster is 0.074 dex in $\log(R_e)$ (Mo99) for the optical V band, 0.08 dex (Djorgovski & Davis 1987; JFK96) and 0.07 dex (de Carvalho & Djorgovski 1992) for the B band and 0.075 dex

Table 1. The photometric and dynamical parameters of early-type galaxies in the Coma cluster of which 28 galaxies have spectroscopic data from the DEIMOS spectrograph. The velocity dispersion for 34 galaxies was extracted from MG05 and for nine galaxies from Co09. The photometric parameters are derived from ACS images taken by the filter $F814W$. μ_0^* is the central surface brightness of each galaxy found by the IRAF task ELLIPSE and μ_0 is the same but for galaxy model created by GALFIT. R_e is the effective radius of the galaxy that encloses half of the total light of the galaxy. n and σ (columns 10 and 11) are the Sérsic index and velocity dispersion, respectively. C (column 12) is the concentration parameter defined as $C = 5 \log_{10}(r_{80}/r_{20})$.

GMP ID (1)	RA J2000 (2)	Dec. J2000 (3)	Mag (AB) $F814W$ (4)	μ_0^* (5)	μ_0 (mag arcsec $^{-2}$) (6)	μ_e (7)	$\langle\mu\rangle_e$ (8)	R_e (kpc) (9)	n (10)	σ (km s $^{-1}$) (11)	C (12)
2489	13:00:44.69	28:06:01.00	15.1	16	16.5	20.7	19.6	1.52	2.1	94 \pm 2	4.11
2510	13:00:42.92	27:57:45.49	14.8	15.5	15.5	22.5	20.8	4.47	6.1	127 \pm 3	4.22
2516	13:00:42.85	27:58:14.90	14.1	15	14.9	21.3	19.7	3.72	5.5	176 \pm 5	3.86
2529	13:00:41.28	28:02:41.00	16.9	18.1	18.6	20.9	20	0.78	1.4	42 \pm 6	3.24
2535	13:00:40.94	27:59:46.19	14.6	15.8	16.3	21.2	20	2.79	3.1	124 \pm 2	3.97
2541	13:00:39.83	27:55:24.40	14.1	15.4	15.6	20.8	19.4	2.88	3.6	177 \pm 4	3.82
2563	13:00:37.30	27:54:41.10	19.4	19.1	21.4	23	22.3	0.66	1.1	25 \pm 10	2.84
2571	13:00:36.60	27:55:52.00	18.5	19.2	20.2	22.5	21.7	0.75	1.3	17 \pm 6	3.05
2585	13:00:35.50	27:56:32.15	17.1	19.2	19.4	22.6	21.7	1.64	1.6	30 \pm 6	3.34
2591	13:00:34.40	27:56:05.00	17.2	19.1	19.7	22.5	21.6	1.56	1.7	53 \pm 9	3.18
2605	13:00:33.30	27:58:49.00	18.3	19.6	20.9	22.9	22.2	1.05	1.2	36 \pm 10	2.89
2654	13:00:28.00	27:57:22.00	14.8	15	15.3	20.4	19	1.38	4.4	144 \pm 7	4.76
2655	13:00:27.90	27:59:16.50	19.2	20.4	21.1	23	22.3	0.71	1.1	45 \pm 5	2.84
2676	13:00:26.20	28:00:32.00	17.8	18.7	20.2	22.1	21.4	0.93	1.1	37 \pm 8	2.94
2692	13:00:24.90	27:55:34.14	16.8	18.4	19.3	23	22	2.16	1.6	41 \pm 6	3.58
2718	13:00:22.70	27:57:55.00	18.6	19.9	20.7	22.8	22.1	0.88	1.3	30 \pm 8	3.01
2736	13:00:21.70	27:53:55.00	16.7	17.4	18.5	21.4	20.4	1.08	1.7	35 \pm 4	3.43
2755	13:00:20.20	27:59:38.00	18.9	20.4	20.9	22.5	21.9	0.69	1	27 \pm 9	2.8
2778	13:00:18.88	27:56:11.75	15.6	17.7	18.2	21.9	20.8	2.47	2.1	57 \pm 4	3.36
2780	13:00:18.70	27:55:13.00	19	19.5	20.7	22.9	22.1	0.69	1.4	63 \pm 12	3.02
2784	13:00:18.62	28:05:48.00	17	19	19.4	21.9	21.1	1.24	1.5	64 \pm 8	3.12
2799	13:00:17.72	27:59:13.29	17.7	18.2	18.4	21.3	20.4	0.65	2.2	45 \pm 9	3.51
2805	13:00:17.11	28:03:48.00	15.1	15.1	15.7	20.2	19	1.22	3.6	129 \pm 3	4.21
2808	13:00:17.00	27:54:16.10	19.6	19.9	20.4	22.1	21.5	0.43	0.9	69 \pm 12	2.85
2839	13:00:14.83	28:02:27.00	14.4	14.5	15.1	19.7	18.3	1.24	3.8	168 \pm 6	4.41
2861	13:00:12.98	28:04:30.00	14.8	15.4	15.6	20.1	18.9	1.37	3.2	123 \pm 2	4.03
2877	13:00:11.40	27:54:36.00	18.6	19.3	20.2	23.1	22.2	0.96	1.8	29 \pm 9	3.23
2879	13:00:11.23	28:03:53.00	16.7	17.9	18.8	21.4	20.5	1.1	1.6	44 \pm 5	3.19
2922	13:00:08.10	28:04:41.00	14.3	15.1	15.6	18.7	17.7	0.93	2.1	185 \pm 4	3.65
2931	13:00:07.10	27:55:52.00	17.6	17.2	17.9	20.7	19.6	0.46	2.3	35 \pm 6	3.71
2960	13:00:05.47	28:01:26.00	15.5	17.6	18	21.8	20.8	2.37	2	60 \pm 5	3.69
3017	13:00:01.05	27:56:41.80	16.9	17.8	18.4	21.9	20.9	1.22	2.4	59 \pm 12	3.69
3018	13:00:01.00	27:59:30.00	18	18.1	20.2	22	21.3	0.81	1.1	38 \pm 5	2.95
3034	12:59:59.56	27:56:24.42	17.1	18.4	19.6	22.7	21.9	1.98	1.9	170 \pm 84	3.15
3068	12:59:56.75	27:55:46.40	15.1	16.1	18	21.4	20.4	2.44	1.8	106 \pm 2	3.5
3080	12:59:55.70	27:55:04.00	17.9	18.8	20.1	21.9	21.1	0.79	1.1	9 \pm 8	2.96
3098	12:59:53.90	27:58:14.00	17.6	18.9	20.3	21.8	21.1	0.93	0.9	33 \pm 7	2.76
3119	12:59:51.50	27:59:35.50	19.6	19.9	21.5	23.6	22.6	0.65	1.2	37 \pm 10	2.79
3131	12:59:50.20	27:54:46.00	17.3	18.9	20.2	22.1	21.4	1.28	1	19 \pm 7	2.79
3133	12:59:50.18	27:55:27.65	15.9	17.1	17.2	20.9	19.9	1.26	2.4	80 \pm 3	3.61
3141	12:59:49.10	27:58:33.90	19.7	19.4	21.5	23.2	22.5	0.58	1	59 \pm 15	2.88
3146	12:59:48.60	27:58:58.00	17.8	20.3	20.9	22.8	22	1.35	1	38 \pm 11	2.78
3166	12:59:46.90	27:59:31.00	17.2	19.3	20	22	21.3	1.16	1.1	38 \pm 6	3
3170	12:59:46.88	27:58:24.04	14.5	15.6	16.1	20.9	19.6	2.35	3.2	143 \pm 4	4.3
3201	12:59:44.47	27:54:43.02	14.1	15.3	16	20	19	2	2.4	170 \pm 4	3.7
3209	12:59:44.20	28:00:47.00	18.5	19.4	20.2	22	21.3	0.66	1	30 \pm 5	2.86
3213	12:59:43.80	27:59:39.15	14.8	15.4	15.8	19.8	18.6	1.15	3.2	139 \pm 5	4.01
3222	12:59:42.38	27:55:27.37	15	14	14.2	19.8	18.1	0.91	7.6	164 \pm 4	4.9
3223	12:59:42.40	28:01:58.60	18.7	19.8	21.2	23.5	22.6	1.19	1.4	22 \pm 9	2.62
3254	12:59:40.36	27:58:03.95	15.6	15.5	16.3	21	19.9	1.45	3.6	101 \pm 5	4.18
3269	12:59:39.73	27:57:12.57	15.3	15.4	15.9	20.5	19.3	1.3	3.2	99 \pm 3	4.18
3292	12:59:38.00	28:00:03.70	16.6	17.1	18.8	21.4	20.5	1.12	1.5	40 \pm 5	3.21
3308	12:59:37.20	27:58:20.00	18	16.3	16.6	18.7	17.6	0.16	2.9	50 \pm 4	3.87
3312	12:59:37.00	28:01:07.00	17.5	17.9	18.3	20.8	19.9	0.56	1.7	31 \pm 4	3.38
3339	12:59:35.37	27:51:47.40	16.1	17.8	17.9	20.7	19.8	1.08	1.7	59 \pm 4	3.17

Table 1 – *continued*

GMP ID	RA J2000	Dec. J2000	Mag (AB) F814W	μ_0^*	μ_0 (mag arcsec $^{-2}$)	μ_e	$\langle\mu\rangle_e$	R_e (kpc)	n	σ (km s $^{-1}$)	C
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)
3367	12:59:32.86	27:58:59.55	14.1	14.7	15.7	20.9	19.5	2.84	3.5	179 \pm 4	4.17
3400	12:59:30.90	27:53:01.74	14	14.5	14.6	19.4	18.1	1.42	3.9	217 \pm 5	3.86
3406	12:59:30.30	28:01:15.10	17.7	18.5	19.2	21.6	20.8	0.76	1.4	36 \pm 3	3.22
3438	12:59:28.50	28:01:09.40	18	18.8	20.1	21.9	21.2	0.78	1.2	26 \pm 8	3
3473	12:59:26.40	27:51:25.00	17.6	18.8	19.2	21.6	20.7	0.78	1.5	33 \pm 4	3.27
3681	12:59:11.64	28:00:31.00	16.6	18.1	18.5	21.1	20.3	1.01	1.6	64 \pm 5	3.31
3707	12:59:09.53	28:02:26.00	16.2	17.5	17.8	21.7	20.6	1.49	2	77 \pm 4	3.77
3780	12:59:04.87	28:03:00.00	16.4	18.1	18.6	22.1	21.2	1.75	1.8	58 \pm 3	3.54
4035	12:58:45.50	27:45:14.00	17.1	19.1	19.7	22.3	21.4	1.49	1.6	26 \pm 4	3.14
4135	12:58:37.30	27:10:35.00	15.1	18.3	19.1	20.6	20.1	1.93	0.7	42 \pm 3	2.55
4215	12:58:31.70	27:23:42.00	17.6	20.5	20.7	23.3	22.6	1.82	1	7 \pm 4	3.11
4381	12:58:15.30	27:27:53.00	17.7	18.4	18.6	21.8	20.8	0.81	2	22 \pm 4	3.56
4430	12:58:20.50	27:25:46.00	16.6	18.8	18.9	21.9	21	1.49	1.4	38 \pm 3	3.17
5102	12:57:04.30	27:31:34.00	15.9	18.1	19.2	21.9	21.2	2.18	1	42 \pm 3	3.18
5364	12:58:33.10	27:21:52.00	15	16.6	18.5	21.8	20.9	3.26	1.5	84 \pm 4	4.11
5365	12:56:34.60	27:13:40.00	15.5	17.6	19.4	21.4	20.7	2.37	1	44 \pm 3	3.01

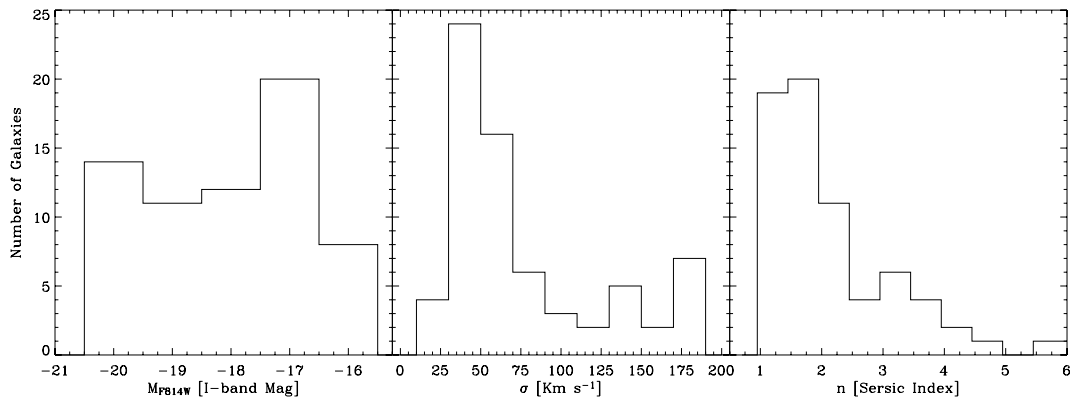


Figure 3. The distribution of galaxies in our sample in terms of the absolute magnitude in the *F814W* band, central velocity dispersion, σ , and the Sérsic parameter, n .

(Lucey et al. 1991) for the *V* band, using other independent samples. The rms scatter about the fitted FPs depends on the sample completeness and the magnitude range of the studied galaxies. We noted that adding fainter galaxies results in larger scatters, indicating that faint dwarf galaxies do not lie on the FP of their brighter counterparts.

The best coefficients of the FP relation for 35 galaxies of our sample, brighter than $M_{814} = -18$, are $A = 0.77 \pm 0.11$ and $B = 0.17 \pm 0.03$, which result in an rms scatter of 0.116 dex in $\log(R_e)$, corresponding to an uncertainty of 28 per cent in distances to individual galaxies. Considering all galaxies in our sample with $M_{814} < -16$ results in an FP with $A = 0.97 \pm 0.16$ and $B = 0.19 \pm 0.02$ with an rms scatter of 0.150 dex in $\log(R_e)$, which corresponds to an uncertainty of 35 per cent in distances to individual galaxies. The coefficients of the FP for faint magnitude cut-offs, $M_{814} < -18$ and $M_{814} < -16$, differ from those of equation (2) for $M_{814} < -20$. This implies that compared to bright ellipticals with $M_{814} < -20$, the faint dEs lie on an entirely different FP with larger scatter. In addition, we found that the fitted FPs for two samples of faint dEs with different faint-end cut-offs (i.e. $M_{814} < -18$ and $M_{814} < -16$) are consistent with each other within the error bars, suggesting the same FP for dwarf galaxies in the magnitude range $-20 < M_{814} < -16$ mag. In Table 2, the results of the best-fitted FPs for different

magnitude ranges are presented. We study the departure of faint dEs from the FP of bright ellipticals in Section 4.

We also studied the FP in κ -space (Bender et al. 1992). κ parameters form an orthogonal coordinate system which combines the central velocity dispersion, the effective radius and the effective surface brightness. In this formalism, κ_1 and κ_3 are proportional to $\log(M)$ and $\log(M/L)$ respectively, in which M is the mass of the galaxy while L is its luminosity. Moreover, κ_2 is proportional to $\log(M/L)/\langle\mu\rangle_e^3$ and measures the galaxy compactness for a given mass. The projection of the galaxies in κ -space is illustrated in Fig. 5. The left panel of Fig. 5 shows that the M/L ratio of faint dEs does not decrease in the same linear trend as brighter ellipticals. The variation of the M/L ratio across our sample galaxies and its physical meanings are studied in Section 5. As seen in the middle panel of Fig. 5, the occupied region of our sample dEs in the κ_1 – κ_2 diagram agrees well with fig. 3 of Maraston et al. (2004). This indicates that dEs are generally form a linear trend perpendicular to the region occupied by bulges and giant elliptical galaxies. The different behaviour of giant and dwarf ellipticals in the κ_1 – κ_2 space indicates that they are distinct families with different formation scenarios. Similar behaviour, albeit with smaller samples of lower-luminosity galaxies, is seen in fig. 9 of Guzman et al. (1993) and fig. 1 of Aguerrí & González-García (2009).

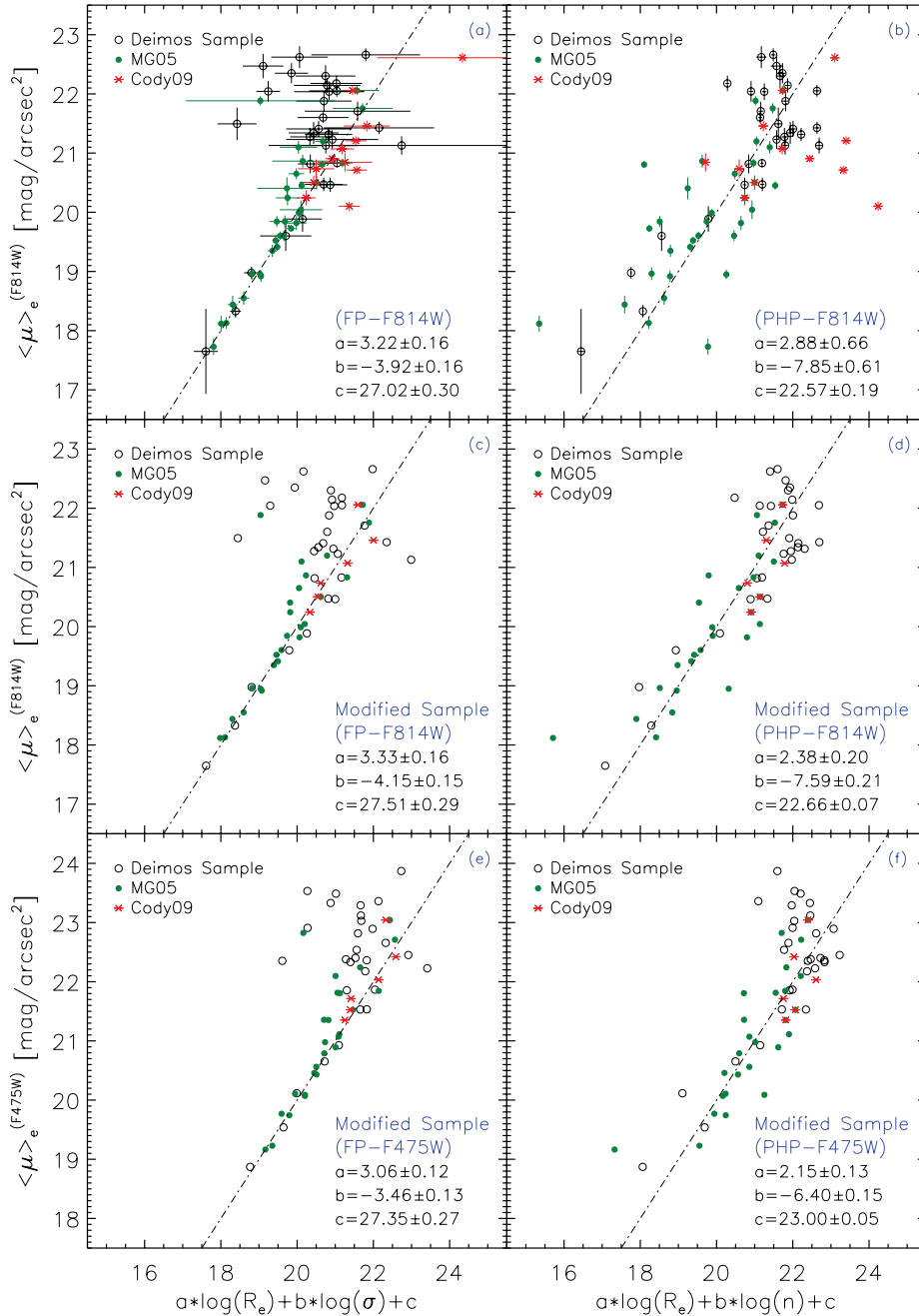


Figure 4. FP and PHP in different filters. (a) The dash-dotted line represents the projection of the FP fitted for galaxies with $\langle \mu \rangle_e < 21$. Most galaxies with $\langle \mu \rangle_e > 21$ tend to have a higher velocity dispersion compared with that expected from the FP of bright galaxies. R_e and σ are in terms of kpc and km s^{-1} , respectively. High errors occur when measuring the velocity dispersion of fainter galaxies (e.g. $\langle \mu \rangle_e > 21$) with very low signal-to-noise ratio in their spectra. Therefore, in this diagram, the horizontal error bars are bigger than the vertical error bars. (b) The dash-dotted line represents the projection of the PHP fitted for galaxies with $\langle \mu \rangle_e < 23$. Most of the outliers are S0 galaxies or galaxies with internal structures. These galaxies typically do not have very good Sérsic fits. All red asterisk belong to dS0 or SB galaxies. (c) Same as panel (a), eliminating the galaxies that are not well modelled by the Sérsic function and have internal structures. The modified sample consisting of dEs is represented in this panel. To avoid any confusion, the error bars are not shown. The typical size of error bars is the same as in (a). (d) Same as the diagram in (b) for the same sample as in panel (c). The typical size of error bars is the same as in panel (b). (e) and (f) Same as (c) and (d) for the $F475W$ band. All photometric properties and profile parameters (e.g. Sérsic index) were derived from $F475W$ images. The resulting FP and PHP in both filters show no discrepancy within 1σ error bars. The difference in coefficient ‘c’ comes from the different magnitude of each galaxy in different passbands. In all panels, open black circles represent dwarf ellipticals from our DEIMOS/Keck observations. For red asterisks and green filled circles, the velocity dispersions were obtained from Cody et al. (2009) and Matković & Guzmán (2005), respectively.

3.2 The Photometric Plane

The Sérsic index and the central velocity dispersion of galaxies are correlated (see Fig. 6) and hence one is able to use the Sérsic

index instead of the velocity dispersion whenever the kinematic parameters of the galaxies are not available. Replacing the velocity dispersion, σ , with the Sérsic index, n , in the FP relation (equation 1), the PHP is obtained more economically than the FP, only

Table 2. Coefficients of the FP for different luminosity cut-offs. The FP relation is represented as $\log(R_e) = A \log(\sigma) + B \langle \mu \rangle_e + C$.

M_{F814W}	A	B	rms
< -20	1.33 ± 0.02	0.32 ± 0.00	0.031
< -18	0.77 ± 0.11	0.17 ± 0.03	0.116
< -16	0.97 ± 0.16	0.19 ± 0.02	0.150

based on the photometric properties of the galaxies. Very similarly, the PHP could be valuably used as a diagnostic tool to study the galaxy evolution and constrain the early-type galaxy luminosity evolution with redshift (La Barbera et al. 2005, hereafter Lb05). In order to provide a local Universe reference and to compare the relation between FP and PHP of the early-type galaxies, we construct the PHP for our sample of dwarf galaxies.

The right-hand panels of Fig. 4 show the projection of best PHP for our sample galaxies regardless of their brightness. We note that most of the outliers in the top right panel have internal spiral structure or a poor Sérsic fit. It appears that while the FP is less sensitive to a detailed morphology of the galaxies, the PHP is able to effectively differentiate the morphological type of the galaxies. In the right middle panel, we have eliminated the outliers of the top panel in order to only have elliptical galaxies with good Sérsic light profiles (hereafter ‘the modified sample’). This reduces the rms scatter in $\langle \mu \rangle_e$ from 1.12 to 0.74 mag. In addition, we have constructed the PHP using the $F475W$ ACS (g band) data in order to examine the wavelength dependence of the PHP. The resulting PHPs for both filters (Figs 4d and f) are compatible with each other within the uncertainties. This is also consistent with the results of a multiband ($R/I/K$ band) study of the PHP of early-type galaxies (Lb05). Considering the following representation for the PHP:

$$\log(R_e) = \alpha \log(n) + \beta \langle \mu \rangle_e + \gamma, \quad (3)$$

Lb05 obtained $\alpha \approx 1$ and $\beta \approx 0.2$, in both the optical and near-infrared wavebands, for a sample of galaxies brighter than $M_I \approx -20$ mag located in the MS1008–1224 galaxy cluster at $z = 0.306$. Lb05 have reported the intrinsic rms dispersion in R_e as 32 per cent. Re-deriving the PHP coefficients for the modified GR02 sample of galaxies in the Virgo and Fornax clusters and using the same

regression method, Lb05 found α and β as 1.14 ± 0.15 and 0.180 ± 0.031 , respectively. This implies no redshift dependence of the PHP of bright elliptical galaxies.

For galaxies brighter than $M_{814} \approx M_I \approx -18$ mag in our modified sample, we found $\alpha = 1.82 \pm 0.10$ and $\beta = 0.27 \pm 0.02$ with an rms scatter of 0.173 dex in $\log(R_e)$, corresponding to an uncertainty of 40 per cent in distances to individual galaxies. Considering all galaxies brighter than $M_I \approx -16$, the results are $\alpha = 2.78 \pm 0.23$ and $\beta = 0.39 \pm 0.04$, with an rms scatter of W dex in $\log(R_e)$ or 61 per cent error in distances to each galaxy.

In this study, we have few bright galaxies and therefore we cannot compare our fitted PHP for bright galaxies with results in the literature. The measured coefficients of the PHP relation for our dEs differ from those of previous studies of the PHP of bright elliptical galaxies (GR02; Lb05). We attribute this discrepancy mainly to the difference in mass range, most of our galaxies have $\sigma < 100 \text{ km s}^{-1}$ whereas previous samples have $\sigma > 100 \text{ km s}^{-1}$. The typical Sérsic indices of our galaxies are less than ~ 2 while those of GR02 and Lb05 are greater (see Fig. 3). As another possibility to explain the discrepancies, we note that our study is based on the *HST*/ACS photometric data and therefore the measured photometric and structural parameters are likely to be more accurate than those measured from ground-based observations.

3.3 Fundamental Plane versus Photometric Plane

As Fig. 4 shows, the deviation of galaxies from the FP and PHP becomes larger as their luminosities, central velocity dispersions and Sérsic indices decrease. In order to quantify the dispersion around the planes, the rms deviation about the fitted planes is listed in Table 3. The deviations are derived in terms of the mean effective surface brightness and the effective radius for three cases: (i) all data points, (ii) brighter galaxies with $\langle \mu \rangle_e < 21$ and (iii) fainter galaxies with $\langle \mu \rangle_e > 21$. The table shows that the scatter around the PHP is less than the corresponding FP when we consider all of our modified sample galaxies (Figs 4c and d) or when galaxies are fainter ($\langle \mu \rangle_e > 21$). For our modified sample (i.e. galaxies with a good Sérsic fit or without any internal structure) the rms scatter in $\langle \mu \rangle_e$ for the FP is 1.04 mag, while it is reduced to 0.72 mag for the PHP. The corresponding deviation in R_e is 0.31 dex for the FP and

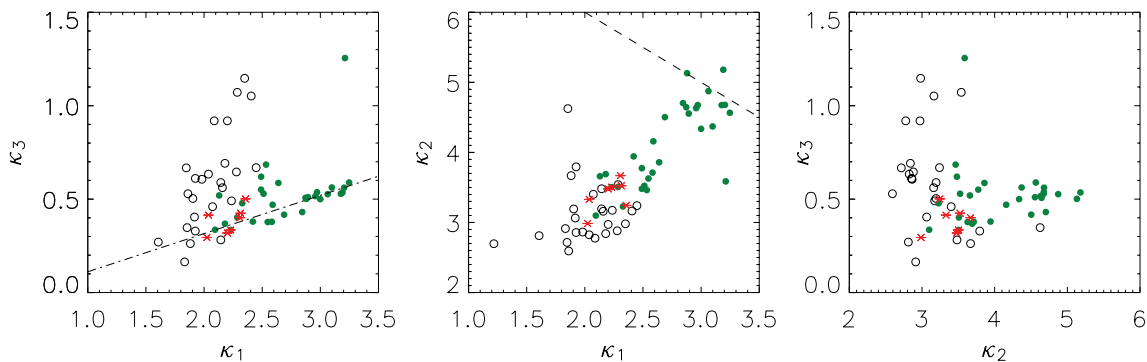


Figure 5. Dwarf elliptical galaxies of the Coma cluster plotted in κ -space (Bender et al. 1992). All galaxies with a bad Sérsic fit, dS0 galaxies and those with internal spiral structure are not plotted in this diagram. κ_1 , κ_2 and κ_3 are proportional to $\log(M)$, $\log(M/L)\langle \mu \rangle_e^3$ and $\log(M/L)$, respectively. The velocity dispersions for black open circles were measured using our DEIMOS data. For red asterisks and filled green circles, the velocity dispersion was extracted from Co09 and MG05, respectively. In the left panel, the dash-dotted line [$\kappa_3 = (0.20 \pm 0.06)\kappa_1 + (-0.09 \pm 0.17)$] is the best fit for galaxies with $\langle \mu \rangle_e < 20$, for which κ_1 is typically greater than 2.50. The estimated linear trend for bright galaxies in our sample is consistent with the results of Treu et al. (2006), who found $\kappa_3 = (0.21 \pm 0.02)\kappa_1 + (0.19 \pm 0.08)$ in the B band, for a large sample of Coma galaxies. This diagram indicates that the M/L ratios of less massive dEs are larger than those predicted by the extrapolated line to lower values of κ_1 [$\kappa_1 \propto \log(M)$]. The dashed diagonal line in the middle panel ($\kappa_1 + \kappa_2 = 8$) represents the border of the ‘zone of avoidance’ not populated by early-type galaxies (Burstein et al. 1997).

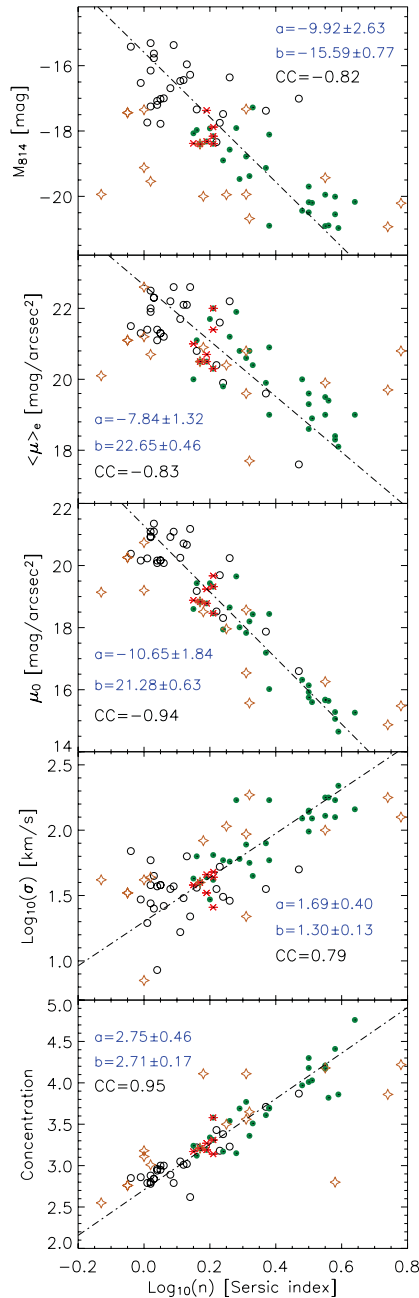


Figure 6. The relation between Sérsic index, n , of the galaxies under study and the photometric and dynamical parameters (i.e. M_{814} , μ_0 , $\langle\mu\rangle_e$, concentration parameter and σ). $\langle\mu\rangle_e$ and μ_0 are the effective and central surface brightness of the galaxies, respectively. For the definition of concentration parameter, refer to Section 2.1. All photometric parameters are extracted from the $F814W$ ACS images. For open circles, the velocity dispersions are derived from DEIMOS spectra in this study and for red asterisks and filled green circles they are extracted from the Co09 and MG05 catalogues, respectively. Open stars belong to the galaxies with a bad Sérsic fit or with internal spiral structures. In each panel, ‘ a ’ and ‘ b ’ are respectively the slope and intercept of the fitted lines using orthogonal regression. CC is the cross-correlation coefficient of the fitting parameters.

0.30 dex for the PHP. If we only consider the fainter galaxies (i.e. $\langle\mu\rangle_e > 21$), the rms deviation in $\langle\mu\rangle_e$ is 1.49 mag for the FP, while for the PHP this decreases to 0.73 mag. The rms scatter of brighter galaxies (i.e. $\langle\mu\rangle_e < 21$) around the PHP increases compared to the

corresponding FP. This implies that the brighter galaxies are still better placed on the FP. We conclude that faint galaxies agree better with the PHP of brighter galaxies, compared to the FP. This implies that the galaxy substructures that are not reflected in the velocity dispersion, and hence in the FP relation, are better reflected in the Sérsic parameter, and hence in the PHP relation. In addition, the larger errors in velocity dispersion measurements for fainter dEs may be responsible for their larger scatter about the conventional FP compared to the PHP.

3.4 Correlations with the Sérsic index

In Fig. 6, we present the correlations between the Sérsic index of the sample galaxies and M_{814} , $\langle\mu\rangle_e$, μ_0 , σ and concentration parameter. We find a linear trend between the Sérsic index and the central surface brightness as $\log_{10}(n) = (2.00 \pm 0.35) - (0.09 \pm 0.02)\mu_0$ with the correlation coefficient of -0.94 . This is entirely consistent with the relation between these quantities found by Graham & Guzmán (2003). Replacing n by σ we find a similar relation [i.e. $\log_{10}(\sigma) = (3.97 \pm 0.56) - (0.12 \pm 0.02)\mu_0$] with a weaker correlation coefficient of -0.80 . We ignore the outliers in the fitting process, which are the galaxies with a poor Sérsic fit, represented by open asterisks. Fig. 6 indicates another correlation between the central velocity dispersion of our sample galaxies, σ , and their Sérsic indices, n , with the correlation coefficient of 0.79 . This correlation enables us to construct a relation between σ and n as $\log_{10}(\sigma) = (0.59 \pm 0.14)\log_{10}(n) - (0.77 \pm 0.20)$, where σ is in km s^{-1} .

The correlation between the light concentration in galaxies and their Sérsic index is also presented in Fig. 6 (bottom). Ignoring the outliers illustrated with open asterisks, this relation is written as $\log_{10}(n) = (0.36 \pm 0.06)C - (0.98 \pm 0.18)$, with the correlation coefficient of 0.95 , where C is the concentration parameter and is defined as $C = 5 \log_{10}(r_{80}/r_{20})$. The process of determining the Sérsic index is model dependent, while the concentration parameter is model independent and well determined using a simple photometric analysis. The C – n relation shows that ‘ C ’ and ‘ n ’ could be interchangeably used.

3.5 The size–surface brightness relation

The linear relation between the effective surface brightness, $\langle\mu\rangle_e$, and half-light radius, R_e , of elliptical galaxies, also known as the Kormendy relation (KR; Kormendy 1977), is represented as

$$\langle\mu\rangle_e = p \log(R_e) + q. \quad (4)$$

Fig. 7 shows the best fit of the KR for our dEs in different luminosity ranges, $M_{814} < -20$, $-20 < M_{814} < -18$ and $-18 < M_{814} < -16$. For our brightest galaxies ($M_{814} < -20$), we find the KR slope as 4.27 ± 0.18 , which is significantly comparable to 2.43 ± 0.15 for a sample of giant galaxies in the Coma cluster with the central velocity dispersion of $\sigma > 200 \text{ km s}^{-1}$ (Ziegler et al. 1999). We attribute this discrepancy to the fact that our sample covers a different range of size and magnitude compared with the galaxies studied by Ziegler et al. (1999). The typical velocity dispersion of our dEs is less than 100 km s^{-1} and the average size of our dEs is $\sim 1.5 \text{ kpc}$. Moreover, the effective radii of our galaxies with $M_{814} < -20$ are less than $\sim 4.5 \text{ kpc}$, while all other studies of the KR cover galaxies with larger sizes ($> 10 \text{ kpc}$) in this magnitude range (e.g. Ziegler et al. 1999; D’Onofrio et al. 2006; La Barbera et al. 2010). The imposed systematic restrictions on our sample of dEs, such as luminosity and size cuts, change the geometric shape of the distribution of galaxies on the $\log(R_e)$ – $\langle\mu\rangle_e$ plane, and therefore results

Table 3. rms scatter of the data points about the best-fitting FP and PHP (see Fig. 4). The scatters are calculated along the effective surface brightness, $\langle\mu_e\rangle$, and effective radius, R_e . For brighter galaxies ($\langle\mu_e\rangle < 21$ in the $F814W$ band), the scatter about the FP is smaller than the corresponding PHP. Considering all data points or only fainter galaxies ($\langle\mu_e\rangle > 21$ in the $F814W$ band) indicates that the PHP reduces the rms scatter in both $\langle\mu_e\rangle$ and R_e .

Fitted plane	Filter	Sample	rms scatter along $\langle\mu_e\rangle$ (mag)			rms scatter along $\log(R_e)$ (dex)		
			All data	$\langle\mu_e\rangle < 21$	$\langle\mu_e\rangle > 21$	All data	$\langle\mu_e\rangle < 21$	$\langle\mu_e\rangle > 21$
FP (a)	$F814W$	All	1.01	0.35	1.50	0.32	0.11	0.47
FP (c)	$F814W$	Modified	1.04	0.28	1.49	0.31	0.08	0.45
FP (e)	$F475W$	Modified	1.04	0.28	1.49	0.34	0.09	0.49
PHP (b)	$F814W$	All	1.10	1.24	0.87	0.38	0.43	0.30
PHP (d)	$F814W$	Modified	0.72	0.72	0.73	0.30	0.30	0.31
PHP (f)	$F475W$	Modified	0.76	0.60	0.92	0.36	0.28	0.43

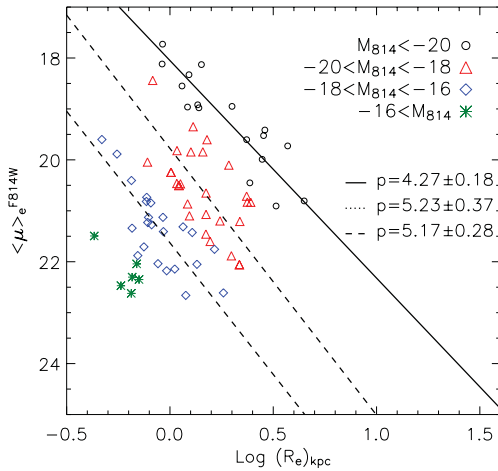


Figure 7. The linear relation between $\langle\mu_e\rangle$ and $\log(R_e)$, also known as the Kormendy relation, for different magnitude ranges in the $F814W$ band. The coefficient ‘ p ’ is the slope of the linear trends: $\langle\mu_e\rangle = p \log(R_e) + q$.

in a different KR slope (Nigoche-Netro, Ruelas-Mayorga & Franco-Balderas 2008). Fitting the KR to our dEs in the magnitude range $-20 < M_{814} < -18$ and $-18 < M_{814} < -16$, we found the KR slope as 5.23 ± 0.37 and 5.17 ± 0.28 , respectively. In agreement with our results, Khosroshahi et al. (2004) have also found the slope of the KR as ~ 5.2 for dEs of 16 nearby galaxy groups with $-14 < M_R < -18$.

The KR is originated from the definition of the effective radius, R_e , which relates the luminosity and effective surface brightness as $L = 2\pi I_e R_e^2$. Theoretically, p equals 5, and any difference from 5 is the result of the change in geometric shape of the distribution of galaxies on the $\log(R_e)$ – $\langle\mu_e\rangle$ plane. Any change in the magnitude range of the galaxies and the shape of the magnitude distribution results in a different slope of the KR (Nigoche-Netro et al. 2008).

In agreement with Khosroshahi et al. (2004) and D’Onofrio et al. (2006), Fig. 7 shows that for fainter galaxies the $\log(R_e)$ – $\langle\mu_e\rangle$ linear trend is steeper than that for the brighter galaxies ($M_{814} < -20$). We also noted that the slopes of the KR for our galaxies in the magnitude range $-20 < M_{814} < -18$ and $-18 < M_{814} < -16$ are consistent within the error bars. Due to the limited R_e range, we did not fit the KR for galaxies with $-16 < M_{814}$.

3.6 The size–luminosity relation

In Fig. 8, we show the size–luminosity relation, $R_e \propto L^\tau$, for our sample dEs. We found $\tau \approx 0.24$, which is consistent with the results of de Rijcke et al. (2005), who found the B -band radius–

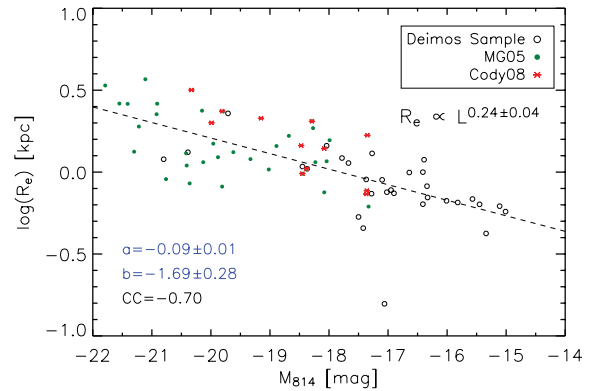


Figure 8. The effective radius against the luminosity of our sample galaxies in the $F814W$ band. ‘ a ’ and ‘ b ’ are the slope and intercept, respectively, of the best linear trend which is found using an orthogonal distance regression (dashed line). As seen, R_e and luminosity are correlated with the correlation coefficient of $CC = -0.70$. The rms scatter about the fitted line is 0.17 dex along the vertical axis, R_e .

luminosity power-law slope between 0.28 and 0.55 for a sample of dEs and dwarf spheroidal galaxies. τ depends on the luminosity range of the galaxies under study. As Desroches et al. (2007) found, τ systematically varies from $\tau \approx 0.5$ at $M_r \approx -20$ to $\tau \approx 0.7$ at $M_r \approx -24$, which are, in general, in agreement with our result if one extrapolates them to lower luminosities despite the structural differences between luminous and faint dEs.

4 DEPARTURE OF GALAXIES FROM THE FP

In order to understand the deviation of the faint dEs from the FP associated with the brighter galaxies, we study the correlations between the departure of galaxies from the FP and other observables (e.g. magnitude, Sérsic index, concentration parameter, velocity dispersion and colour). In this study, we define the departure of galaxies from the FP as $\Delta_{FP} = |\langle\mu_e\rangle - \langle\mu_e^*\rangle|$, where $\langle\mu_e\rangle$ is the effective surface brightness of the sample galaxies and $\langle\mu_e^*\rangle$ is the expected effective surface brightness from the FP of bright ellipticals.

4.1 Dependence on galaxy light profile

In Fig. 9, the correlation between Δ_{FP} and magnitude, M_{814} , Sérsic index (n) and the concentration parameter is studied. The data points are binned and the error bars are based on the 1σ rms scatter of Δ_{FP} within each bin. As the different panels of the figure show, Δ_{FP} is anti-correlated with the luminosity of galaxies, their Sérsic indices and their light concentration. Although these three parameters are

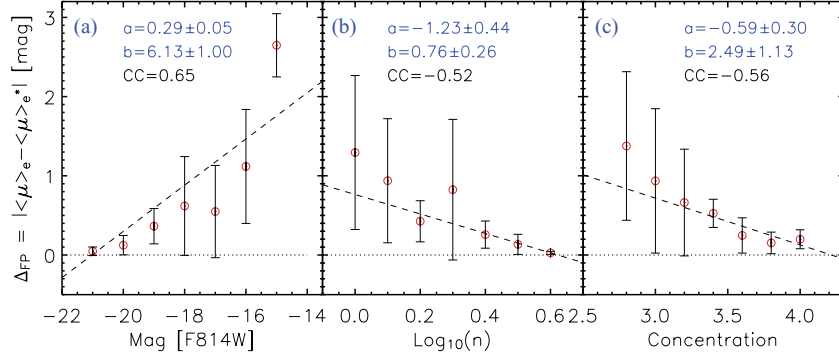


Figure 9. The deviation of the real effective surface brightness, $\langle \mu \rangle_e$, from that predicted by the FP, $\langle \mu \rangle_e^*$, against different dynamical and photometric parameters. Panel (a) shows the tight correlation between the deviation from the FP and the luminosity of the galaxies. Panel (b) shows weaker correlations in terms of n . In panel (c), the deviation is drawn in terms of the concentration parameter which is model independent and defined as $5 \log_{10}(r_{80}/r_{20})$. For all cases, we used the least-squares method to find the best linear trend (dashed lines) by minimizing the residuals in Δ_{FP} . ‘ a ’ and ‘ b ’ are the slope and intercept of the fitted lines, respectively.

correlated with each other (see Fig. 6), the figure shows that the maximum correlation is seen between the luminosity of galaxies and Δ_{FP} with a correlation coefficient of 0.65.

In this study, we obtain the partial correlation coefficients between Δ_{FP} and other observables, defined as $P_{cc}(\Delta_{FP}, \alpha; \beta)$. Here, P_{cc} measures the correlation between Δ_{FP} and one of the observables, α , in the case where the influence of the third variable, β , is eliminated. We obtained the partial correlation coefficients between Δ_{FP} and luminosity by taking out the effects of σ and n as $P_{cc}(\Delta_{FP}, M_{814}; \sigma) = 0.64$ and $P_{cc}(\Delta_{FP}, M_{814}; n) = 0.45$. In addition, we found that Δ_{FP} is very weakly correlated with σ with the correlation coefficient of -0.30 . This is consistent with the fact that σ itself contributes to the FP relation and hence no strong correlation between Δ_{FP} and σ is expected. Furthermore, taking out the effect of σ (or n), when calculating the $P_{cc}(\Delta_{FP}, M_{814}; \sigma \text{ or } n)$, does not significantly change the correlation coefficient between Δ_{FP} and M_{814} . Furthermore, the correlation coefficient between Δ_{FP} and n is -0.52 and after taking out the effect of galaxy magnitudes, we obtain $P_{cc}(\Delta_{FP}, n; M_{814}) = -0.19$. We also obtained $P_{cc}(\Delta_{FP}, n; \sigma) = -0.46$. These all imply that the departure from the FP is correlated more tightly with luminosity than with the Sérsic index.

The departure from the FP, Δ_{FP} , is also correlated with the light concentration in the sample galaxies, C , with the cross-correlation factor of -0.56 . As the concentration is expected to be correlated with the intrinsic luminosity of the galaxies, we found partial correlation to be $P_{cc}(\Delta_{FP}, C; M_{814}) = -0.17$. This is similar to the Δ_{FP} correlation with the Sérsic index, as both indicate roughly the same in the galaxy light profile. This can also be inferred from the tight correlation between the Sérsic index and the concentration parameter (see Fig. 6).

We have also examined the correlation between Δ_{FP} and effective radius, R_e , of our sample galaxies, as a representative of their size. In this case, the correlation coefficient is obtained as -0.30 . This implies that Δ_{FP} is almost size independent. This is also expected from the contribution of R_e in the FP relation.

4.2 Dependence on galaxy central excess light

Another possible parameter that may be correlated with Δ_{FP} is the central excess light (CEL) relative to their best single Sérsic model (see Fig. 2). This CEL or a nuclear star cluster occurs in the majority of dwarf galaxies in the Coma cluster (den Brok et al. 2011), and is

somewhat bluer than the main body of the galaxy. The excess light is an imprint of the formation history of galaxies and its colour and distribution indicate that it is a result of some dissipational process. A probable explanation is that it is the result of the last major merger, in which the gas content of the gas-rich progenitor falls into the core and undergoes star formation, generating the excess light (e.g. HCH08; Kormendy et al. 2009). Numerical simulations (Makino & Hut 1997) show that the galaxy merger rate scales as N^2/σ^3 , where N is the density of galaxies and σ is velocity dispersion. Hence, in the Coma cluster, with 3.5 times the density and σ 1.2 times greater, mergers should occur seven times more frequently than in the Virgo cluster. den Brok et al. (2011) conclude that the observed relationship between the Sérsic index and the strength of the colour gradient also points towards a history of wet mergers in the history of dwarf galaxies.

Alternatively, the nuclear excess could be formed from gas already in the dwarf galaxies as they fall into the cluster. Ram-pressure stripping will remove gas from the outer parts and tidal interactions could drive the residual gas into the core, where it forms stars. In this case, we would expect an environmental dependence of the CEL properties. Such a dependence will be complex, as both the stripping which removes gas and the tidal interactions driving residual gas to the core will be strongest in the centre of the cluster; unfortunately, due to the failure of ACS in 2007, and the consequent truncation of the *HST* survey, the sample of dwarfs with high spatial resolution images outside the cluster core is small.

An inspection of the radial surface brightness profile of our sample galaxies confirmed that most parts of the galaxies, beyond the aperture with 0.40 arcsec diameter, are well modelled by the Sérsic function. Therefore, we measured the CEL of our sample galaxies within apertures with a diameter of 0.4 arcsec (i.e. 0.19 kpc at the Coma distance), which were co-centred with the target galaxies. Here, we represent the CEL as

$$\text{CEL} = \left(\frac{F_{\text{gal}} - F_{\text{mod}}}{F_{\text{mod}}} \right)_{\text{aper}}, \quad (5)$$

where F_{gal} is the flux of the target galaxy within the defined central aperture, and F_{mod} is the flux of the fitted Sérsic model, within the same aperture. The extra flux is plotted against the deviation of galaxies from the FP (Δ_{FP}) in Fig. 10. Even though the Δ_{FP} –CEL correlation coefficient is 0.12, we did not find any clear relation between Δ_{FP} and the CEL. In general, the most deviant galaxies (the labelled points) tend to have more CEL. As Fig. 10 shows, the

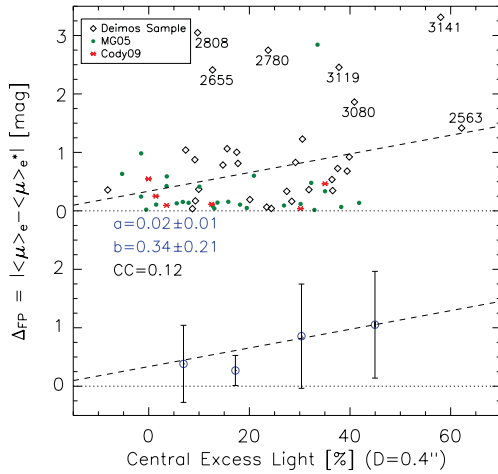


Figure 10. Deviation from the FP against the relative CEL (see the text). CEL is measured within an aperture of 0.4 arcsec diameter (i.e. ~ 0.19 kpc). The black open diamonds show the galaxies in our DEIMOS sample. For red asterisks and filled green circles, the velocity dispersion was extracted from Co9 and MG05, respectively. The most deviant galaxies are labelled with their GMP IDs (1983).

Table 4. The cross-correlation coefficient of the colour of our sample galaxies and other photometric and dynamical parameters. Please refer to Section 4.3 for the definition of colour parameters. Fig. 11 also represents Δ_{FP} versus the colour parameters.

Correlation coefficients (1)	Δ_{FP} (2)	M_{814} (3)	$\log_{10}(\sigma)$ (4)	$\log_{10}(n)$ (5)
$g - r$	-0.48	-0.73	0.57	0.72
Δm	-0.54	-0.78	0.58	0.74
$\Delta m(04)$	-0.58	-0.78	0.61	0.74
$\Delta m(10)$	-0.59	-0.81	0.62	0.78

relative CEL of all galaxies with $\Delta_{FP} > 1$ mag is more than 10 per cent.

It is likely that these galaxies have extra luminosity at their central regions due to star formation imposed by either wet mergers or gas driven to their centres by tidal interactions. The star formation activity expels out the luminous baryonic matter from the galaxy and changes its mass-to-light ratio, M/L . Any change in the M/L ratio results in the deviation of the galaxy from the FP. We study the M/L ratio of our sample dEs in Section 5.

4.3 Dependence on galaxy colour

The location of early-type galaxies relative to the FP is a result of their age and stellar population properties (Forbes, Ponman & Brown 1998; Terlevich & Forbes 2002; Graves, Faber & Schiavon 2009; Graves & Faber 2010). Here, we study the colour- Δ_{FP} relation, in order to assess the stellar population dependence of the deviation of faint dEs from the FP of bright ellipticals.

The colour of our dEs is calculated in four ways. First, we used the SDSS DR7 data to construct $g - r$ for all of our galaxies. Secondly, we used our *HST*/ACS images in both F814W (*I* band) and F475W (*g* band) to calculate $\Delta m = m_{475} - m_{814}$. We also estimated the galaxy colours within two co-centred apertures with diameters of 0.4 and 1.0 arcsec (equal to 0.30 and 0.75 kpc), located at the centre of each galaxy [hereafter $\Delta m(04)$ and $\Delta m(10)$, respectively]. In Table 4, we present the correlation factor for each pair of defined

colours of our sample galaxies and other photometric and dynamical parameters (i.e. M_{814} , σ , n and Δ_{FP}).

Table 4 shows that Δ_{FP} is more tightly correlated with the colour of galaxies compared to their CEL (see Section 4.2). In all cases, $\Delta m(10)$ displays a stronger correlation compared to all other defined colour parameters. Moreover, $\Delta m(10)$ is strongly correlated with the galaxy magnitudes, M_{814} , implying that fainter galaxies are bluer. A strong correlation also exists between $\Delta m(10)$ and $\log_{10}(n)$, which indicates that bluer galaxies have smaller Sérsic indices and therefore are less concentrated.

In Fig. 11, the deviation of galaxies from the FP is plotted against different colour indicators. As seen, the bluer galaxies have larger Δ_{FP} values. In agreement with our results, Bernardi et al. (2003) have also found a correlation between the residuals to the FP and colour of their sample of bright early-type galaxies. Fig. 3 of their paper indicates that bluer galaxies display a greater Δ_{FP} .

This implies that the deviation of dEs from the FP is correlated to their stellar population. The supernova activities, as a consequence of star formation in bluer galaxies, have swept away the luminous baryonic matter from these galaxies. Thus, the supernova-driven winds truncate star formation and modify the M/L ratio of the bluer galaxies, increasing their scatter about the FP. This scenario also explains the location of the faintest galaxies of our sample on the $L-\sigma$ diagram (see Paper I), indicating that these galaxies have larger velocity dispersions compared to the trend of brighter dEs. We noted that these galaxies are bluer than the other galaxies in our sample and display larger deviation from the FP. The star formation has ejected the luminous matter of these galaxies and shifted them to the faint end of the $L-\sigma$ diagram, while their total mass (mostly the dark matter content), and hence their internal velocity dispersions, remains the same.

5 THE M/L RATIOS

The behaviour of the M/L ratio may account for the FP tilt (e.g. Cappellari et al. 2006; ZGZ06; Graves & Faber 2010) and/or deviation of galaxies from the FP (e.g. Reda et al. 2005). Based on the virial theorem, the dynamical mass, M , within the effective radius, R_e , can be derived as

$$M \equiv c \frac{\sigma^2 R_e}{G}, \quad (6)$$

where σ is the galaxy central velocity dispersion and G is the gravitational constant. Assuming that our sample is homologous, c is almost constant for our galaxies and depends on the profiles of the dark and luminous matter. We adopt $c \equiv 5$ (Cappellari et al. 2006), although our conclusion is independent of the exact value of c as long as it is uniformly applied to all galaxies. In this study, we assumed that the absolute magnitudes of the Sun (in the AB system) in the F475W and F814W bands are 5.14 and 4.57 mag, respectively.

In Fig. 12, we show the variation of the dynamical mass-to-light ratio, M/L , as a function of the galaxy dynamical mass, M . For galaxies brighter than $M_{814} = -17.5$ mag, the correlation between M/L and M (the dotted line) is obtained as $M/L \propto M^{0.23 \pm 0.04}$, in agreement with the results of other studies for bright ellipticals. As an example, Mo99 found $M/L \propto M^{0.18 \pm 0.01}$ in the *K* band and $M/L \propto M^{0.23 \pm 0.01}$ in the optical *V* band, for galaxies in the mass range of 10^{10} – $10^{13} M_{\odot}$. In addition, studying a sample

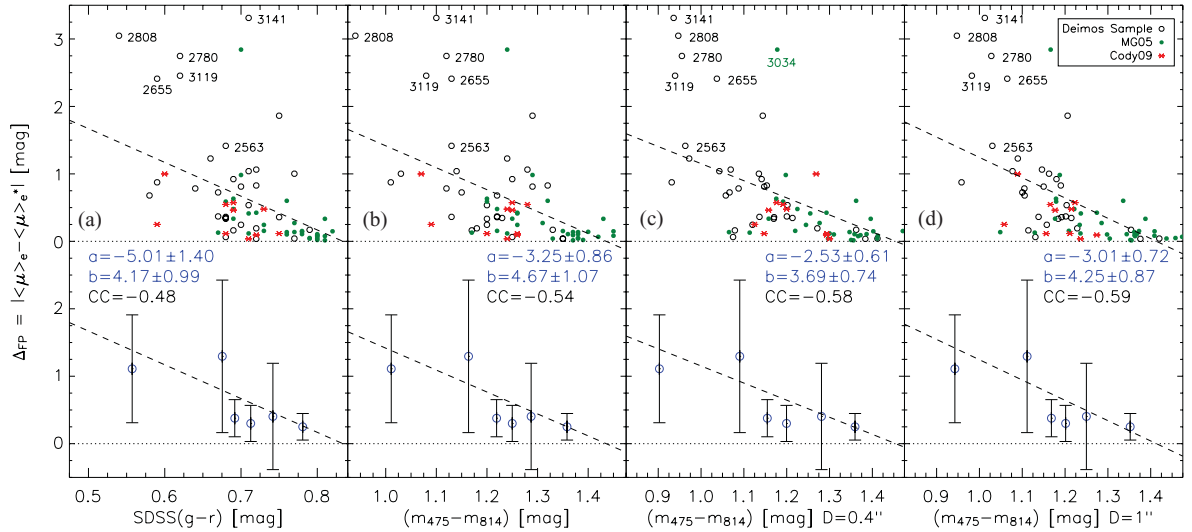


Figure 11. Deviation of galaxies from the FP (the F814W band), Δ_{FP} , against their colours. The colours of galaxies in panel (a) were calculated based on SDSS DR7. In panel (a), colours were derived from our analysis on ACS F814W/F475W images. Panels (c) and (d) are the same as panel (b) except for colours which were derived within the apertures of 0.4 and 1 arcsec diameter at the centre of galaxies. At the bottom of each panel, all data points were binned. Each bin contains the same number of galaxies. For all cases we used the least-squares method to find the best linear trend (dashed lines) by minimizing the residuals in Δ_{FP} . ‘ a ’ is the slope and ‘ b ’ is the intercept of the fitted lines. CC is the cross-correlation coefficient of the fitting parameters. This diagram shows that the bluer galaxies are more deviant from the FP specifically when considering the colour of galaxies at their 0.4 or 1.0 arcsec central regions. The most deviant galaxies are labelled with their GMP IDs.

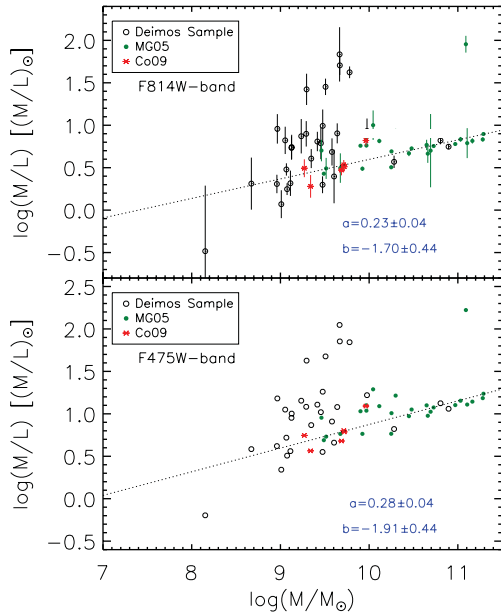


Figure 12. The mass-to-light ratio (M/L) of our sample galaxies in terms of their dynamical mass. The top and bottom panels are based on the ACS F814W-band and F475W-band photometric data, respectively. As seen, the slopes of the (M/L)– M lines do not depend on the two passbands (F814W and F475W). Open circles represent the galaxies with the DEIMOS data for which their σ were measured in Paper I. Red asterisks and green filled circles represent dEs from Co09 and MG05, respectively. The dotted line is the best linear fit for brighter galaxies with $M_{814} > -17.5$. ‘ a ’ and ‘ b ’ are the slope and the intercept of the dotted lines, respectively. In the top panel, the typical error bars of M/L are illustrated. The mass, M , and M/L ratios are in solar units.

of 9000 local early-type galaxies of SDSS in the redshift range of $0.01 < z < 0.3$, Bernardi et al. (2003) found $M/L \propto M^{-0.2}$.

In our sample of dEs, the best linear relation between $\log_{10}(\sigma)$ and $\langle\mu\rangle_e$ is found to be $\log_{10}(\sigma) = (4.94 \pm 0.75) - (0.16 \pm 0.02)\langle\mu\rangle_e$ with the correlation coefficient of -0.66 . This relation is translated to $I_e \propto \sigma^{2.5 \pm 0.3}$, where I_e is effective surface brightness in flux units, in the F814W band. Using the results of Binggeli, Sandage & Tarenghi (1984), we expect $I_e \propto \sigma^{-2.5}$ for elliptical galaxies. Using the virial theorem for spherical systems, $\sigma^2 \propto GM_e/R_e$, and the relation between luminosity and effective surface brightness, $L = 2\pi I_e R_e^2$, the M/L ratio is obtained as $M/L \propto \sigma^2 / \sqrt{I_e L}$ (Co09). Substituting our derived Faber–Jackson relation (Paper I), $L \propto 1.99 \pm 0.14$, and the dependence of I_e upon σ , we obtain $M/L \propto \sigma^{-0.25 \pm 0.33} \propto L^{-0.13 \pm 0.17} \propto M^{-0.15 \pm 0.22}$ for all galaxies in our sample. Our result is more consistent with Co09, who found $M_{dyn}/L \propto M^{0.09 \pm 0.06}$ for a sample with almost the same luminosity range as ours.

As seen in Figs 12(a) and (b), for $M_{814} > -17.5$, some galaxies have larger M/L ratios with respect to the faintward extrapolation of the linear trend of brighter galaxies, while a few faint galaxies follow the M/L – M relation of bright galaxies. The most deviant galaxies in the M/L – M diagram (see Fig. 12) are the bluer galaxies, too. This suggests that the formation of the most deviant dwarf galaxies and those following the trend of bright ellipticals can be explained by different scenarios. Nevertheless, we still need more data points to study the scatter of faint dEs about the trends of brighter ellipticals and to examine different formation mechanisms at the faint regime. In agreement with our results (see Fig. 12), studying the dwarf galaxies with $-16 < M_V < -12$ in the core of the Perseus cluster, Penny et al. (2009) found that fainter dwarfs have larger M/L ratios. In addition, Geha et al. (2002) showed that fainter galaxies have larger M/L ratios compared to their brighter counterparts.

We have also investigated the waveband dependence of the (M/L)– M relation, to explore whether it is governed by the change

in metallicity of galaxies and/or their stellar population. Comparing Figs 12(a) and (b), which are based on the ACS *F814W*- and *F475W*-band photometry, we find the same behaviour of the M/L ratio in different wavebands. We noted that the (M/L) - M relation of galaxies fainter than $M_{814} = -17.5$ is also independent of the passband. Therefore, the stellar population is not the only parameter responsible for the change of M/L ratios (Mo99).

As a conclusion, we found that the M/L ratio is not constant over all of our sample dEs and varies with the mass and the luminosity of the galaxies. The variation in the M/L ratio is responsible for the deviation of our fainter dEs from the FP. Since the fainter galaxies in our sample are bluer than the other galaxies and have larger M/L ratios, we attribute their deviation from the FP to their recent star formation activities (see also Section 4.3).

6 DISCUSSION

In this paper, we studied the Fundamental and Photometric Planes of a sample of 71 dEs in the core of the Coma cluster, the nearest massive elliptical-rich cluster down to a luminosity of $M_{814} < -15.3$. Taking advantage of the DEIMOS high-resolution spectrograph, which enables us to measure the internal velocity dispersion of dEs, and high-resolution imaging of *HST*/ACS, which allows an accurate surface brightness modelling, we were able to extend the FP of galaxies to ~ 1 mag fainter than the previous studies.

We obtained the FP for a subsample of 12 galaxies brighter than $M_{814} = -20$ as $R_e \propto \sigma^{1.33 \pm 0.02} (I_e)^{-0.80 \pm 0.01}$, which is consistent with the previous studies of bright galaxies in the Coma cluster (JFK96; Mo99). Studying the FP of 141 early-type galaxies in the Shapley supercluster at $z = 0.049$, Gargiulo et al. (2009, hereafter Ga09) found that the FP follows the relation $R_e \propto \sigma^{1.35} (I_e)^{-0.81}$ for galaxies with $\sigma > 100 \text{ km s}^{-1}$ and $M_R < -18.7$. When including all galaxies in their sample, including low-mass galaxies down to $\sigma \sim 50 \text{ km s}^{-1}$, Ga09 found a shallower exponent for σ . The FP relation of our dEs displays even shallower exponents for σ and (I_e) than in Ga09, due to further extension to fainter galaxies.

In Figs 4 and 11, the faintest galaxies and the most deviant data points are based on our DEIMOS observations. One could argue that our σ measurements might be overestimated. In Paper I, we have considered all sources of uncertainties in measuring the velocity dispersions, namely the statistical errors, template mismatch and other systematic uncertainties. We have also checked the sensitivity of our measured velocity dispersions to different kinds of stellar templates. Moreover, using different sets of mixed stellar templates, covering a vast range of spectral types, when measuring the velocity dispersions, confirmed our measurements. The presented uncertainty in the measurement includes all such sources and therefore we can rule out the overestimation argument.

In this study, the spectroscopy was carried out using the slits with the size of 0.7 arcsec covering the body and the central region of the galaxies. Due to the size of slits, for galaxies with net rotation, the measured velocity dispersion differs from the central velocity dispersion. The kinetic energy (KE) per unit mass for a spheroid is (e.g. Busarello, Longo & Feoli 1992)

$$\text{KE} = \frac{1}{2} \langle v \rangle = \frac{1}{2} V_{\text{rot}}^2 + \frac{3}{2} \sigma_0^2 \simeq \frac{3}{2} \sigma_m^2, \quad (7)$$

where V_{rot} is the rotation velocity, σ_0 is the central velocity dispersion and σ_m is the measured velocity dispersion. Thus, we have

$$\left(\frac{\sigma_m}{\sigma_0} \right)^2 \simeq 1 + \frac{1}{3} \left(\frac{V_{\text{rot}}}{\sigma_0} \right)^2. \quad (8)$$

The observed values of V_{rot}/σ_0 range from 0.01 to ~ 1.0 for galaxies with $-14 < M_R < -26$ (Davies et al. 1983; Pedraz et al. 2002; Simien & Prugniel 2002; Geha et al. 2003; van Zee, Skillman & Haynes 2004), implying that the difference between σ_m and σ_0 is less than ~ 15 per cent of the central velocity dispersion. Such uncertainty is smaller than the total estimated uncertainty in measuring the velocity dispersions of faint dEs ($\lesssim 50 \text{ km s}^{-1}$). Based on the relation for an oblate, isotropic galaxy flattened by rotation (Davies et al. 1983), the mean isophotal ellipticity of four galaxies in our sample, GMP2780, GMP2563, GMP3119 and GMP3141, is less than 0.1, corresponding to $V_{\text{rot}}/\sigma_0 \simeq 0.3$. Using equation (8), the uncertainty in the measured velocity dispersion due to the rotation would be ~ 2 per cent, which is significantly smaller than the uncertainties due to template mismatch and other systematics discussed in Paper I. The uncertainty due to the rotation is ~ 6 per cent for GMP2655 and GMP2808, for which the mean ellipticity is ~ 0.25 . Thus, rotation cannot be solely responsible for the deviation of faint dEs from the faintward extrapolation of the Faber–Jackson relation and the FP of bright ellipticals.

The non-universality of the FP is claimed by ZGZ06, where they suggest that the spheroids ranging from dEs to the galaxy clusters lie on a curved surface, in $(\sigma, R_e, \langle \mu \rangle_e)$ space. Describing the M/L ratio as a parabolic function of the internal velocity dispersion, ZGZ06 could explain the change in the coefficients of the FP for different spherical systems. They predicted that for $\sigma \sim 50 \text{ km s}^{-1}$, M/L is almost a flat function of σ , with a large scatter, which tends to increase for smaller values of σ . In agreement with Co09 and ZGZ06, we found $M/L \propto \sigma^{-0.25 \pm 0.33}$ for all galaxies in our sample. In addition, we found larger M/L ratios for the faintest dEs in our sample with respect to the (M/L) - M trend of brighter galaxies (see Fig. 12).

A mixture of formation mechanisms has been considered to model dE formation. In the ‘wind model’, dEs are primordial objects that lost their gas in a supernova-driven galactic wind (DS86; Yoshii & Arimoto 1987, hereafter YA87; Chiosi & Carraro 2002, hereafter CC02). In contrast to massive ellipticals, dEs have very long star formation histories. Since giant galaxies hold on very strongly to their gas content, their gas is almost completely converted into stars in a single burst. On the other hand, supernova activities disperse the gas and switch off further star formation in dwarf galaxies. As the gas cools, it sinks back in and again undergoes a new phase of star formation followed by supernova explosions which disperse the gas, and this cycle can be repeated many times (Bender & Nieto 1990; de Rijcke et al. 2005).

An alternative formation path for dwarf galaxies involves stripping of larger galaxies by gravitational and gas dynamical processes. As galaxies fall into the cluster, they are stripped of much of their gas by ram-pressure effects (Gunn & Gott 1972; van Zee et al. 2004); this process can be seen at work in both $\text{H}\alpha$ and ultraviolet imaging of the Coma cluster (Smith et al. 2010; Yagi et al. 2010). Stars and dark matter can then be removed by interactions with more massive galaxies and the cluster potential, termed ‘harassment’ (Moore et al. 1996; Moore, Lake & Katz 1998; Mayer et al. 2001; Mastropietro et al. 2005), or by tidal encounters between galaxies of near-equal mass (Richstone 1976; Aguerri & González-García 2009). Both harassment and tidal encounters can create bars (Mastropietro et al. 2005; Aguerri & González-García 2009) which could play a part in driving gas not stripped by the ram pressure of the intergalactic medium inwards, triggering star formation. The resulting galaxy is likely to have a higher central surface brightness than its progenitor and would probably resemble a nucleated dE which rotates quite rapidly and still displays some memory of its

former state. Although harassment tends to increase internal velocity dispersion, it is unable to disrupt rotational motion and therefore is not obviously reconciled with the low rotational velocities found by Simien & Prugniel (2002), Geha et al. (2002, 2003) and MG05.

Another possible way to explain the formation of dwarf galaxies is presented by Kroupa (1998) and Duc, Bournaud & Masset (2004). The merging process of two gas-rich disc galaxies results in many massive star clusters which coalesce to produce a small number of dwarf galaxies. These tidally formed dwarf galaxies contain negligible amounts of dark matter and hence cannot explain a significant fraction of the dE population.

Our findings appear to be most consistent with the ‘wind model’. In the wind-stripping model, the shallow potential well of dwarf galaxies, with the velocity dispersion of less than 100 km s^{-1} , does not allow intense star formation without stripping most of the gas (Schaeffer & Silk 1988). We found the effective radius–luminosity and the Faber–Jackson relation as $R_e \propto L^{0.24}$ and $L \propto \sigma^{2.0}$ in the $F814W$ band, respectively. These two relations are in agreement with the calculations of de Rijcke et al. (2005) based on the wind models of YA87 and CC02. They showed that, although YA87 models do not predict a tight L – R_e relation, CC02 models can reproduce it very nicely.

Using the accurate measurement of the M/L ratios for the SAURON galaxy sample (Bacon et al. 2001), the FP tilt of early-type galaxies can be attributed to the variation of the M/L ratio which is governed by the stellar population and dark matter properties (Cappellari et al. 2006). Moreover, in agreement with Cappellari et al. (2006), considering the resulting M/L ratios based on different star formation models and mass profiles, Allanson et al. (2009) and Grillo & Gobat (2010) concluded that non-homology is not a major driver of the FP tilt while the mass sensitivity of the M/L ratio is the primary source of the FP behaviour in different mass and luminosity regimes. We find that the fainter dEs are bluer, and scatter more about the FP, than the brighter ellipticals. Star formation activity, together with the wind model, can explain how the bluer dwarfs at a given mass have different M/L ratios and thus follow different scaling relations.

7 SUMMARY

The results of this study can be summarized as follows.

(i) The FP of bright ellipticals with $M_{814} < -20$ in our sample is fully consistent with the previous studies of the Coma bright galaxies. We found that the scatter about the FP depends on the faint-end luminosity cut-off, which is significantly increased. We noted that the FP tilt of faint dEs with $M_{814} > -20$ differs from that of giant ellipticals.

(ii) The FP is not colour dependent, which implies that the stellar population plays a insignificant role in the scatter about the FP and changing its tilt. We studied the departure of faint dEs from the FP of bright ellipticals, Δ_{FP} , which is found to be correlated with the luminosity and the galaxy brightness profile. Less concentrated faint dEs, smaller Sérsic indices, show a larger deviation from the FP. We also found a relation between Δ_{FP} and the CEL of galaxies as an imprint of their formation history. In addition, we found a correlation between Δ_{FP} and the colour of galaxies, which indicates that the bluer galaxies in our sample dEs display larger deviation from the FP of the bright ellipticals. On the other hand, we have already shown that in the L – σ relation discussed in Paper I, fainter dEs seem to be more massive (e.g. dynamical mass estimated using the internal velocity dispersion) than predicted from their luminosity. Although

tidal effects can be important in removing luminous matter from the galaxy, we believe that supernova-driven winds are more probably the dominant process.

(iii) Replacing the central velocity dispersion with the Sérsic index in the FP relation, we studied the PHP of our sample galaxies. The differences in the coefficients of the PHP relation for our dEs and those of previous studies in local elliptical galaxies (GR02) mainly arise from the fainter luminosity coverage of our sample. We noted that most of the outliers in the projected PHP diagram have internal morphological features or a poor Sérsic fit. Our conclusion is that while the FP is less sensitive to a detailed morphology of galaxies, the PHP is able to effectively differentiate between different morphologies. In other words, the scatter about the FP increases as we extend the sample to fainter dwarf galaxies. In the PHP, however, the scatter is driven by asymmetry and unsmooth galaxy brightness distribution.

(iv) We studied the correlations between the Sérsic index, n , and other photometric or kinematic parameters (i.e. M_{814} , $\langle \mu \rangle_e$, μ_0 , σ and the concentration). The light concentration and the Sérsic index are strongly correlated with the correlation coefficient of 0.95. The process of determining n is model dependent. On the other hand, the concentration parameter is model independent and is determined using a simple photometric analysis. Therefore, based on the n –concentration correlation, one is able to economically replace n with the concentration parameter when processing the scaling relations.

The best radius–luminosity relation for our sample dEs is obtained as $R_e \propto L^{0.24}$, which is consistent with the results of de Rijcke et al. (2005), who reported a B -band radius–luminosity power-law slope between 0.28 and 0.55 for a sample of dwarf ellipticals and dwarf spheroidal galaxies.

(v) We found the KR slopes as 5.23 ± 0.37 and 5.17 ± 0.28 for galaxies in the luminosity ranges of $-20 < M_{814} < -18$ and $-18 < M_{814} < -16$, respectively. This is in agreement with the results of Khosroshahi et al. (2004), who found the KR slope as ~ 5.2 for dEs of 16 nearby galaxy groups with $-18 < M_R < -14$. Furthermore, in agreement with Khosroshahi et al. (2004), we found the steeper linear trends for faint dEs compared to that of the brighter galaxies.

(vi) We obtained the mass-to-light ratio as $M/L \propto \sigma^{-0.25 \pm 0.33} \propto L^{-0.13 \pm 0.17} \propto M^{-0.15 \pm 0.22}$ for our sample dEs, indicating that the M/L ratio is almost a flat function of mass, luminosity and velocity dispersion. Our finding agrees well with that of Co09, who reported $M/L \propto M^{0.09 \pm 0.06}$ for a sample with nearly the same luminosity range. Galaxies with $M_{814} > -17.5$ seem to have higher M/L ratios with respect to the linear extrapolation of the same associated with brighter galaxies. The variation in the M/L ratio is responsible for the deviation of the fainter dEs from the FP. Since the fainter galaxies in our sample are bluer than the other galaxies and have higher M/L ratios, we attribute their deviation from the FP to their recent star formation activities.

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